Introduction to high-energy astrophysical neutrinos 2/3

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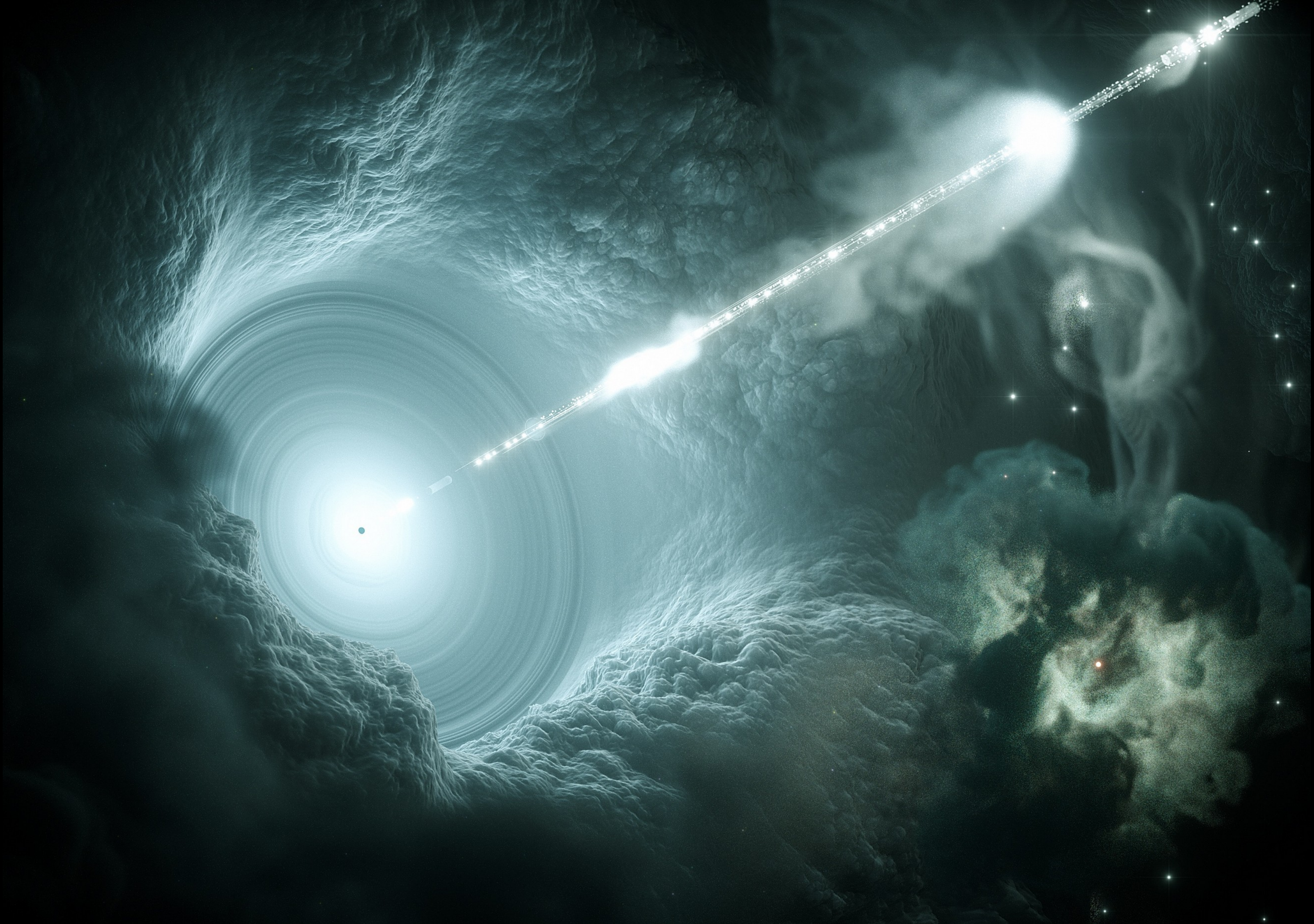
EuCAPT Astroneutrino Theory Workshop
Prague, September 16–20, 2024

UNIVERSITY OF COPENHAGEN

VILLUM FONDEN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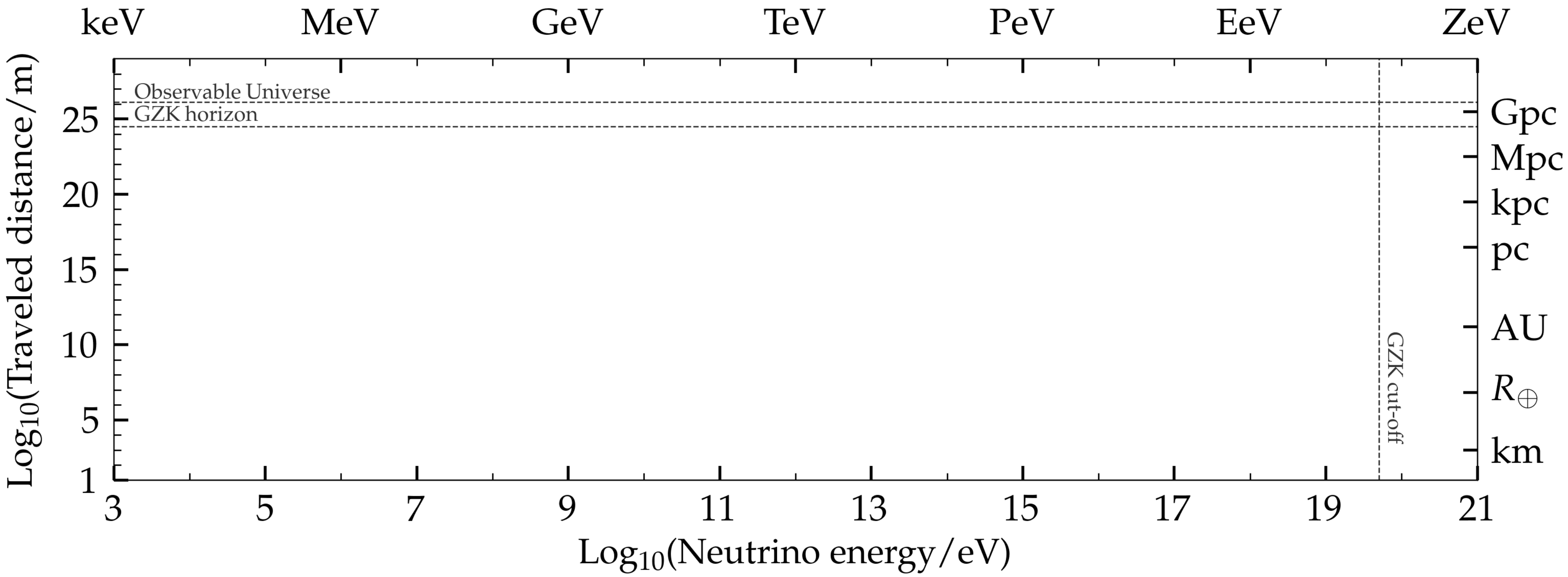
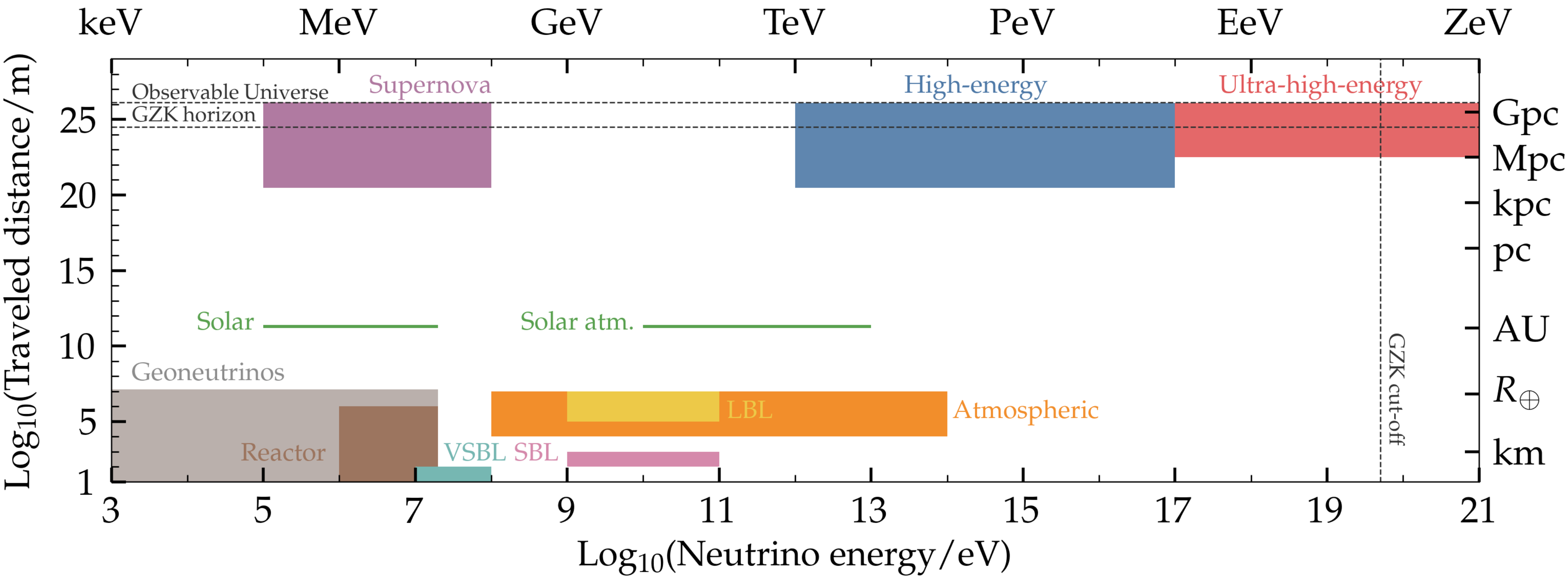


Ackermann, **MB**, *et al.*, Astro2020 Decadal Survey (1903.04333), adapted



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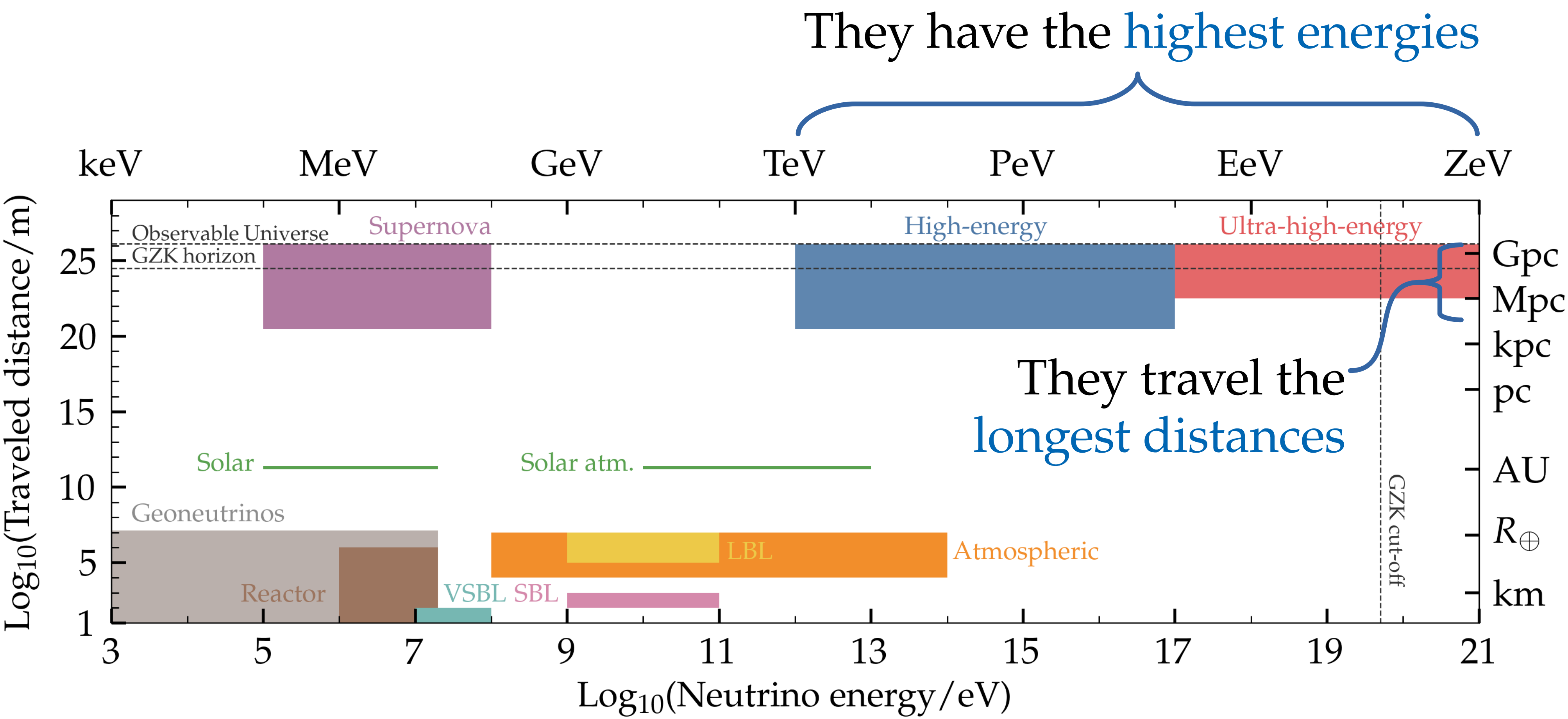

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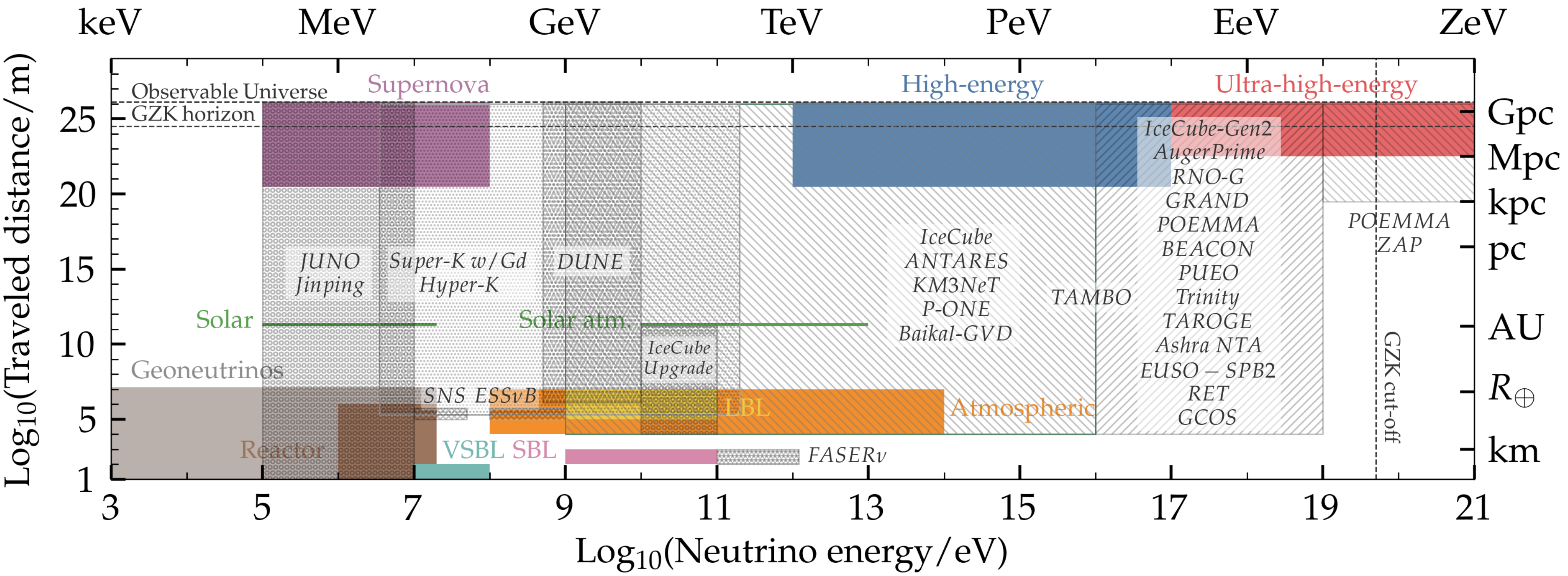

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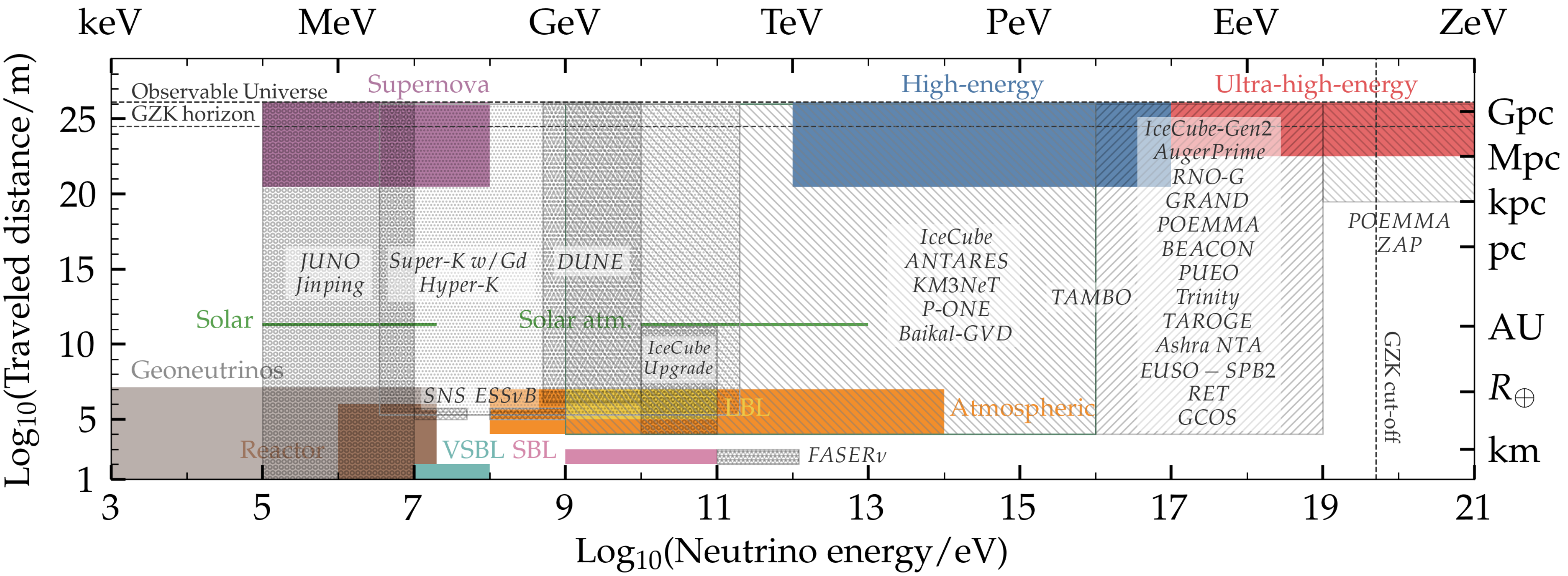
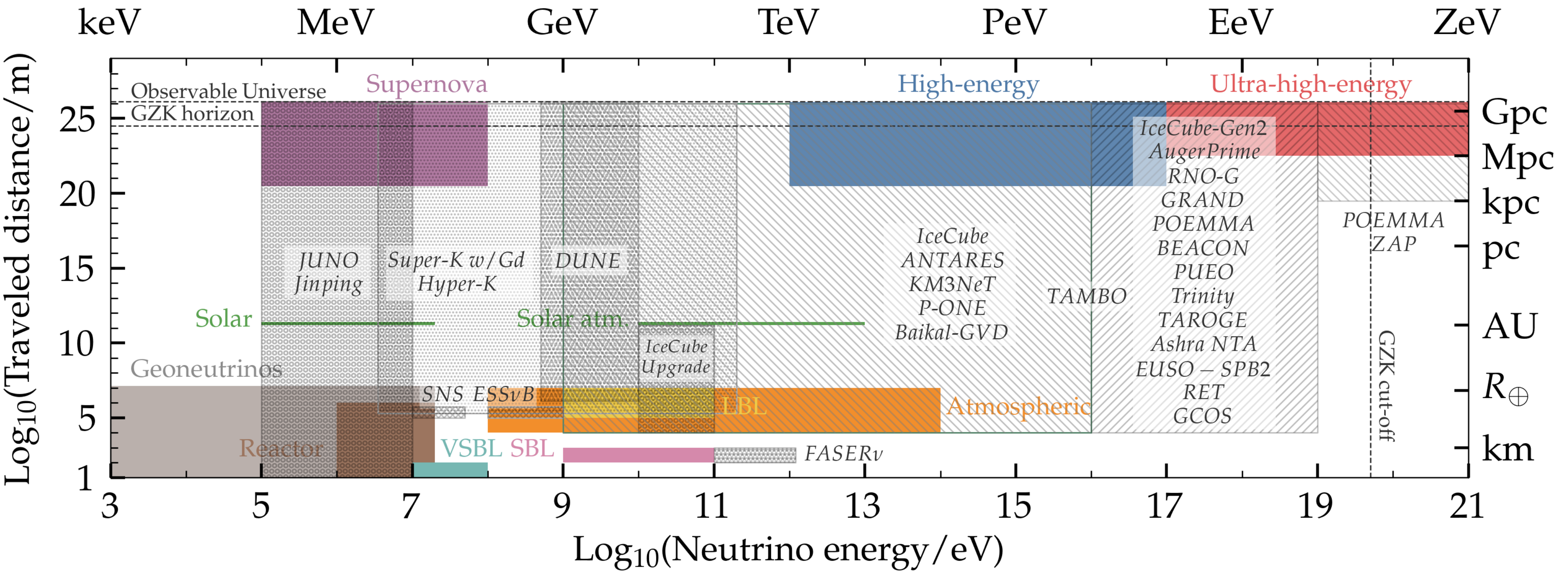


Synergies with lower energies

Ackermann, **MB**, *et al.*, Astro2020 Decadal Survey (1903.04333), adapted

Discovered in 2013
by IceCube



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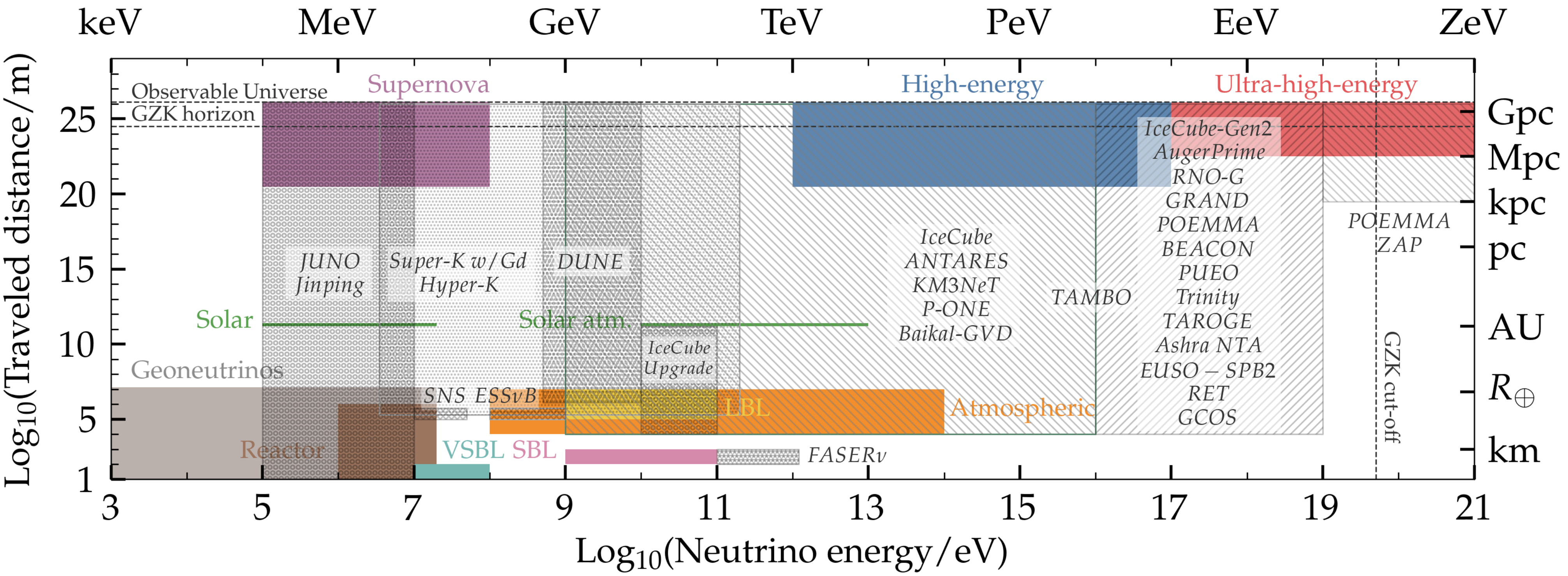

Synergies with lower energies

Increase TeV–PeV
ν statistics

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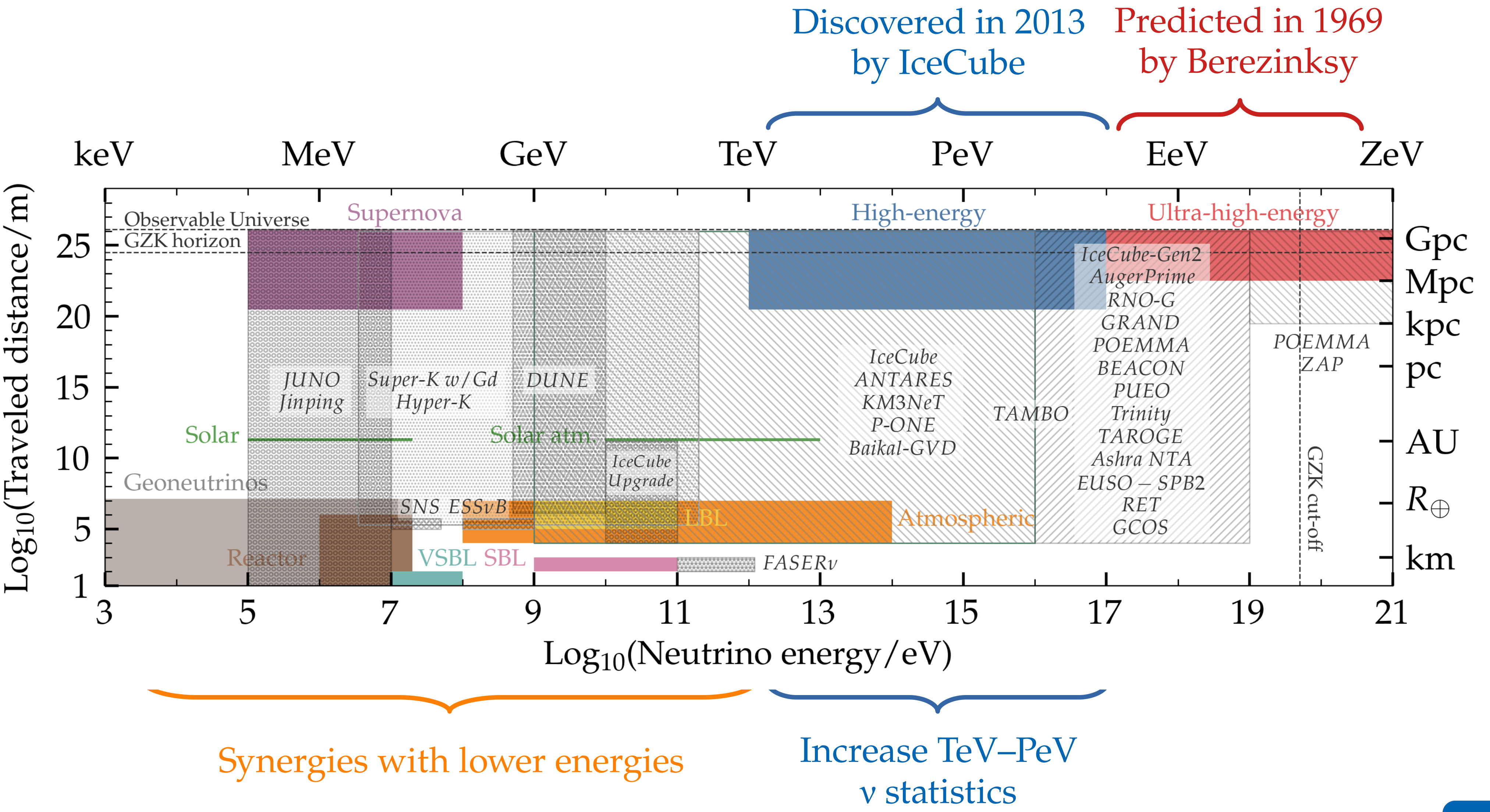

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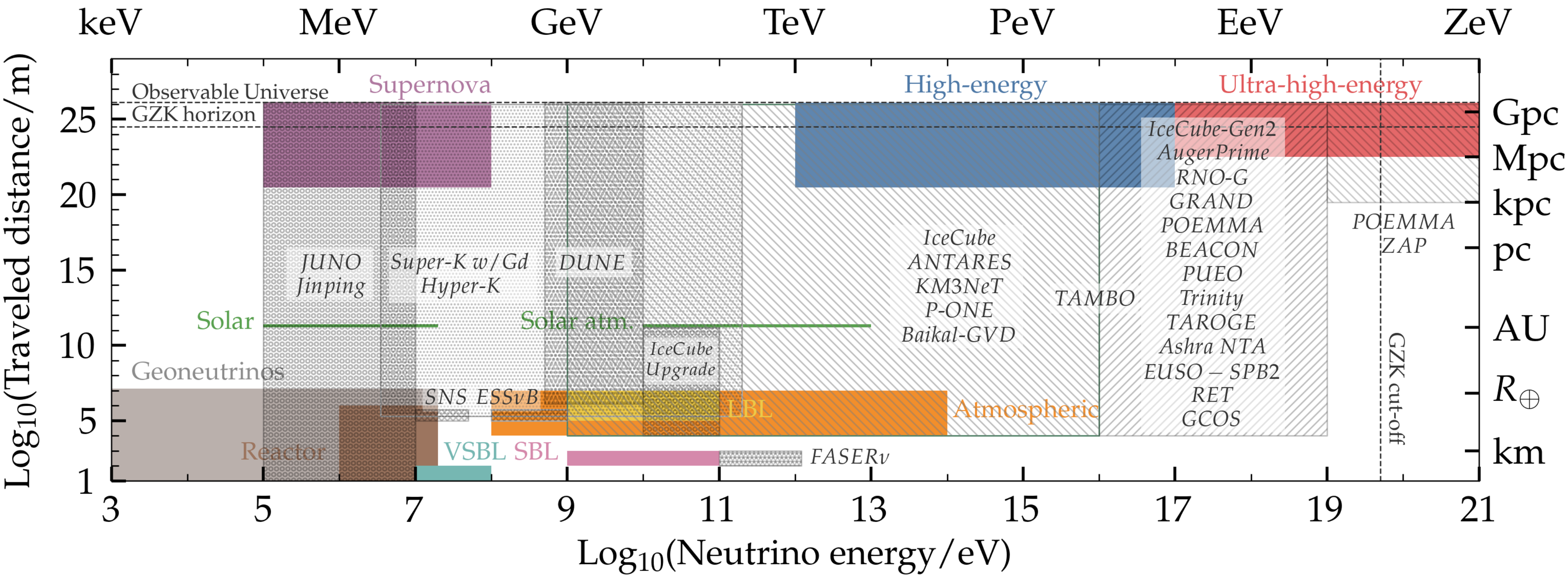
Discovered in 2013 by IceCube

Predicted in 1969 by Berezinksy



Synergies with lower energies

Increase TeV–PeV ν statistics

Discover > EeV ν

Ackermann, **MB**, *et al.*, Astro2020 Decadal Survey (1903.04333), adapted

Today

TeV–PeV ν

Next decade

> 100-PeV ν

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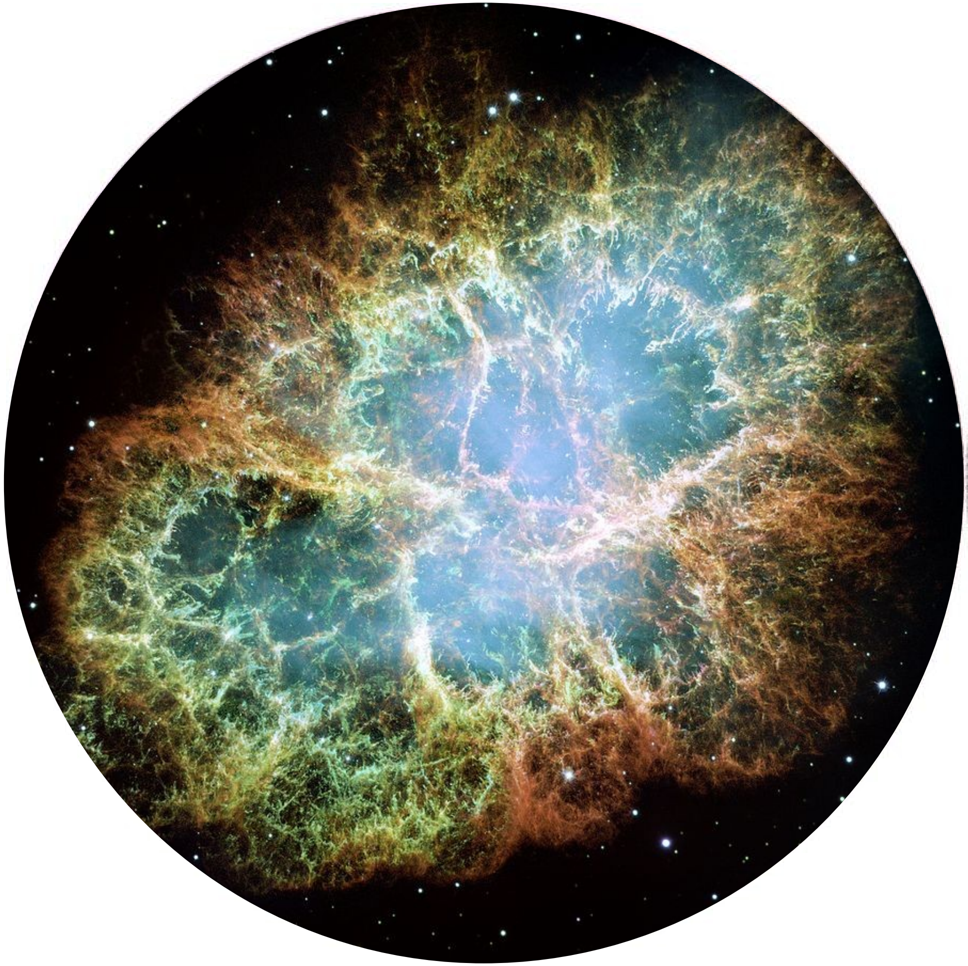

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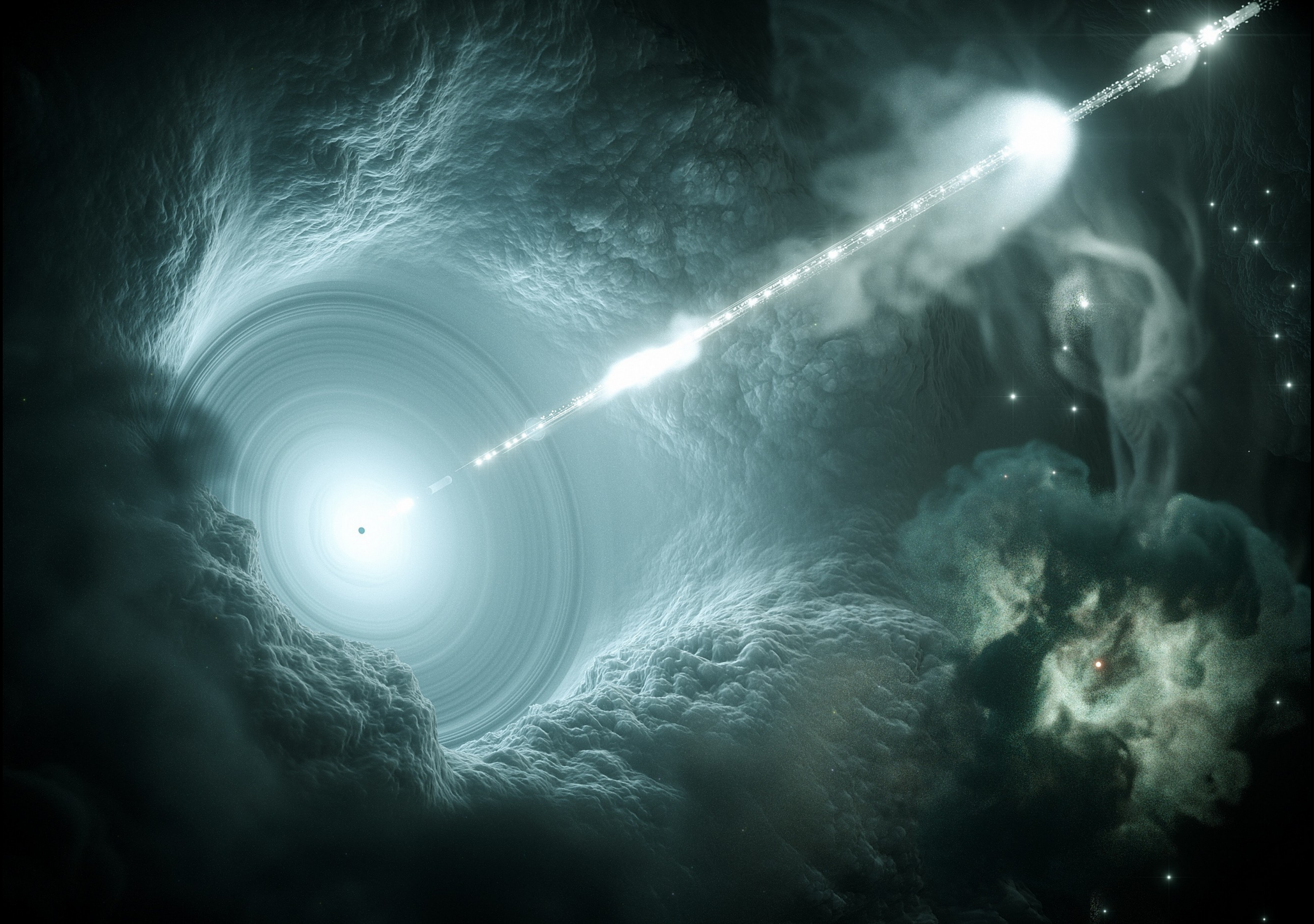

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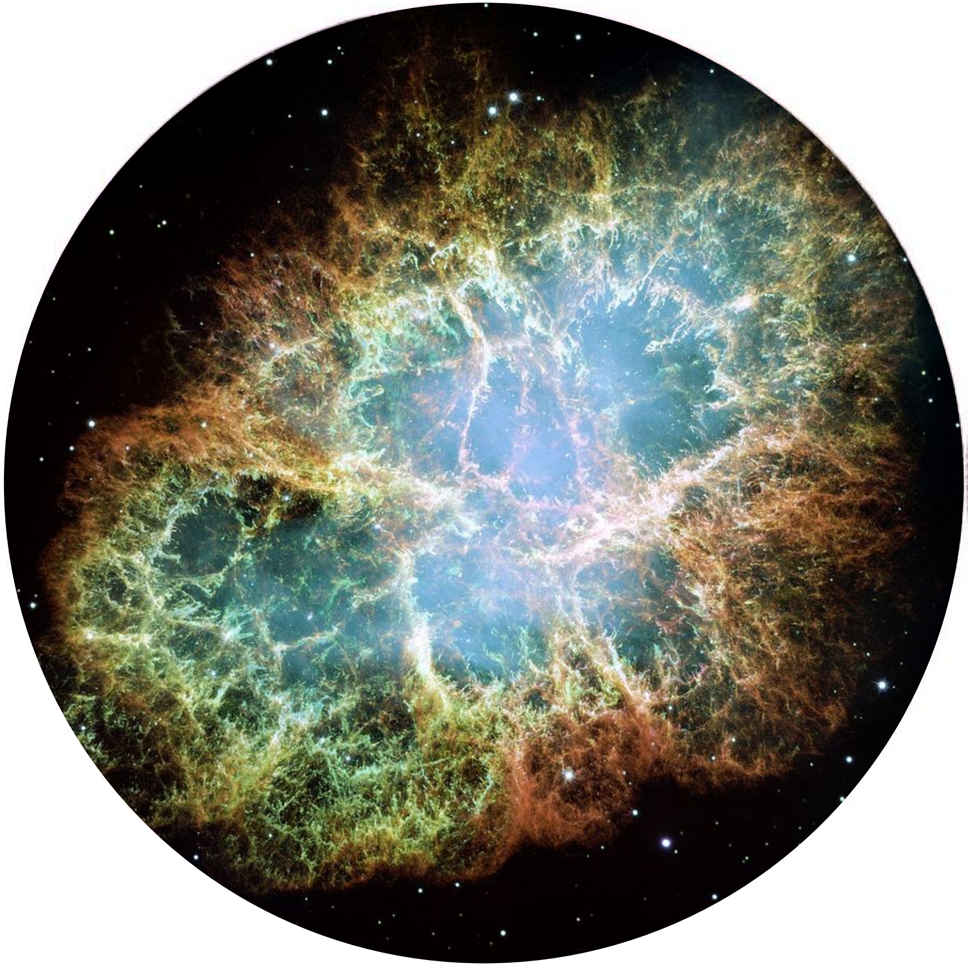

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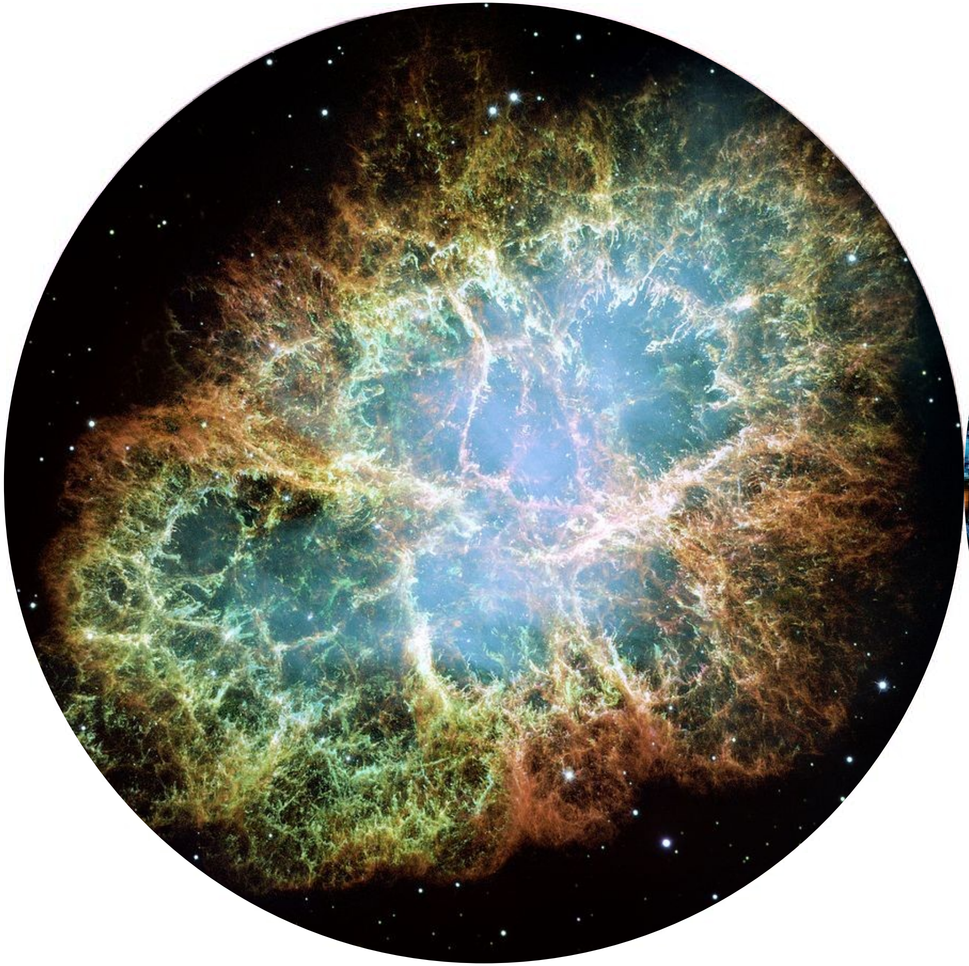

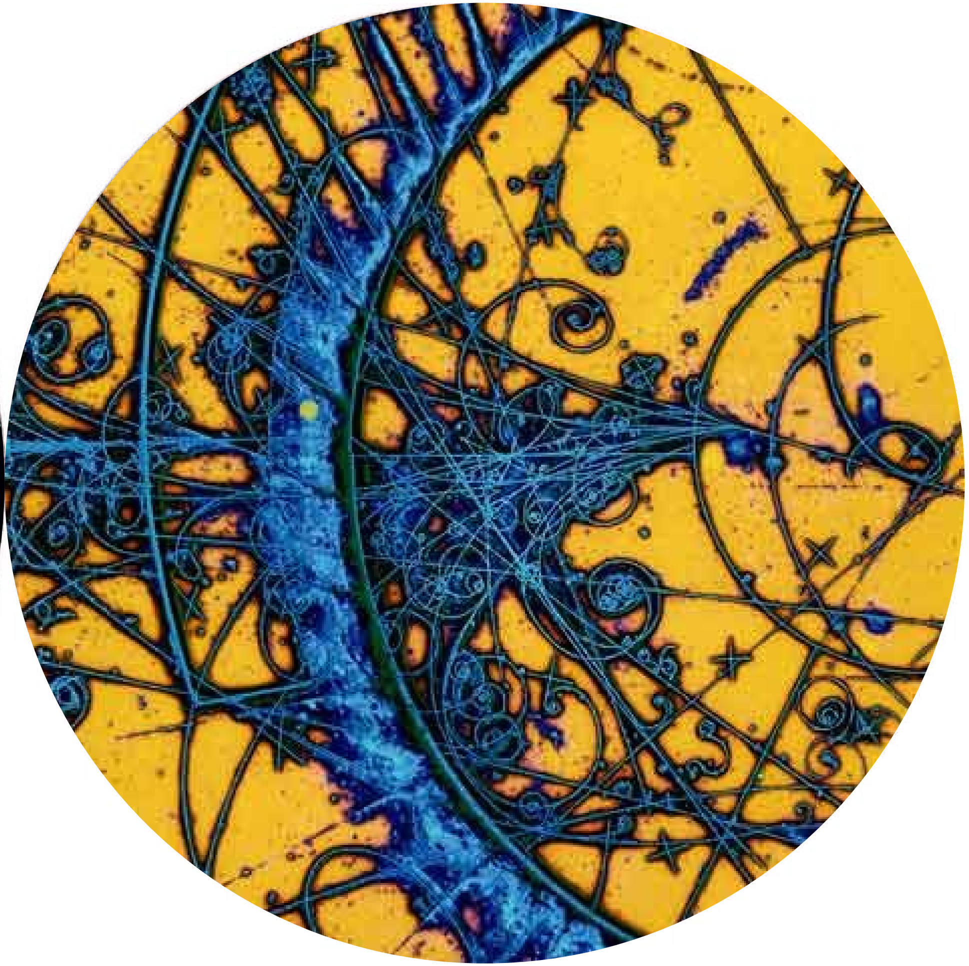

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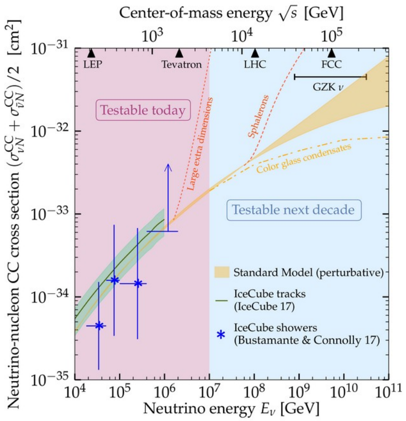
TeV–EeV ν cross sections



MB & Connolly, *PRL* 2019

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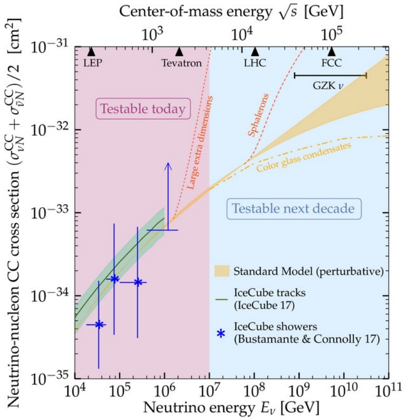
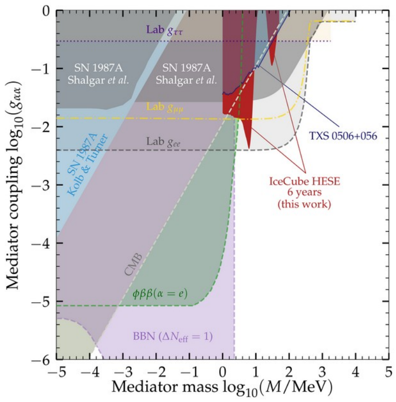


MB & Connolly, *PRL* 2019

ν self-interactions



MB, Rosenstrøm, Shalgar, Tamborra, *PRD* 2020

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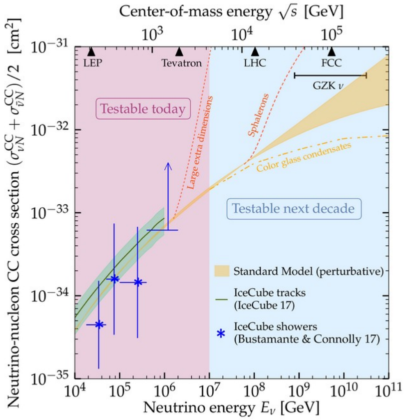

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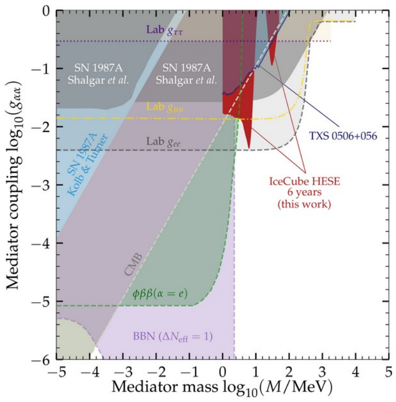
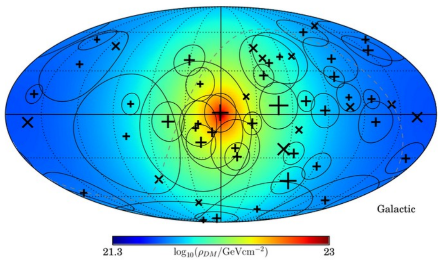


MB, Rosenstrøm, Shalgar, Tamborra, *PRD* 2020

ν scattering on Galactic DM



Argüelles, Kheirandish, Vincent, *PRL* 2017

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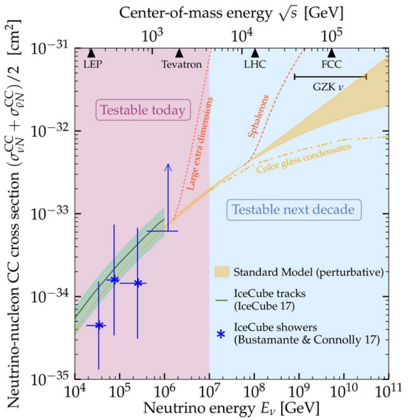

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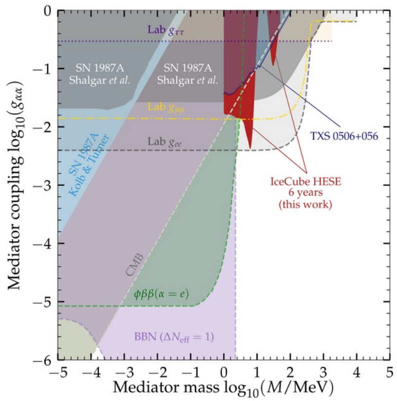

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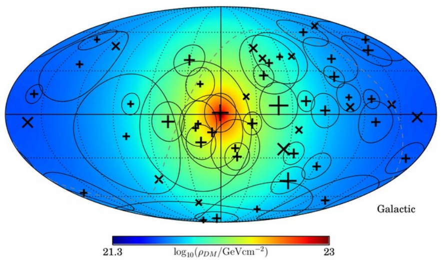

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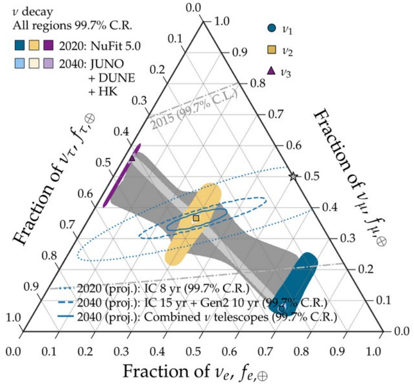
ν decay



Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

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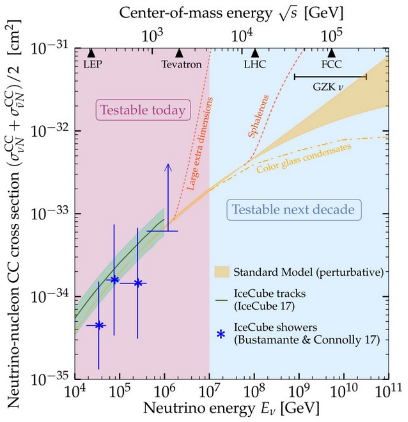

MB & Connolly, *PRL* 2019

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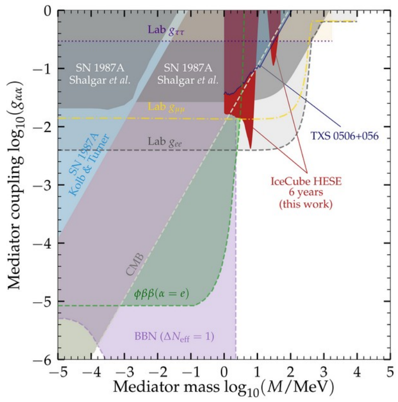

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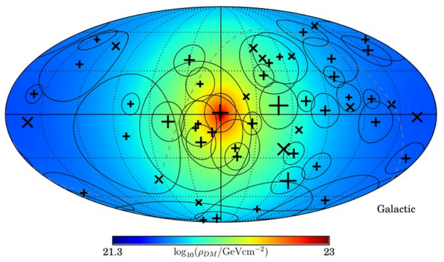

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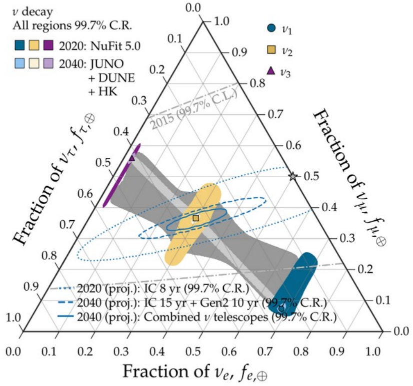
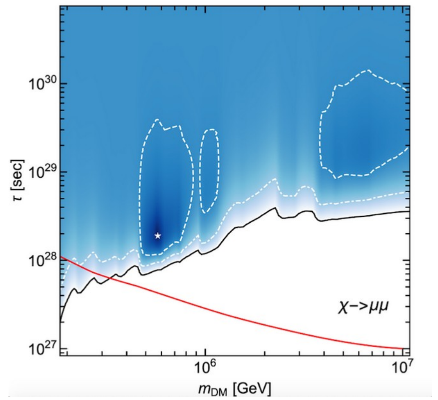


Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

Dark matter decay



Chianese, Fiorillo, Miele, Morisi, Pisanti, *JCAP* 2019

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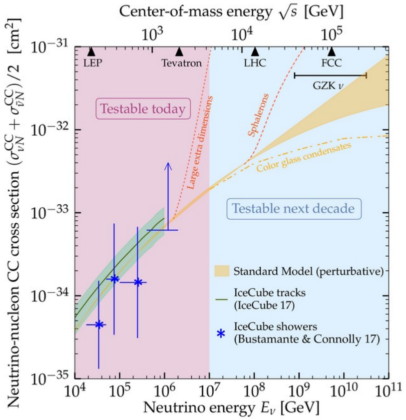

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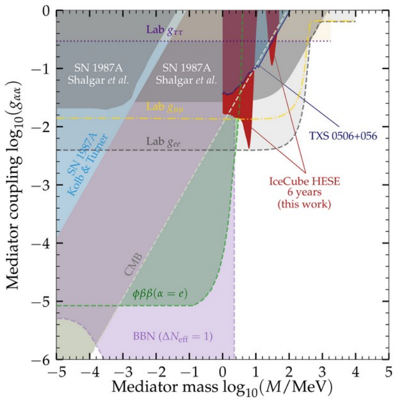

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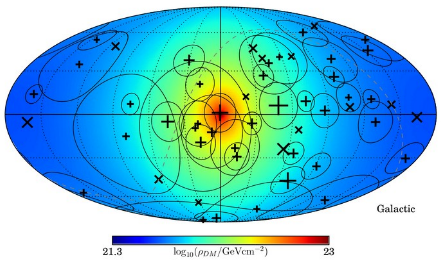

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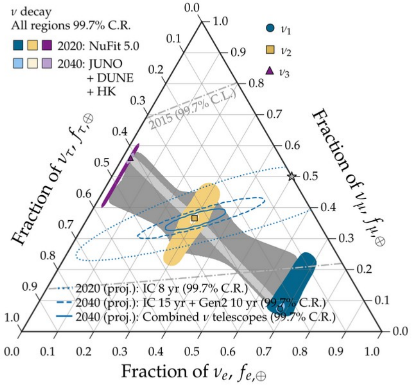

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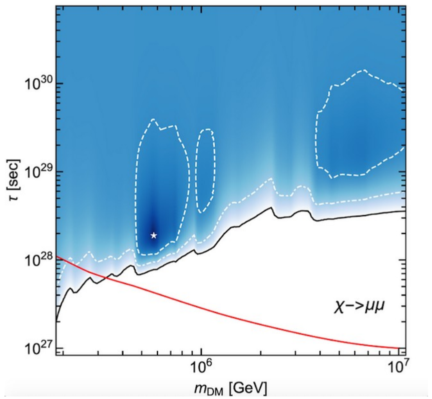

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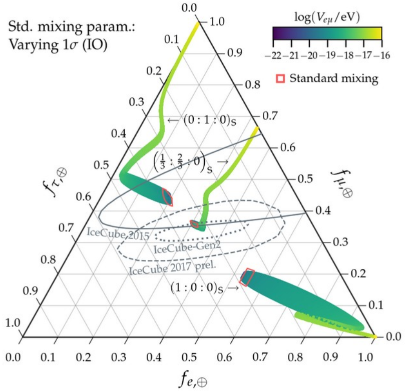
ν-electron interaction



MB & Agarwalla, *PRL* 2019

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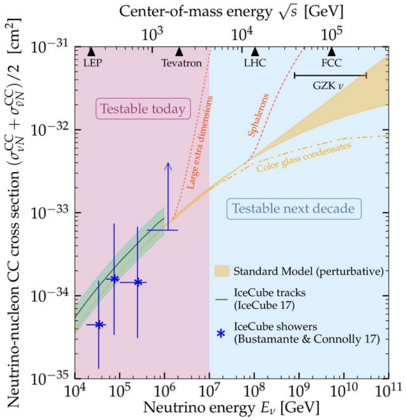

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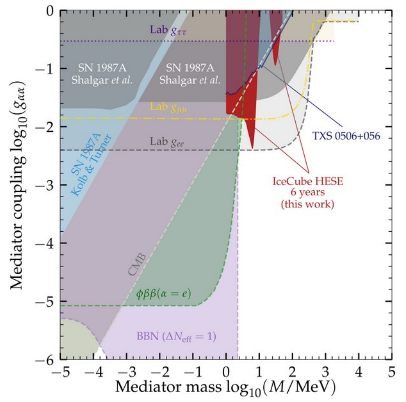

MB, Rosenstrøm, Shalgar, Tamborra, *PRD* 2020

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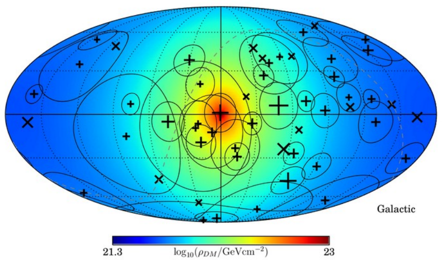

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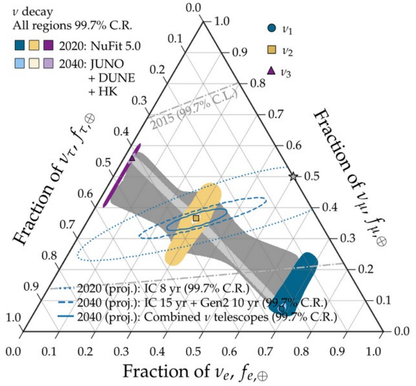

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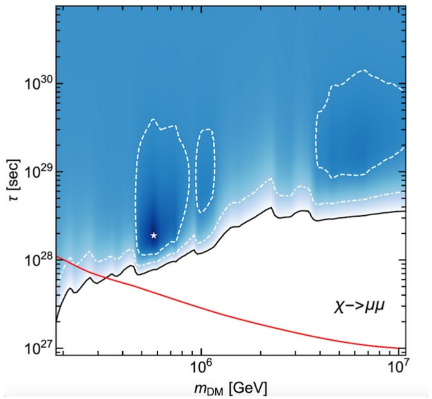

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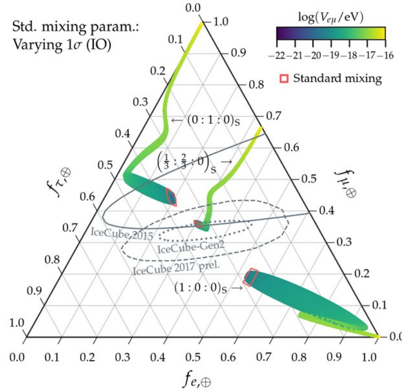
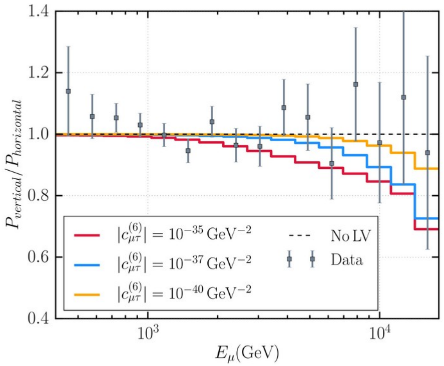


MB & Agarwalla, *PRL* 2019

Lorentz-invariance violation



IceCube, *Nature Phys.* 2018

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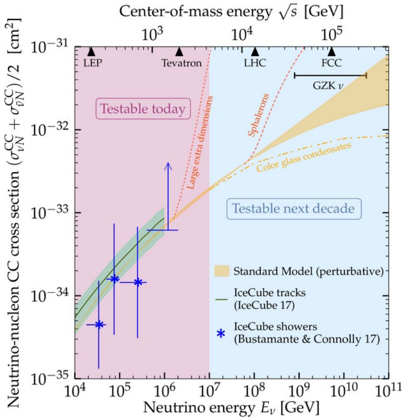

MB & Connolly, *PRL* 2019

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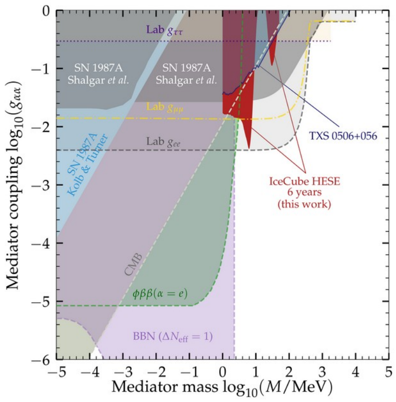

MB, Rosenstrøm, Shalgar, Tamborra, *PRD* 2020

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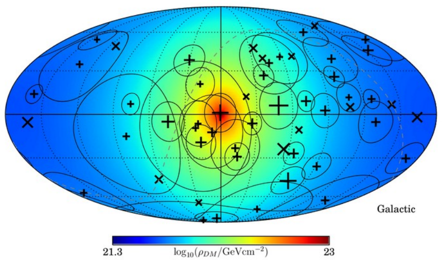

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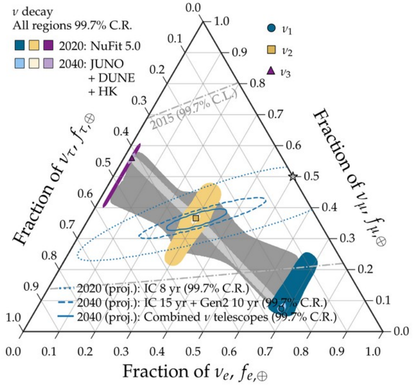

Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

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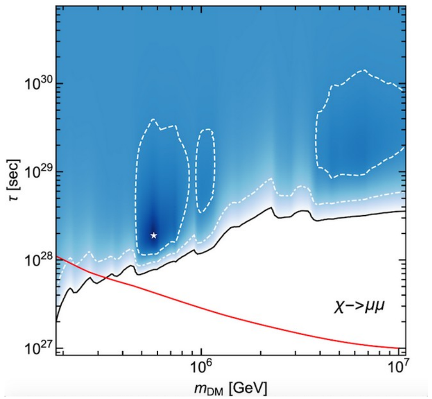

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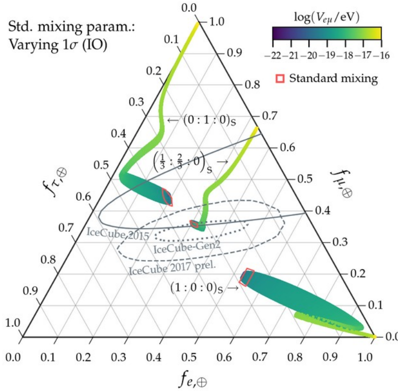

MB & Agarwalla, *PRL* 2019

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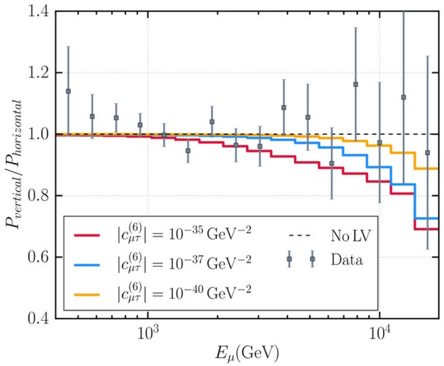
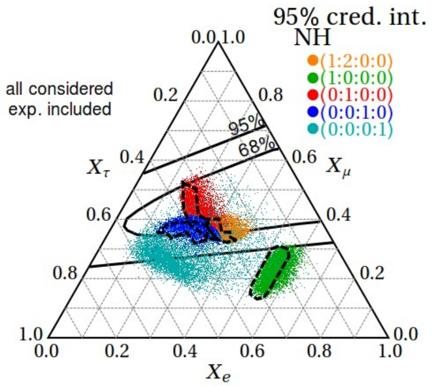


IceCube, *Nature Phys.* 2018

Sterile neutrinos



Brdar, Kopp, Wang, *JCAP* 2017

Making high-energy astrophysical neutrinos: a toy model

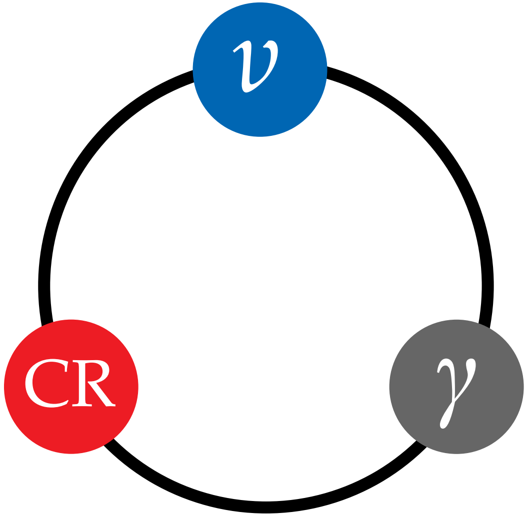
(or $p + p$)

$$p + \gamma_{\text{target}} \longrightarrow \Delta^+ \longrightarrow \begin{cases} p + \pi^0, & \text{Br} = 2/3 \\ n + \pi^+, & \text{Br} = 1/3 \end{cases}$$

$$\pi^0 \longrightarrow \gamma + \gamma$$

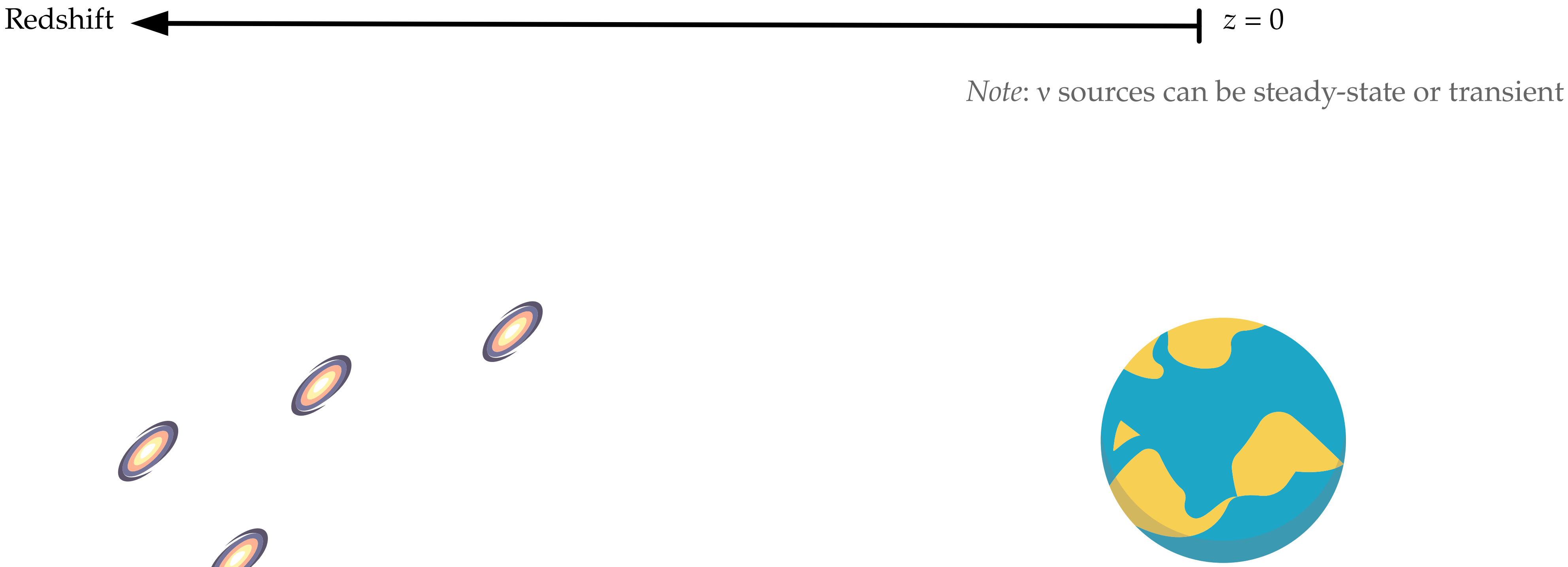
$$\pi^+ \longrightarrow \mu^+ + \nu_\mu \longrightarrow \bar{\nu}_\mu + e^+ + \nu_e + \nu_\mu$$

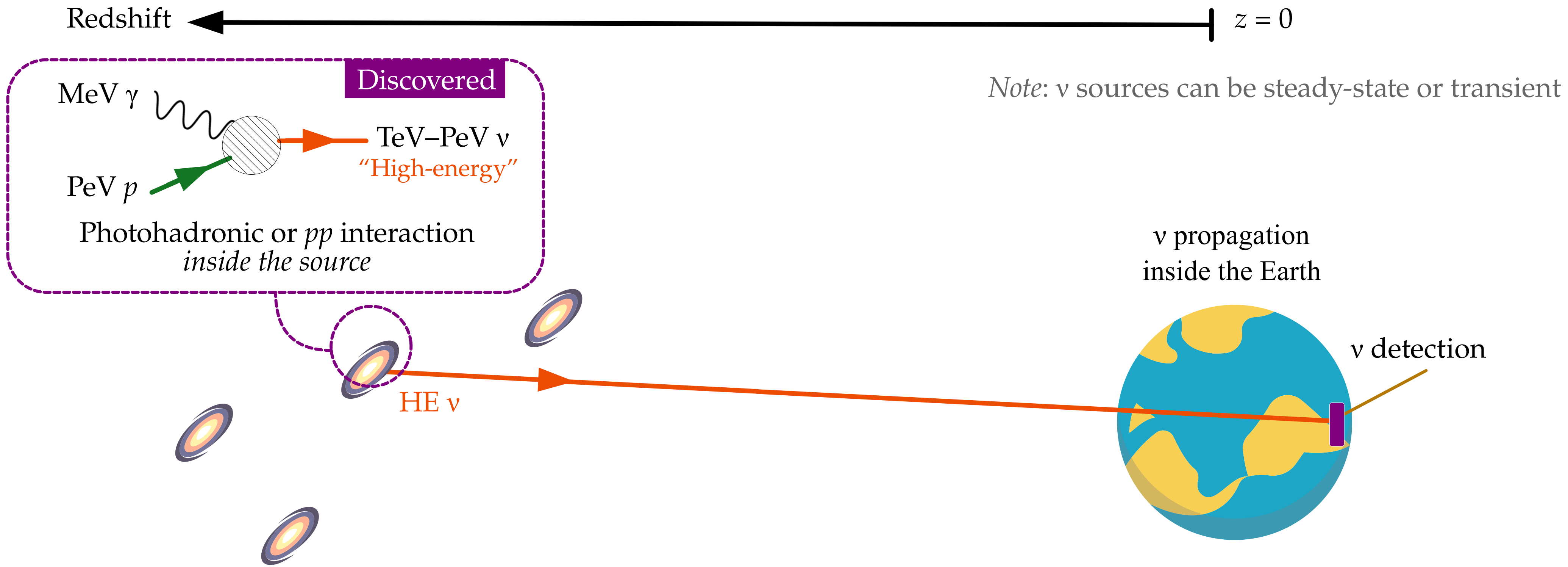
$$n \text{ (escapes)} \longrightarrow p + e^- + \bar{\nu}_e$$



Neutrino energy = Proton energy / 20

Gamma-ray energy = Proton energy / 10


Redshift
$z = 0$
Note: ν sources can be steady-state or transient



Redshift
z = 0
Discovered
MeV γ
TeV–PeV ν
"High-energy"
PeV p
Photohadronic or pp interaction
inside the source
Note: ν sources can be steady-state or transient
ν propagation
inside the Earth
ν detection
HE ν

How many neutrinos? The Waxman-Bahcall bound

- Energy production rate of extragalactic cosmic-ray protons in the energy range 10^{19}–10^{20} eV:

$$\dot{\varepsilon}_{\rm CR}^{[10^{19},10^{21}]} \sim 5 \cdot 10^{44}~{\rm erg~Mpc^{-3}~yr^{-1}}$$

- So, the energy-dependent generation rate of cosmic rays is $E_{\rm CR}^2 \frac{d\dot{N}_{\rm CR}}{dE_{\rm CR}} = \frac{\dot{\varepsilon}_{\rm CR}^{[10^{19},10^{21}]}}{\ln(10^{21}/10^{19})} \approx 10^{44}~{\rm erg~Mpc^{-3}~yr^{-1}}$

- Protons lose a fraction $\epsilon < 1$ in photohadronic production of pions in the sources

- Present-day energy density of $\nu_\mu+\bar{\nu}_\mu$: $E_\nu^2 \frac{dN_\nu}{dE_\nu} \approx \frac{1}{4}\epsilon~t_{\rm H}~E_{\rm CR}^2 \frac{d\dot{N}_{\rm CR}}{dE_{\rm CR}}$

$\mathrm{Br}(p + \gamma \to \pi^+) = 0.5$ × Fraction of π energy going to $\nu_\mu+\bar{\nu}_\mu$

Hubble time: $t_{\rm H} \sim 10^{10}$ yr

- Maximum neutrino intensity is for $\epsilon = 1$: $I_{\rm max} \approx \frac{1}{4}\xi_z~t_{\rm H}~\frac{c}{4\pi} E_{\rm CR}^2 \frac{d\dot{N}_{\rm CR}}{dE_{\rm CR}} \approx 1.5 \cdot 10^{-8} \xi_z~{\rm GeV~cm^{-2}~s^{-1}~sr^{-1}}$

- So the expected neutrino flux is $E_\nu^2 \Phi_{\nu_\mu} \equiv \frac{c}{4\pi} E_\nu^2 \frac{dN_\nu}{dE_\nu} = \frac{1}{2}\epsilon I_{\rm max}$

Waxman & Bahcall, *PRD* 1999

Waxman-Bahcall bound: $E_\nu^2 \Phi_{\nu_\mu} \approx 0.75 \cdot 10^{-8} \xi_z~\epsilon~{\rm GeV~cm^{-2}~s^{-1}~sr^{-1}}$

The need for km-scale detectors

Predicted by Waxman-Bahcall 1998

- Neutrino flux at TeV–PeV: $E^2 \cdot \Phi \sim 10^{-8}$ GeV cm^{-2} s^{-1} sr^{-1}

- Neutrino-nucleon cross section: $\sigma_{\nu p} \sim 10^{-35}\ \mathrm{cm}^2\ (E/\mathrm{GeV})^{0.36}$

At center-of-mass energy of 1 GeV:
$\sigma_{pp} \sim 10^{-28}$ cm^2
$\sigma_{\gamma p} \sim 10^{-29}$ cm^2

- Number of detected neutrinos from half the sky in 1 yr:

$$N = (n_{\mathrm{nucl}} \cdot V_{\mathrm{det}}) \cdot (2\pi) \cdot (1\ \mathrm{yr}) \cdot \int_{100\ \mathrm{TeV}} \Phi(E) \cdot \sigma_{\nu p}(E)\, \mathrm{d}E$$

- To detect $N > 10$ neutrinos, we need

$$V_{\mathrm{det}} > 1\ \mathrm{km}^3$$

The need for km-scale detectors

Predicted by Waxman-Bahcall 1998

- Neutrino flux at TeV–PeV: $E^2 \cdot \Phi \sim 10^{-8}$ GeV cm^{-2} s^{-1} sr^{-1}

- Neutrino-nucleon cross section: $\sigma_{\nu p} \sim 10^{-35}\ \text{cm}^2\ (E/\text{GeV})^{0.36}$

At center-of-mass energy of 1 GeV:
$\sigma_{pp} \sim 10^{-28}$ cm^2
$\sigma_{\gamma p} \sim 10^{-29}$ cm^2

- Number of detected neutrinos from half the sky in 1 yr:

$$N = (n_{\text{nucl}}) \cdot (V_{\text{det}}) \cdot (2\pi) \cdot (1\ \text{yr}) \cdot \int_{100\ \text{TeV}} \Phi(E) \cdot \sigma_{\nu p}(E)\, dE$$

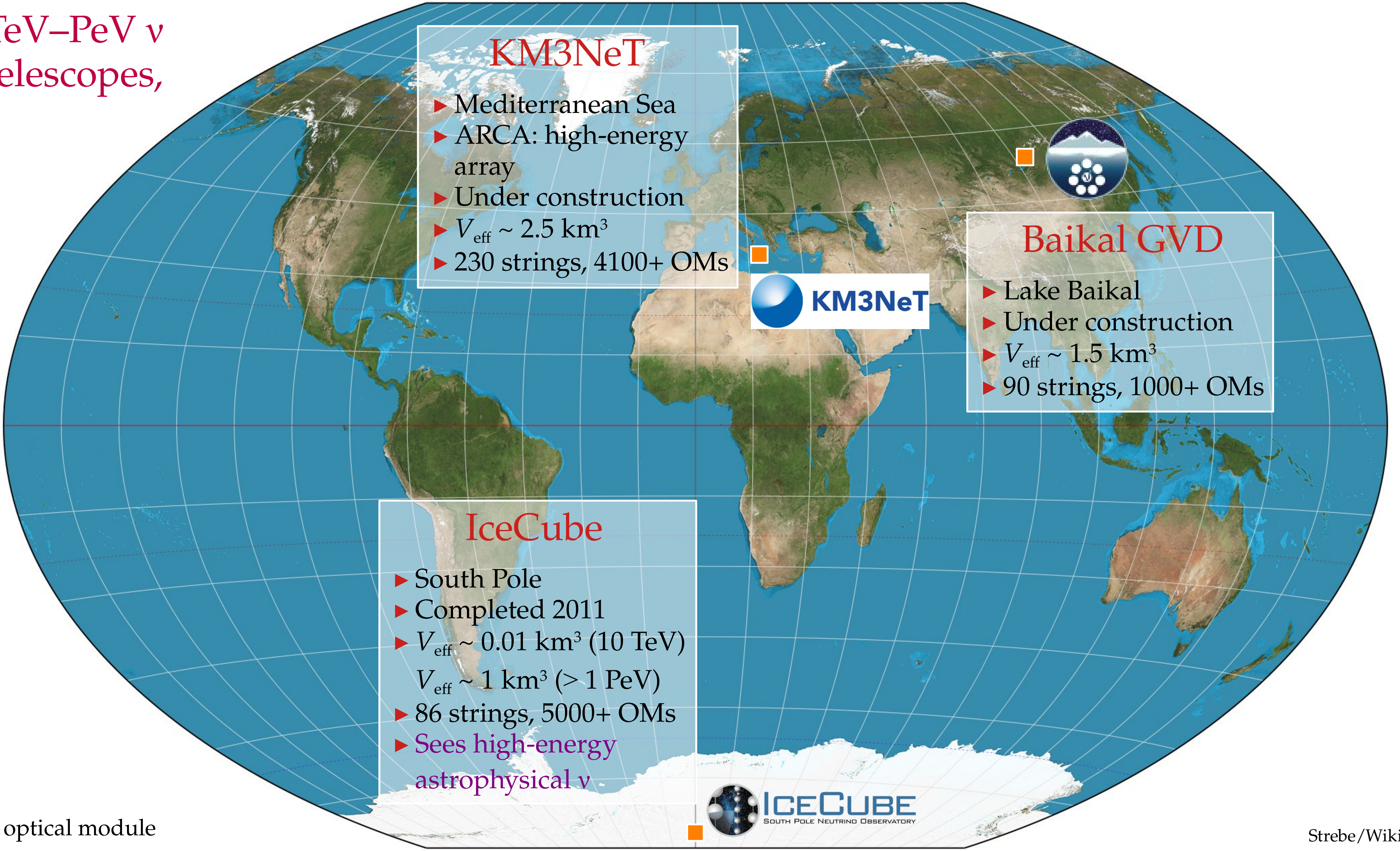
Number density of nucleons: $\sim N_{\text{Av}}$ cm^3

Detector volume

- To detect $N > 10$ neutrinos, we need

$$V_{\text{det}} > 1\ \text{km}^3$$

TeV–PeV ν telescopes,



KM3NeT

- Mediterranean Sea
- ARCA: high-energy array
- Under construction
- $V_{\text{eff}} \sim 2.5\ \text{km}^3$
- 230 strings, 4100+ OMs

KM3NeT

Baikal GVD

- Lake Baikal
- Under construction
- $V_{\text{eff}} \sim 1.5\ \text{km}^3$
- 90 strings, 1000+ OMs

IceCube

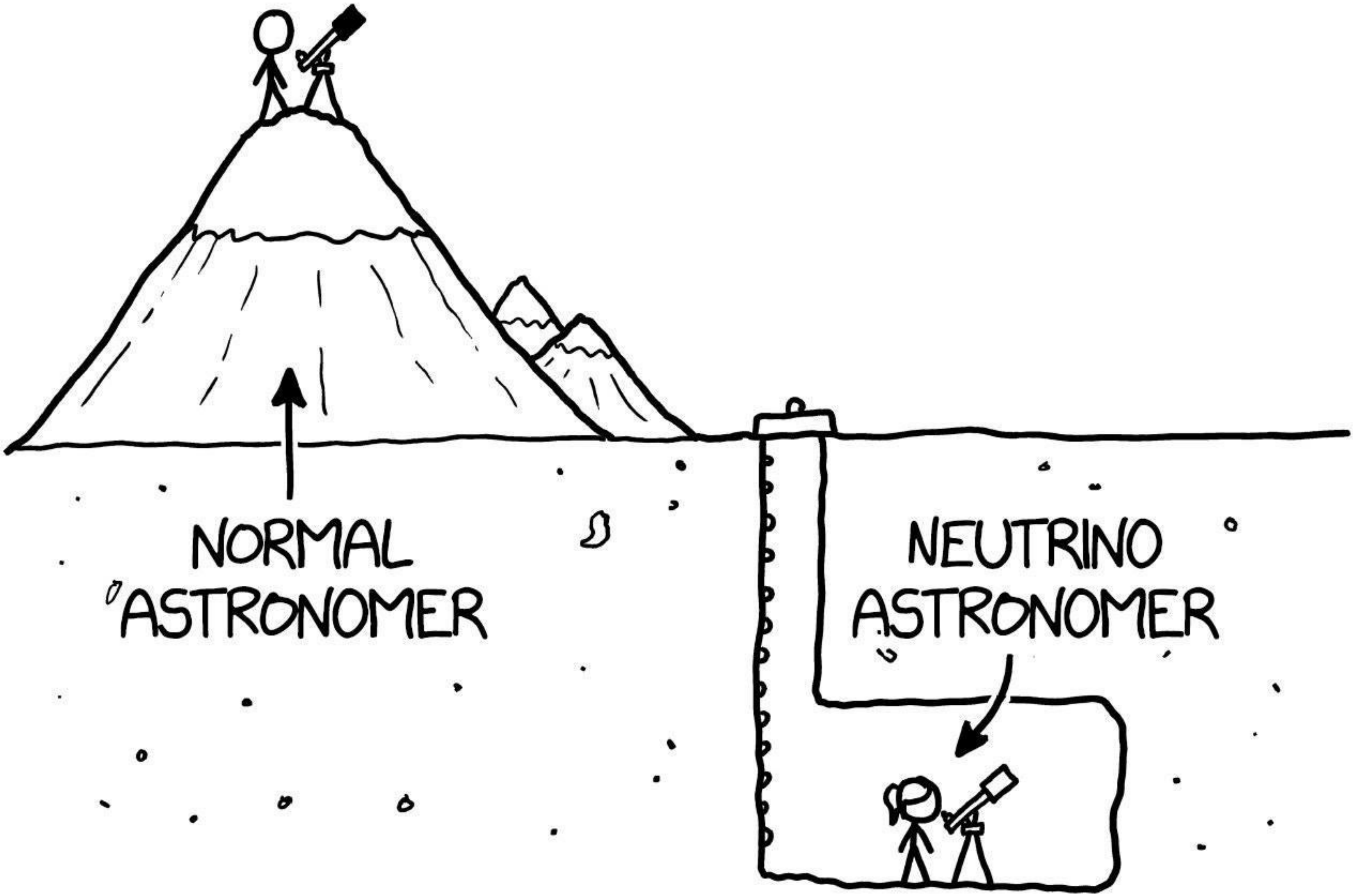
- South Pole
- Completed 2011
- $V_{\text{eff}} \sim 0.01\ \text{km}^3$ (10 TeV)
  $V_{\text{eff}} \sim 1\ \text{km}^3$ (> 1 PeV)
- 86 strings, 5000+ OMs
- Sees high-energy astrophysical ν

ICECUBE
South Pole Neutrino Observatory

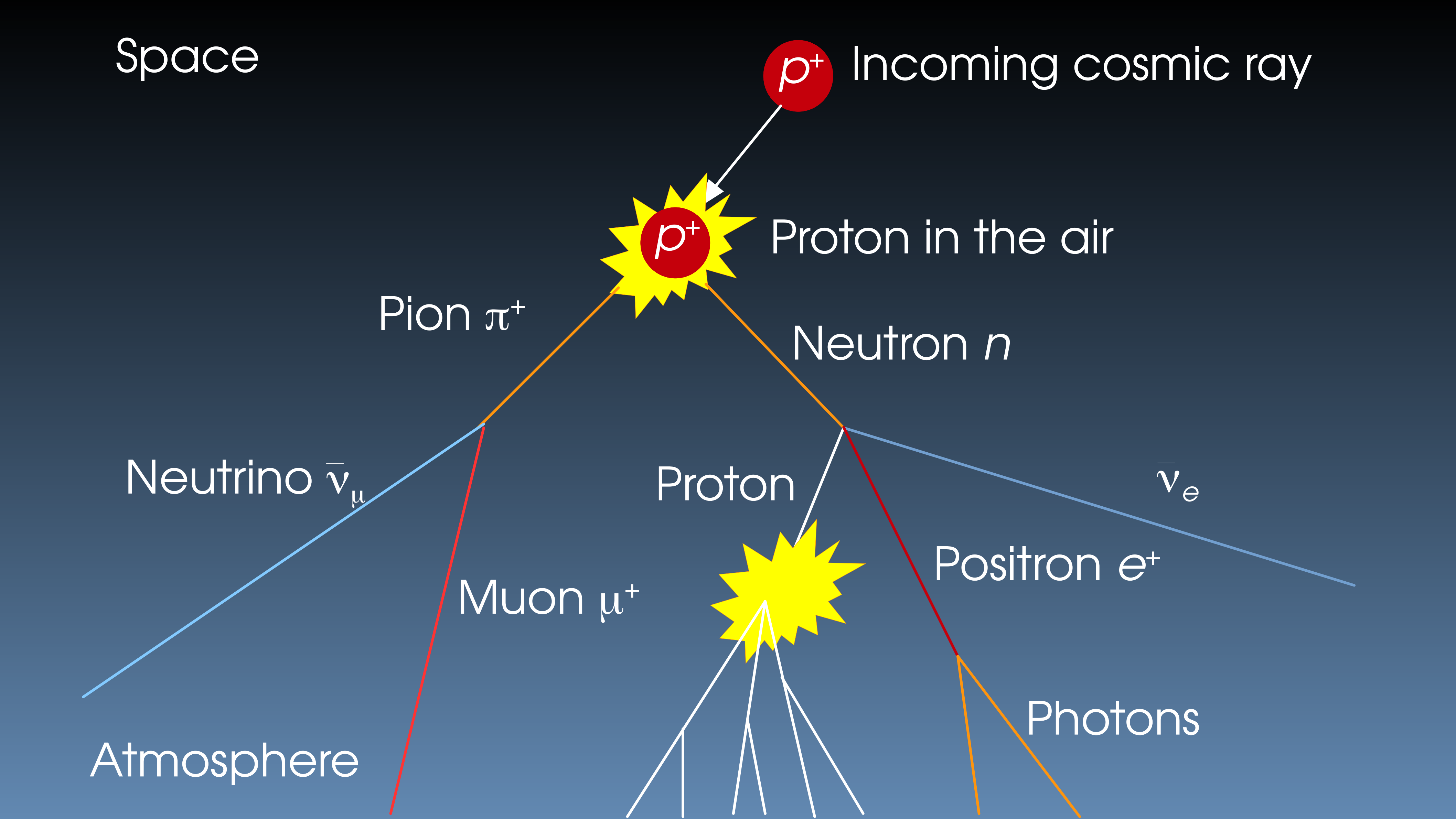
OM: optical module

Strebe/Wikipedia

Yuya Makino, IceCube/NSF

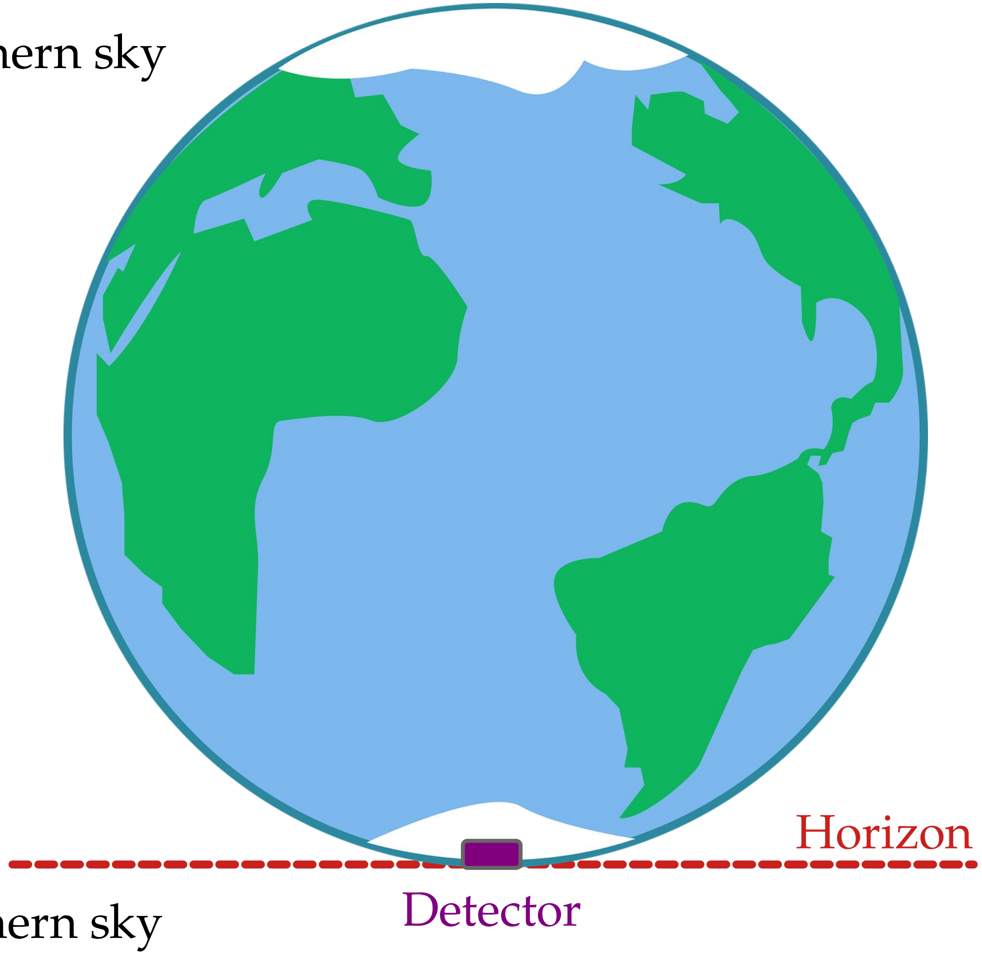


xkcd


Space
p^+ Incoming cosmic ray
p^+
Proton in the air
Pion π^+
Neutron n
Neutrino $\bar{\nu}_\mu$
Proton
$\bar{\nu}_e$
Positron e^+
Muon μ^+
Photons
Atmosphere

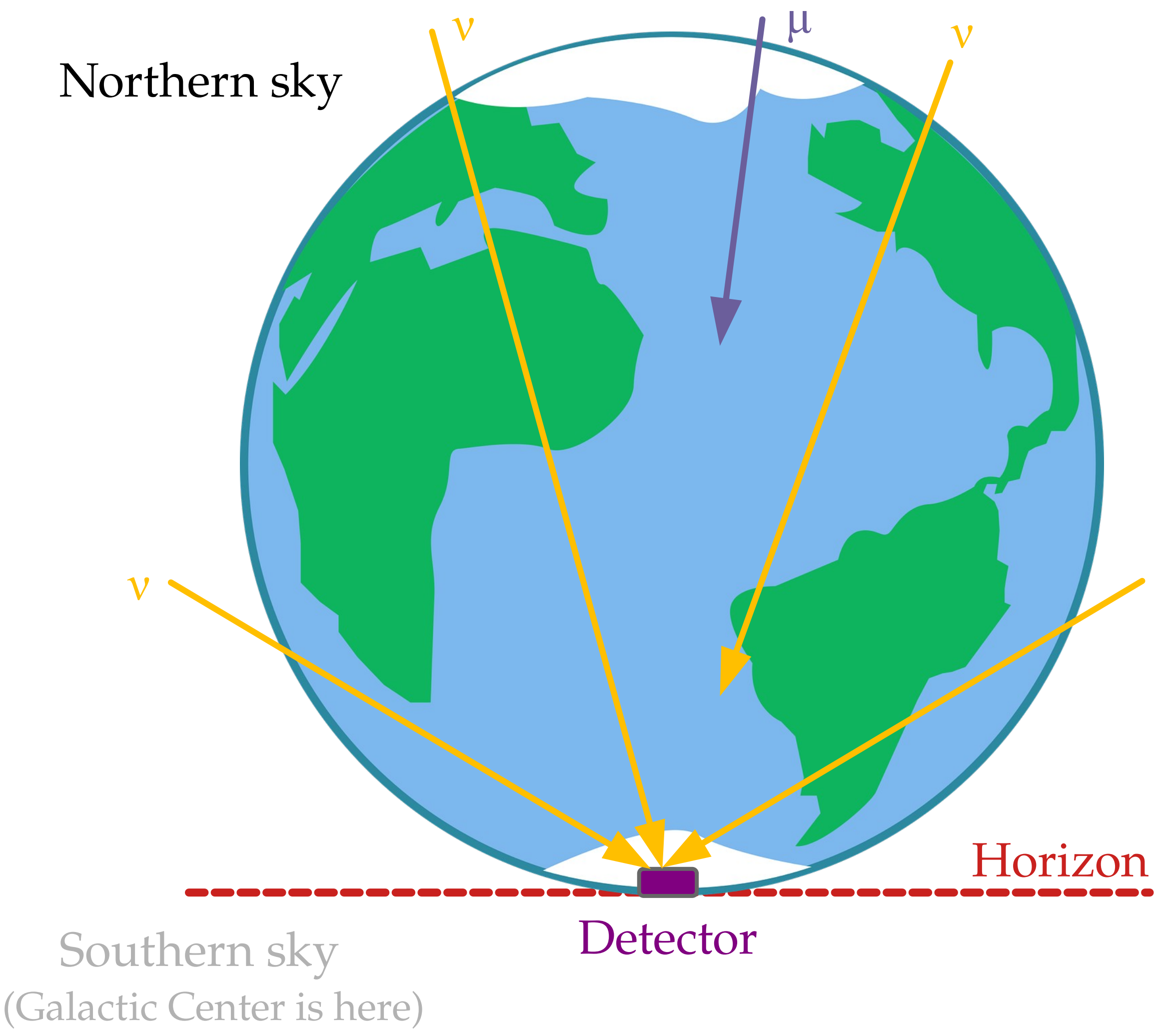
Upgoing *vs.* downgoing neutrinos

Northern sky



Southern sky
(Galactic Center is here)

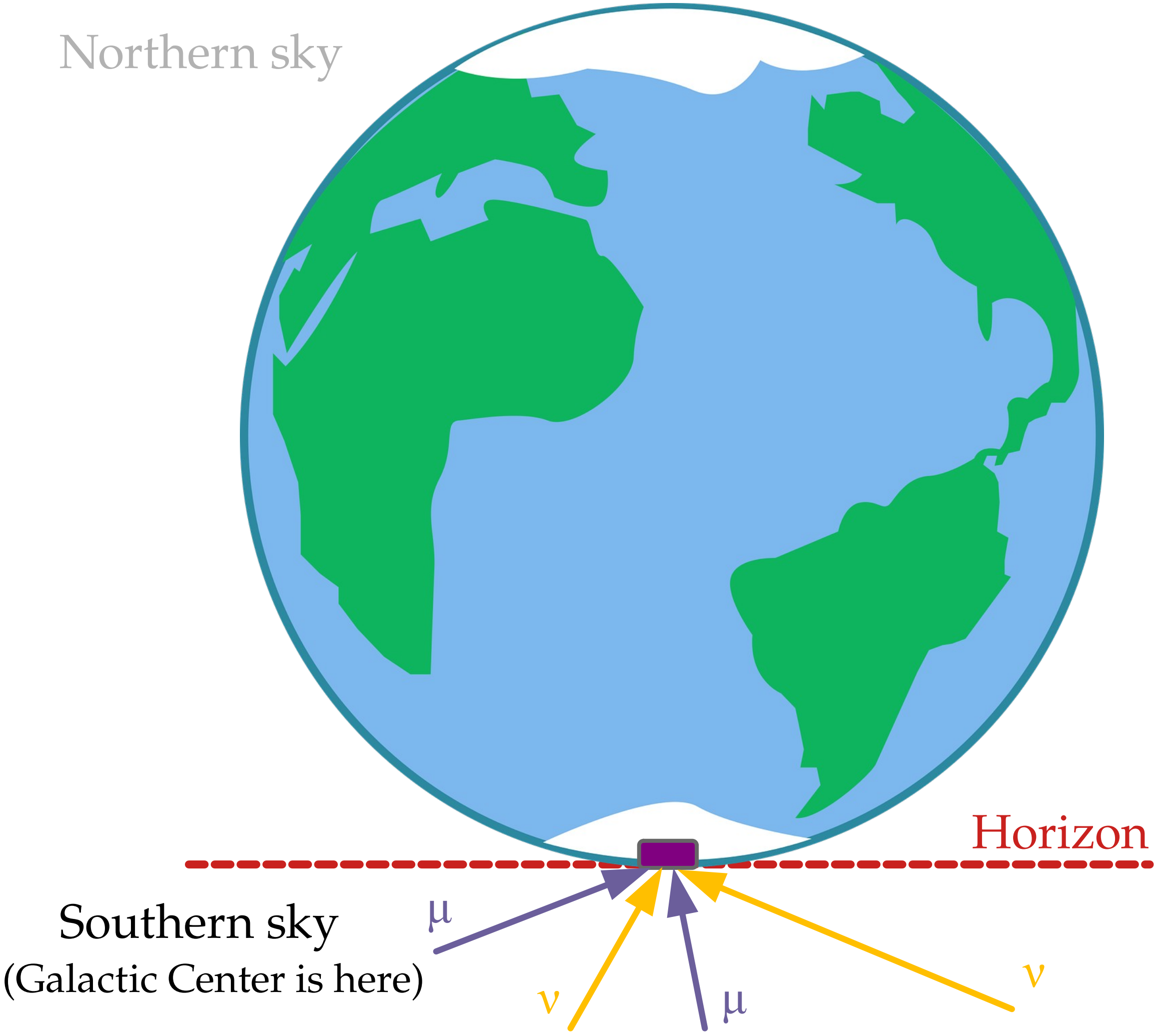
Upgoing *vs.* downgoing neutrinos



Neutrinos from the Northern sky
≡
Upgoing neutrinos

- Atmospheric muons stopped
- Dominated by atmospheric ν
- High-energy ν flux attenuated
- High statistics
- Good for finding sources with through-going muon tracks

Upgoing *vs.* downgoing neutrinos



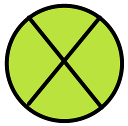
Neutrinos from the Southern sky
≡
Downgoing neutrinos

- Need to mitigate atmospheric muons and ν:
  - Use higher-energy events
  - Use starting a self-veto
- Dominated by astrophysical ν (*after* event selection)
- Low statistics
- Good for measuring the diffuse flux of astrophysical ν

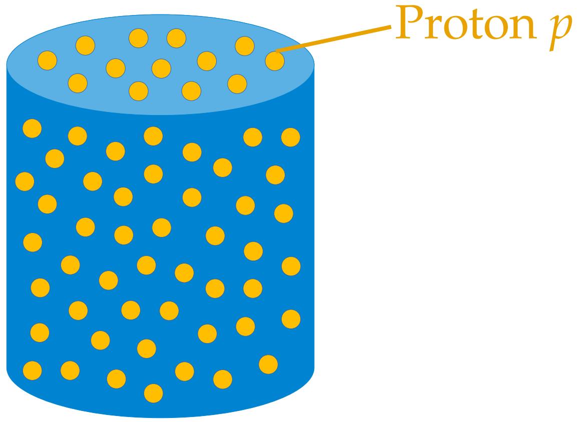
Detecting the undetectable

Detecting the undetectable

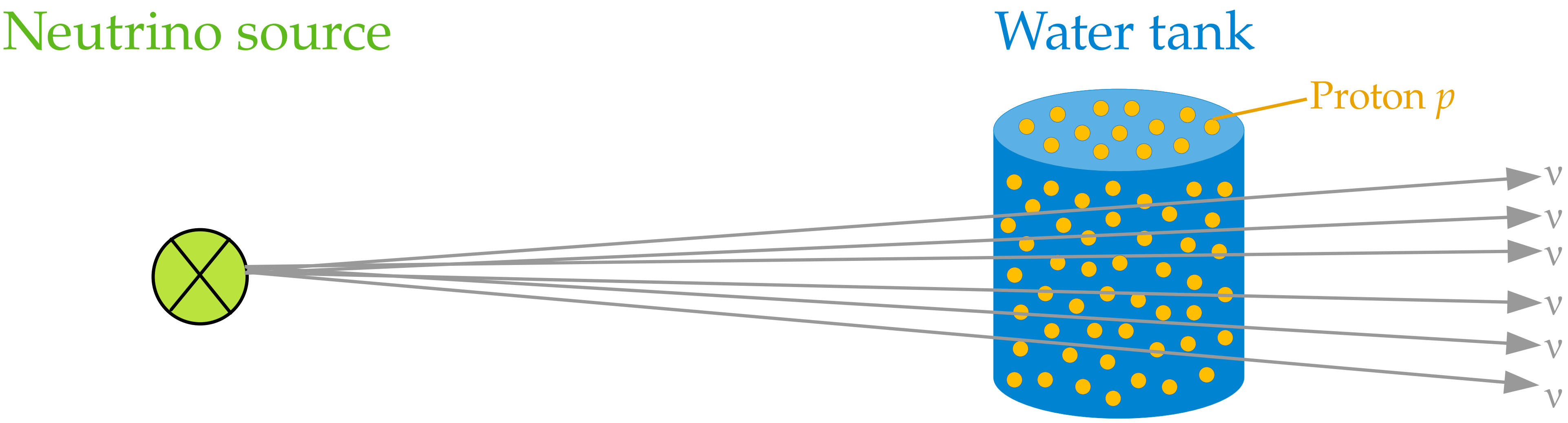
Neutrino source



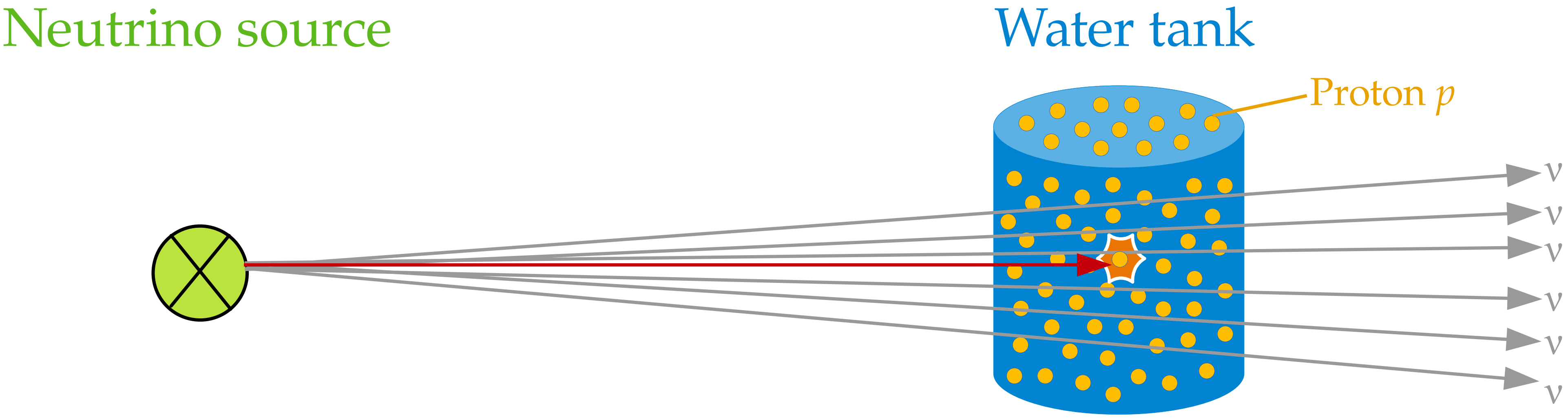
Water tank



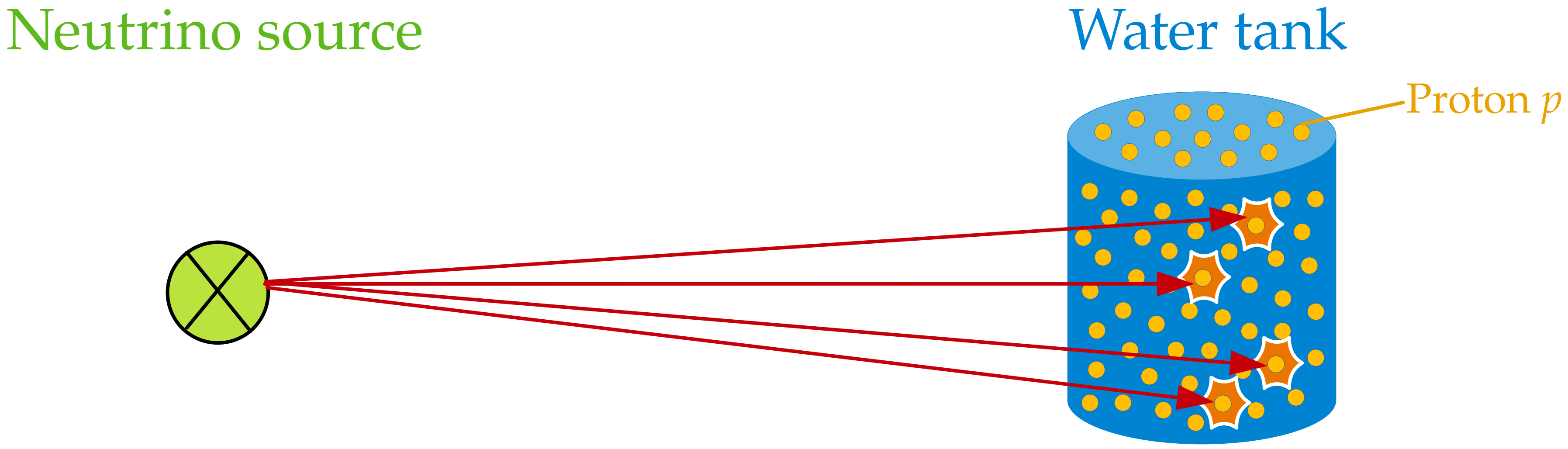
Detecting the undetectable



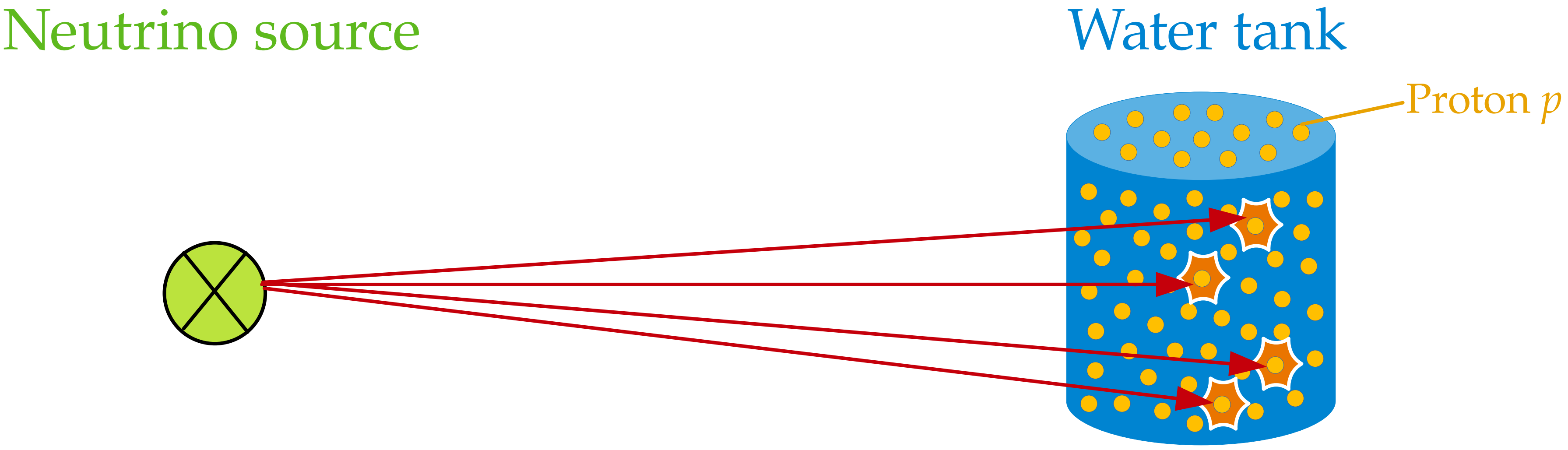
Detecting the undetectable



Detecting the undetectable

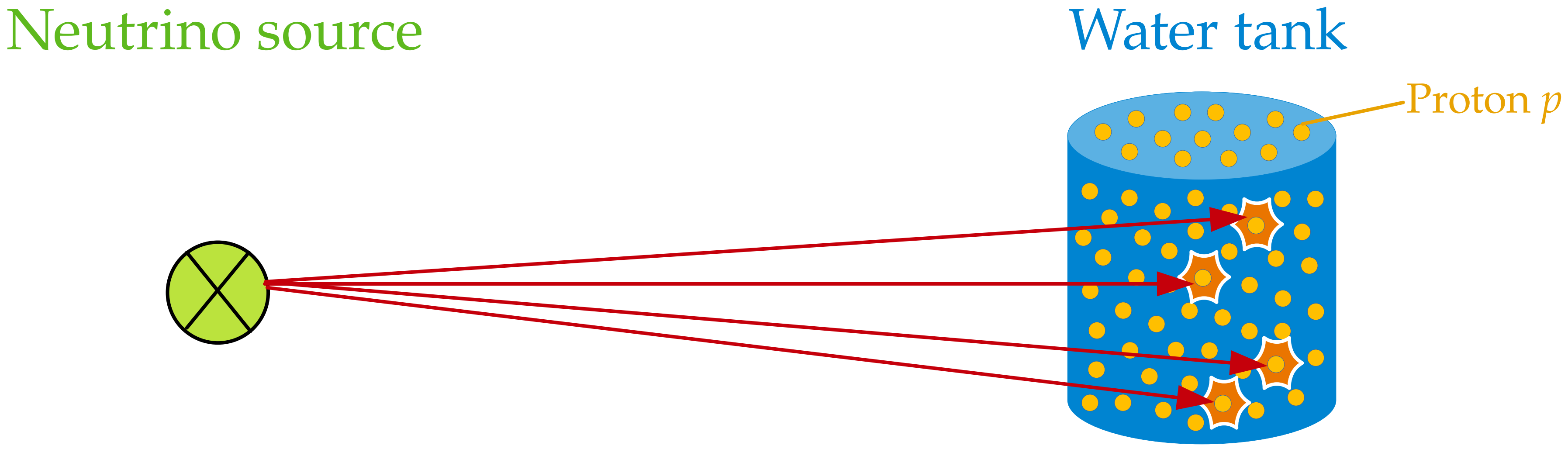


Detecting the undetectable



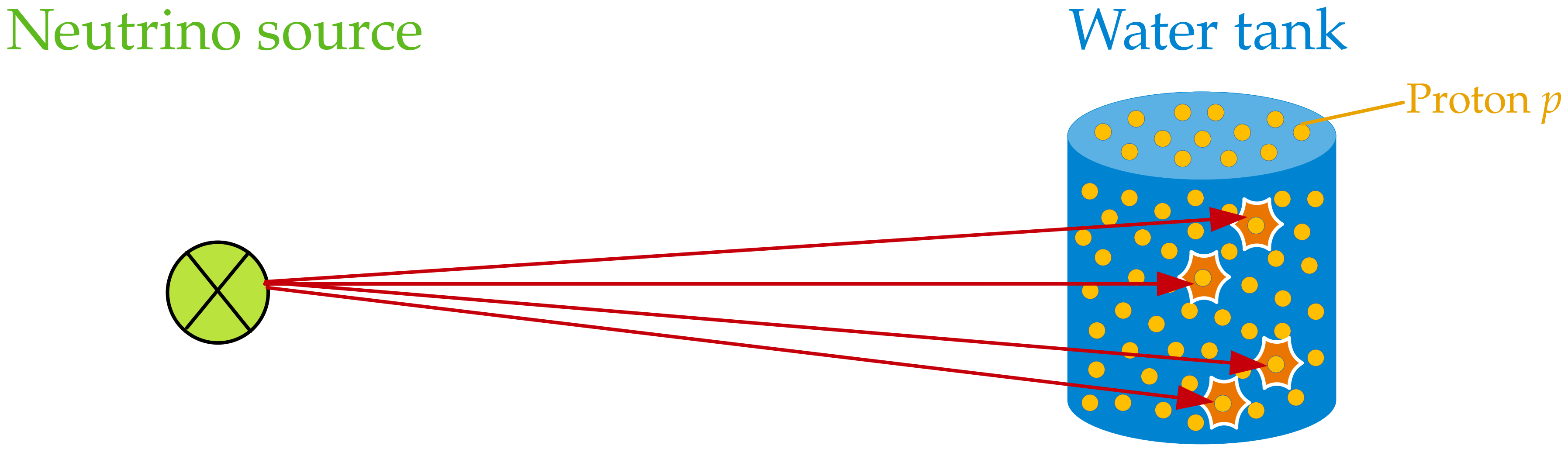
Number of interacting ν =

Detecting the undetectable



Number of interacting ν = Chance that one ν interacts with one p

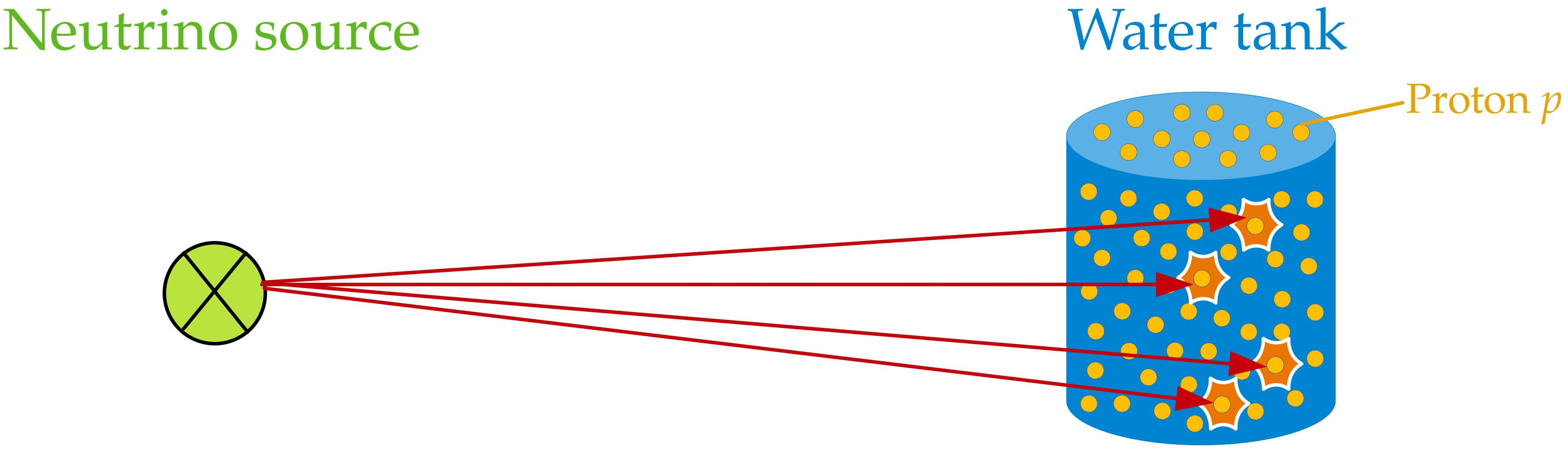
Detecting the undetectable



Number of interacting ν = Chance that one ν interacts with one p

Fixed by Nature (weak interactions): *neutrino-proton cross section*

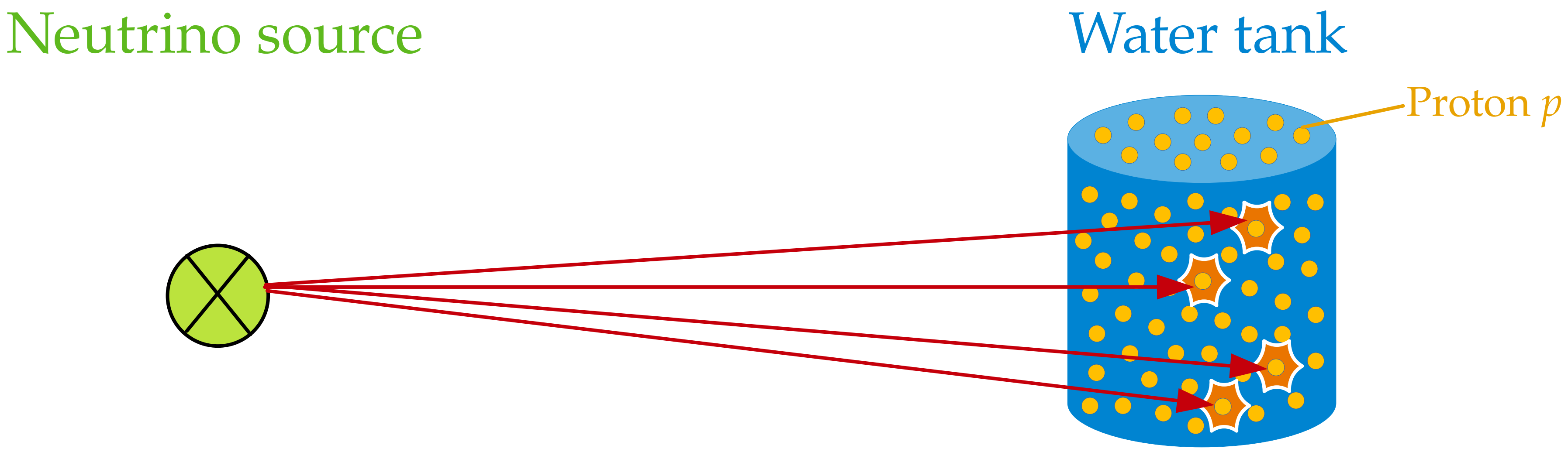
Detecting the undetectable



Number of interacting ν = Chance that one ν interacts with one p × Number of ν that reach the tank

Fixed by Nature (weak interactions): *neutrino-proton cross section*

Detecting the undetectable

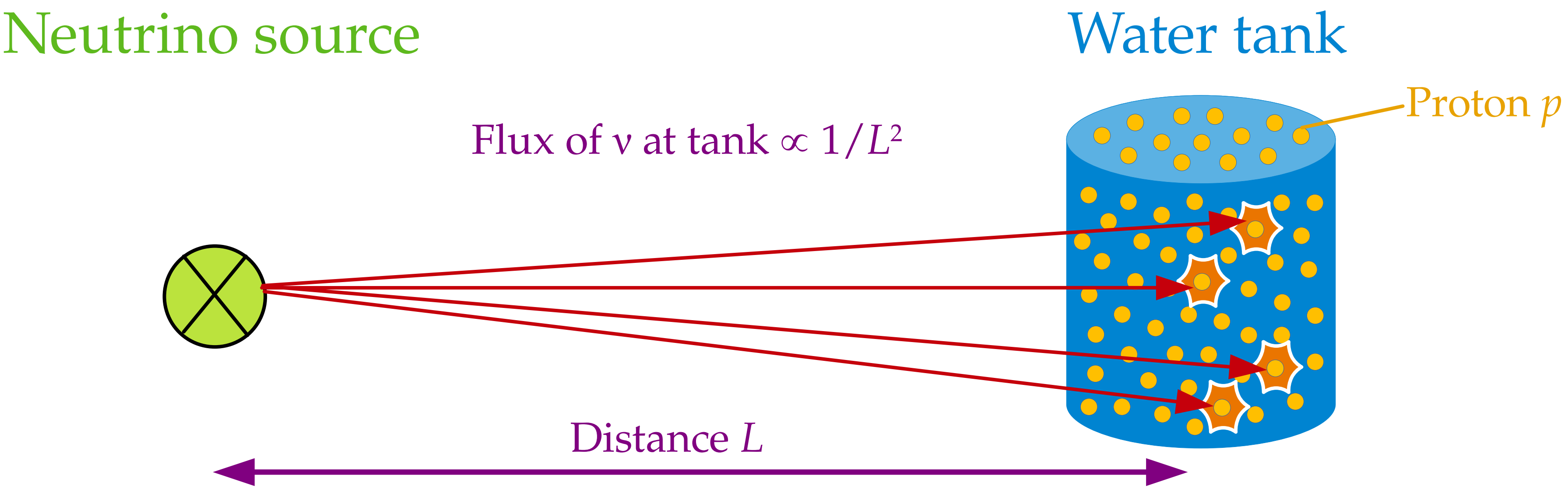


Number of interacting ν = Chance that one ν interacts with one p × Number of ν that reach the tank

Fixed by Nature (weak interactions): *neutrino-proton cross section*

Use an intense source, place the tank close to it, and be patient

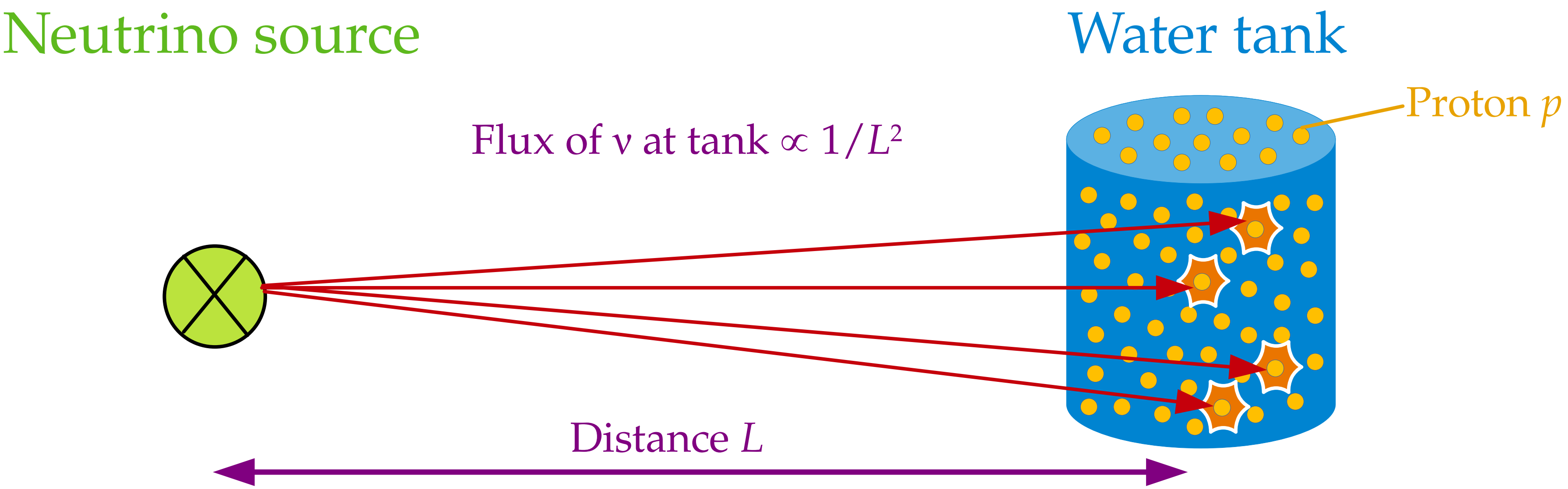
Detecting the undetectable



Number of interacting ν = Chance that one ν interacts with one p × Number of ν that reach the tank

- Chance that one ν interacts with one p: Fixed by Nature (weak interactions): *neutrino-proton cross section*
- Number of ν that reach the tank: Use an intense source, place the tank close to it, and be patient

Detecting the undetectable

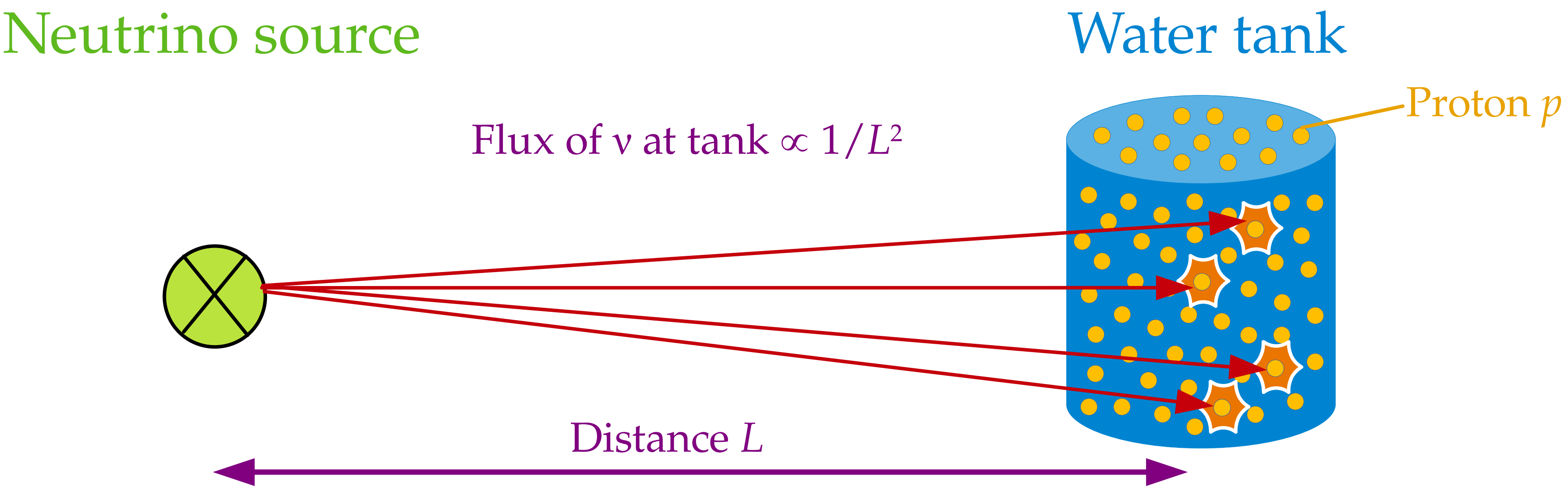


Number of interacting ν = Chance that one ν interacts with one p × Number of ν that reach the tank × Number of p in the tank

Chance that one ν interacts with one p: Fixed by Nature (weak interactions): *neutrino-proton cross section*

Number of ν that reach the tank: Use an intense source, place the tank close to it, and be patient

Detecting the undetectable

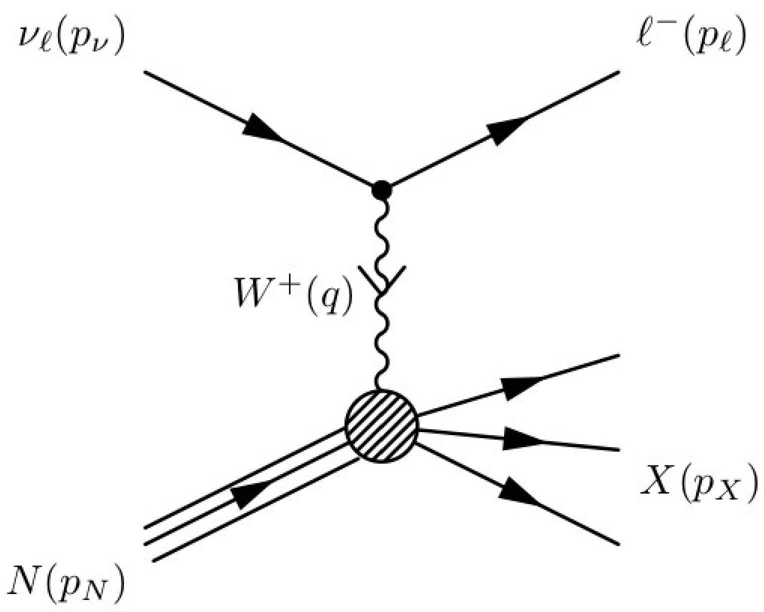


Number of interacting ν = Chance that one ν interacts with one p × Number of ν that reach the tank × Number of p in the tank

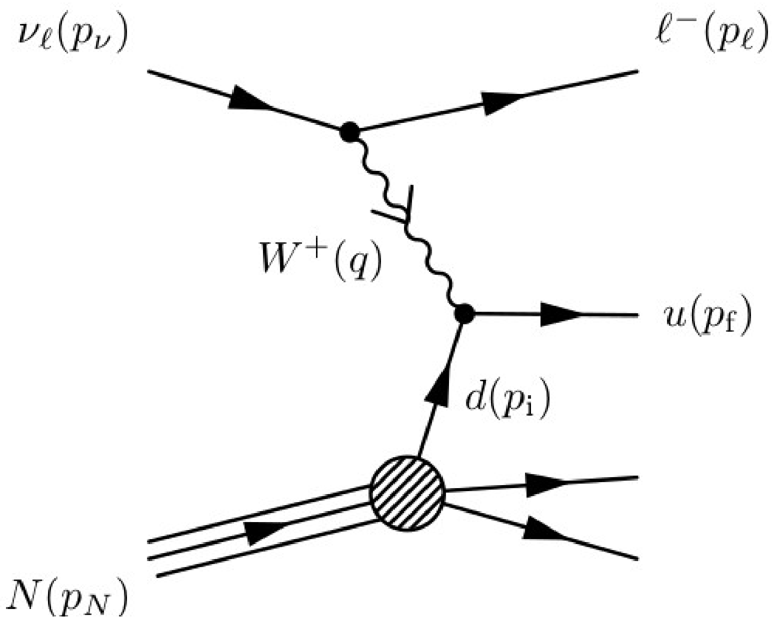
- Chance that one ν interacts with one p: Fixed by Nature (weak interactions): *neutrino-proton cross section*
- Number of ν that reach the tank: Use an intense source, place the tank close to it, and be patient
- Number of p in the tank: Build as big a water tank as possible

Neutrino-nucleon deep inelastic scattering

What you see

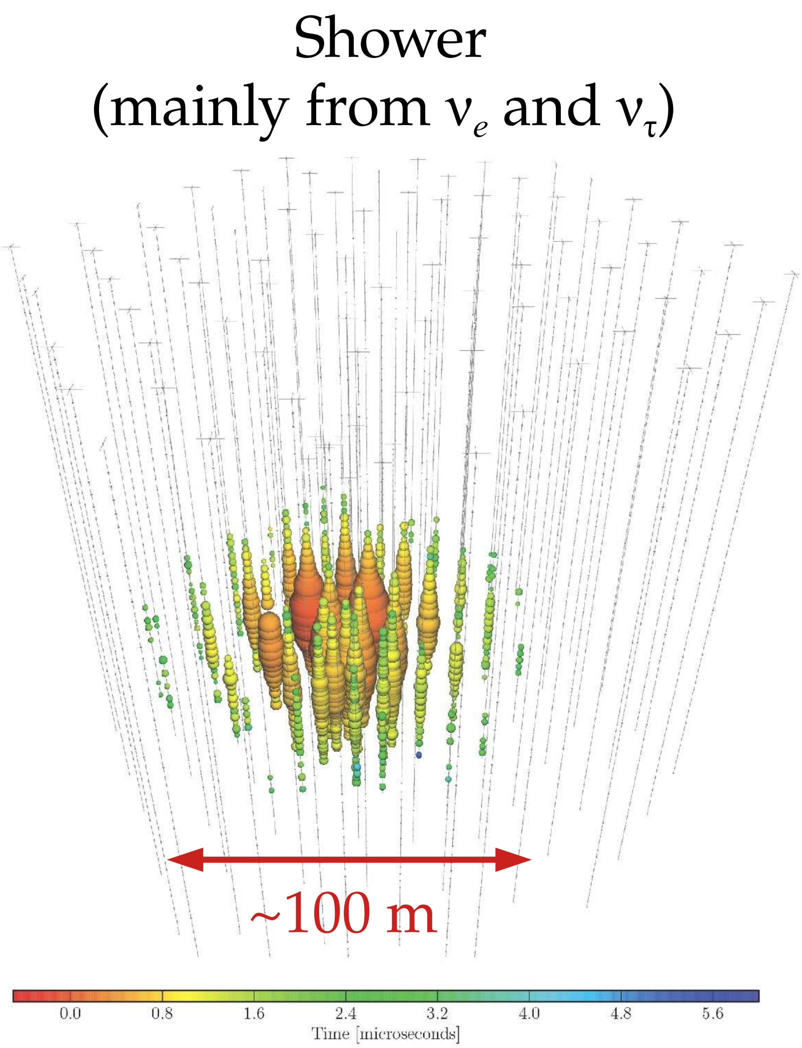


Beneath the hood

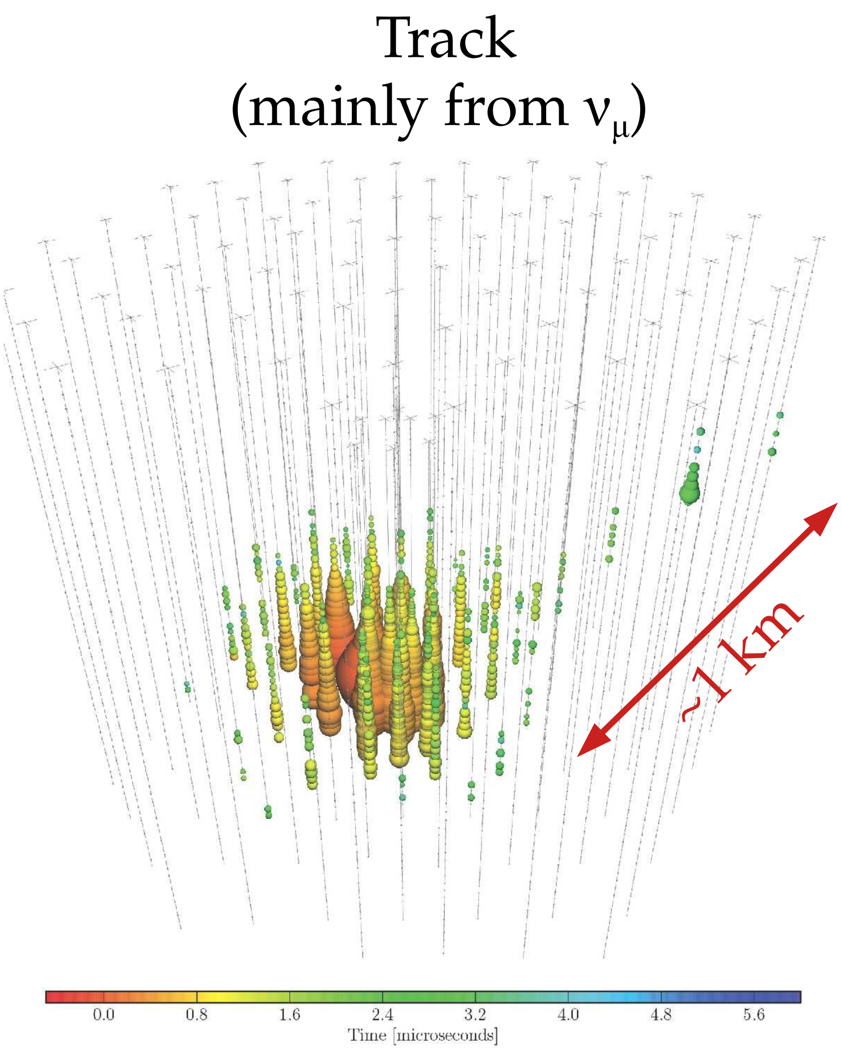


(Plus the equivalent neutral-current process (Z-exchange))

Giunti & Kim, *Fundamentals of Neutrino Physics & Astrophysics*



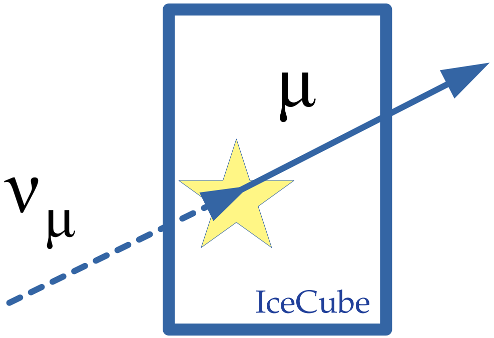
Poor angular resolution: ~10°



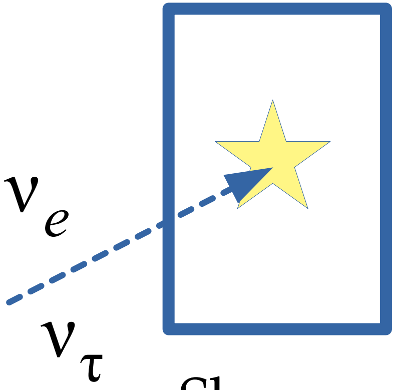
Angular resolution: < 1°

Contained *vs.* uncontained events

Contained events



Starting track

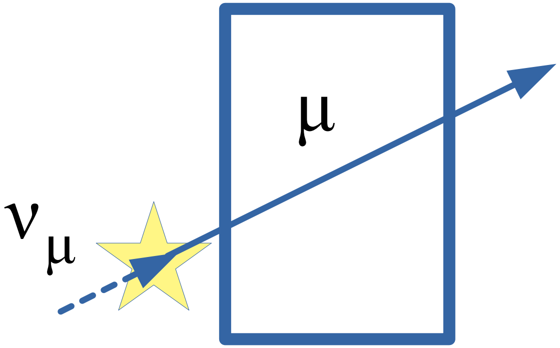


Shower

Pro: Clean determination of E_ν

Con: Few events (~100 in 10 yr)

Through-going muons

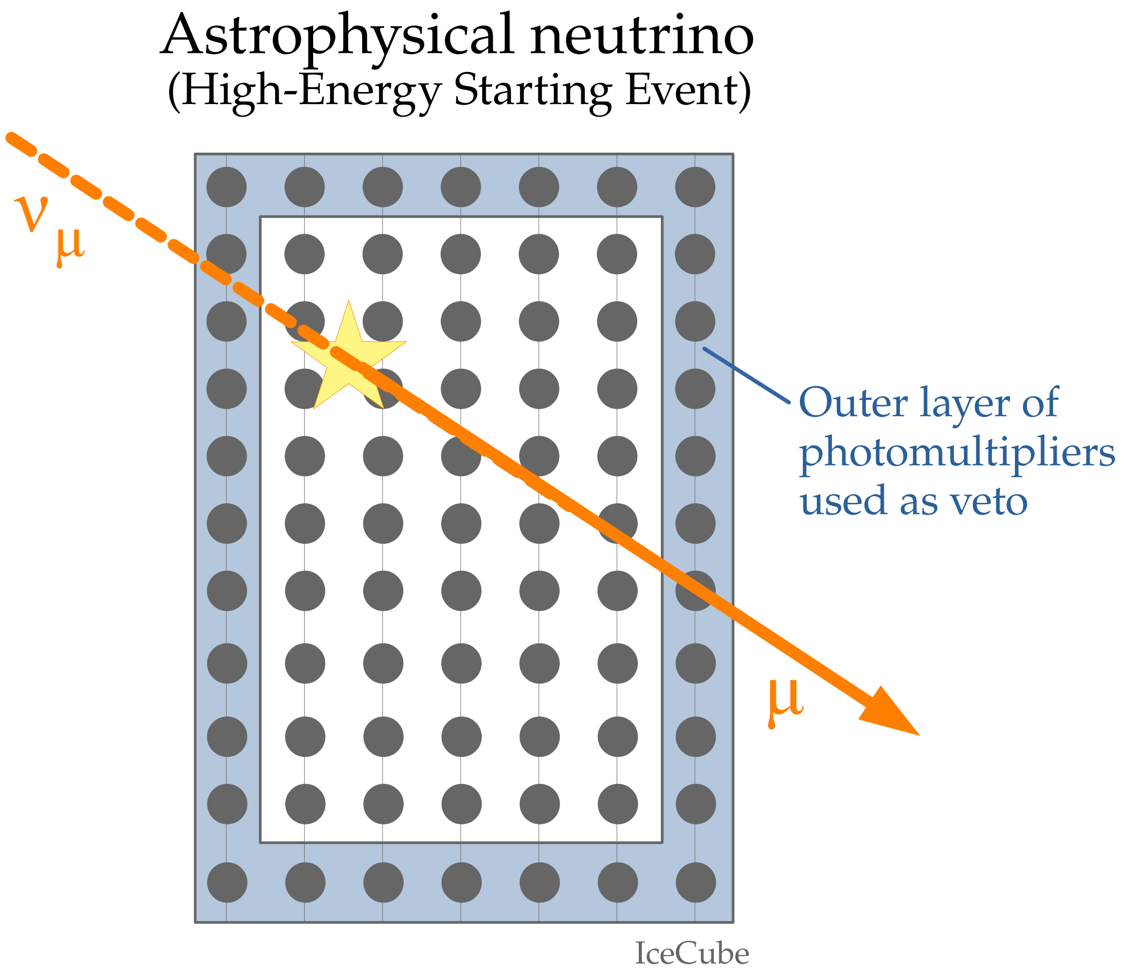


Through-going muon

Pro: Lots of events (few 100k)

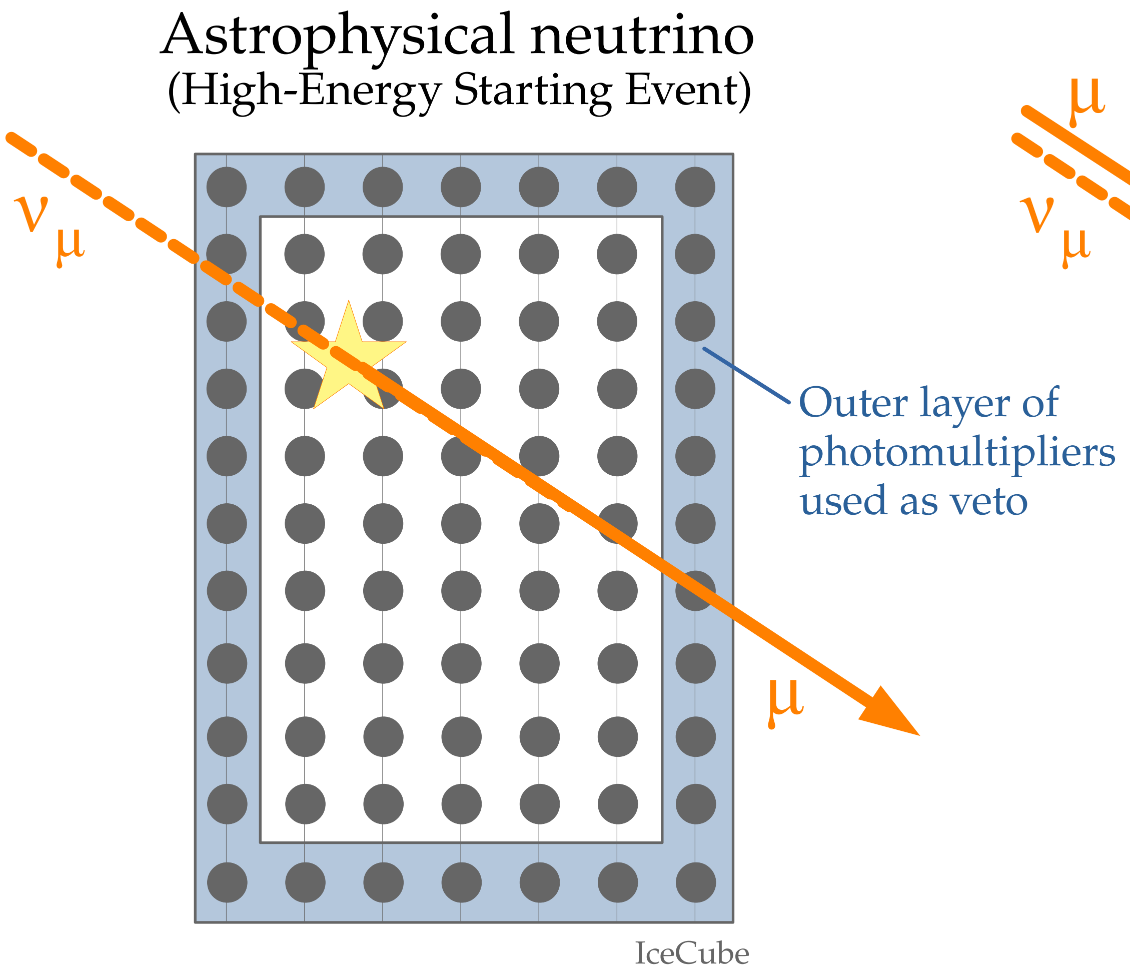
Con: Uncertain estimates of E_ν

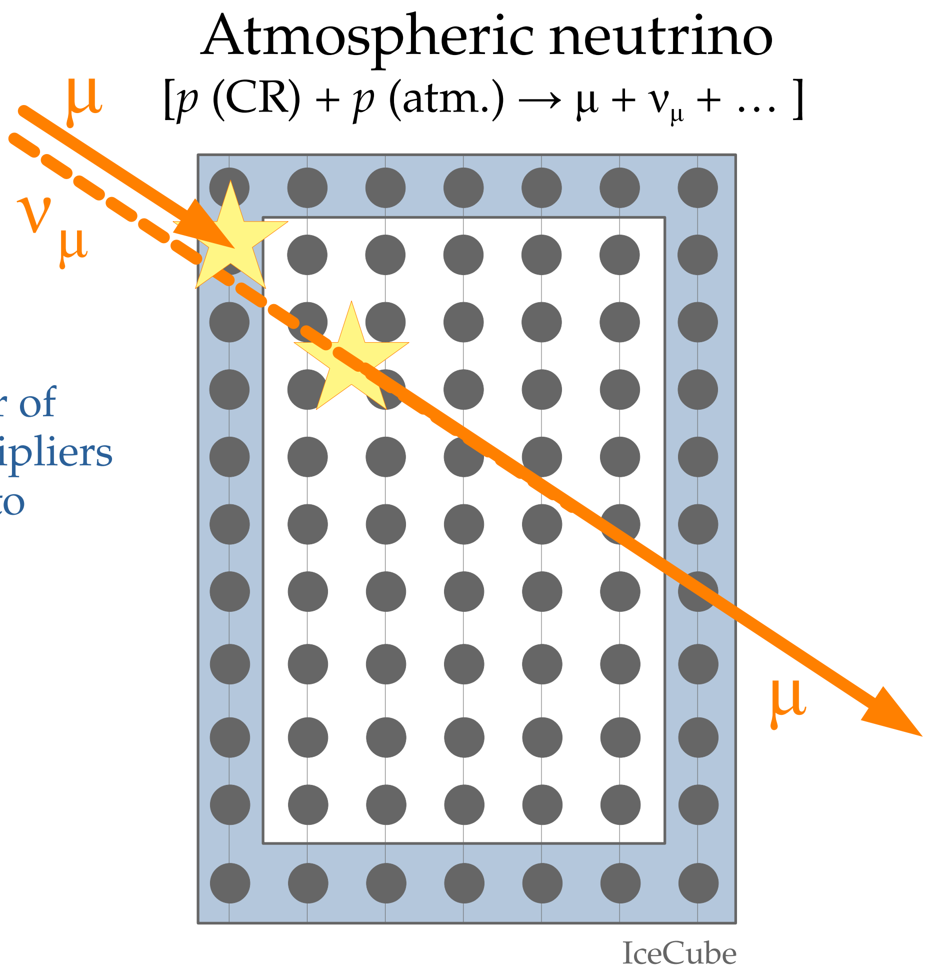
IceCube self-veto: High-Energy Starting Events (HESE)



Schönert, Gaisser, Resconi, Schulz, *PRD* 2008
Gaisser, Jero, Karle, van Santen, *PRD* 2014

IceCube self-veto: High-Energy Starting Events (HESE)





Schönert, Gaisser, Resconi, Schulz, *PRD* 2008
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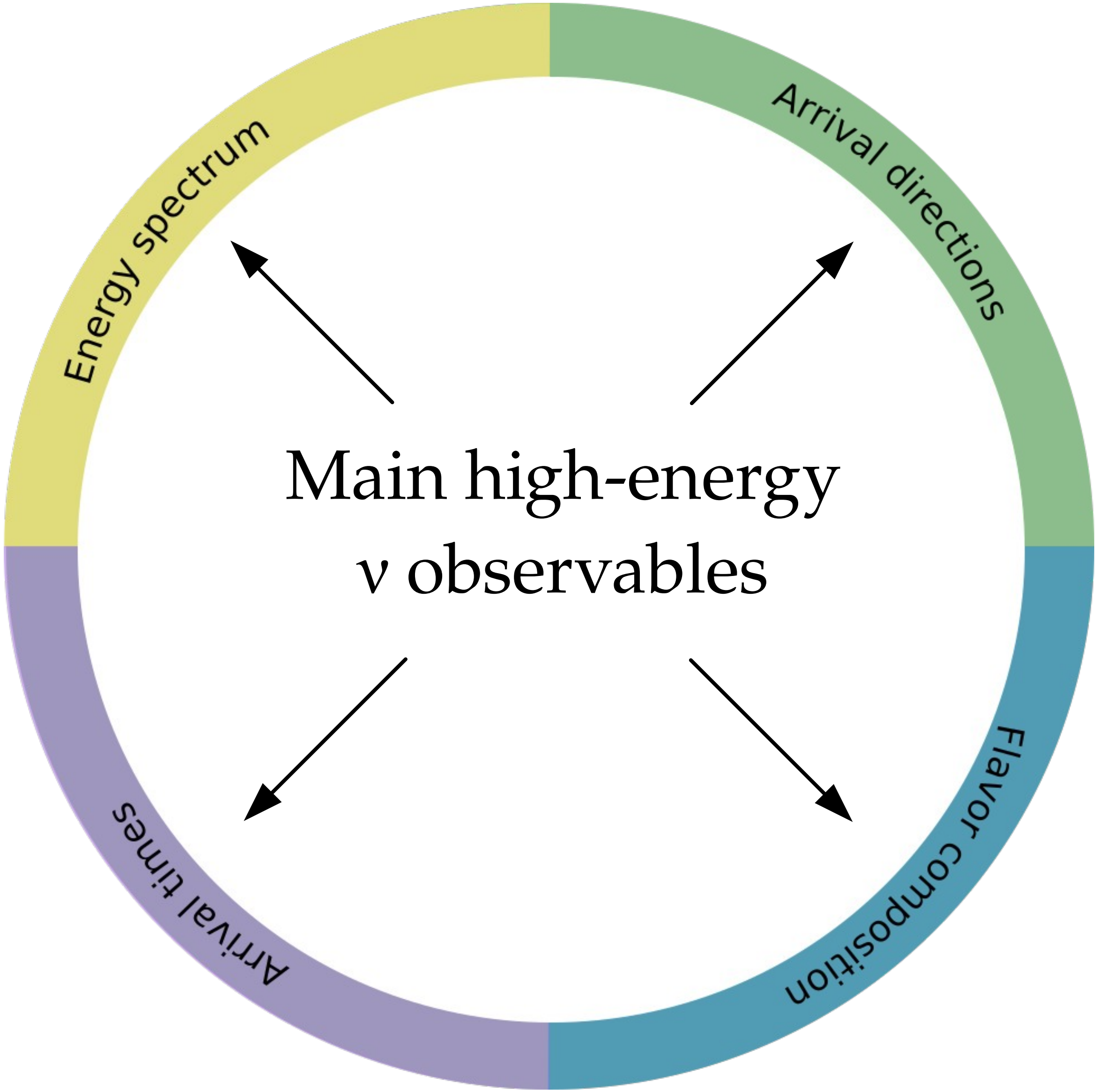
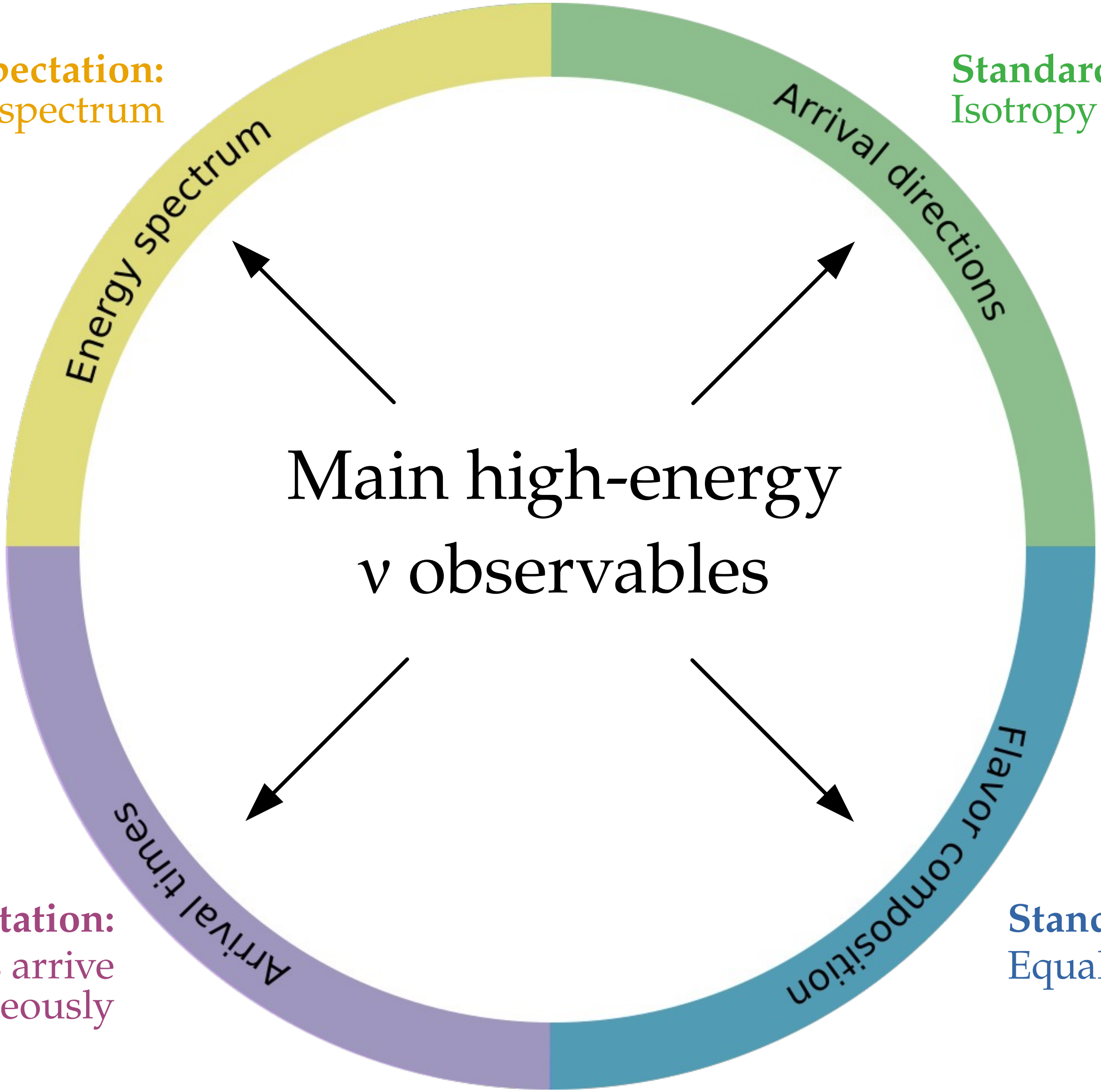

Energy spectrum
Arrival directions
Main high-energy
ν observables
Arrival times
Flavor composition

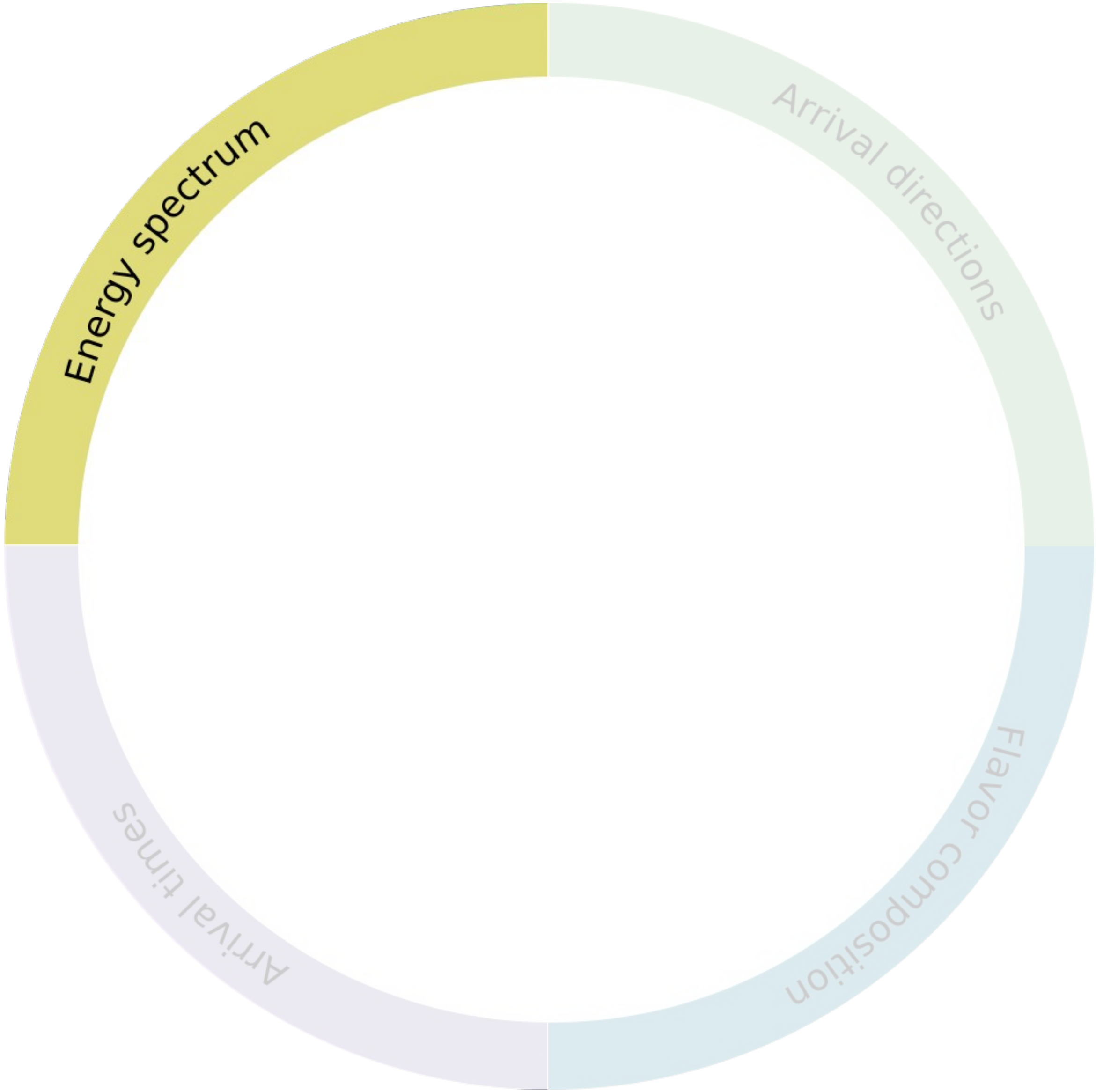
Standard expectation:
Power-law energy spectrum

Standard expectation:
Isotropy (for diffuse flux)



Standard expectation:
ν and γ from transients arrive simultaneously

Standard expectation:
Equal number of ν_e, ν_μ, ν_τ


Energy spectrum
Arrival directions
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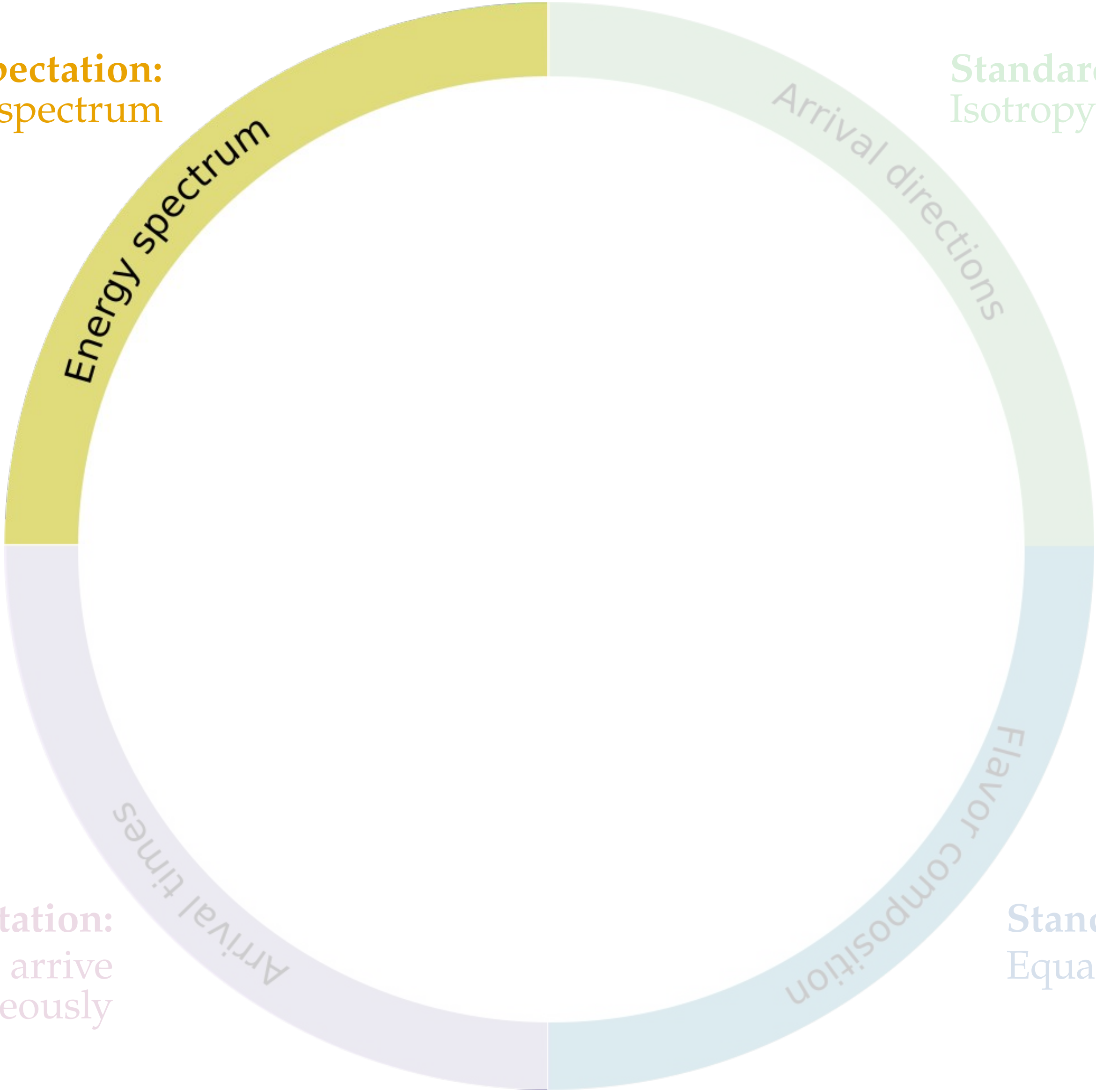
Isotropy (for diffuse flux)

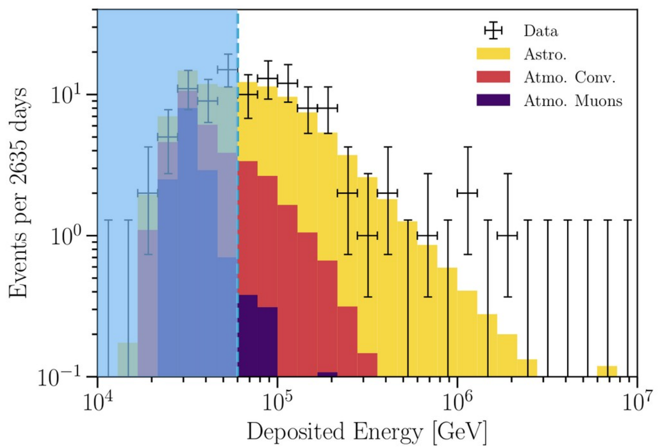


Standard expectation:
ν and γ from transients arrive simultaneously

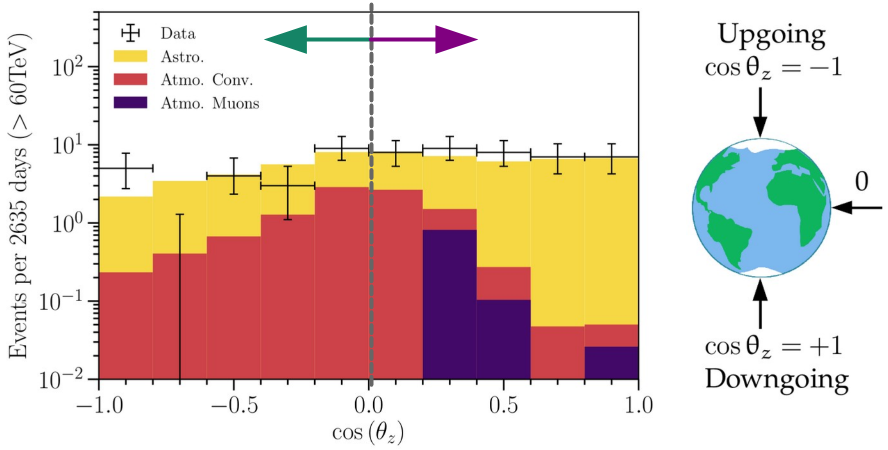
Standard expectation:
Equal number of ν_e, ν_μ, ν_τ

Energy spectrum (7.5 yr)

100+ contained events above 60 TeV:

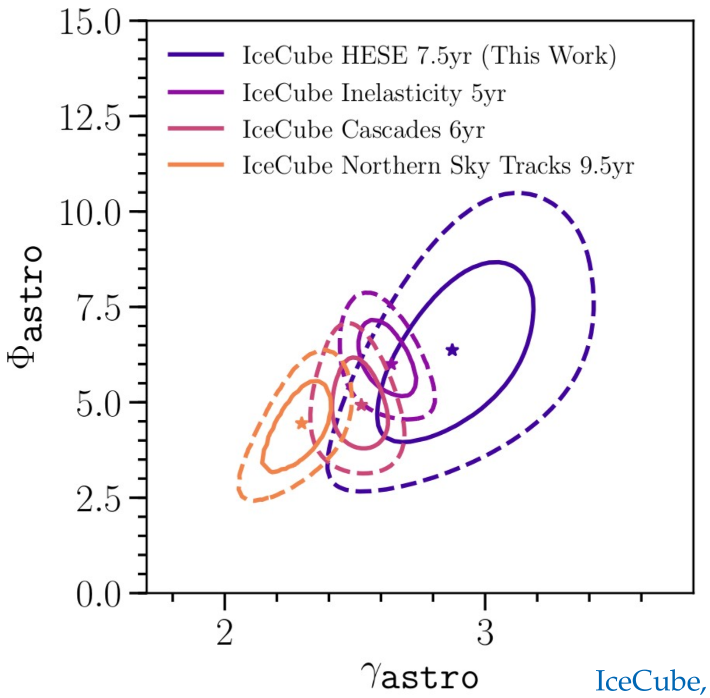


ν attenuated by Earth   Atm. ν and μ vetoed



Data is fit well by a single power law:

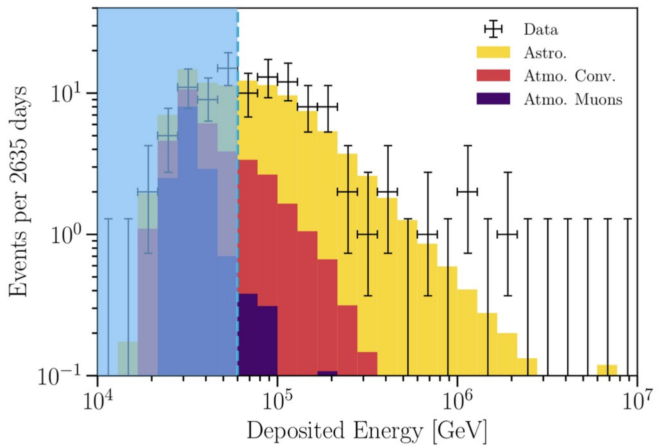
$$\frac{d\Phi_{6\nu}}{dE_\nu} = \Phi_{\text{astro}} \left(\frac{E_\nu}{100\ \text{TeV}} \right)^{-\gamma_{\text{astro}}} \cdot 10^{-18}\ \text{GeV}^{-1}\ \text{cm}^{-2}\ \text{s}^{-1}\ \text{sr}^{-1}$$



IceCube, *PRD* 2021

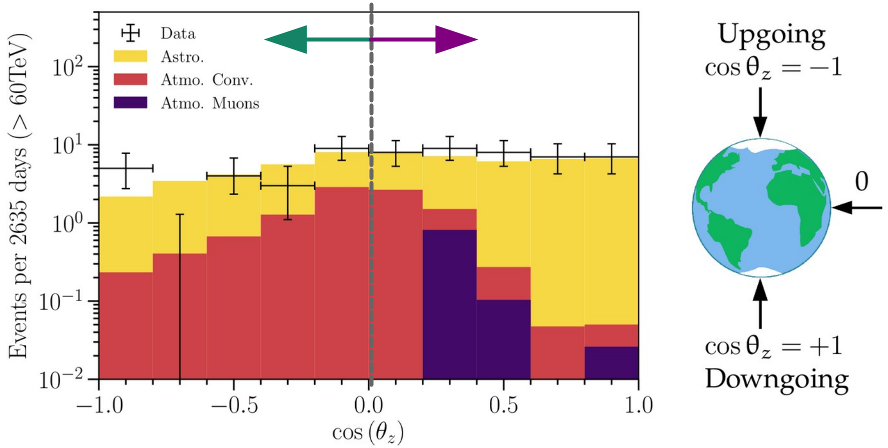
Energy spectrum (7.5 yr)

100+ contained events above 60 TeV:



ν attenuated by Earth

Atm. ν and μ vetoed



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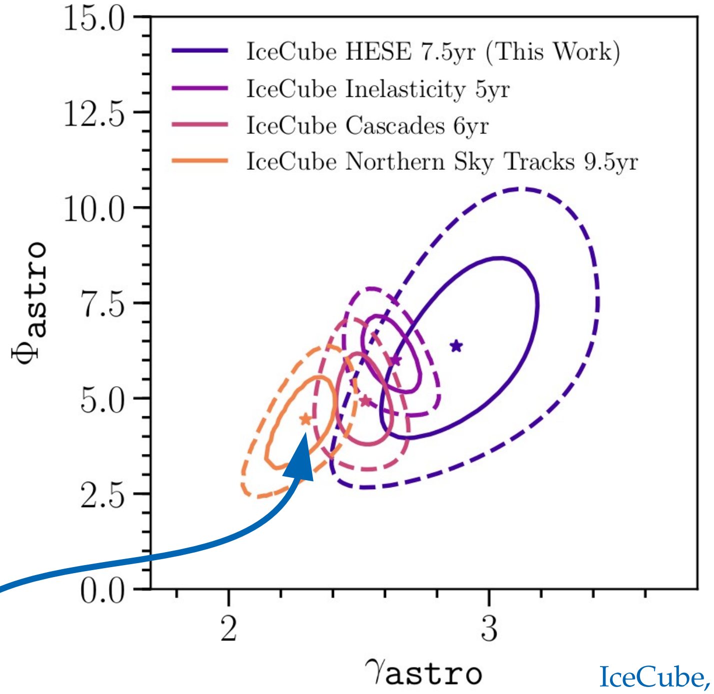
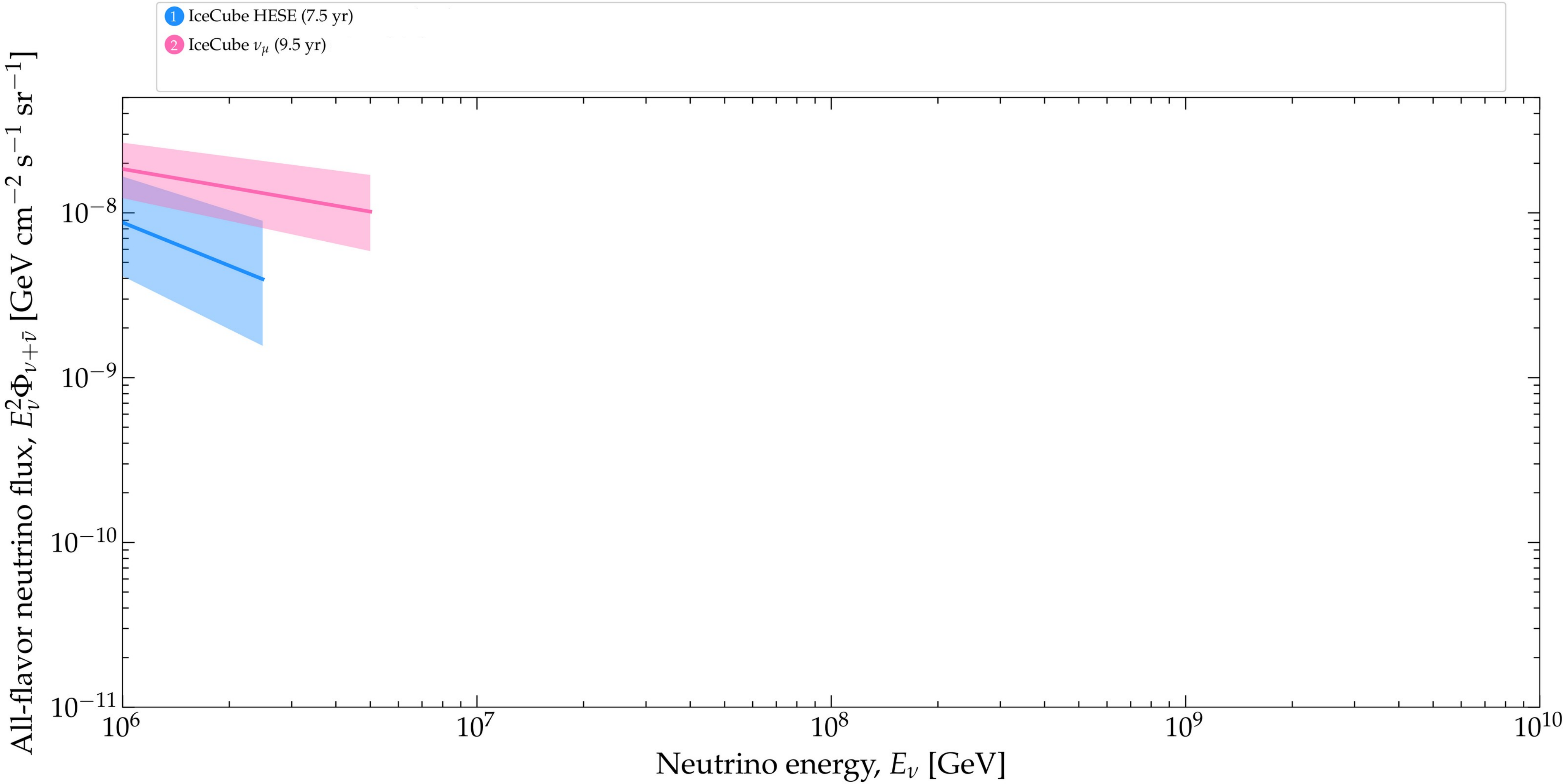
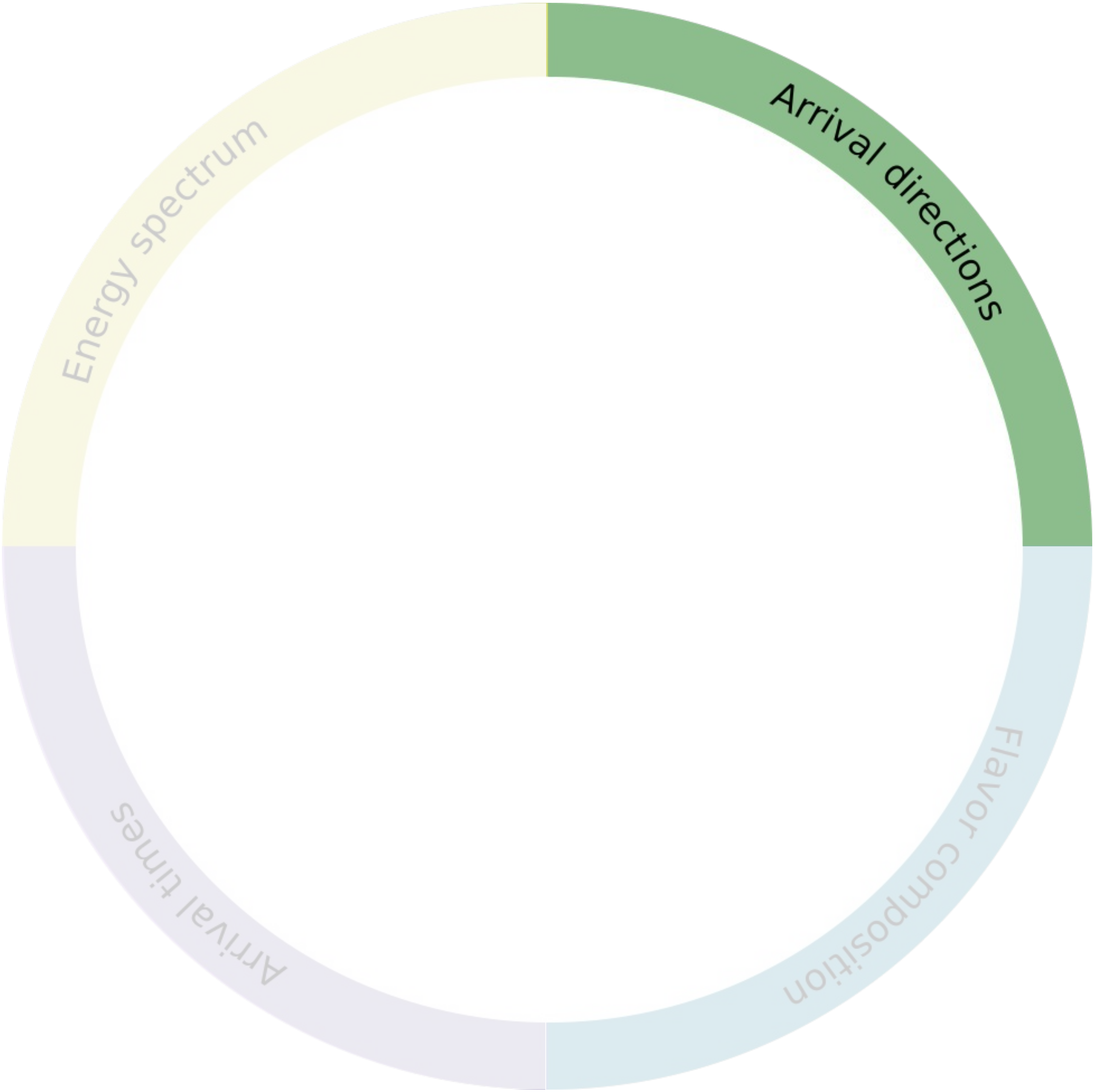


IceCube, *PRD* 2021

Spectrum looks harder for through-going ν_μ

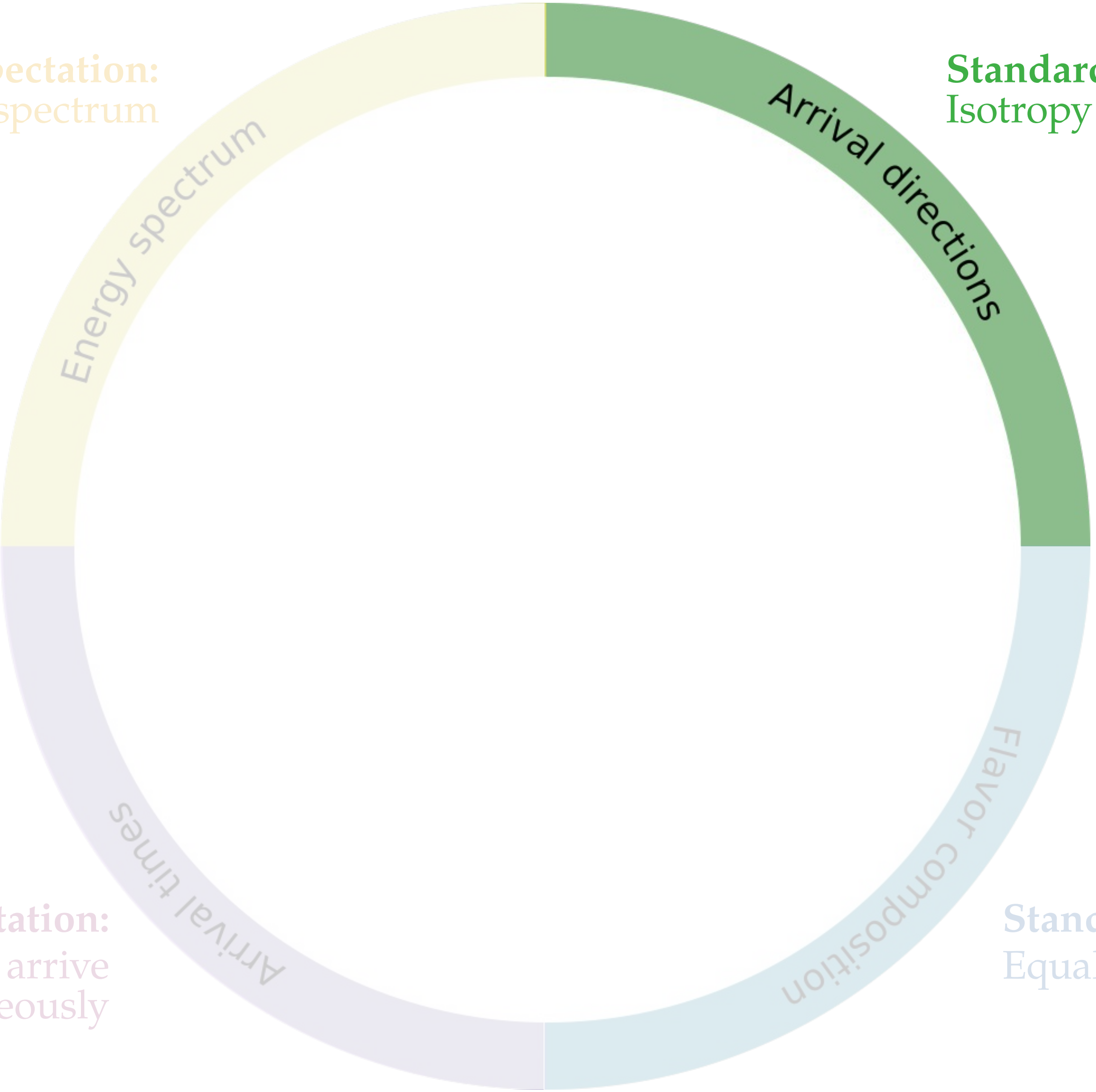


1 IceCube HESE (7.5 yr)
2 IceCube ν_μ (9.5 yr)
All-flavor neutrino flux, $E_\nu^2 \Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^6
10^7
10^8
10^9
10^{10}
Neutrino energy, E_ν [GeV]


Arrival directions
Energy spectrum
Flavor composition
Arrival times

Standard expectation:
Power-law energy spectrum

Standard expectation:
Isotropy (for diffuse flux)

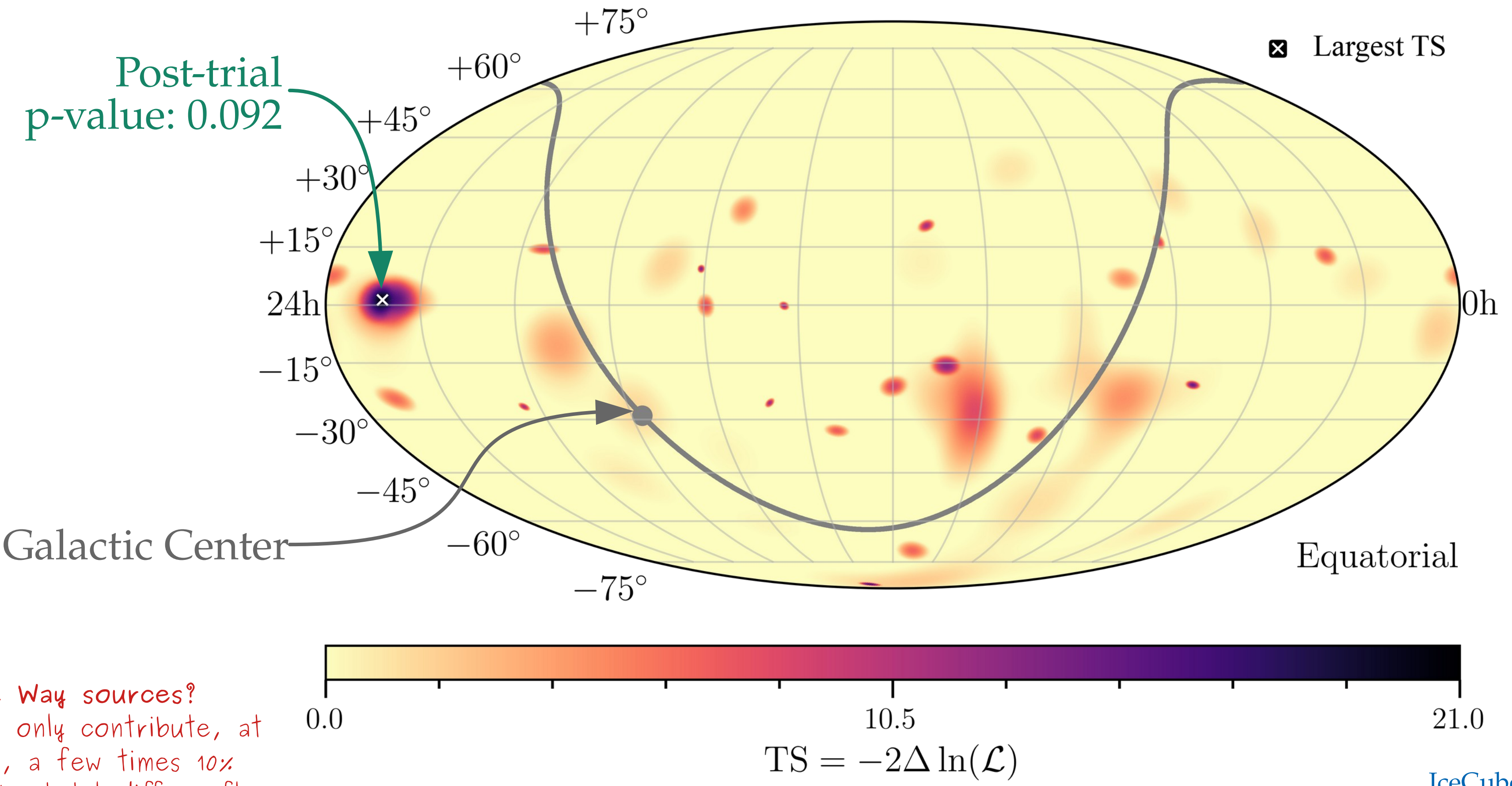


Standard expectation:
ν and γ from transients arrive simultaneously

Standard expectation:
Equal number of ν_e, ν_μ, ν_τ

Arrival directions (7.5 yr)

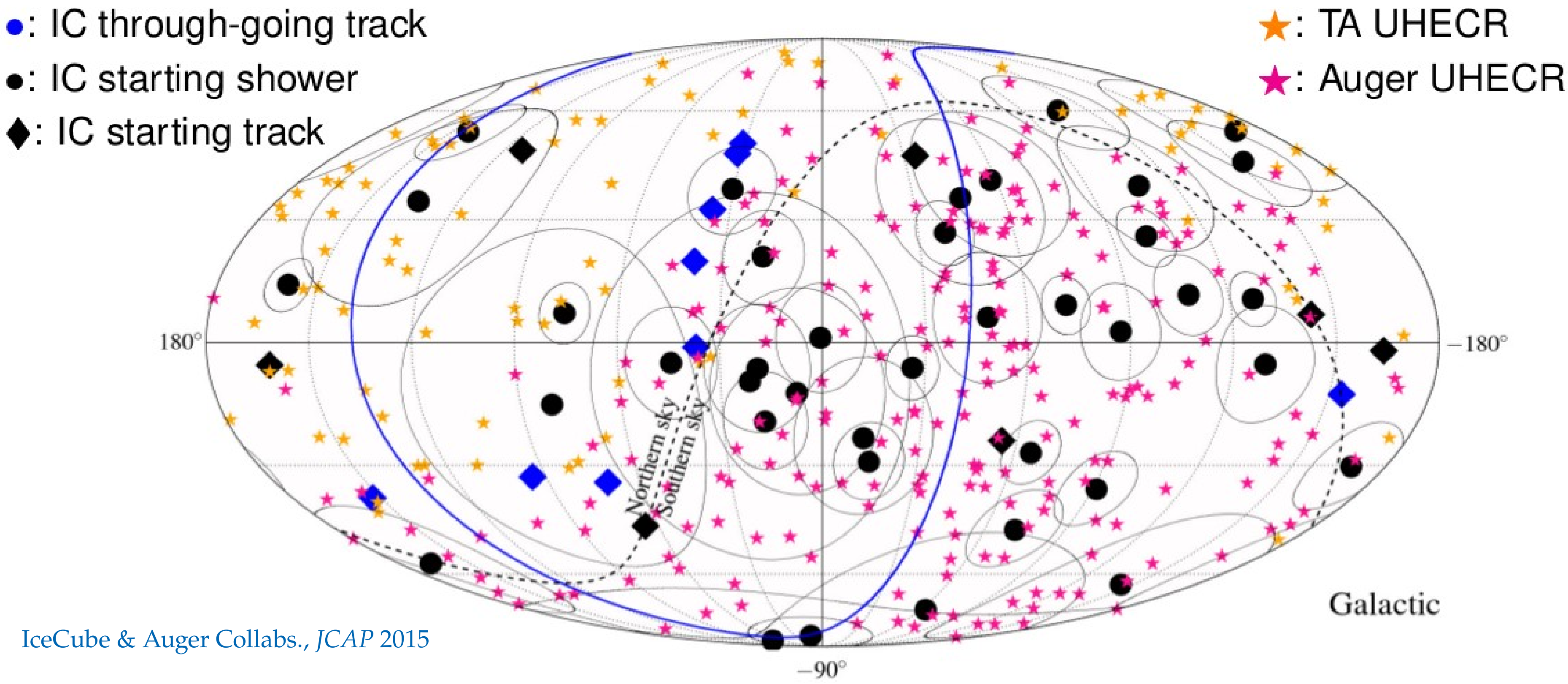
No significant excess in the neutrino sky map:



Milky Way sources?
They only contribute, at most, a few times 10% of the total diffuse flux

IceCube, *PRD* 2021

Neutrino–UHECR angular correlation?



IceCube & Auger Collabs., *JCAP* 2015

No significant correlation with UHECRs (<3.3σ)

A null neutrino-UHECR correlation *makes sense*

UHECRs trace sources within $\lambda_{\mathrm{GZK}} \approx 100$ Mpc

Neutrinos come from anywhere inside the Hubble horizon $D_{\mathrm{H}} \approx 4$ Gpc

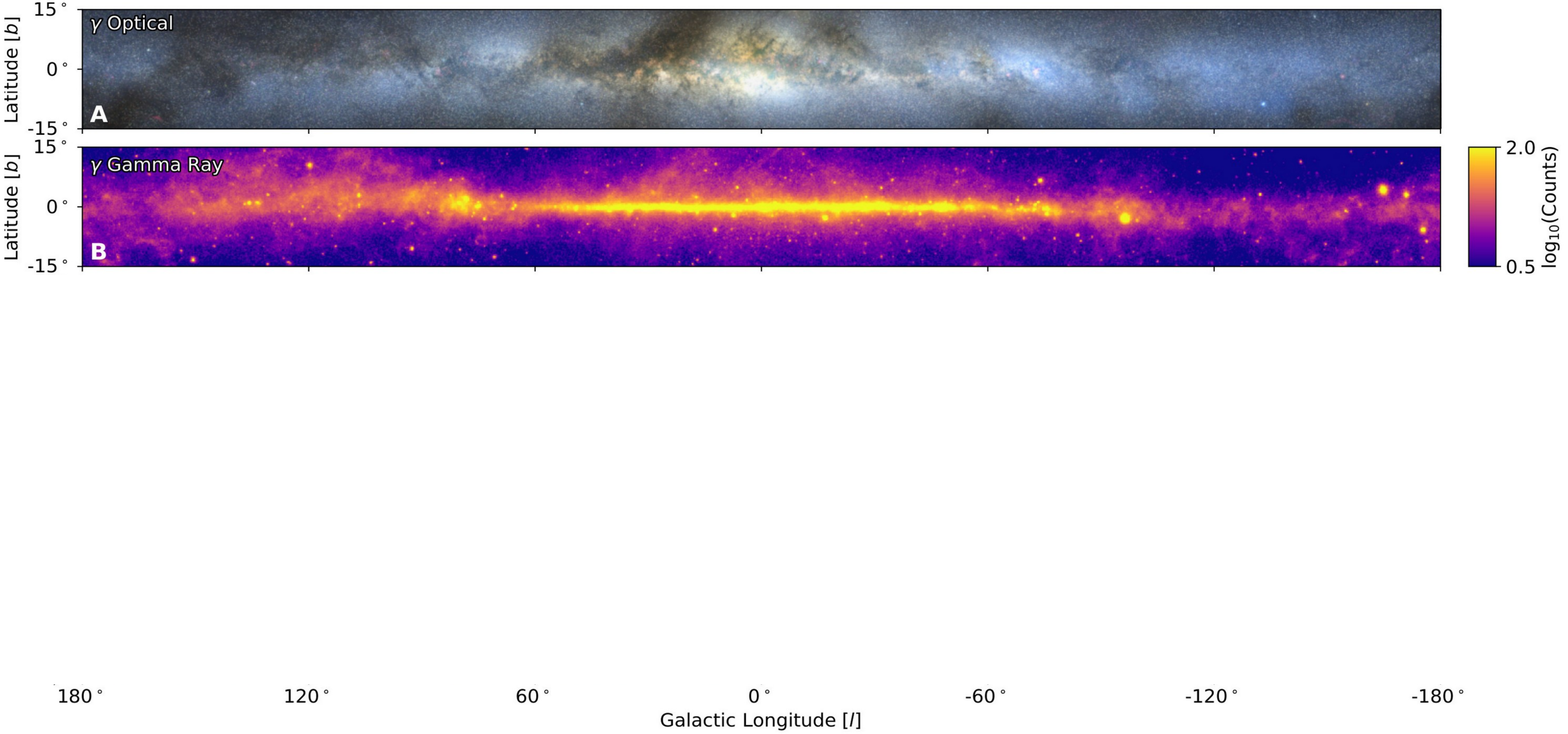
So the maximum possible correlation is $\frac{\lambda_{\mathrm{GZK}}}{D_{\mathrm{H}}} \approx 2.5\%$

Current number of IceCube high-energy starting tracks (HESE): ~100

∴ Expected UHECR correlation with only ~3 neutrinos

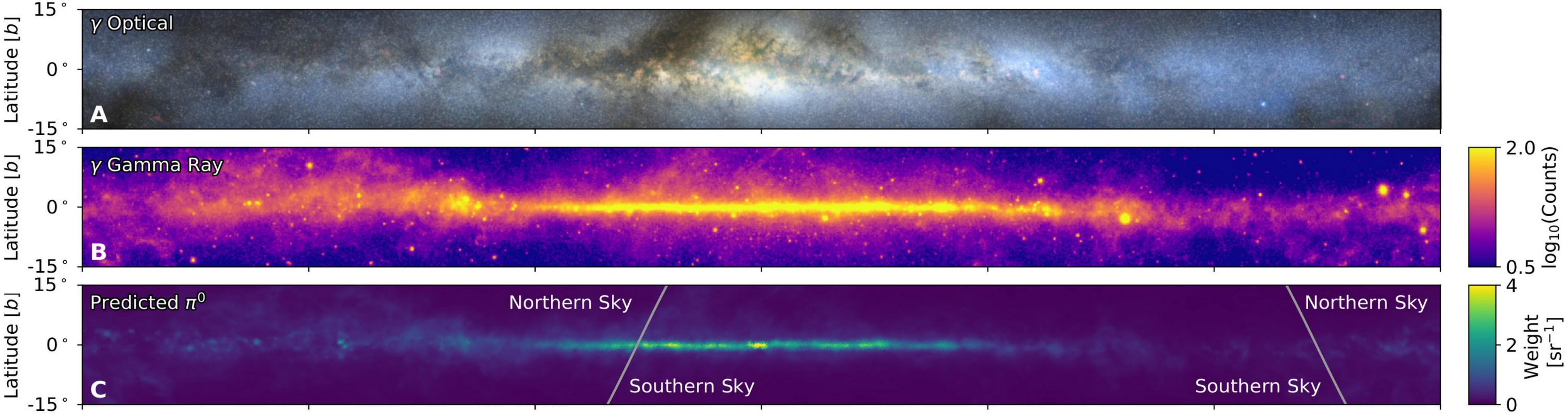
(Also, potential correlation is weakened by magnetic deflection, angular resolution, *etc.*)

High-energy neutrinos from the Galactic Plane



IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane



$$p + p \rightarrow \pi_0 \dashrightarrow \pi_0 \rightarrow \gamma + \gamma$$

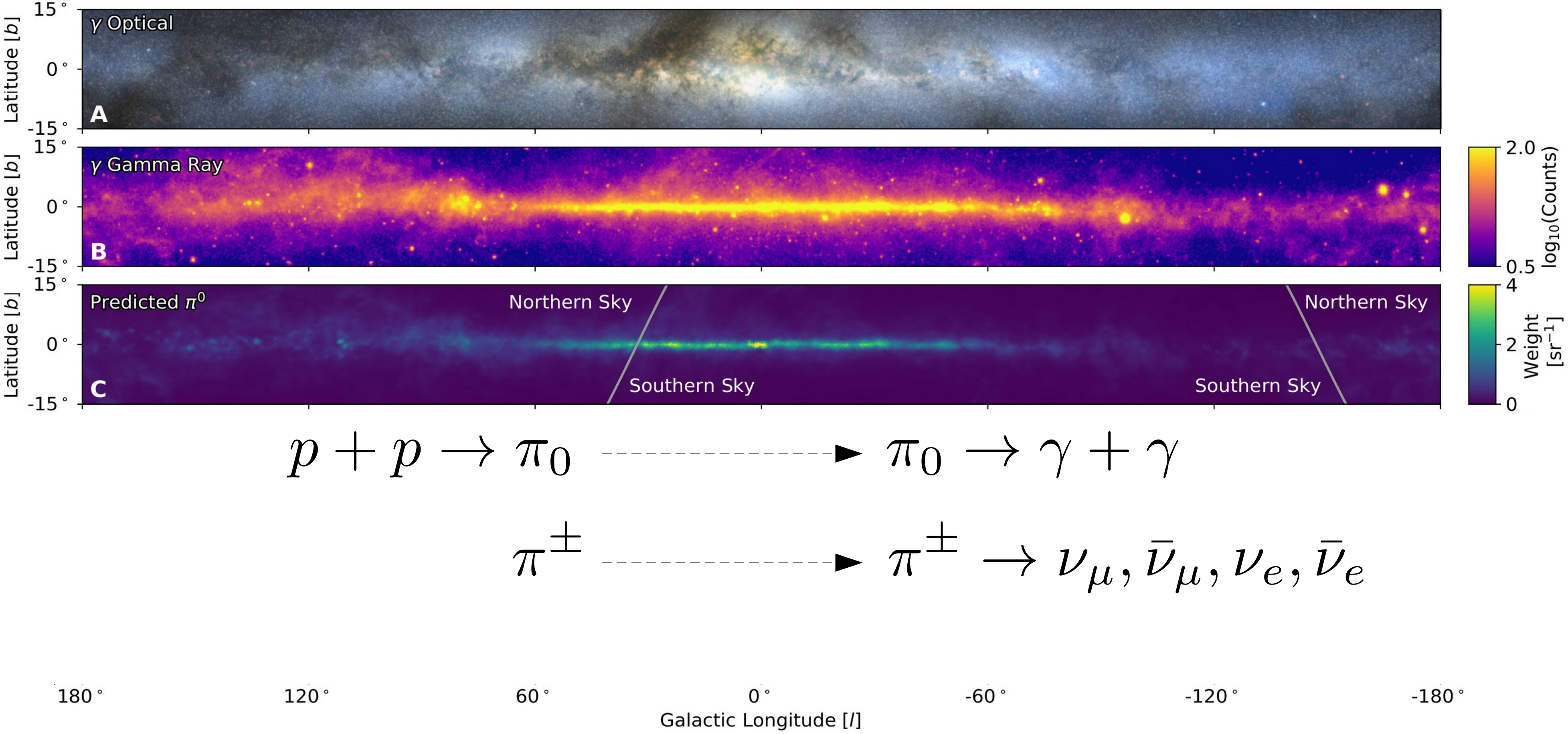
Cosmic rays

Interstellar medium (aligned with GP)

Gamma rays from the GP

IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane

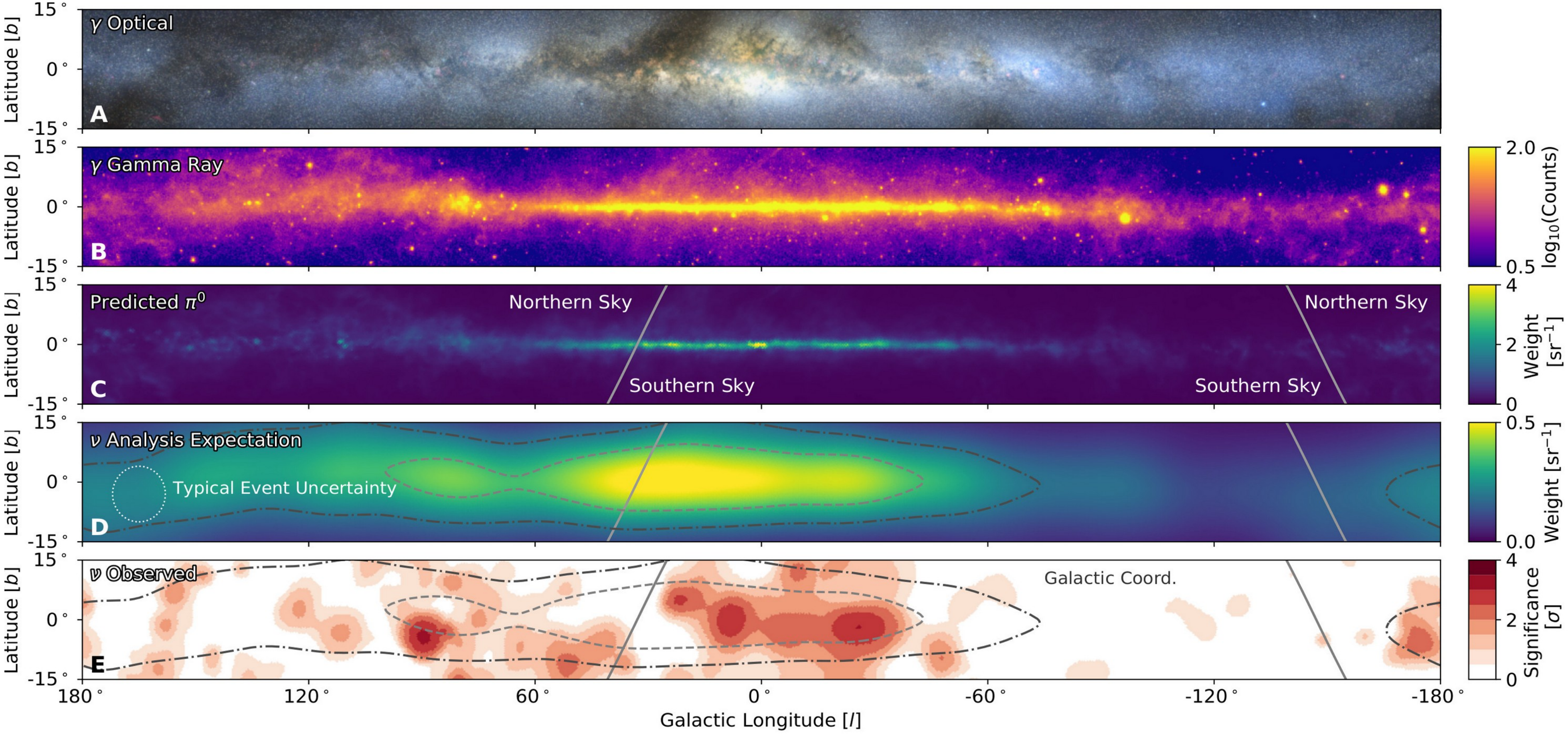


$$p + p \rightarrow \pi_0 \dashrightarrow \pi_0 \rightarrow \gamma + \gamma$$

$$\pi^{\pm} \dashrightarrow \pi^{\pm} \rightarrow \nu_\mu, \bar{\nu}_\mu, \nu_e, \bar{\nu}_e$$

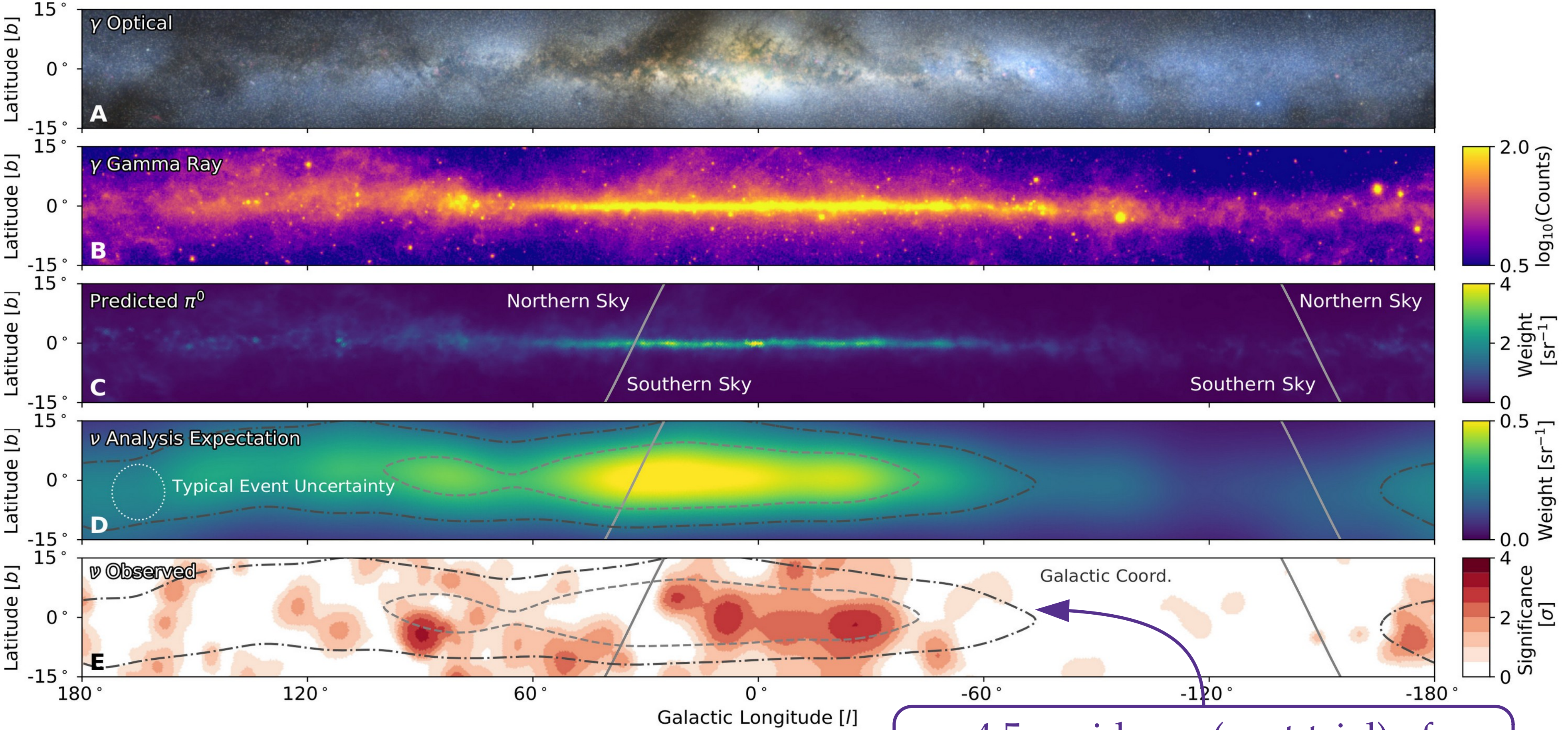
IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane



IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane

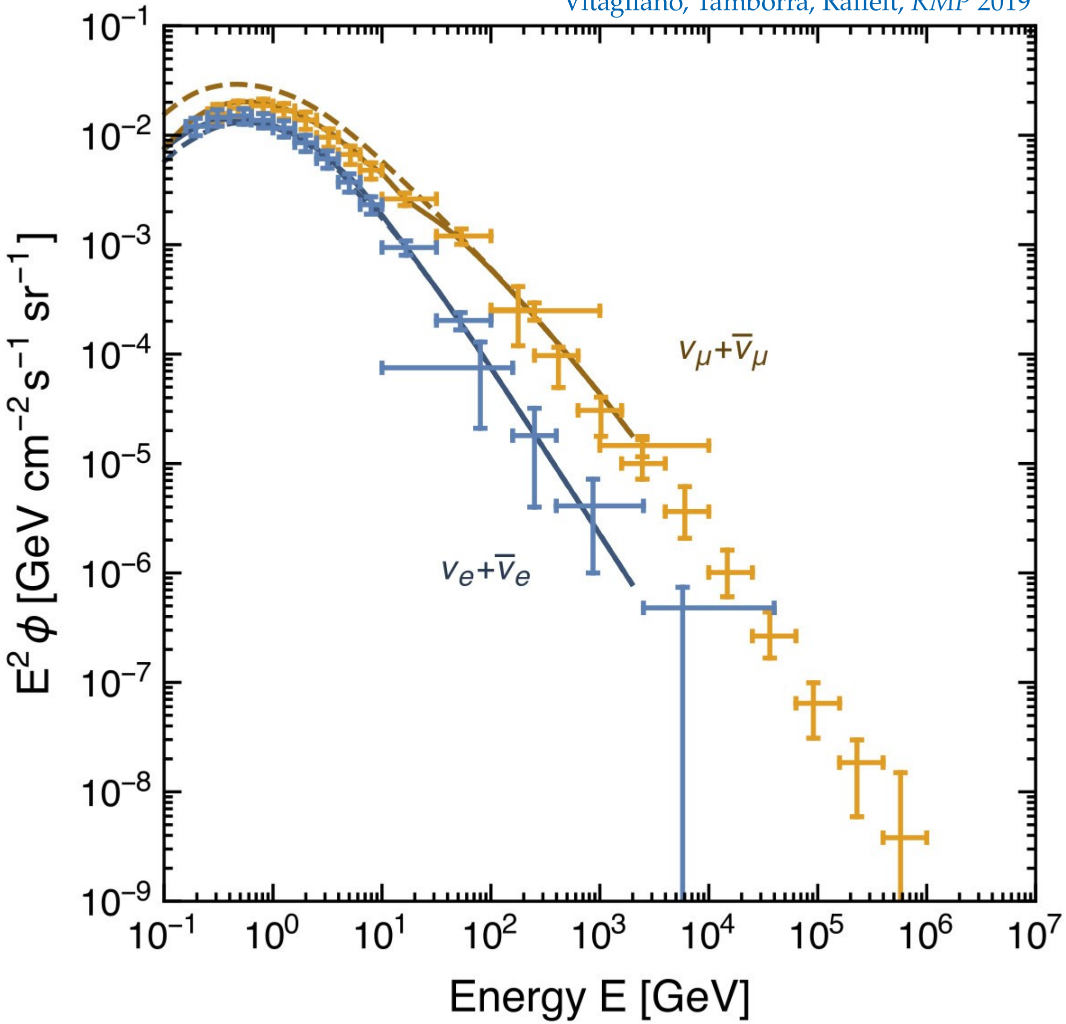


4.5σ evidence (post-trial) of diffuse flux of > TeV ν from the GP

IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane

Vitagliano, Tamborra, Raffelt, *RMP* 2019



See also: Beacom & Candia, *JCAP* 2004

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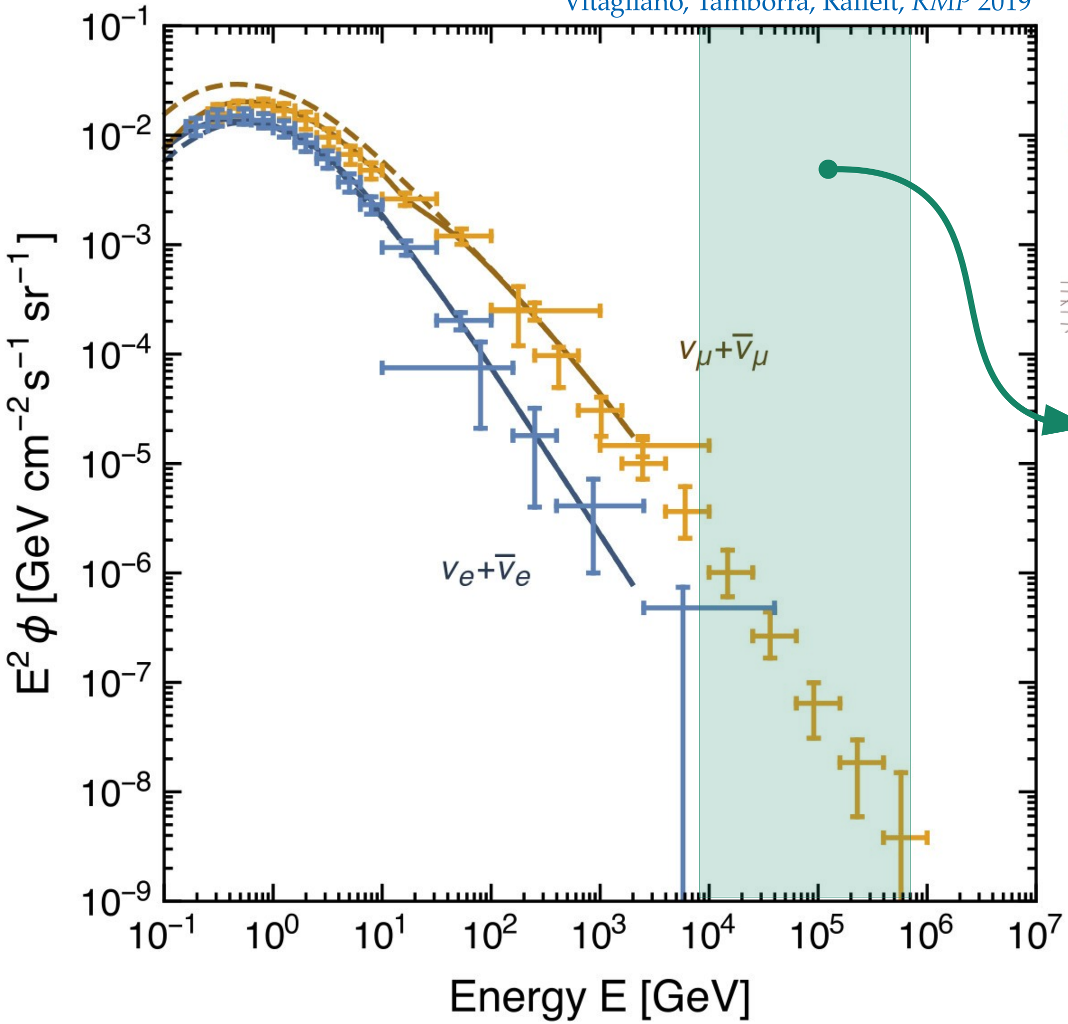
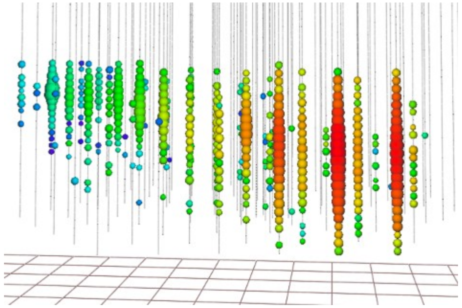


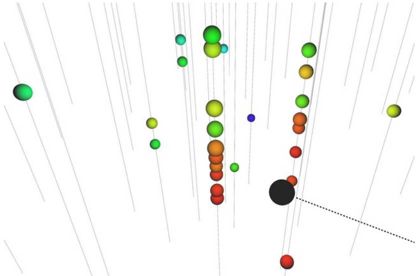


Search for **>10-TeV** astrophysical ν

- Use **muon tracks**
- Pointing accuracy: ~1°
- Atm. bg. is mostly ν_μ
- Self-veto screens for atm. muons to cut ν bg.

See also: Beacom & Candia, *JCAP* 2004

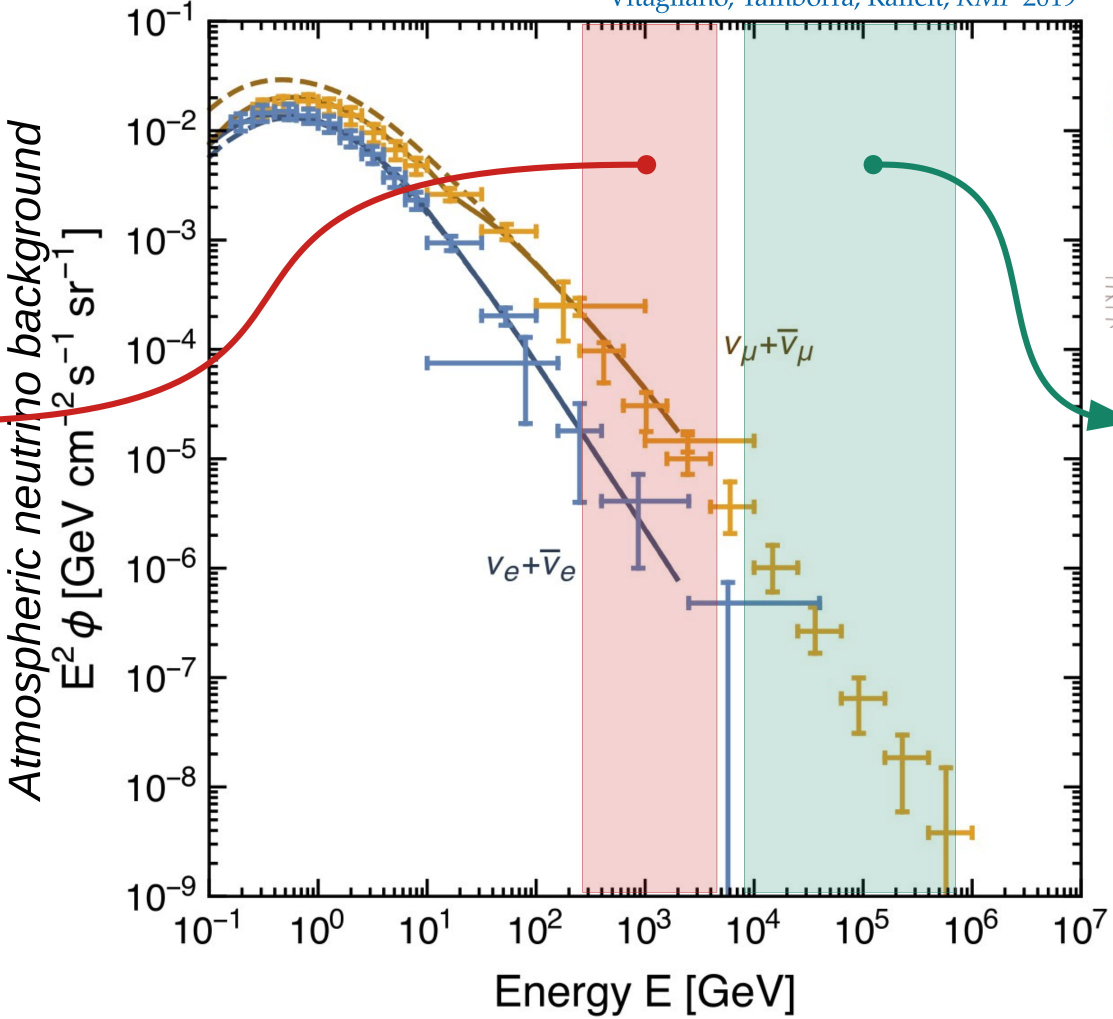
High-energy neutrinos from the Galactic Plane

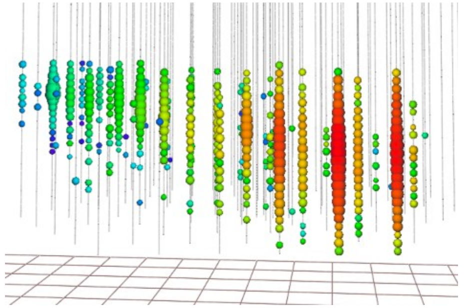


Search for **TeV** astrophysical ν

- But GP ν are TeV
- Use **cascades**
- Atm. ν_e bg. ×10 lower
- Bg.-to-signal: 10^8:1
- *Deep learning retains 20 times more events, ×2 better angular res.*

Vitagliano, Tamborra, Raffelt, *RMP* 2019





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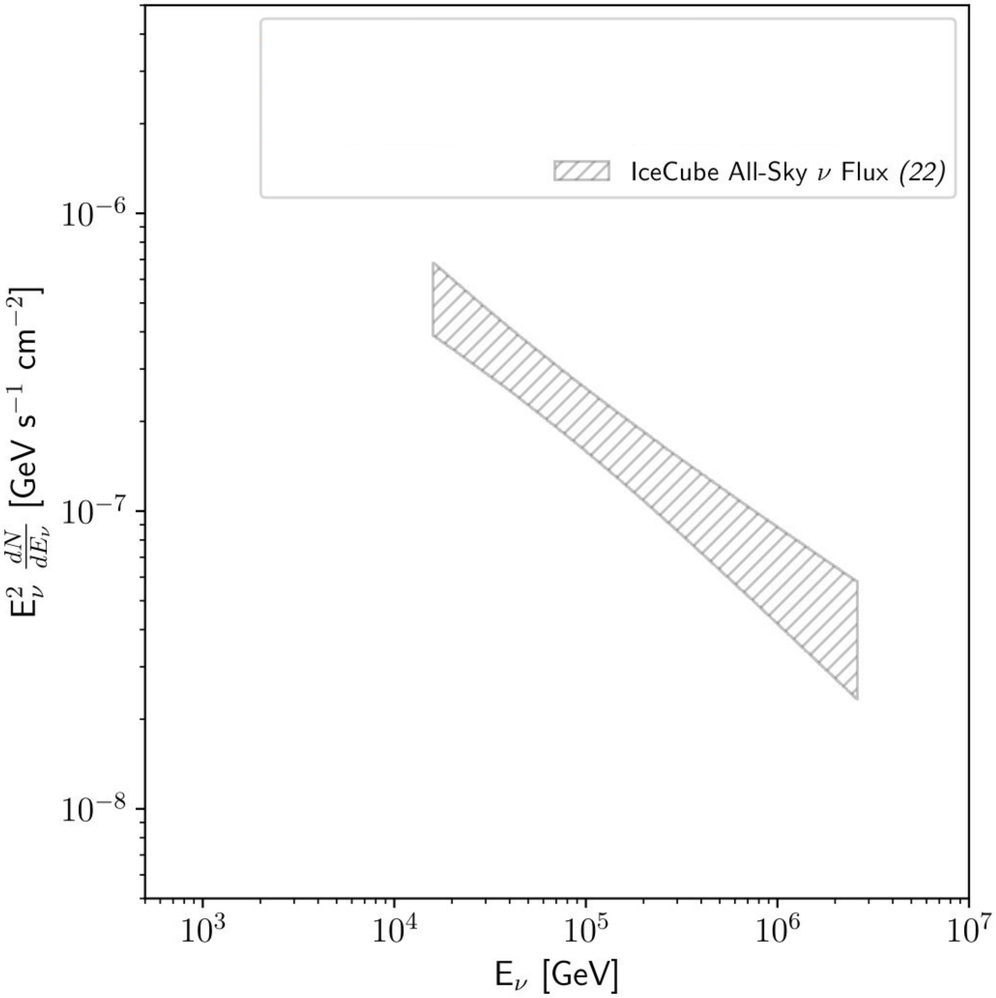
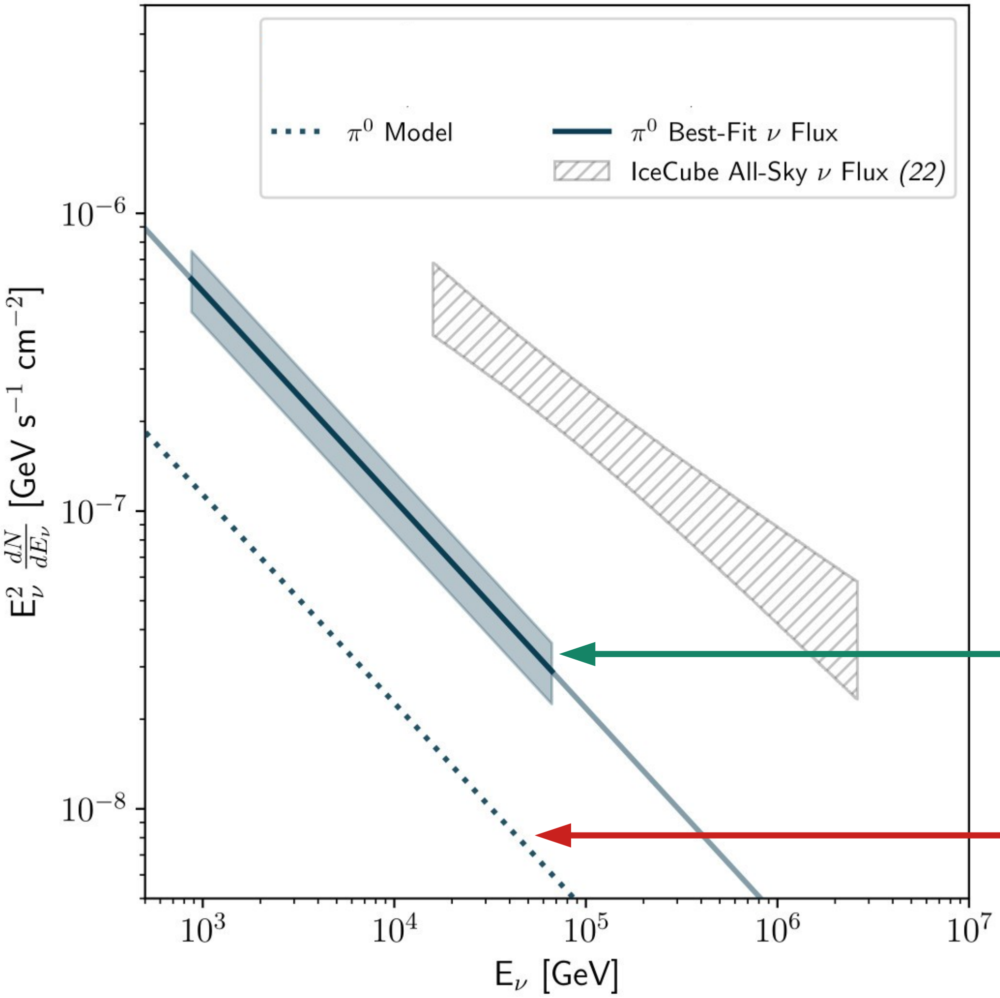


IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane



Three models of Galactic diffuse ν:

π^0: MeV–GeV π^0 template inferred from gamma rays extrapolated to TeV

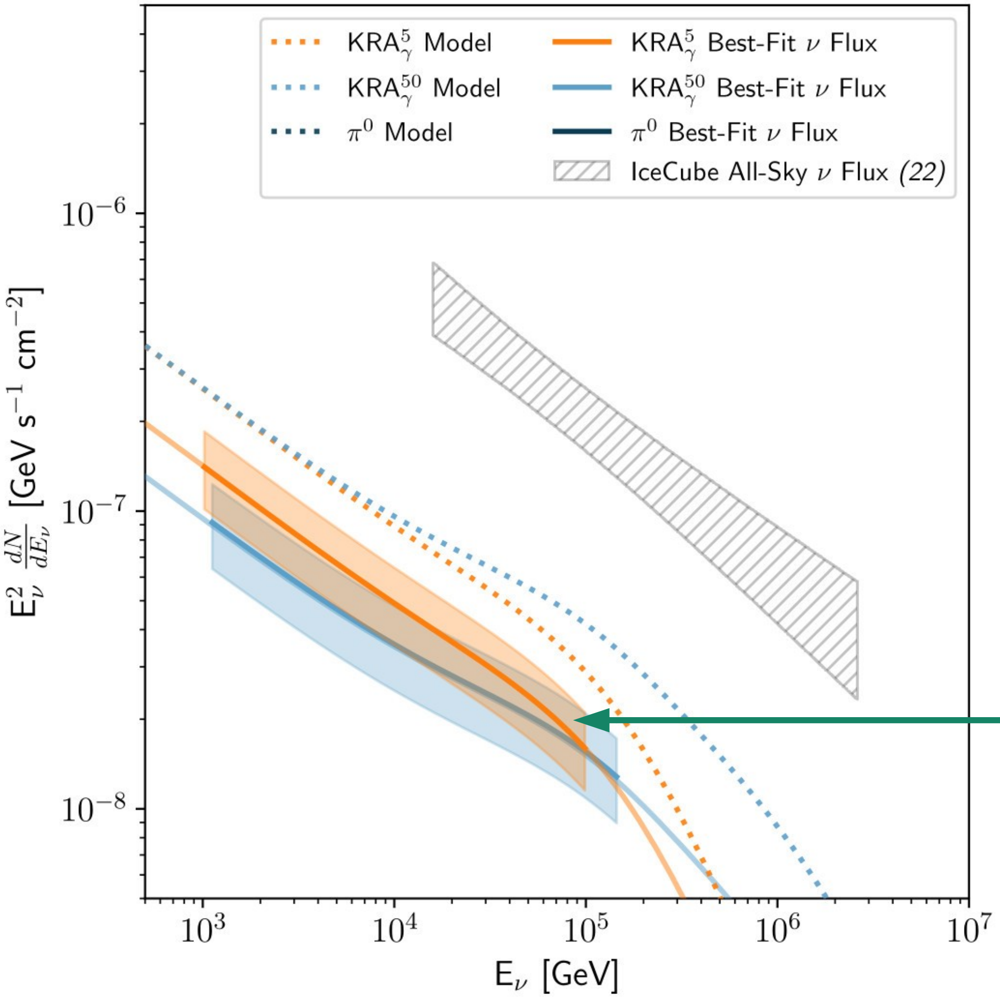
Observed (× 5 model)

Consistent with 100-TeV observations by Tibet Air Shower Array

Model

IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane



Three models of Galactic diffuse ν:

π^0: MeV–GeV π^0 template inferred from gamma rays extrapolated to TeV

KRA$_\gamma^5$: Spectrum varies spatially, harder ν spectrum, cut-off at 5 PeV in CR energy

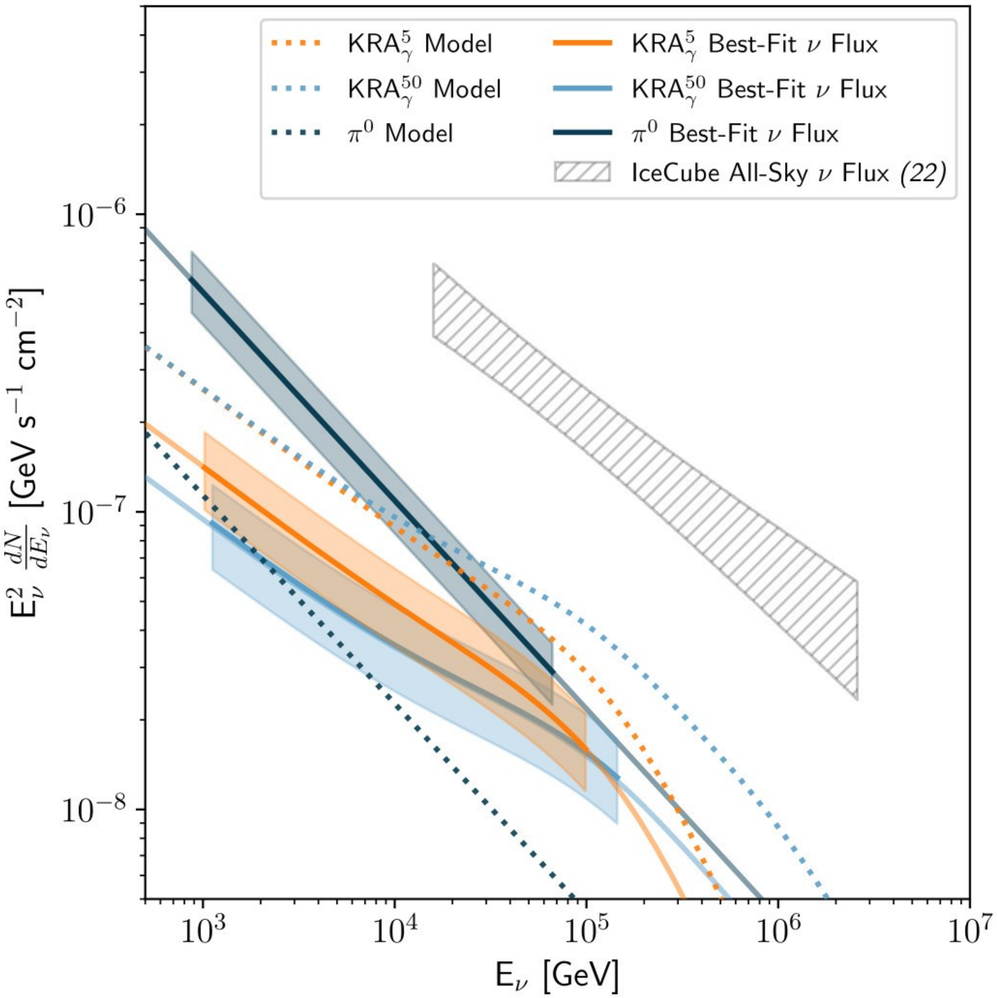
KRA$_\gamma^{50}$: Cut-off at 50 PeV in CR energy

Observed ($\times$ 0.5 model)

Cut-off energy could be different from the 5 and 50 PeV tested

IceCube Collab., *Science* 2023

High-energy neutrinos from the Galactic Plane



IceCube Collab., *Science* 2023

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KRA$_\gamma^5$: Spectrum varies spatially, harder ν spectrum, cut-off at 5 PeV in CR energy

KRA$_\gamma^{50}$: Cut-off at 50 PeV in CR energy

None of the models matched data

(caveat: there are relatively simple models)

No Galactic ν source identified

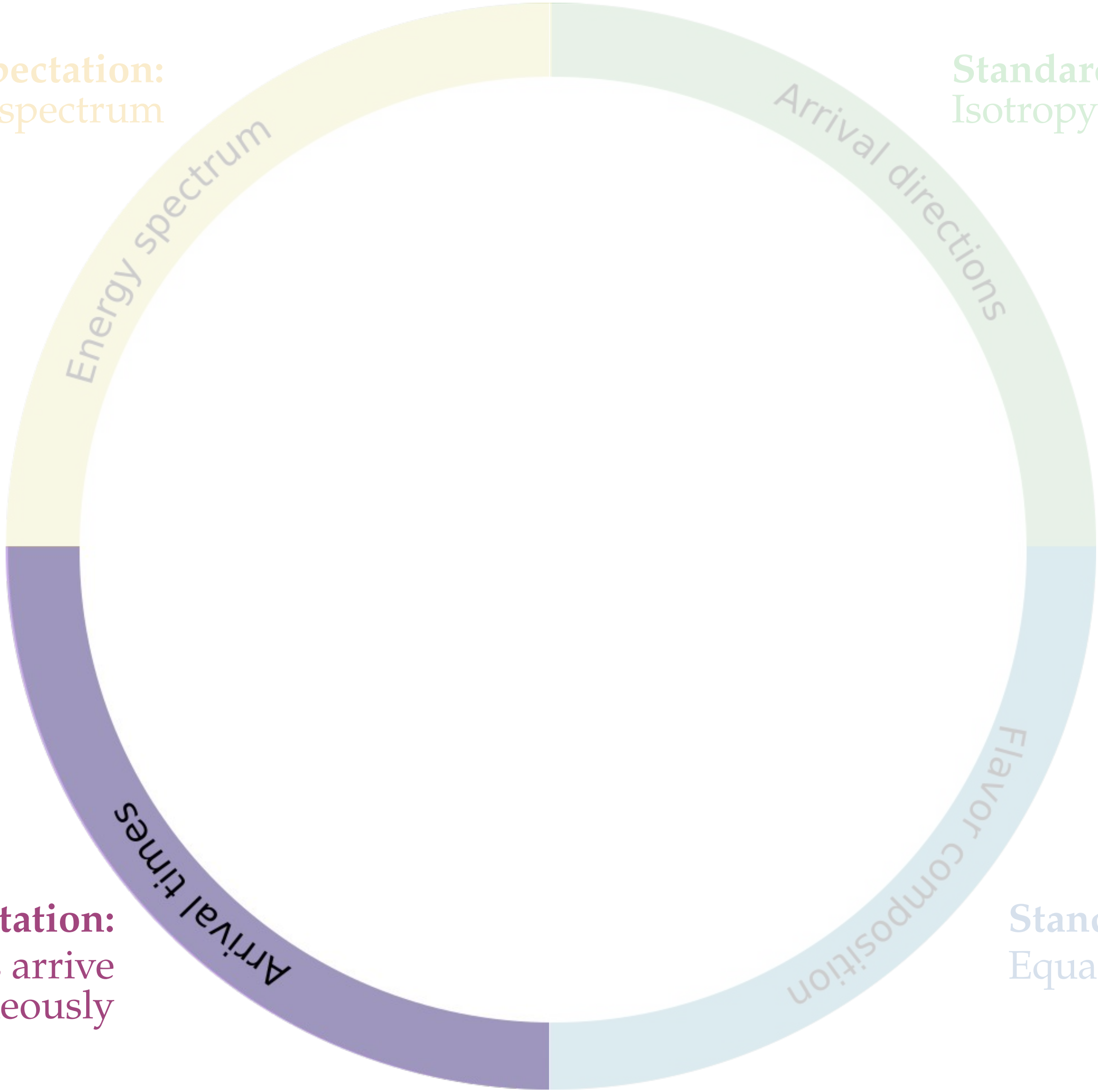
(likely diffuse + source: Fang & Murase, 2307.02905)

GP flux is 6–13% of all-sky at 30 TeV


Energy spectrum
Arrival directions
Flavor composition
Arrival times

Standard expectation:
Power-law energy spectrum

Standard expectation:
Isotropy (for diffuse flux)



Standard expectation:
ν and γ from transients arrive simultaneously

Standard expectation:
Equal number of ν_e, ν_μ, ν_τ

Bright in gamma rays, bright in high-energy neutrinos

Energy in neutrinos $\propto$ energy in gamma rays

$$\int_0^\infty \mathrm{d}E_\nu E_\nu F_\nu(E_\nu) = \frac{1}{8}\left[1 - \left(1 - \langle x_{p\to\pi}\rangle\right)^{\tau_{p\gamma}}\right] \frac{f_p}{f_e} \int_{1\ \mathrm{keV}}^{10\ \mathrm{MeV}} \mathrm{d}E_\gamma E_\gamma F_\gamma(E_\gamma)$$

E. Waxman & J. Bahcall, *PRL* 1997
D. Guetta *et al.*, *Astropart. Phys.* 2004

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Fraction of total *p* energy given to pions

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Fraction of total *p* energy given to pions

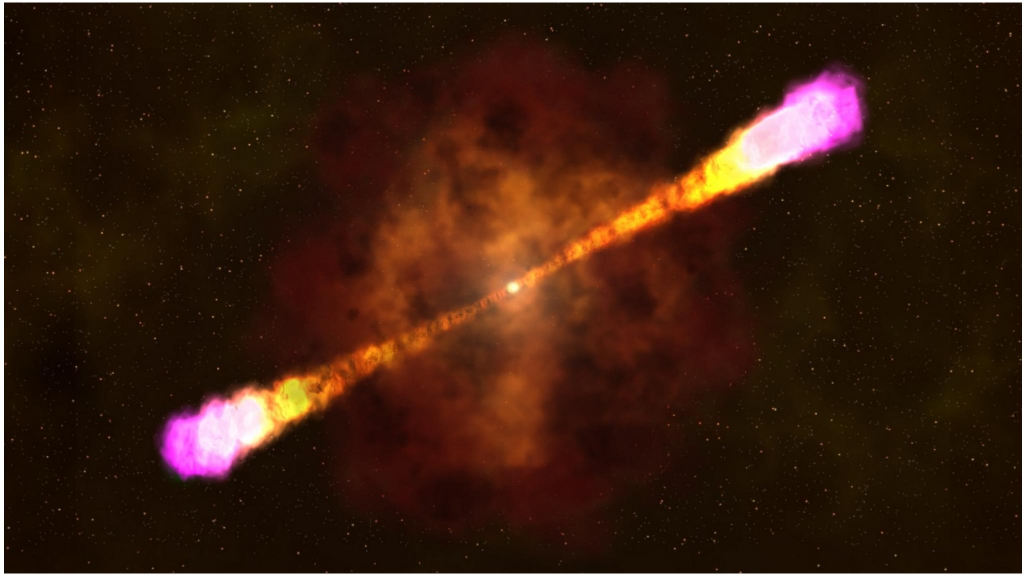
Baryonic loading

Optical depth to *p*γ: $\tau_{p\gamma} = \left(\frac{L_\gamma^{\mathrm{iso}}}{10^{52}\mathrm{ergs}^{-1}}\right)\left(\frac{0.01}{t_{\mathrm{v}}}\right)\left(\frac{300}{\Gamma}\right)^4\left(\frac{\mathrm{MeV}}{\epsilon_{\gamma,\mathrm{break}}}\right)$

E. Waxman & J. Bahcall, *PRL* 1997
D. Guetta *et al.*, *Astropart. Phys.* 2004

Gamma-ray bursts and blazars – *not* dominant

Gamma-ray bursts



Blazars

Gamma-ray bursts and blazars – *not* dominant

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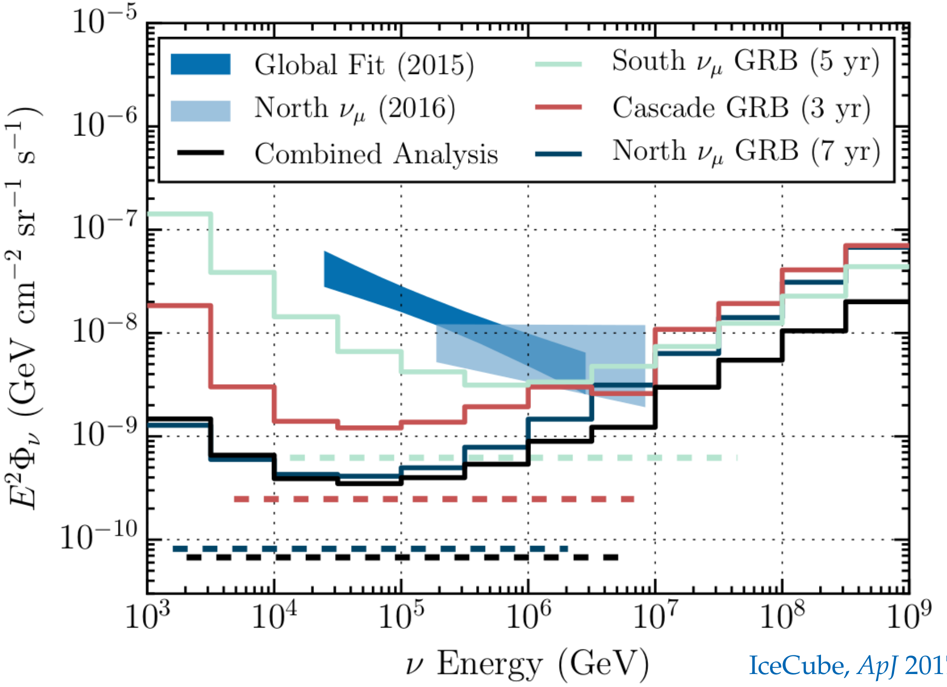
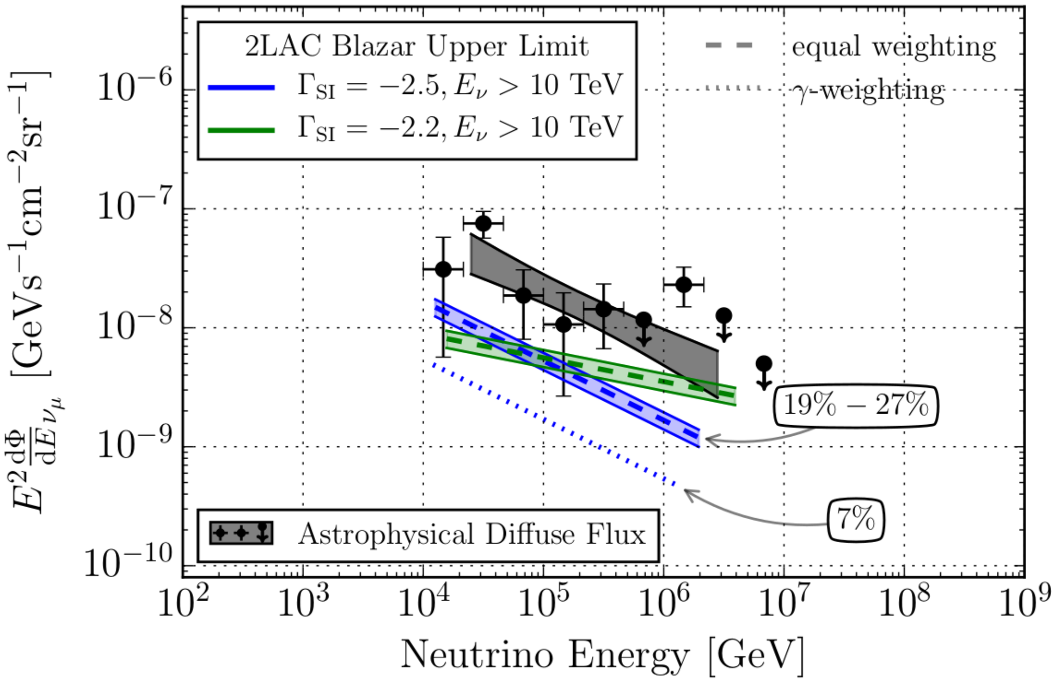


IceCube, *ApJ* 2017

1172 GRBs inspected, no correlation found

< 1% contribution to diffuse flux

Blazars



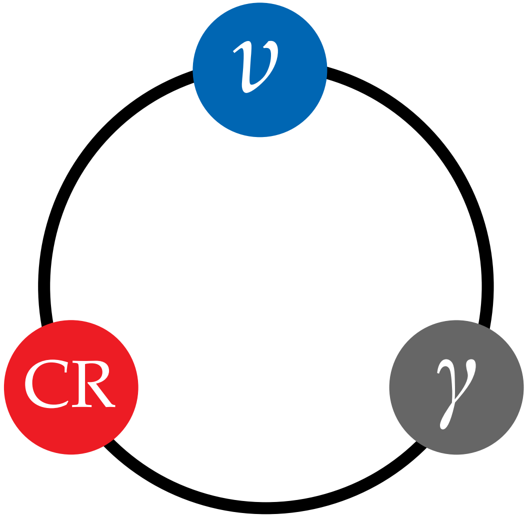
862 blazars inspected, no correlation found

< 27% contribution to diffuse flux

Bright in gamma rays, bright in high-energy neutrinos (?)

Energy in neutrinos ∝ energy in gamma rays

Waxman & Bahcall, *PRL* 1997



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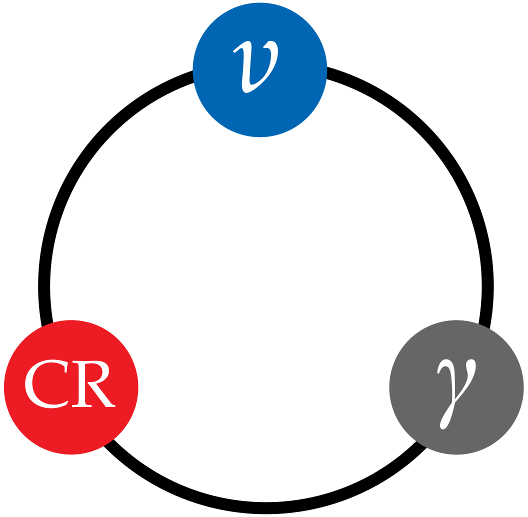
Energy in neutrinos $\propto$ energy in gamma rays

Waxman & Bahcall, *PRL* 1997

Fudge factors:
Source properties (*e.g.*, baryonic loading)
Particle effects (*e.g.*, ν-producing channels)



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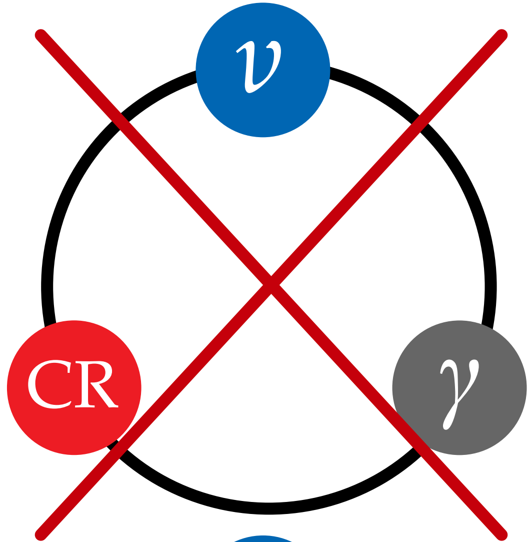
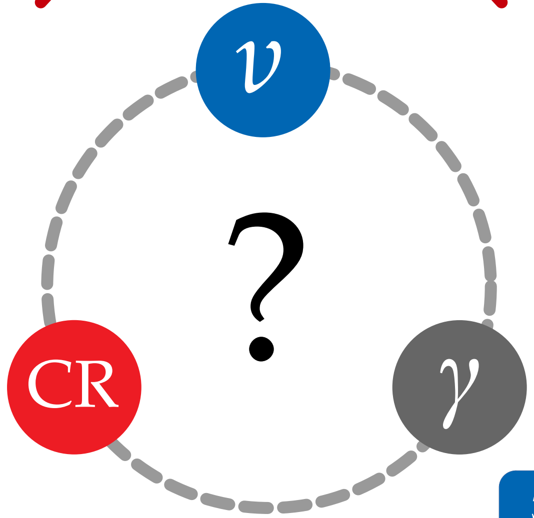


But the correlation between ν and γ may be more nuanced:

Gao, Pohl, Winter, *ApJ* 2017



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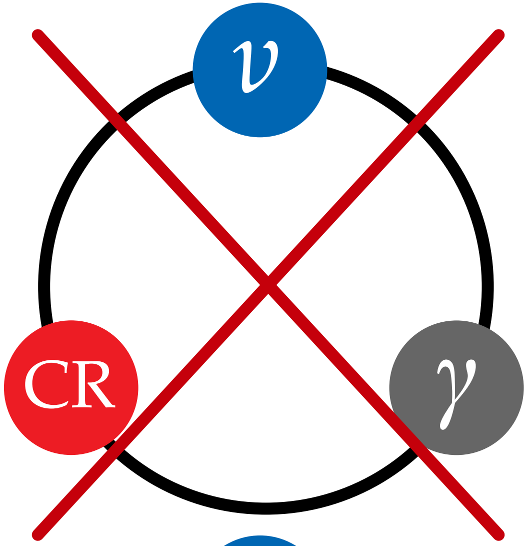


But the correlation between ν and γ may be more nuanced:

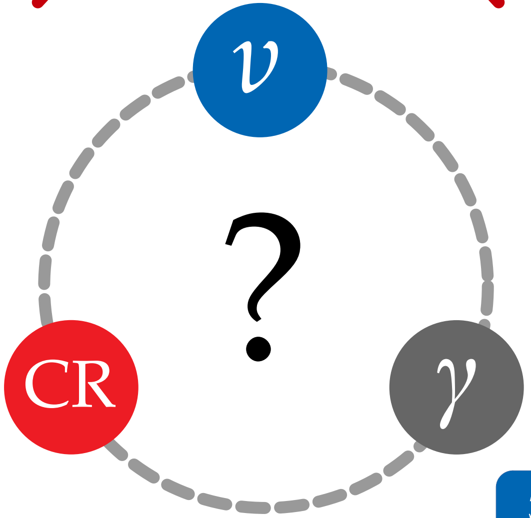
Gao, Pohl, Winter, *ApJ* 2017

Sources that make neutrinos via $p\gamma$
may be opaque to 1–100 MeV gamma rays

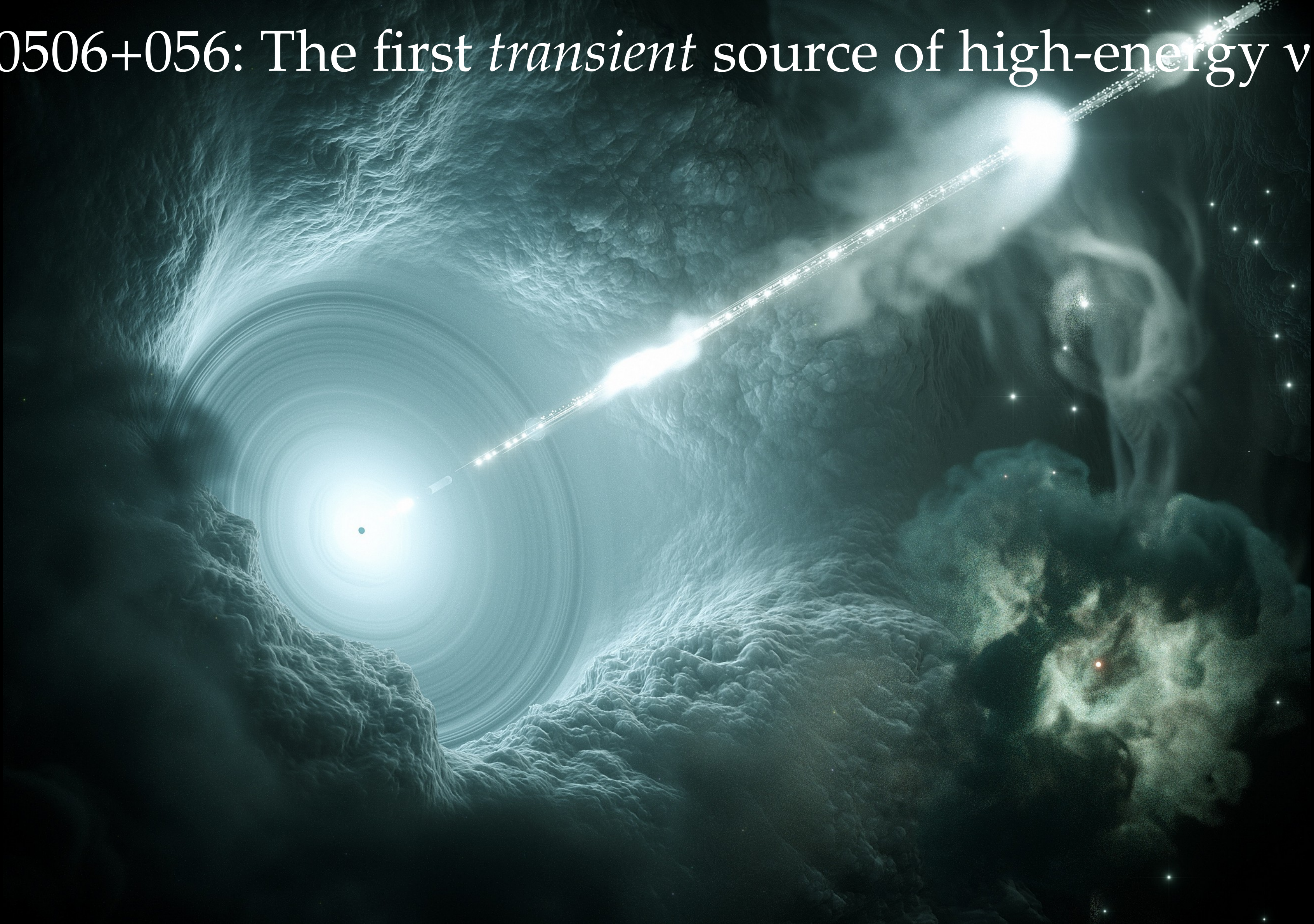
Murase, Guetta, Ahlers, *PRL* 2016

Modeling of $p\gamma$ interactions & nuclear cascading
in the sources is complex and uncertain

Morejon, Fedynitch, Boncioli, Winter, *JCAP* 2019
Boncioli, Fedynitch, Winter, *Sci. Rep.* 2017



TXS 0506+056: The first *transient* source of high-energy ν

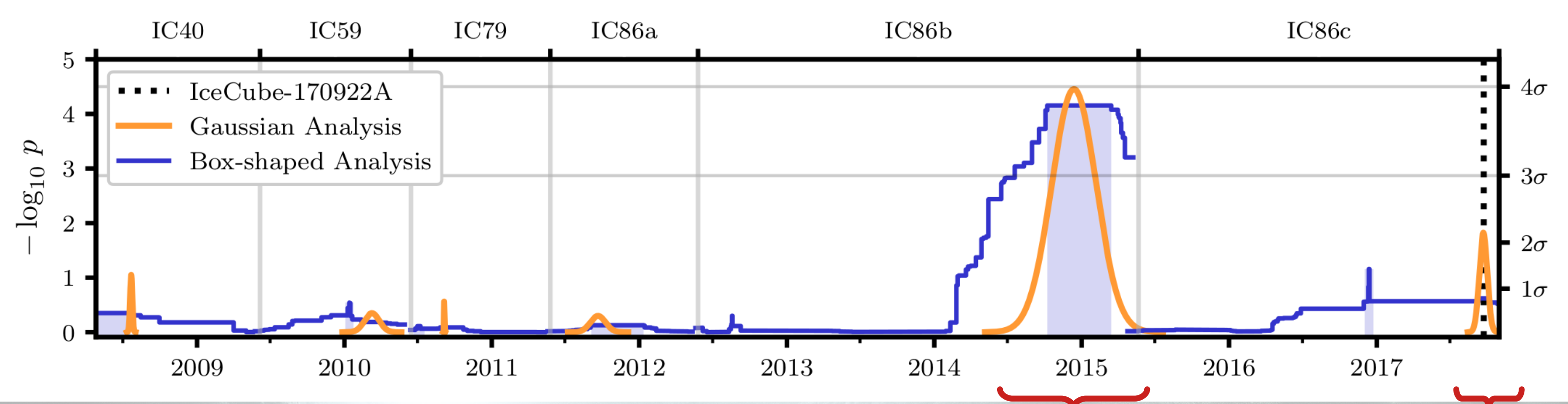


DESY

TXS 0506+056: The first *transient* source of high-energy ν

Blazar TXS 0506+056:

IceCube, *Science* 2018



After re-analysis (2101.09836), significance dropped from $p=7\times10^{-5}$ to $p=8\times10^{-3}$

2014–2015: 13±5 ν flare, no X-ray flare
3.5σ significance of correlation (post-trial)

2017: one 290-TeV ν + X-ray flare
1.4σ significance of correlation

Combined (pre-trial): 4.1σ

Tidal disruption events

Solar-mass star disrupted by SMBH (>10^5 $M_\odot$)

~50% of the debris bound to the SMBH, creates a flare (occasionally a jet)

NASA

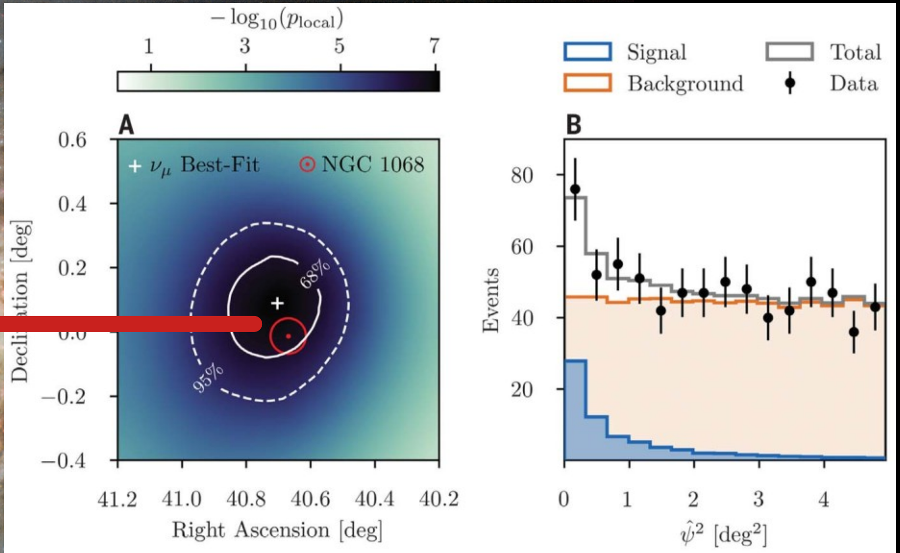
NGC1068: The first *steady-state* source of high-energy ν

Active galactic nucleus

Brightest type-2 Seyfert

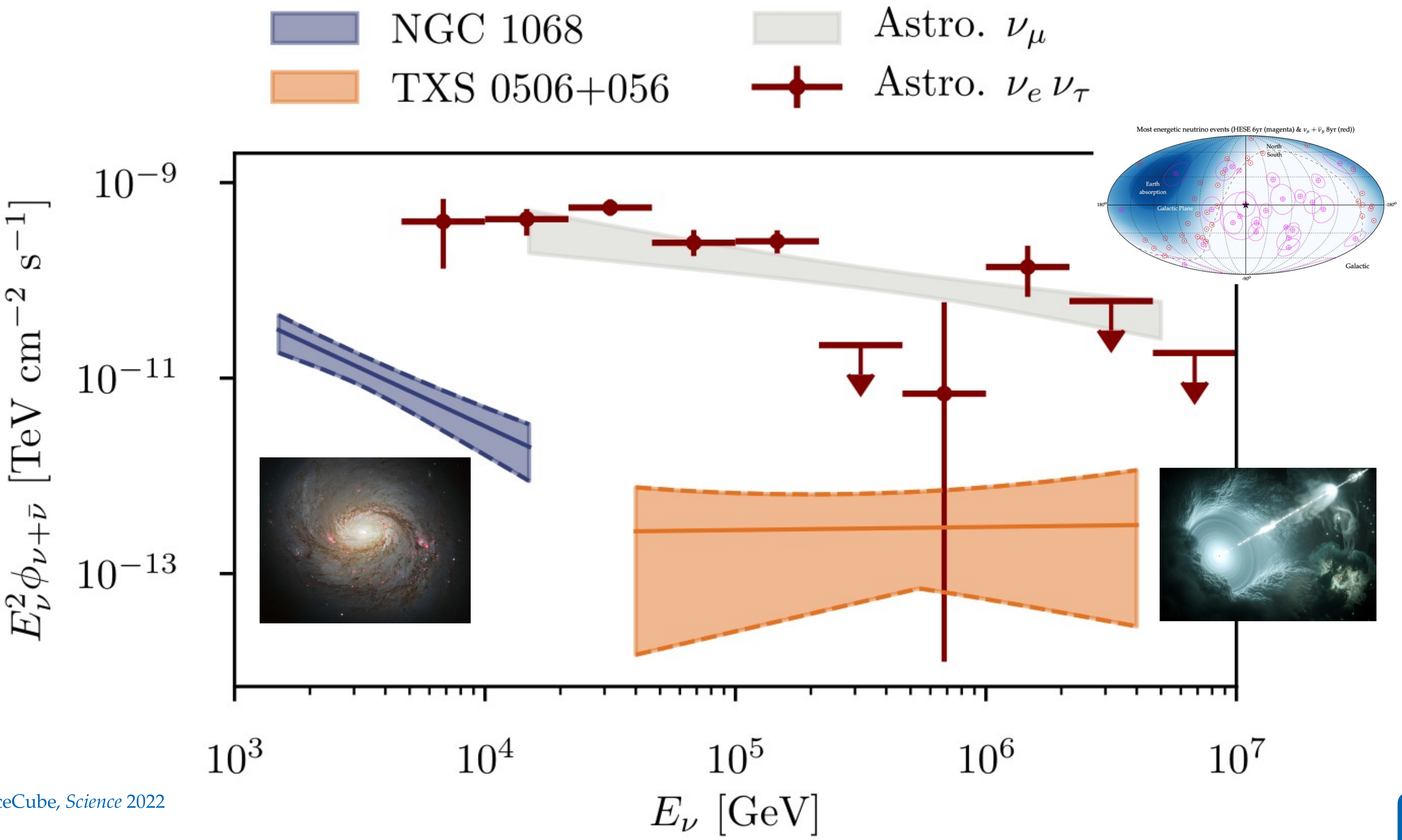
79^{+22}_{-20} ν of TeV energy

Significance: 4.2σ (global)



IceCube Collab., *Science* 378, 538 (2022)

Hubble Space Telescope,
NASA, ESA & A. van der Hoeven

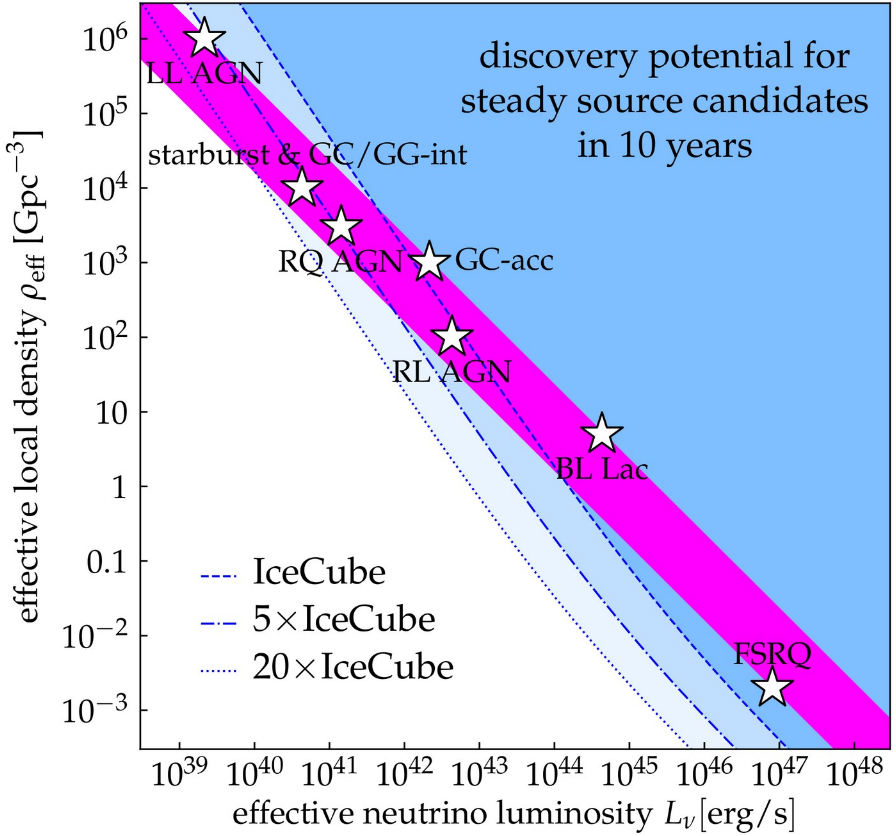


IceCube, *Science* 2022

Source discovery potential: today and in the future

Accounts for the observed diffuse ν flux (lower/upper edge: rapid/no redshift evolution)

Closest source with $E^2\phi_{\nu_\mu+\bar{\nu}_\mu} = 10^{-9}\ \mathrm{GeV cm^{-2} s^{-1}}$

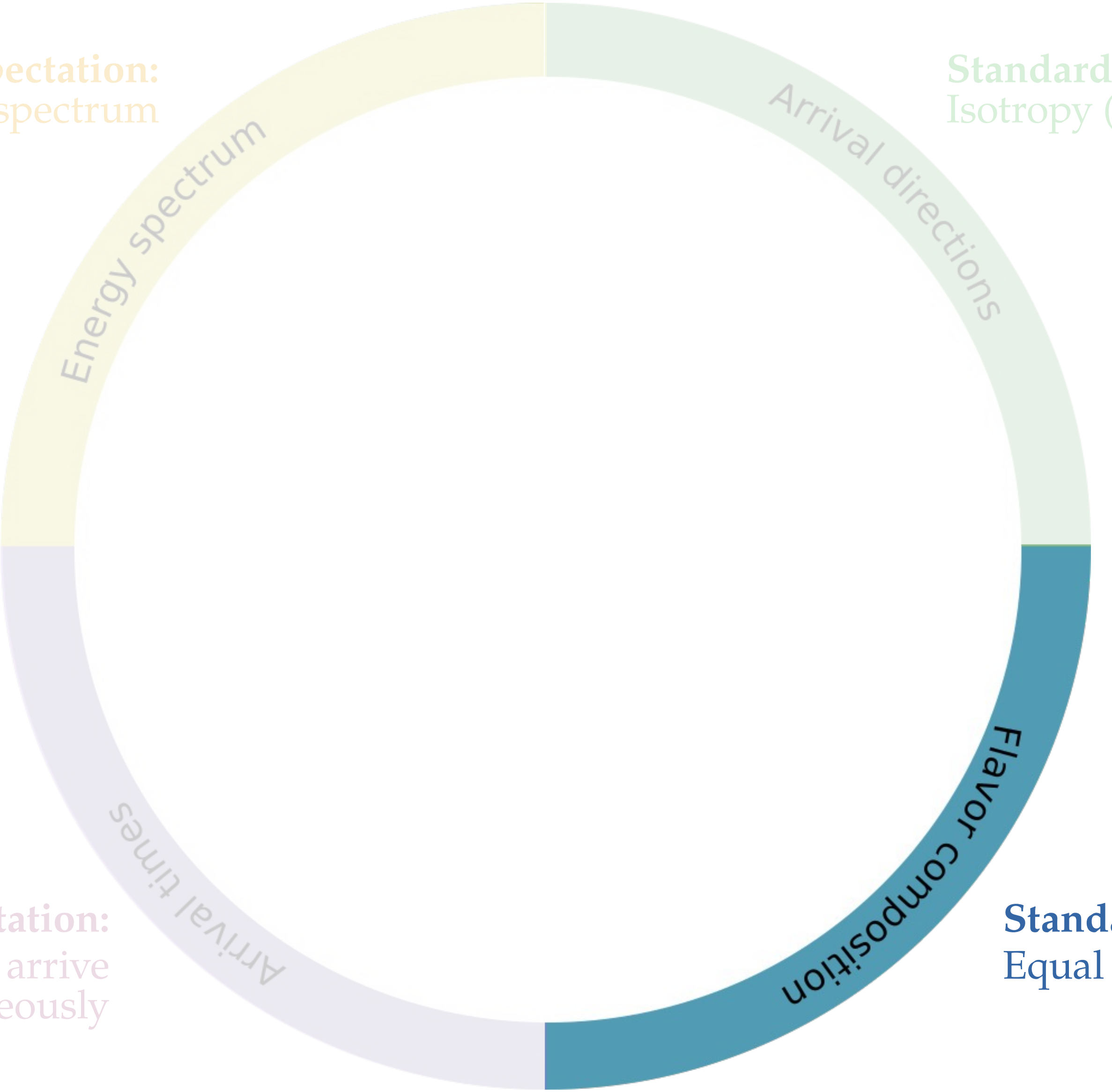


Ackermann, **MB** *et al.*, Astro2020 Survey (1903.04333) – See also: Silvestri & Barwick, *PRD* 2010; Murase & Waxman, *PRD* 2016


Energy spectrum
Arrival directions
Flavor composition
Arrival times

Standard expectation:
Power-law energy spectrum

Standard expectation:
Isotropy (for diffuse flux)



Standard expectation:
ν and γ from transients arrive simultaneously

Standard expectation:
Equal number of ν_e, ν_μ, ν_τ

Flavor-transition probability

- In matrix form: $$\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix} = \begin{pmatrix}U^*_{e1} & U^*_{e2} & U^*_{e3}\\ U^*_{\mu1} & U^*_{\mu2} & U^*_{\mu3}\\ U^*_{\tau1} & U^*_{\tau2} & U^*_{\tau3}\end{pmatrix}\begin{pmatrix}\nu_1\\ \nu_2\\ \nu_3\end{pmatrix}$$

- Pontecorvo-Maki-Nakagawa-Sakata matrix ($c_{ij} = \cos\theta_{ij}$, $s_{ij} = \sin\theta_{ij}$):

$$U = \underbrace{\begin{pmatrix}1 & 0 & 0\\ 0 & c_{23} & s_{23}\\ 0 & -s_{23} & c_{23}\end{pmatrix}}_{\text{Atmospheric}} \underbrace{\begin{pmatrix}c_{13} & 0 & s_{13}e^{-i\delta}\\ 0 & 1 & 0\\ -s_{13}e^{i\delta} & 0 & c_{13}\end{pmatrix}}_{\text{Cross mixing}} \underbrace{\begin{pmatrix}c_{12} & s_{12} & 0\\ -s_{12} & c_{12} & 0\\ 0 & 0 & 1\end{pmatrix}}_{\text{Solar}} \underbrace{\begin{pmatrix}e^{i\alpha_1/2} & 0 & 0\\ 0 & e^{i\alpha_2/2} & 0\\ 0 & 0 & 1\end{pmatrix}}_{\text{Majorana CP phases}}$$

- Probability for $\nu_\alpha \to \nu_\beta$: $$P_{\nu_\alpha\to\nu_\beta} = \delta_{\alpha\beta} - 4\sum_{i>j}\mathrm{Re}(U^*_{\alpha i}U_{\beta i}U_{\alpha j}U^*_{\beta j})\sin^2\left(\Delta m^2_{ij}\frac{L}{4E}\right) + 2\sum_{i>j}\mathrm{Im}(U^*_{\alpha i}U_{\beta i}U_{\alpha j}U^*_{\beta j})\sin\left(\Delta m^2_{ij}\frac{L}{2E}\right)$$

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$\theta_{23} \approx 48°$
$\theta_{13} \approx 9°$
$\theta_{12} \approx 34°$
$\delta \approx 222°$

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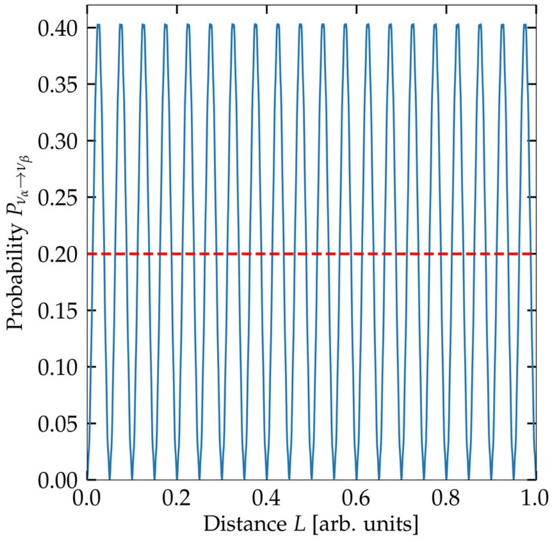
… But high-energy neutrinos oscillate *fast*



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Oscillation length for 1-TeV ν: $2\pi \times 2E/\Delta m^2 \sim 0.1$ pc

~ 8% of the way to Proxima Centauri
≪ Distance to Galactic Center (8 kpc)
≪ Distance to Andromeda (1 Mpc)
≪ Cosmological distances (few Gpc)

We cannot resolve oscillations, so we use instead the average probability:

$$\langle P_{\nu_\alpha \to \nu_\beta} \rangle = \sum_{i=1}^{3} |U_{\alpha i}|^2 \, |U_{\beta i}|^2$$

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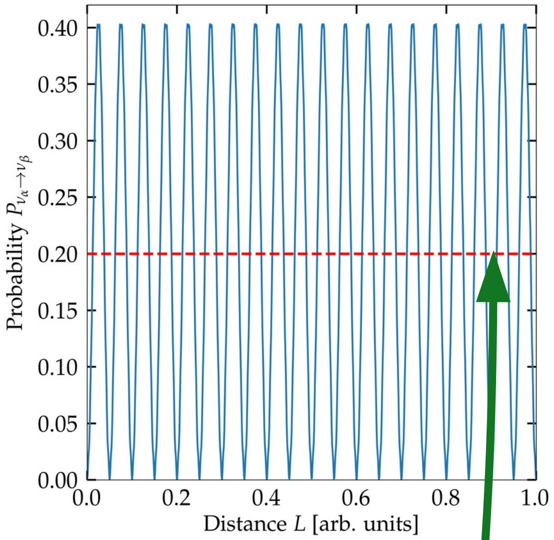

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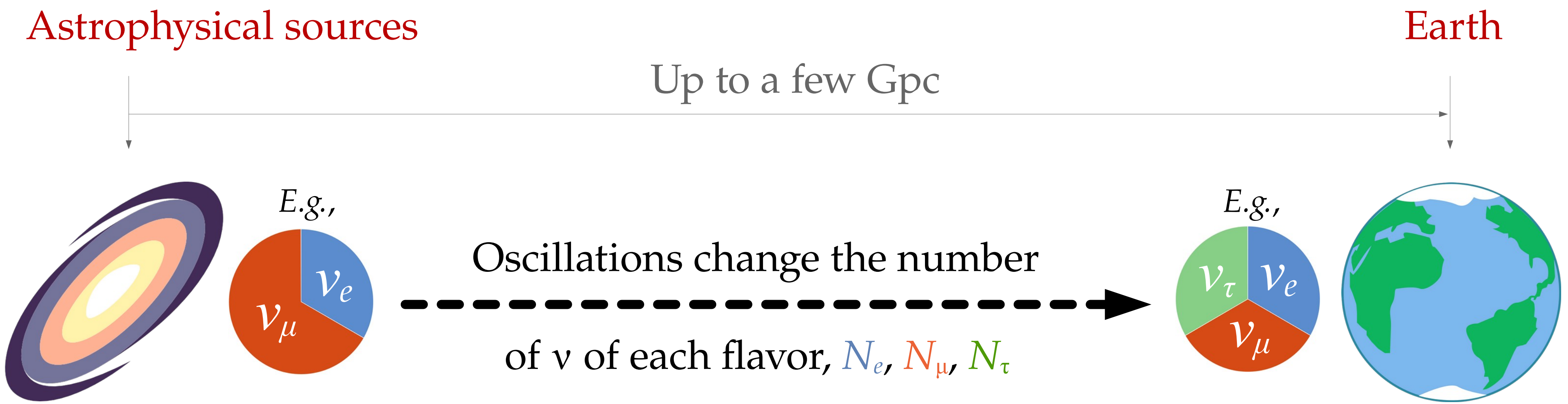
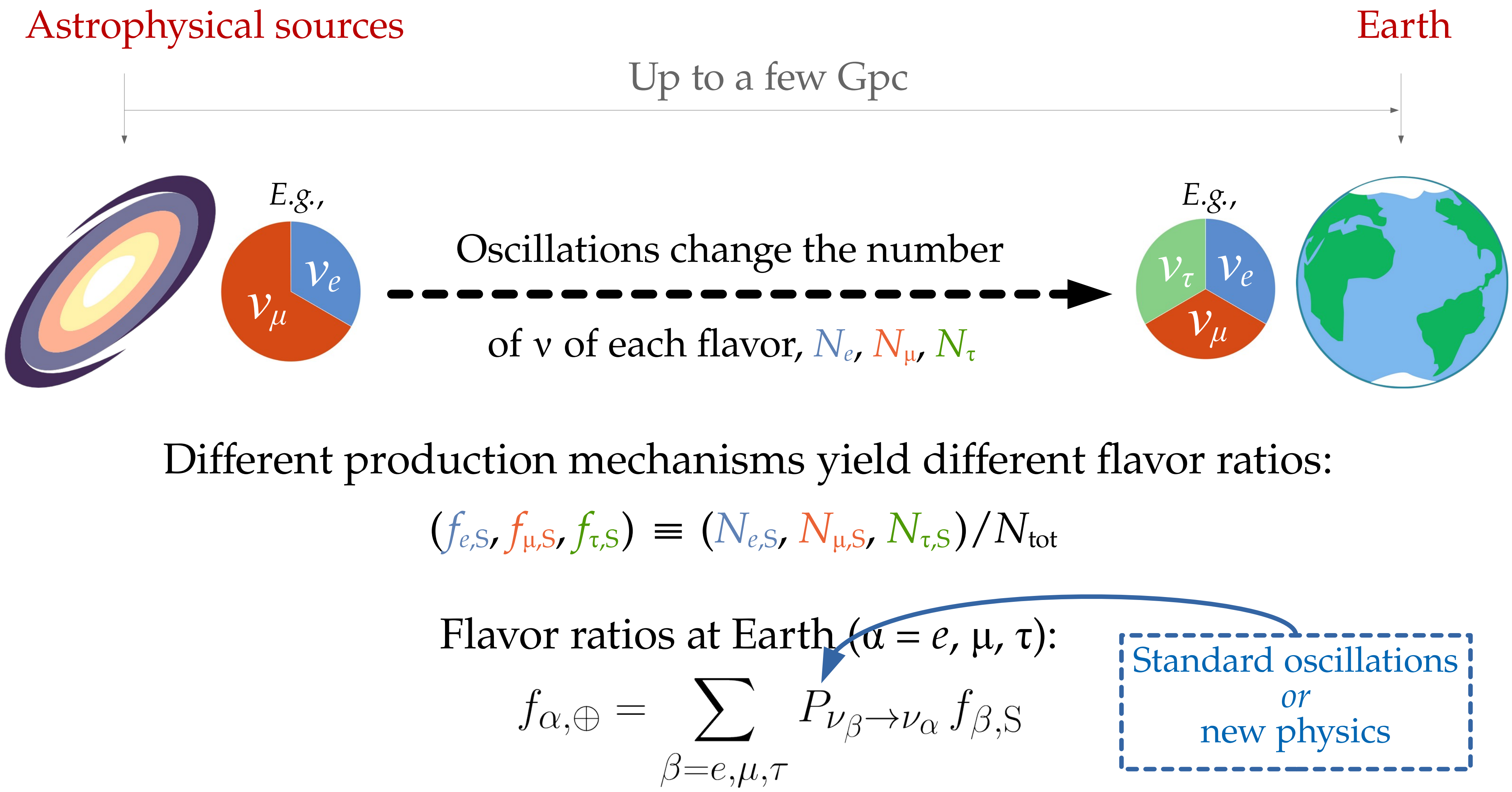


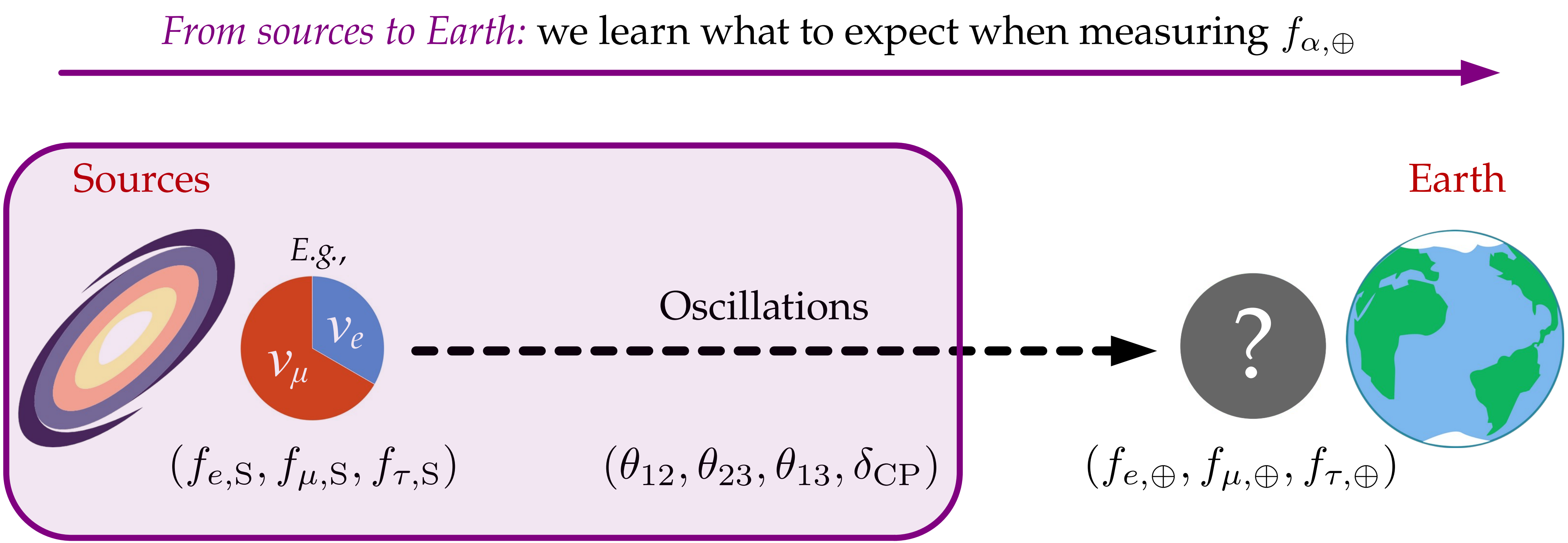
Different production mechanisms yield different flavor ratios:

$$(f_{e,\mathrm{S}}, f_{\mu,\mathrm{S}}, f_{\tau,\mathrm{S}}) \equiv (N_{e,\mathrm{S}}, N_{\mu,\mathrm{S}}, N_{\tau,\mathrm{S}})/N_{\mathrm{tot}}$$

Flavor ratios at Earth ($\alpha = e, \mu, \tau$):

$$f_{\alpha,\oplus} = \sum_{\beta=e,\mu,\tau} P_{\nu_\beta \to \nu_\alpha} f_{\beta,\mathrm{S}}$$


Astrophysical sources
Earth
Up to a few Gpc
E.g.,
ν_e
ν_μ
Oscillations change the number
of ν of each flavor, N_e, N_μ, N_τ
E.g.,
ν_τ
ν_e
ν_μ
Different production mechanisms yield different flavor ratios:
$(f_{e,S}, f_{\mu,S}, f_{\tau,S}) \equiv (N_{e,S}, N_{\mu,S}, N_{\tau,S})/N_{tot}$
Flavor ratios at Earth ($\alpha = e, \mu, \tau$):
$f_{\alpha,\oplus} = \sum_{\beta=e,\mu,\tau} P_{\nu_\beta \to \nu_\alpha} f_{\beta,S}$
Standard oscillations
or
new physics



From sources to Earth: we learn what to expect when measuring $f_{\alpha,\oplus}$
Sources
E.g.,
ν_e
ν_μ
Oscillations
Earth
?
$(f_{e,\mathrm{S}}, f_{\mu,\mathrm{S}}, f_{\tau,\mathrm{S}})$
$(\theta_{12}, \theta_{23}, \theta_{13}, \delta_{\mathrm{CP}})$
$(f_{e,\oplus}, f_{\mu,\oplus}, f_{\tau,\oplus})$

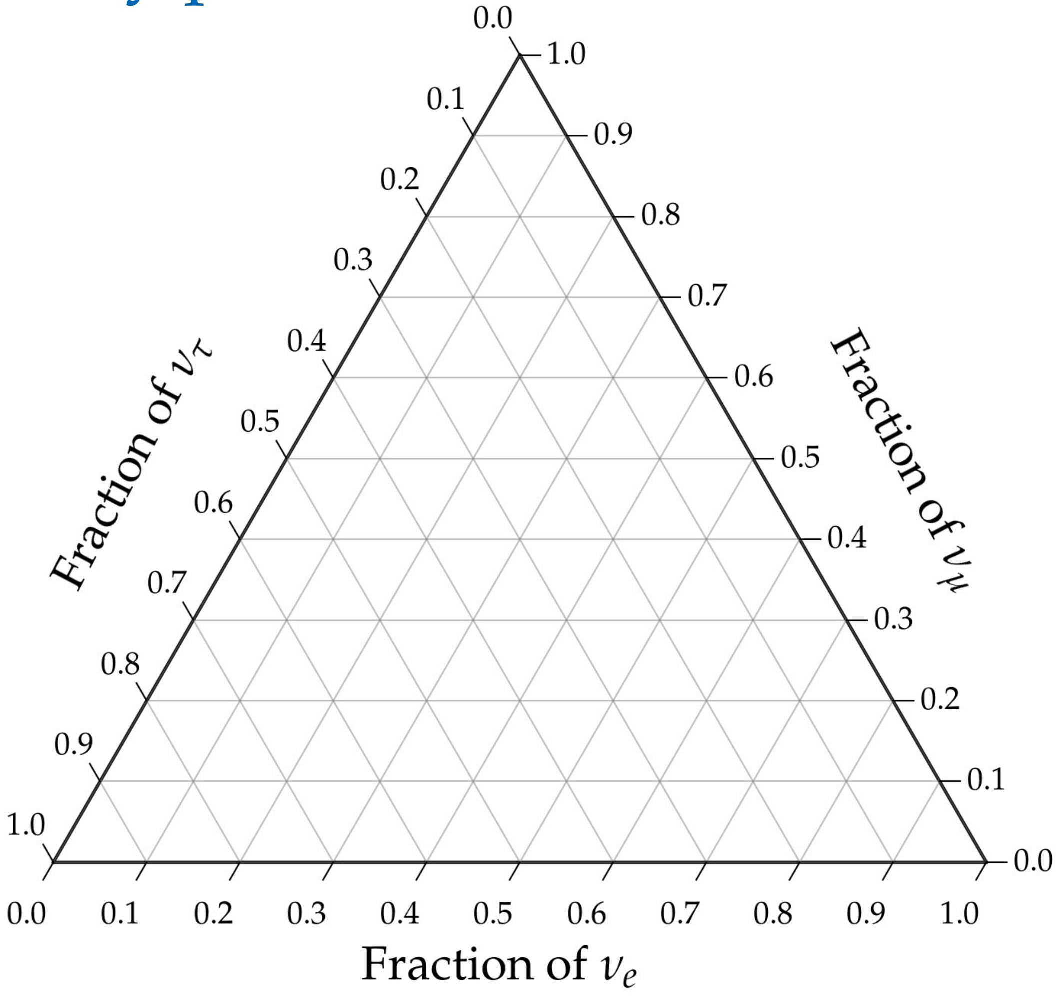
Quick aside: how to read a ternary plot

Assumes underlying unitarity –
sum of projections on each axis is 1

How to read it:

Follow the tilt of the tick marks

Always in this order: (f_e, f_μ, f_τ)



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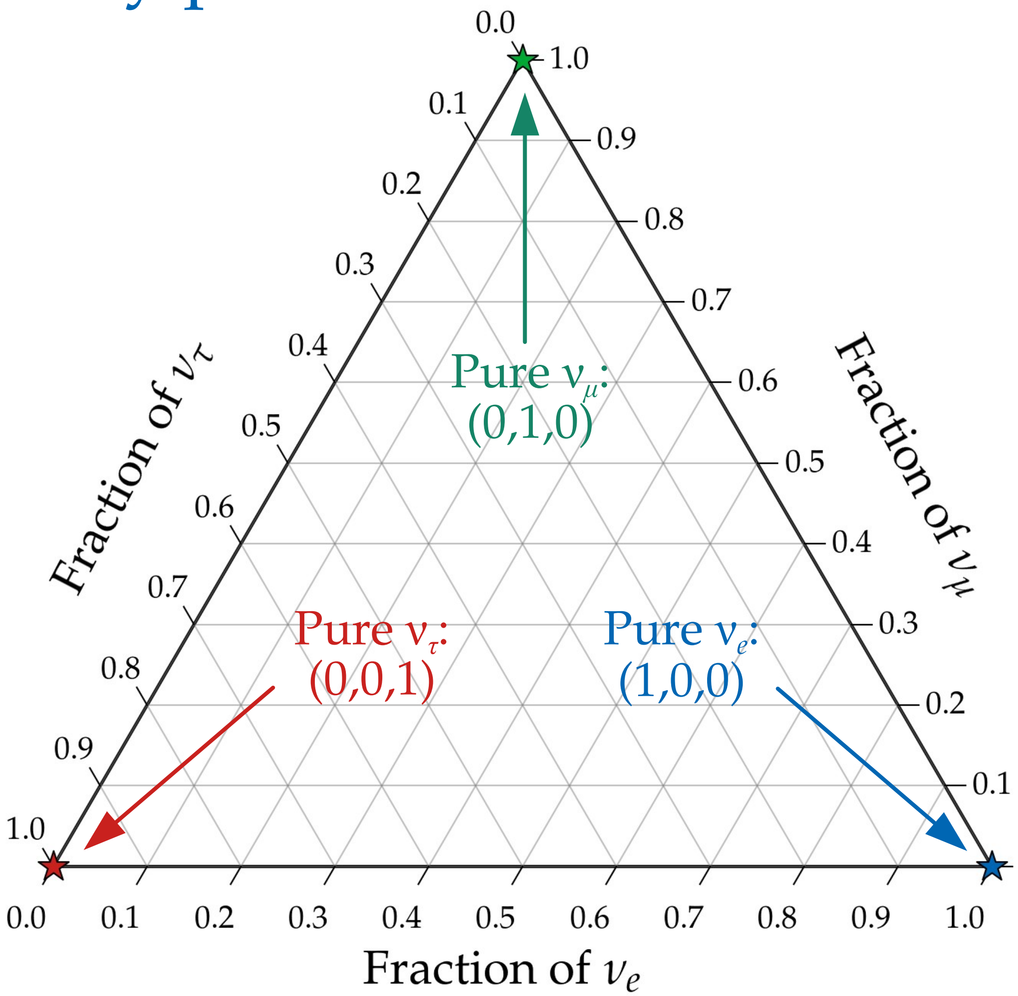

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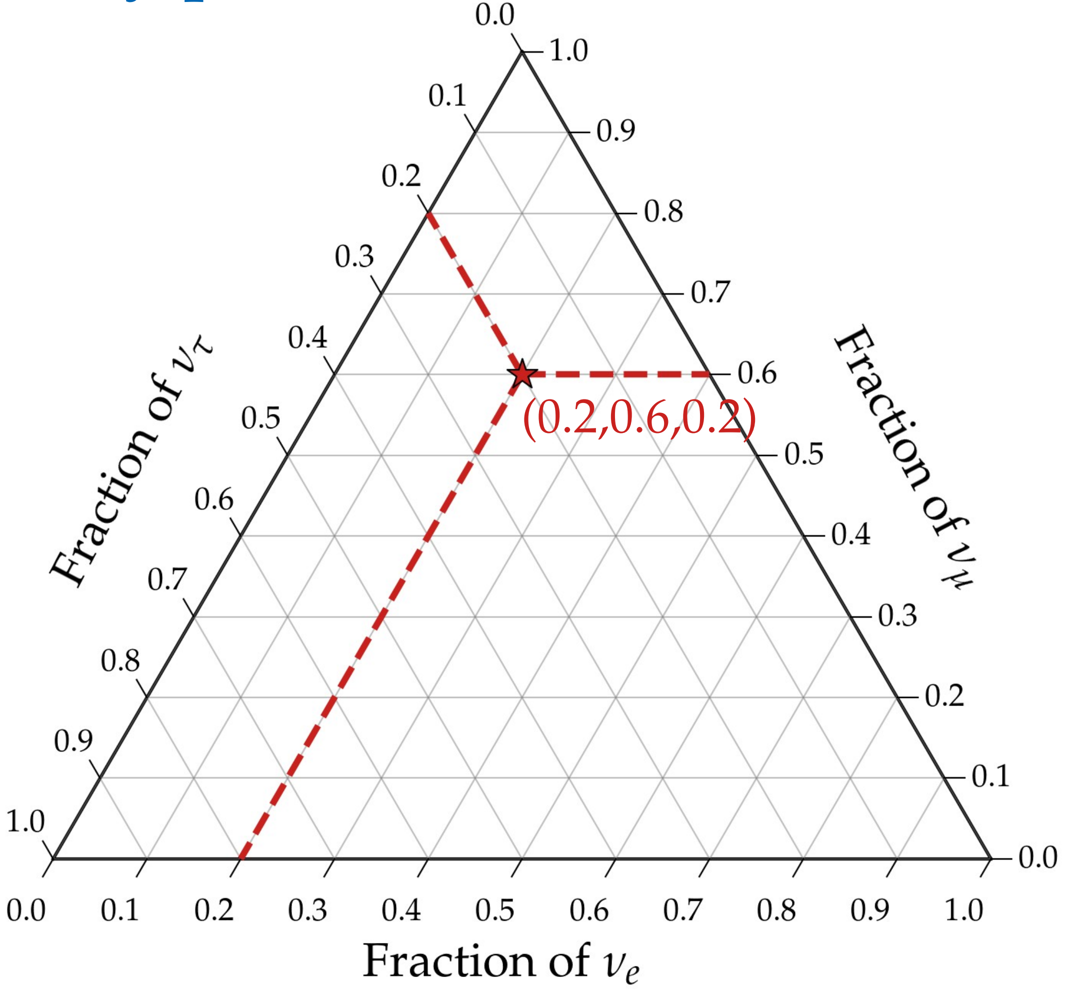

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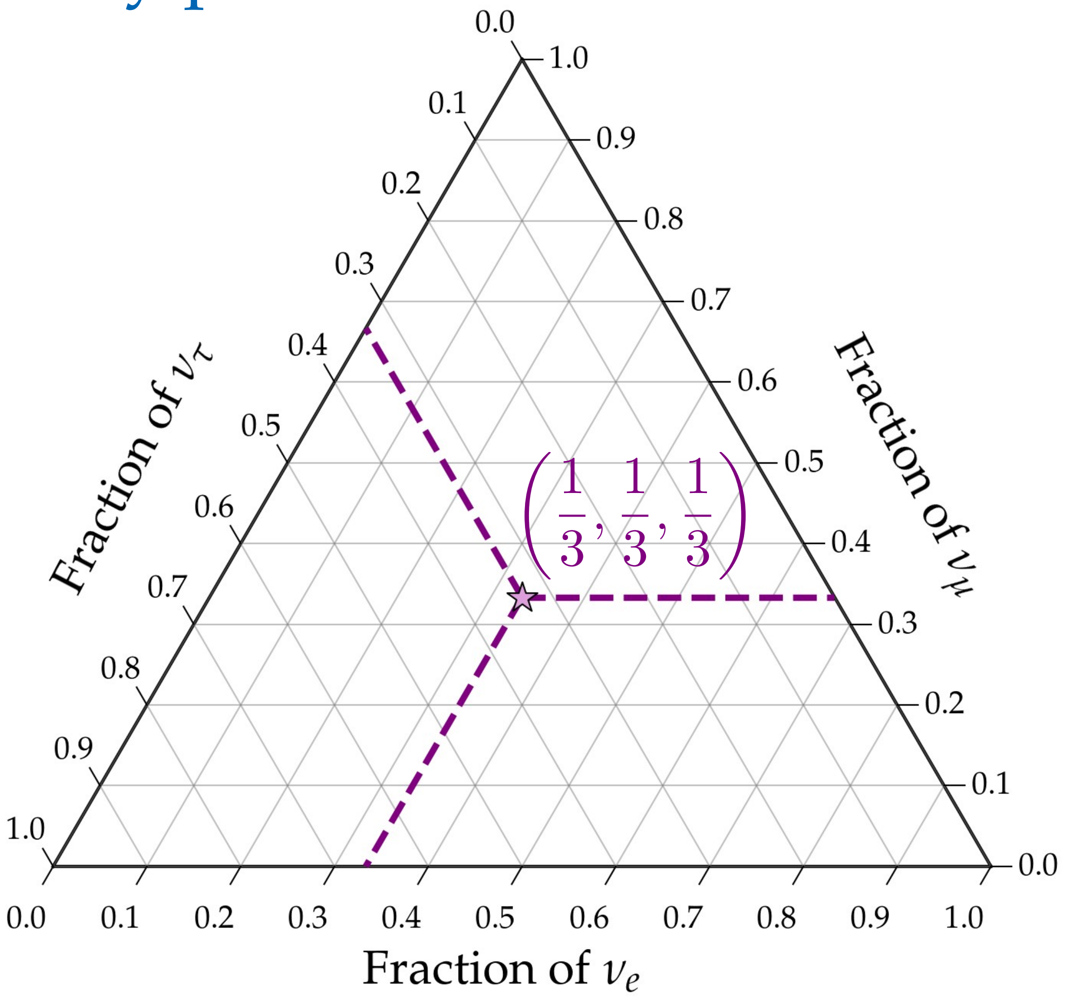

One likely TeV–PeV ν production scenario:

$p + \gamma \to \pi^+ \to \mu^+ + \nu_\mu$ followed by $\mu^+ \to e^+ + \nu_e + \bar{\nu}_\mu$

Full π decay chain

$(1/3 : 2/3 : 0)_\mathrm{S}$

Note: ν and $\bar{\nu}$ are (so far) indistinguishable in neutrino telescopes

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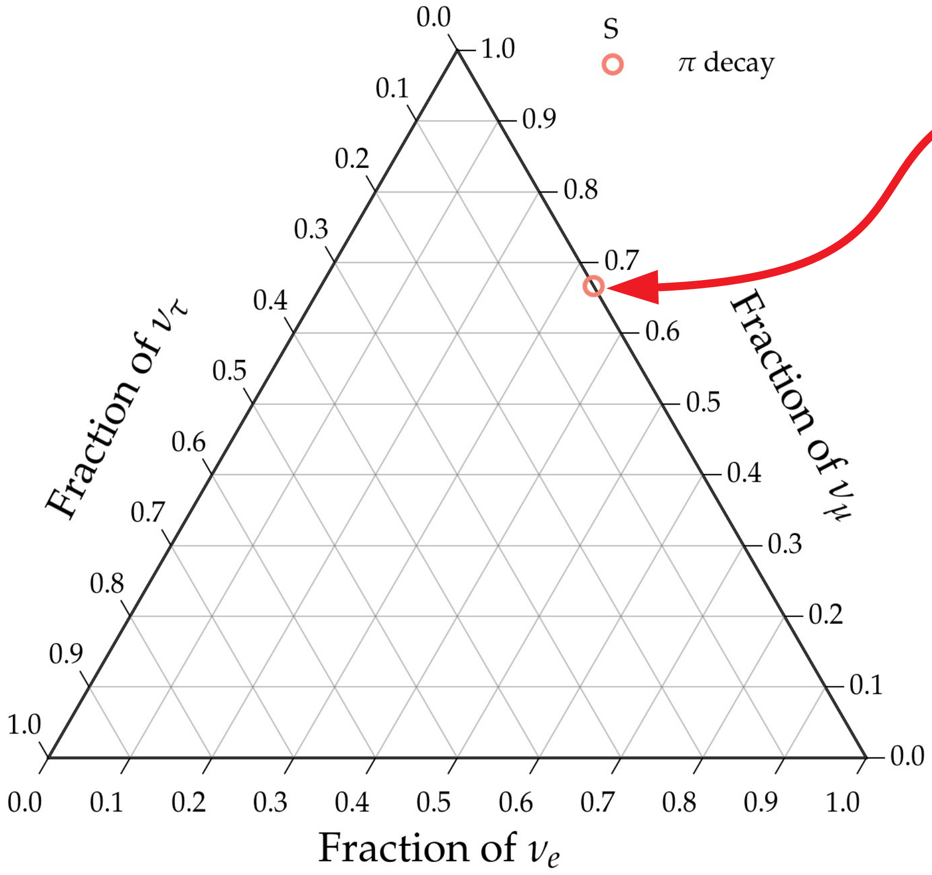

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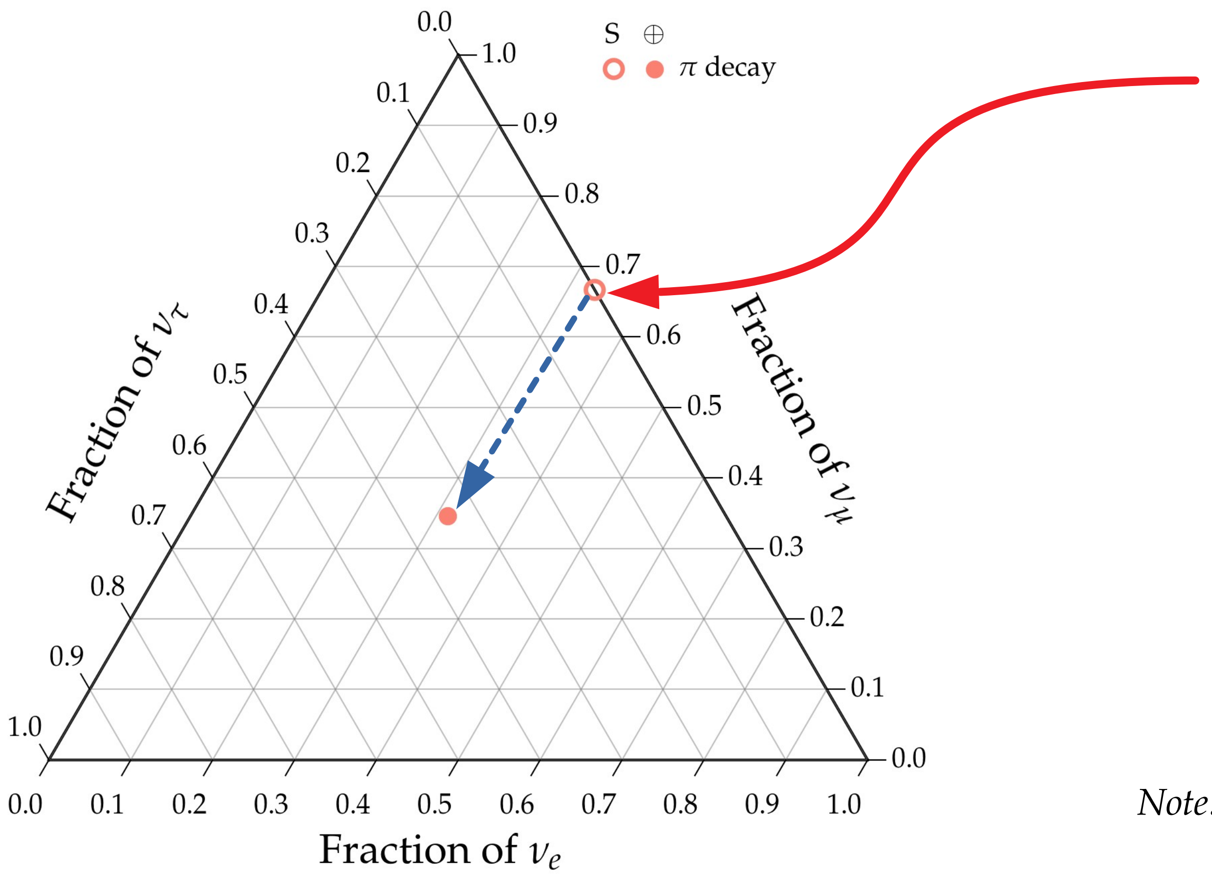

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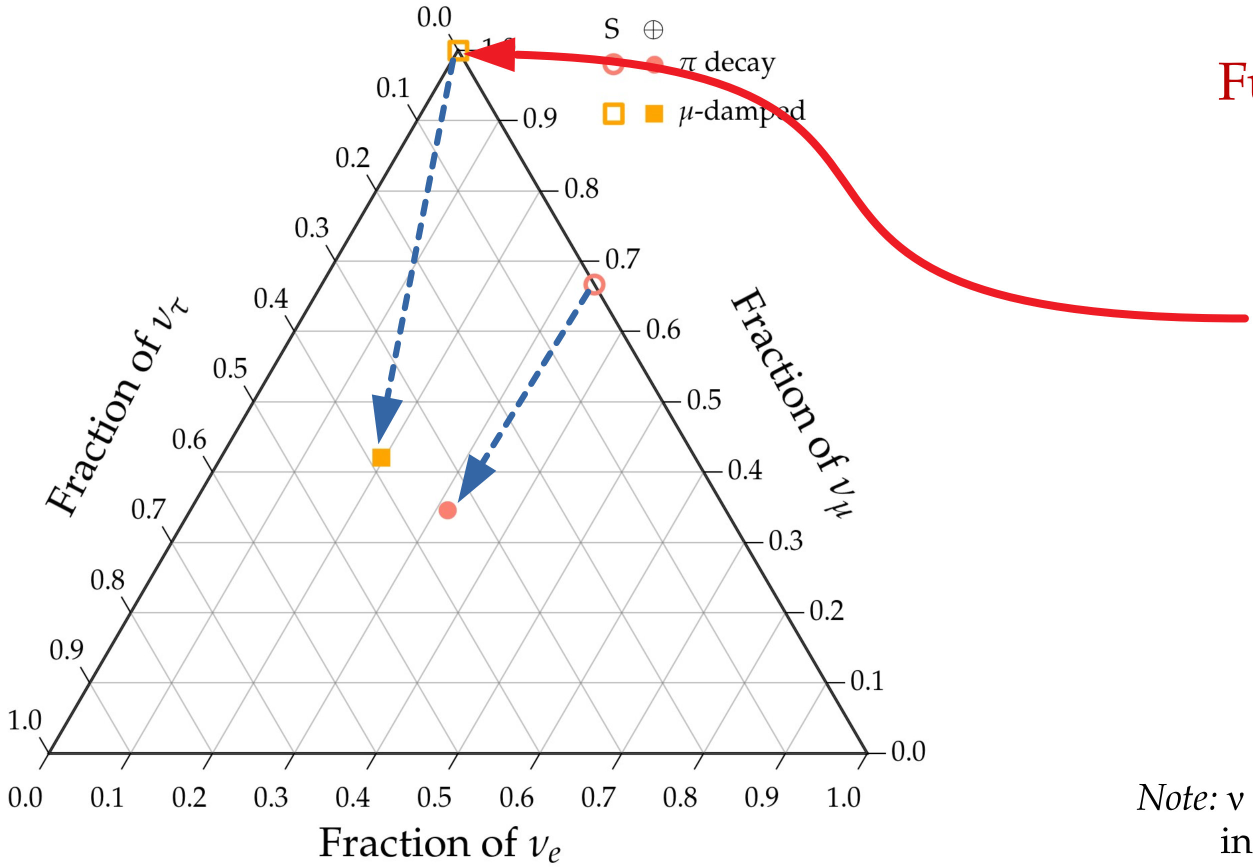

Full π decay chain

$(1/3:2/3:0)_S$

Muon damped

$(0:1:0)_S$

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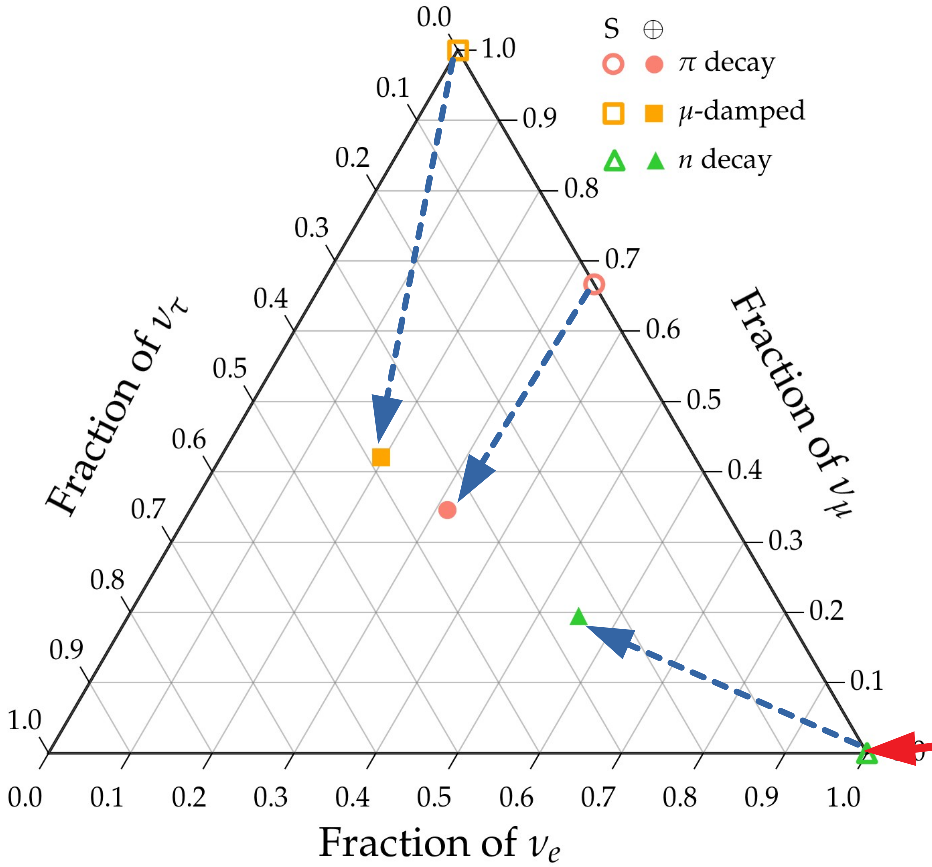


Full π decay chain

$(1/3:2/3:0)_S$

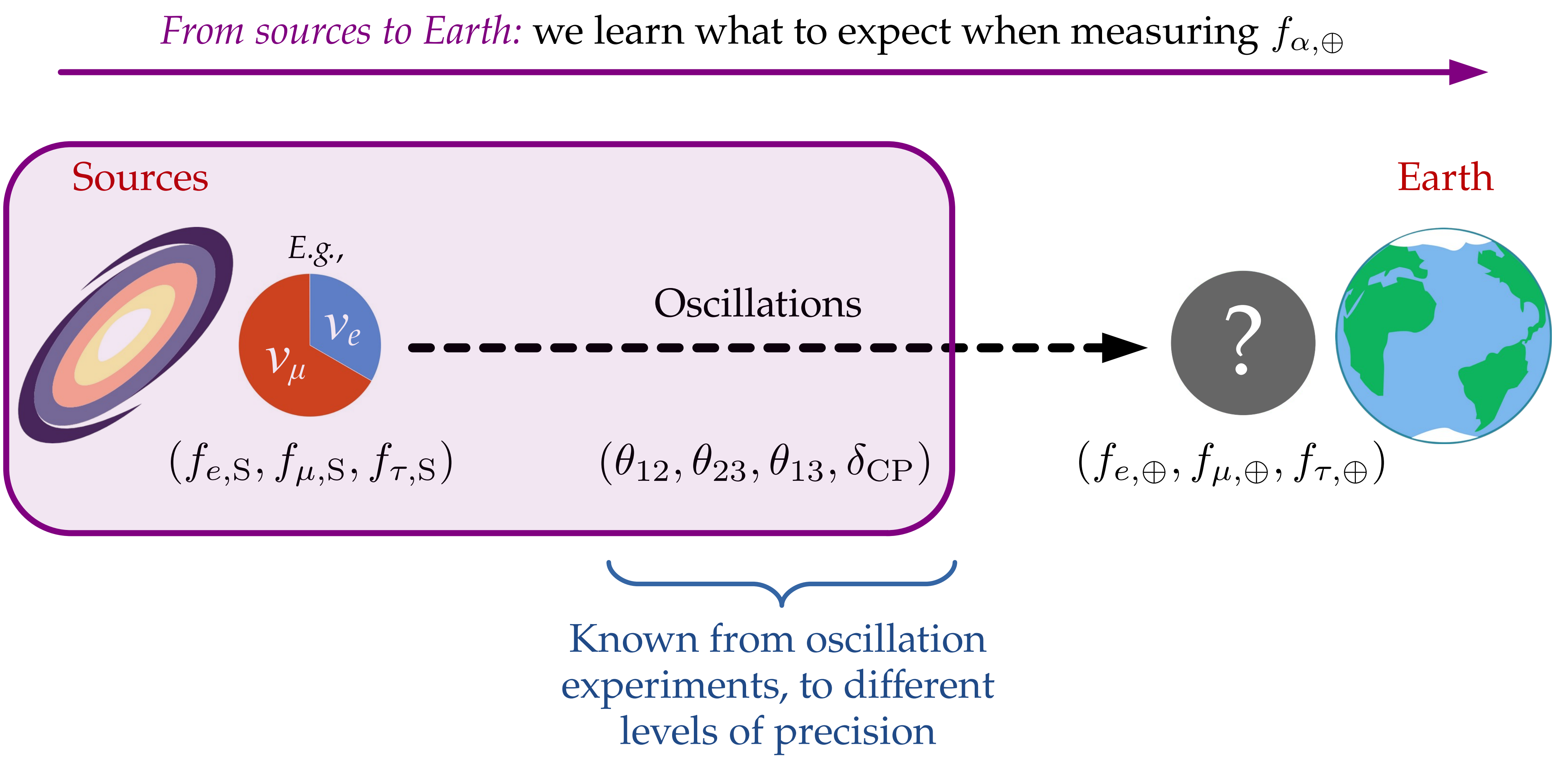
Muon damped

$(0:1:0)_S$

Neutron decay

$(1:0:0)_S$

Note: ν and $\bar{\nu}$ are (so far) indistinguishable in neutrino telescopes



From sources to Earth: we learn what to expect when measuring $f_{\alpha,\oplus}$
Sources
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E.g.,
ν_e
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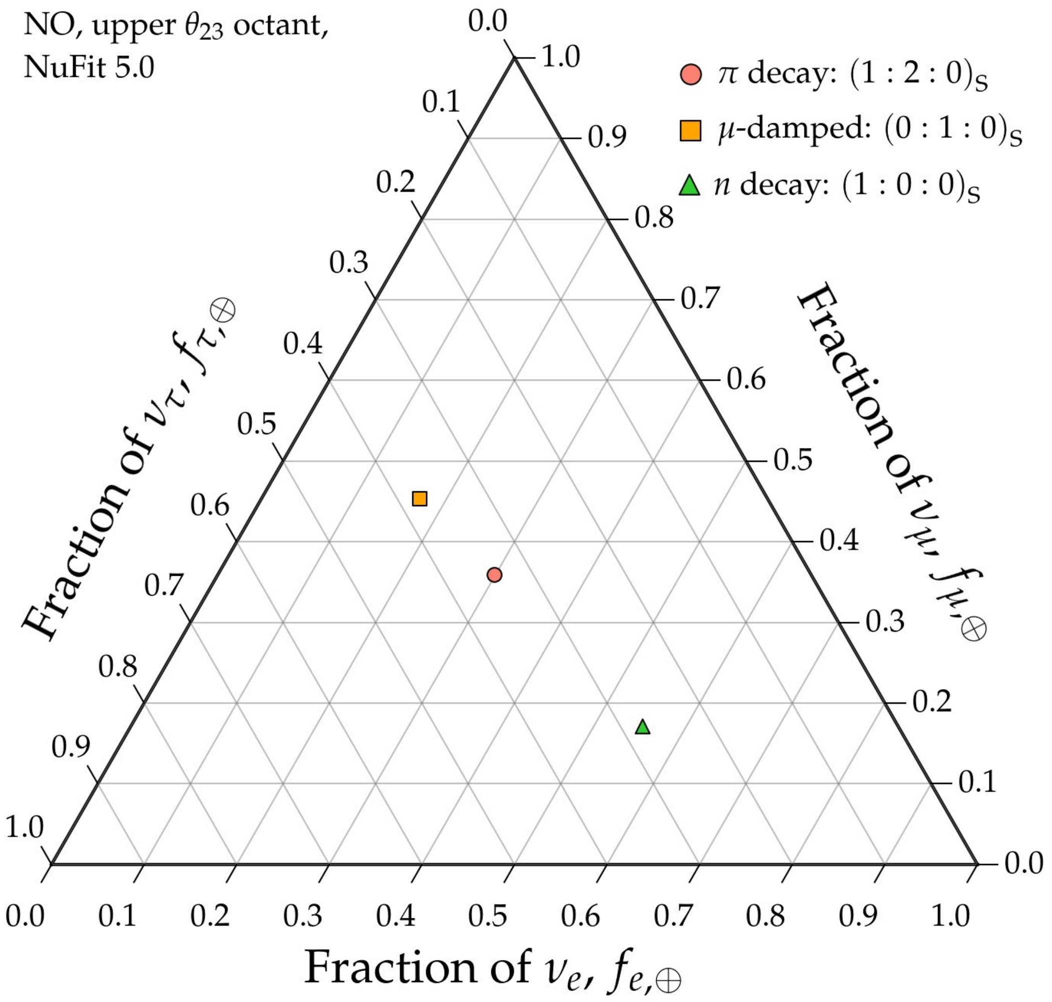
$(f_{e,\mathrm{S}}, f_{\mu,\mathrm{S}}, f_{\tau,\mathrm{S}})$
$(\theta_{12}, \theta_{23}, \theta_{13}, \delta_{\mathrm{CP}})$
$(f_{e,\oplus}, f_{\mu,\oplus}, f_{\tau,\oplus})$
Known from oscillation experiments, to different levels of precision

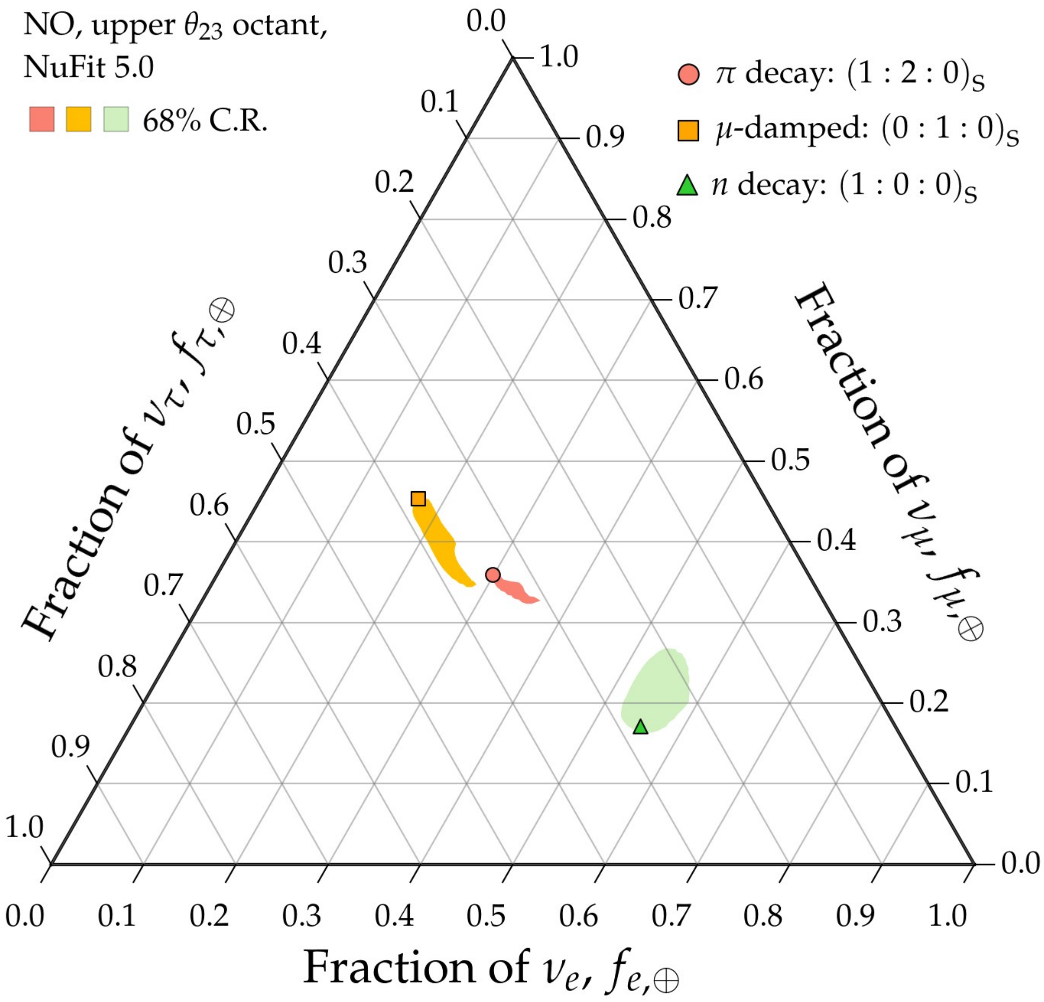
Theoretically palatable regions: today



Note:
All plots shown are for normal neutrino mass ordering (NO); inverted ordering looks similar

Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021
See also: **MB**, Beacom, Winter, *PRL* 2015

Theoretically palatable regions: today



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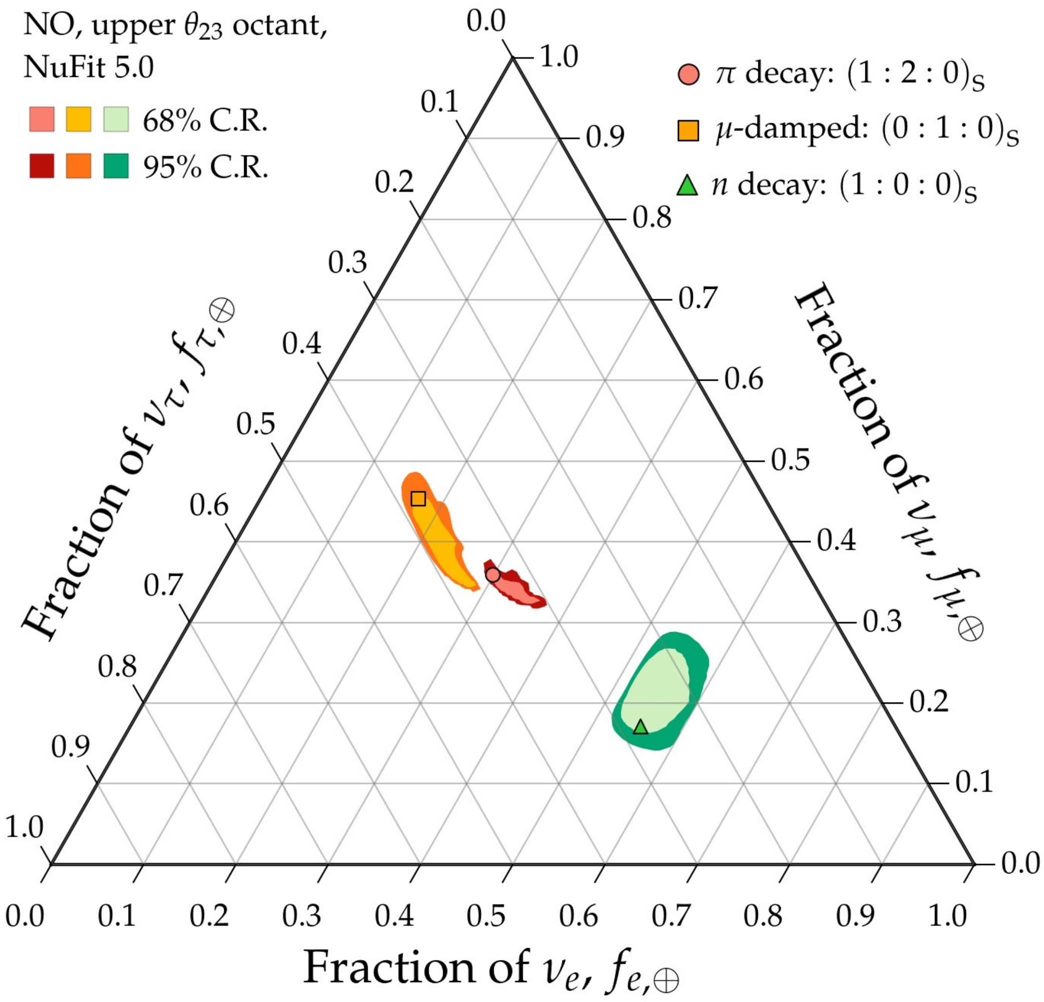

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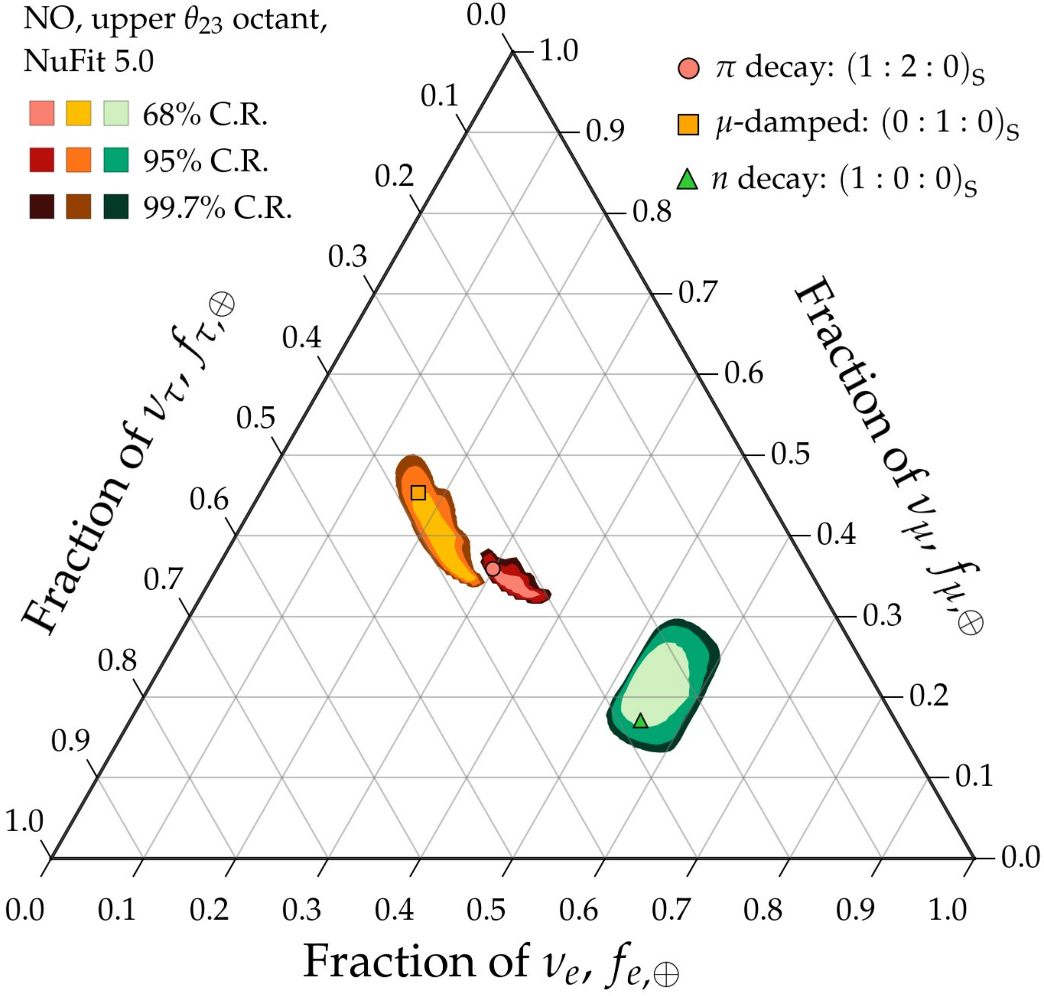

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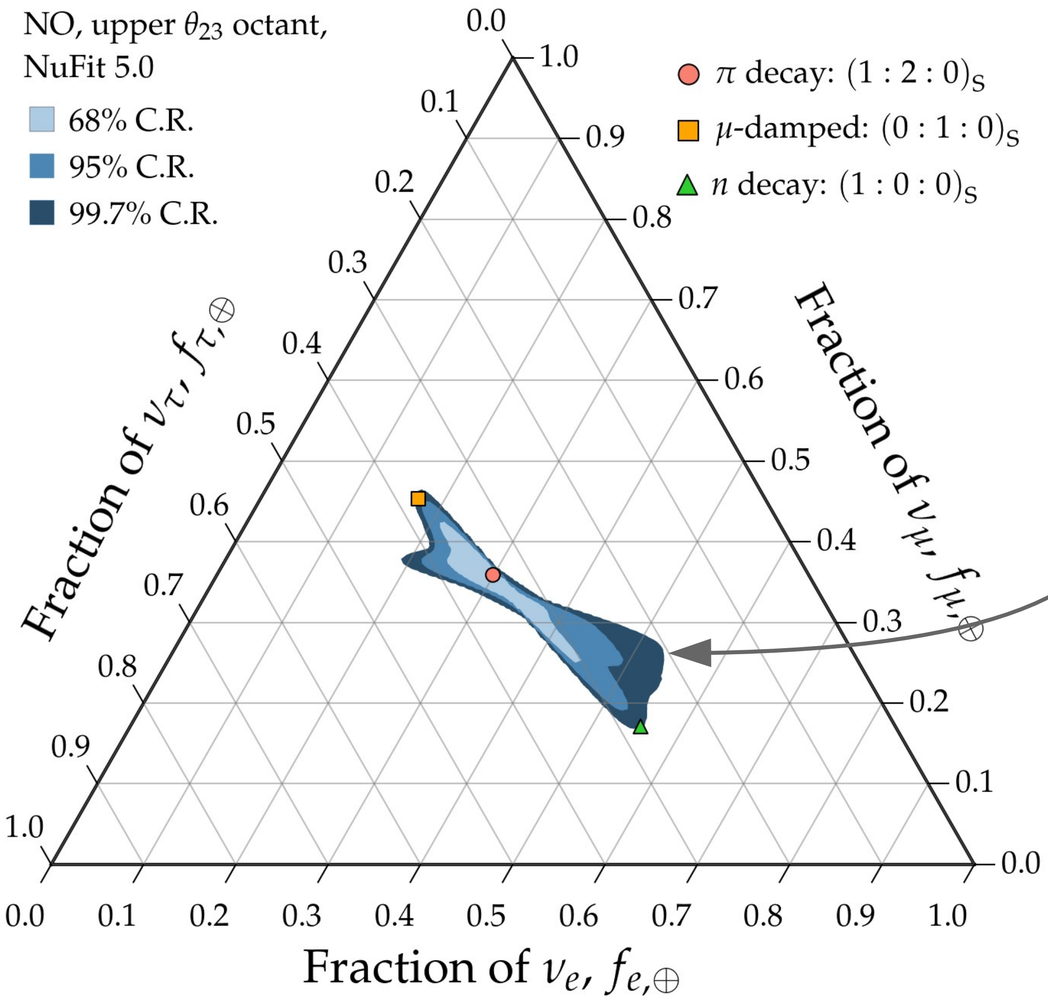
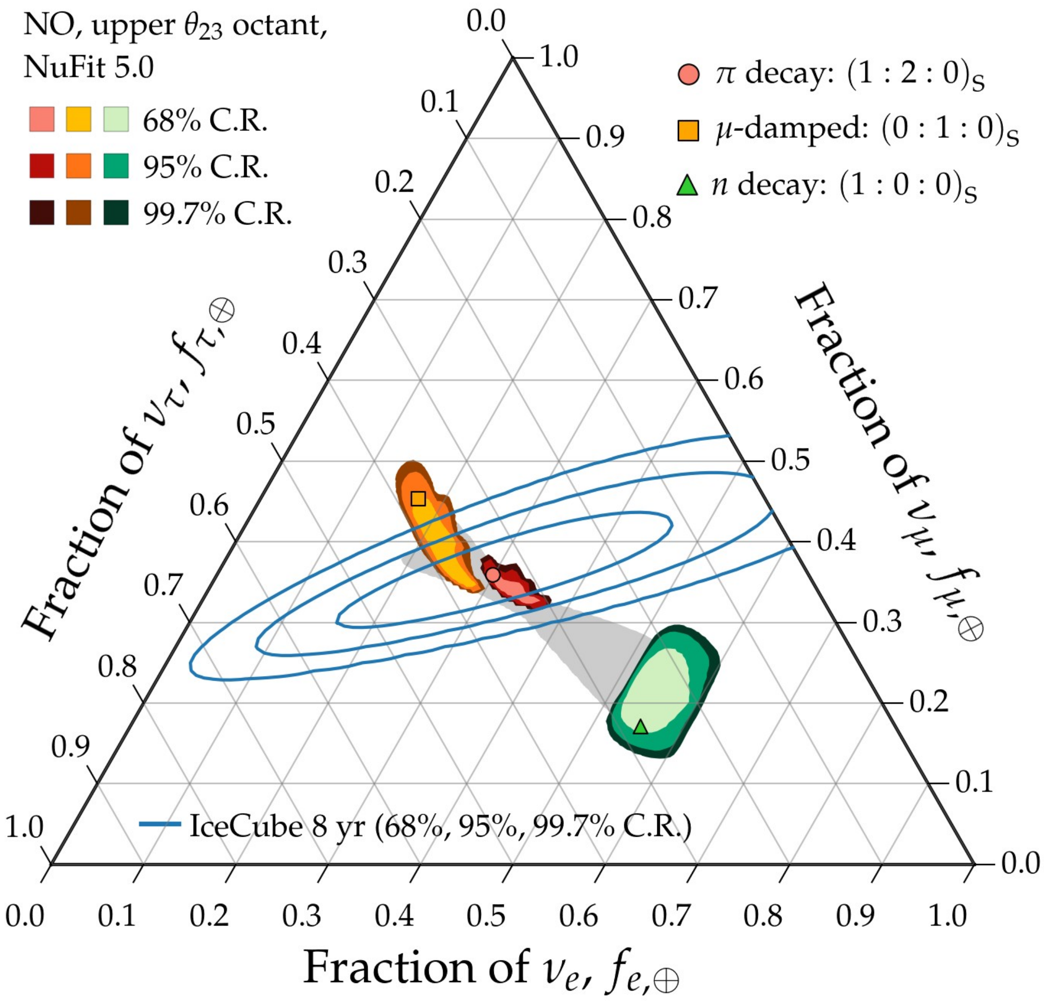


Varying over all possible flavor ratios at the source

Note:
All plots shown are for normal neutrino mass ordering (NO); inverted ordering looks similar

Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021
See also: **MB**, Beacom, Winter, *PRL* 2015

Theoretically palatable regions: today



Note:
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Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021
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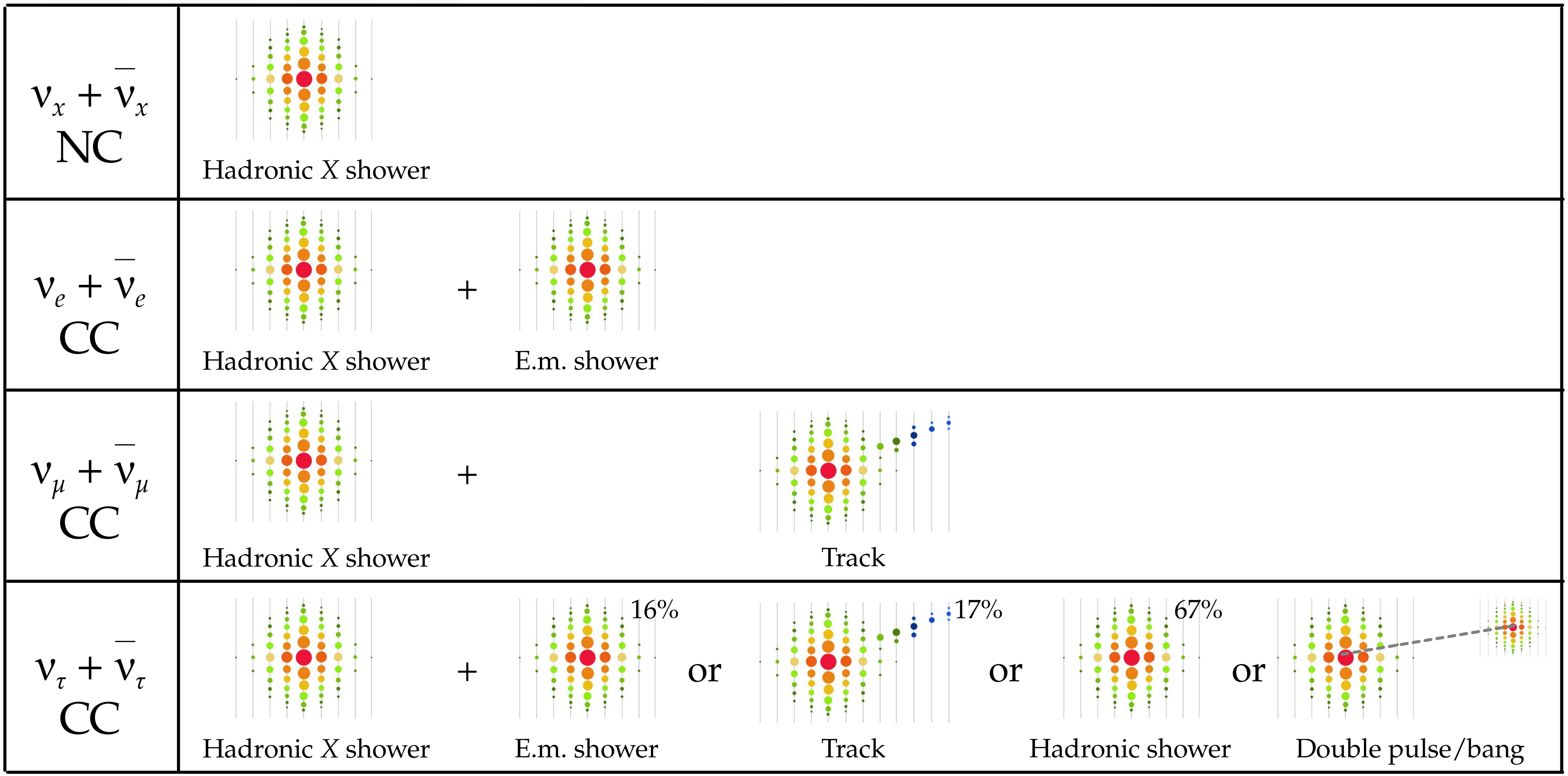
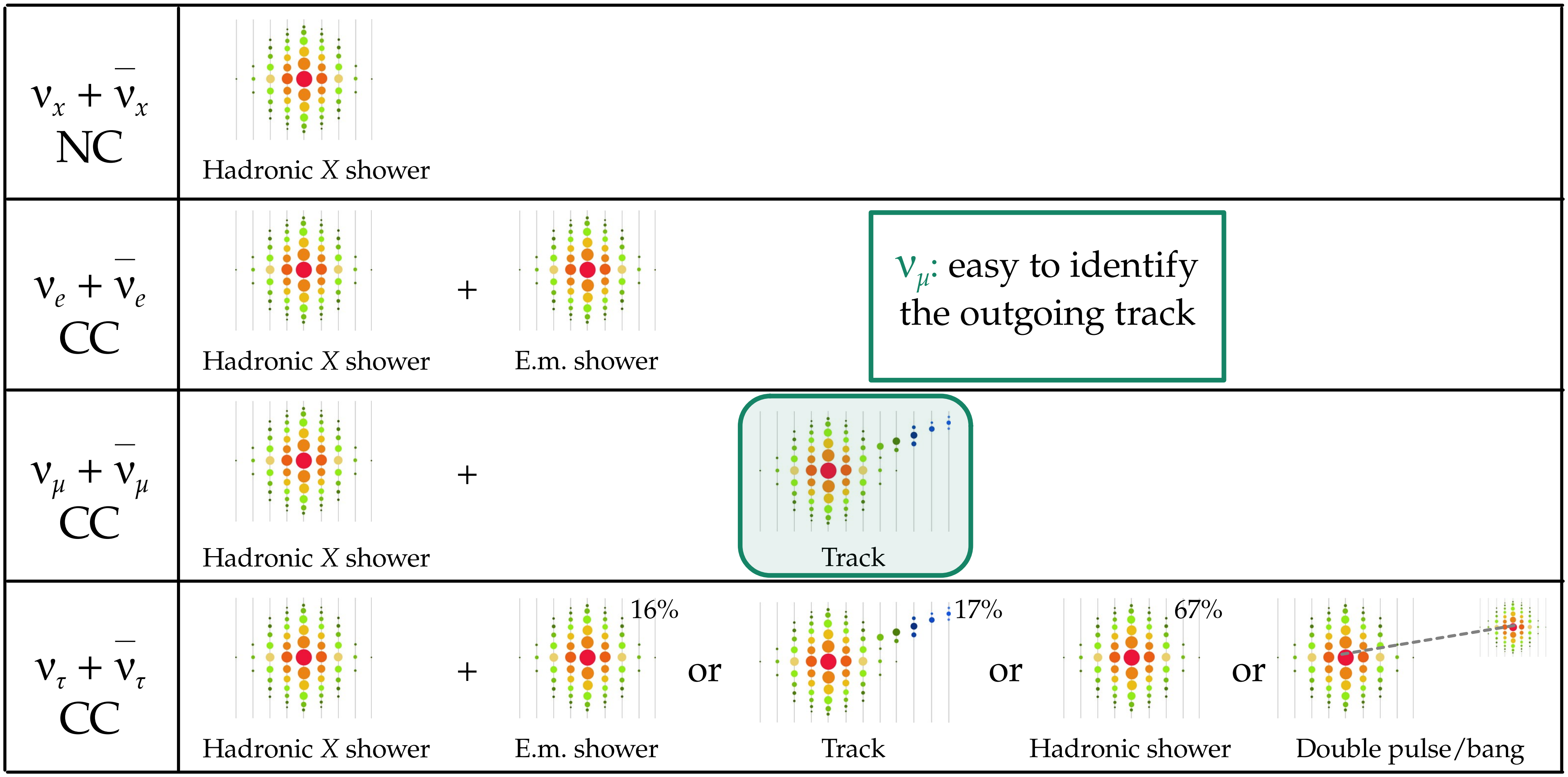
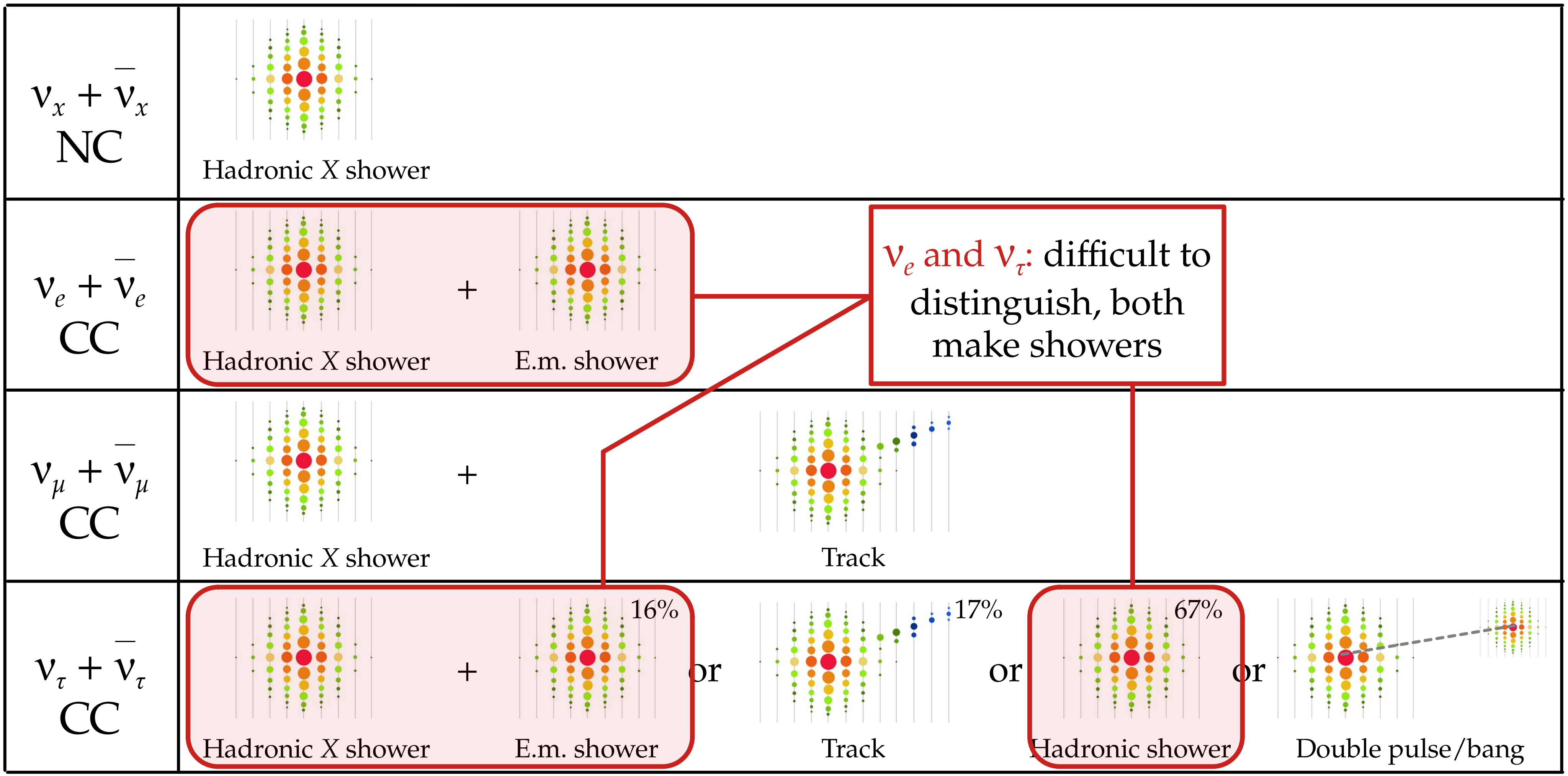
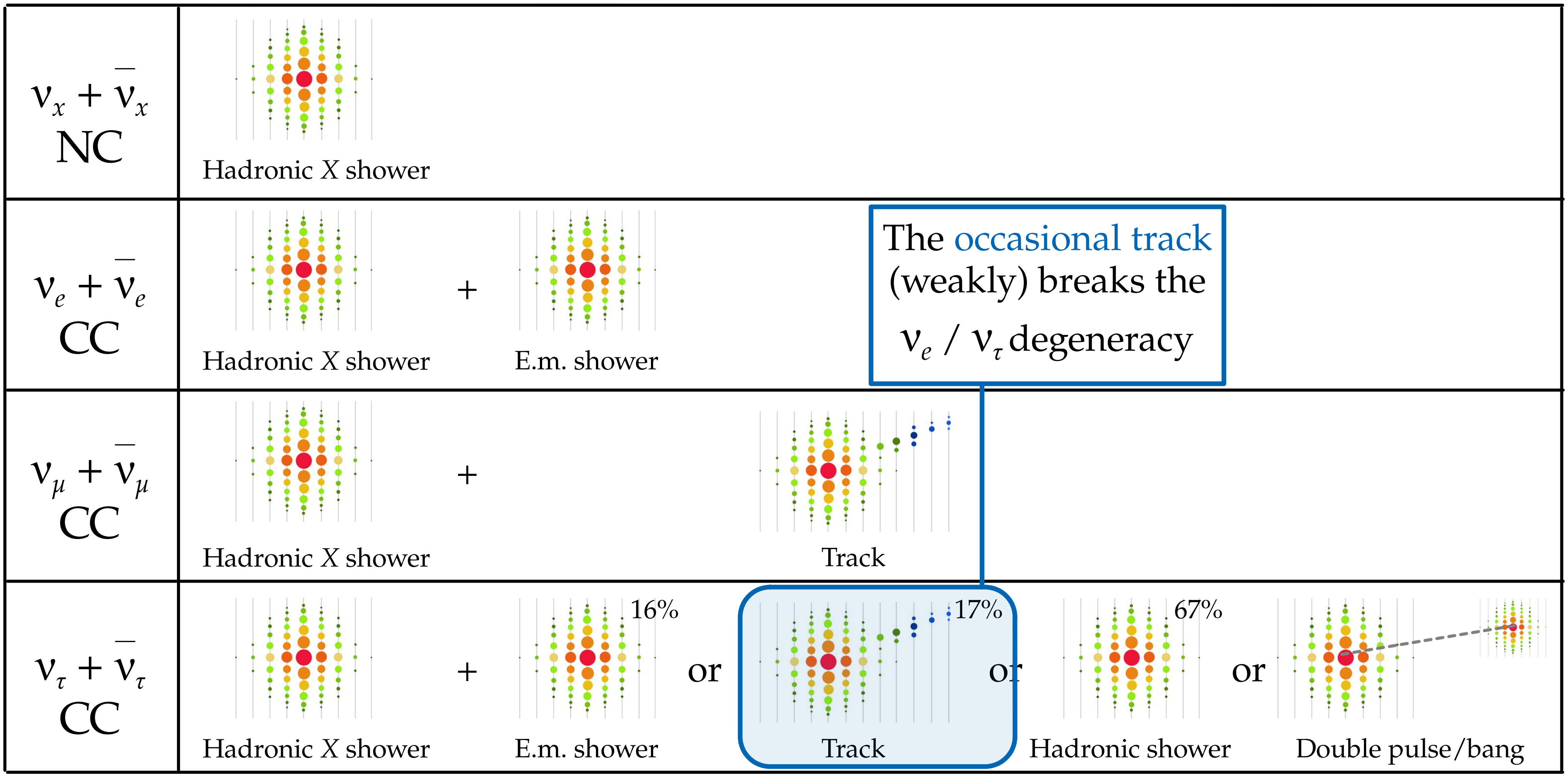

$\nu_x + \bar{\nu}_x$ NC
Hadronic X shower
$\nu_e + \bar{\nu}_e$ CC
Hadronic X shower
+
E.m. shower
$\nu_\mu + \bar{\nu}_\mu$ CC
Hadronic X shower
+
Track
$\nu_\tau + \bar{\nu}_\tau$ CC
Hadronic X shower
+
E.m. shower
16%
or
Track
17%
or
Hadronic shower
67%
or
Double pulse/bang



$\nu_x + \bar{\nu}_x$ NC
Hadronic X shower
$\nu_e + \bar{\nu}_e$ CC
Hadronic X shower
+
E.m. shower
ν_μ: easy to identify the outgoing track
$\nu_\mu + \bar{\nu}_\mu$ CC
Hadronic X shower
+
Track
$\nu_\tau + \bar{\nu}_\tau$ CC
Hadronic X shower
+
16%
E.m. shower
or
17%
Track
or
67%
Hadronic shower
or
Double pulse/bang


$\nu_x + \bar{\nu}_x$ NC
Hadronic X shower
$\nu_e + \bar{\nu}_e$ CC
Hadronic X shower
+
E.m. shower
ν_e and ν_τ: difficult to distinguish, both make showers
$\nu_\mu + \bar{\nu}_\mu$ CC
Hadronic X shower
+
Track
$\nu_\tau + \bar{\nu}_\tau$ CC
Hadronic X shower
+
E.m. shower
16%
or
Track
17%
or
Hadronic shower
67%
or
Double pulse/bang


$\nu_x + \bar{\nu}_x$ NC
Hadronic X shower
$\nu_e + \bar{\nu}_e$ CC
Hadronic X shower
+
E.m. shower
The occasional track (weakly) breaks the ν_e / ν_τ degeneracy
$\nu_\mu + \bar{\nu}_\mu$ CC
Hadronic X shower
+
Track
$\nu_\tau + \bar{\nu}_\tau$ CC
Hadronic X shower
+
16%
E.m. shower
or
17%
Track
or
67%
Hadronic shower
or
Double pulse/bang

Measuring flavor composition: 2015–2040

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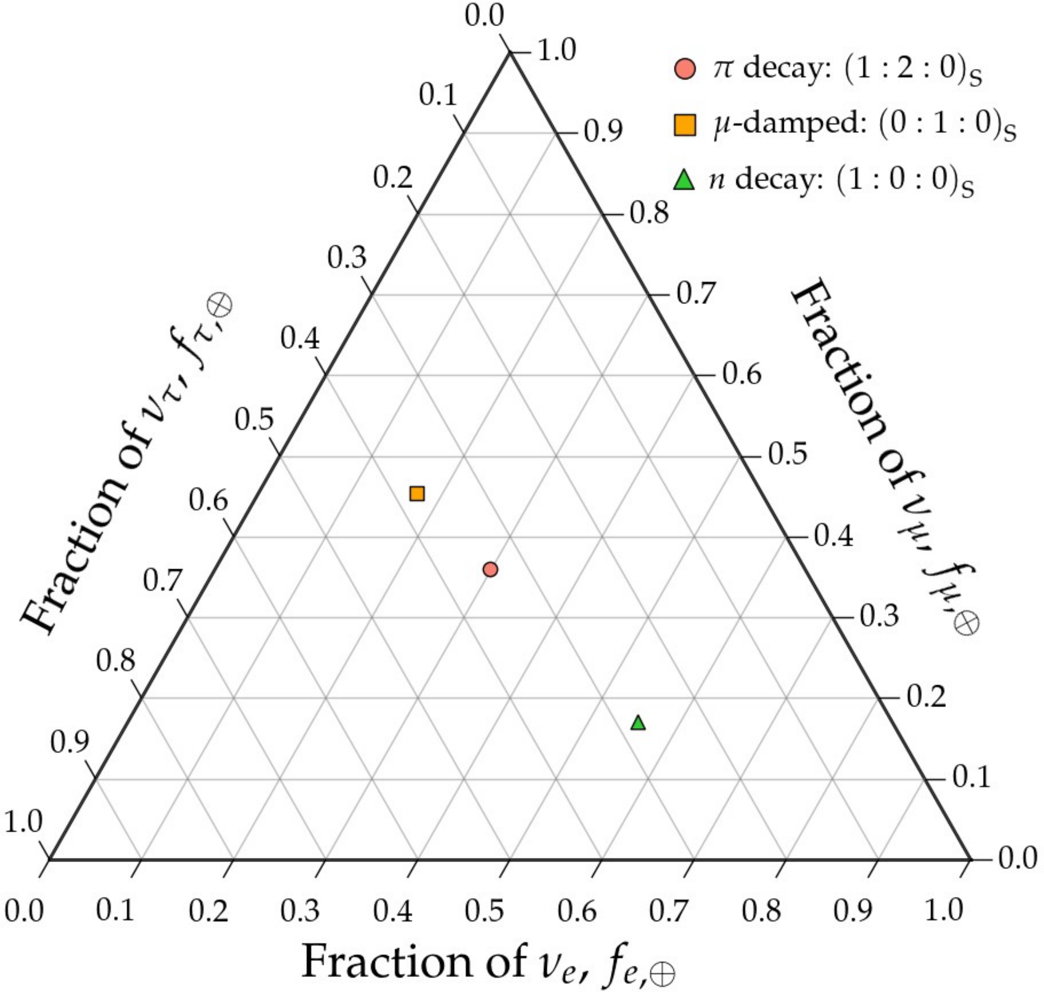
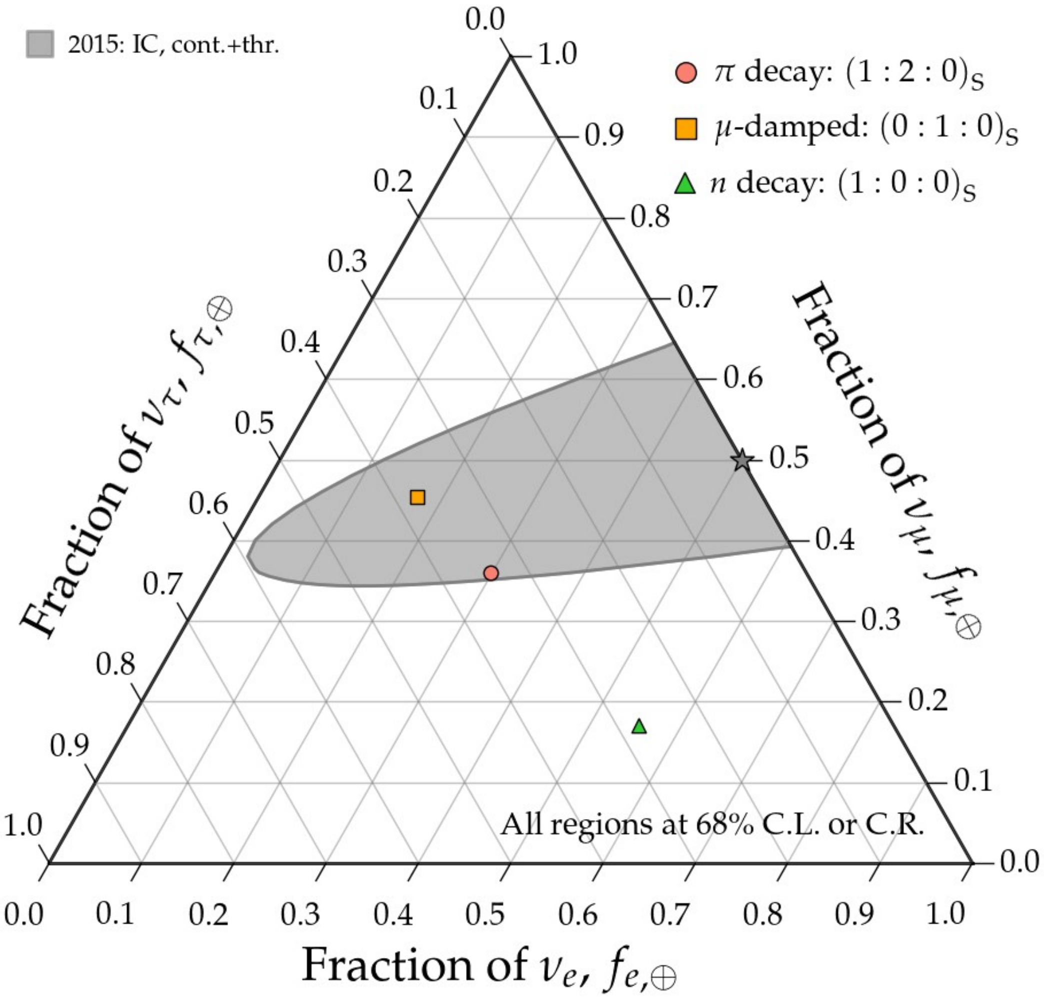


Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

Measuring flavor composition: 2015–2040



Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

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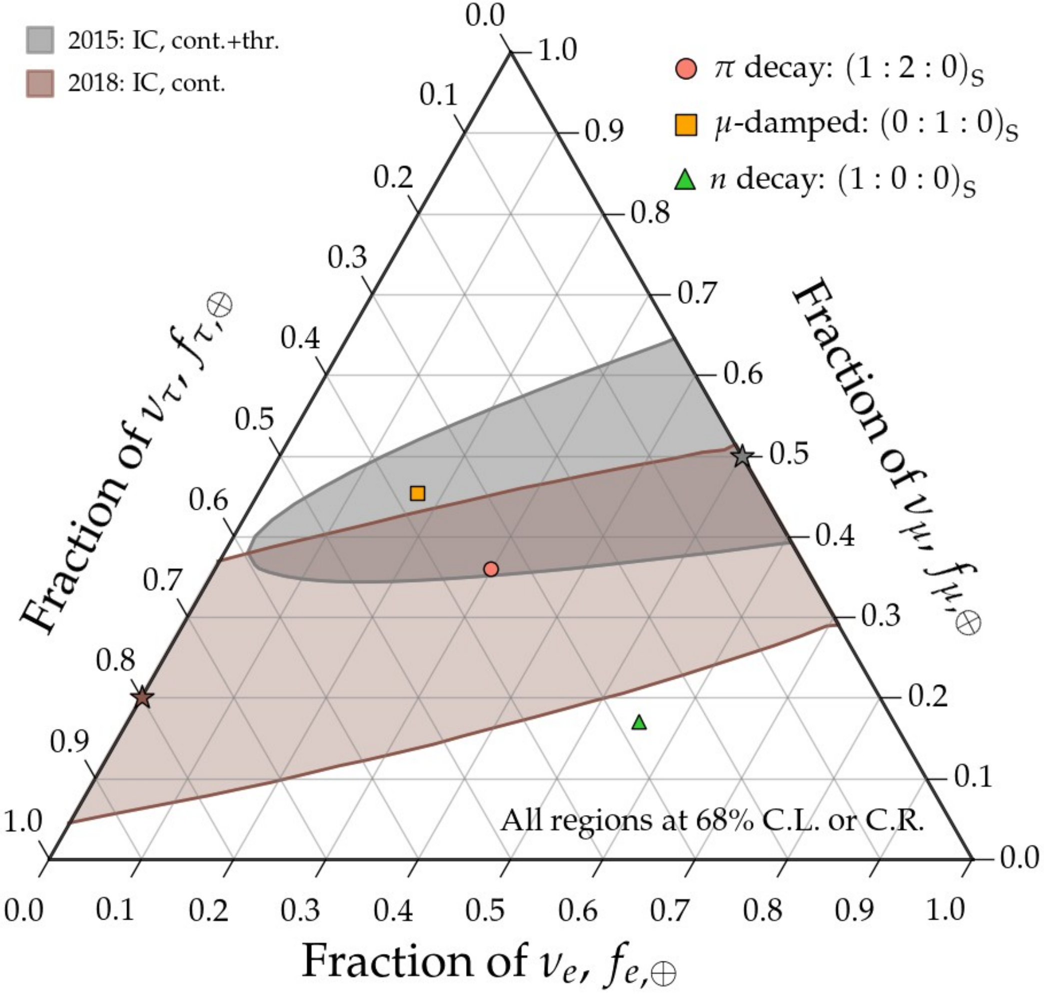
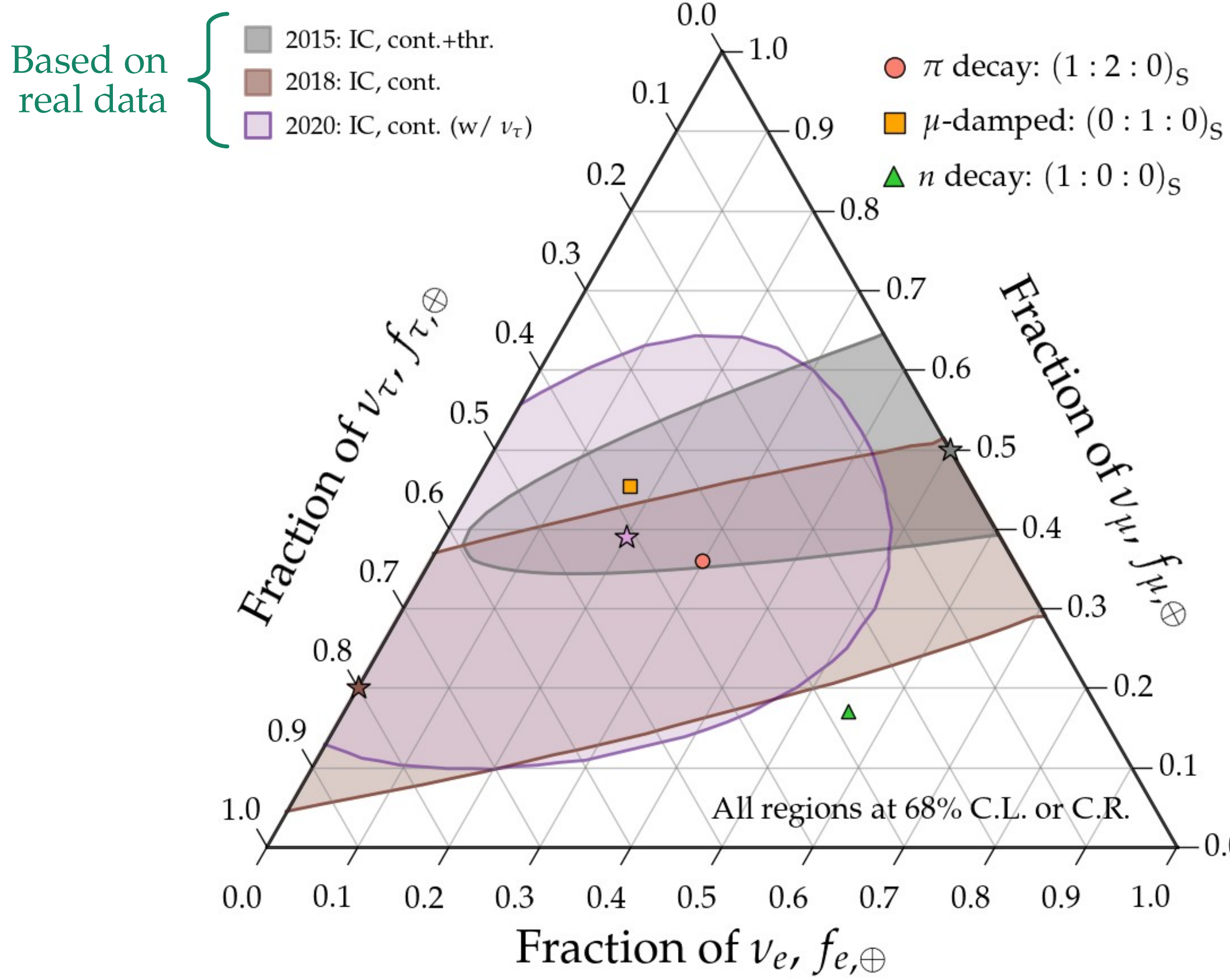


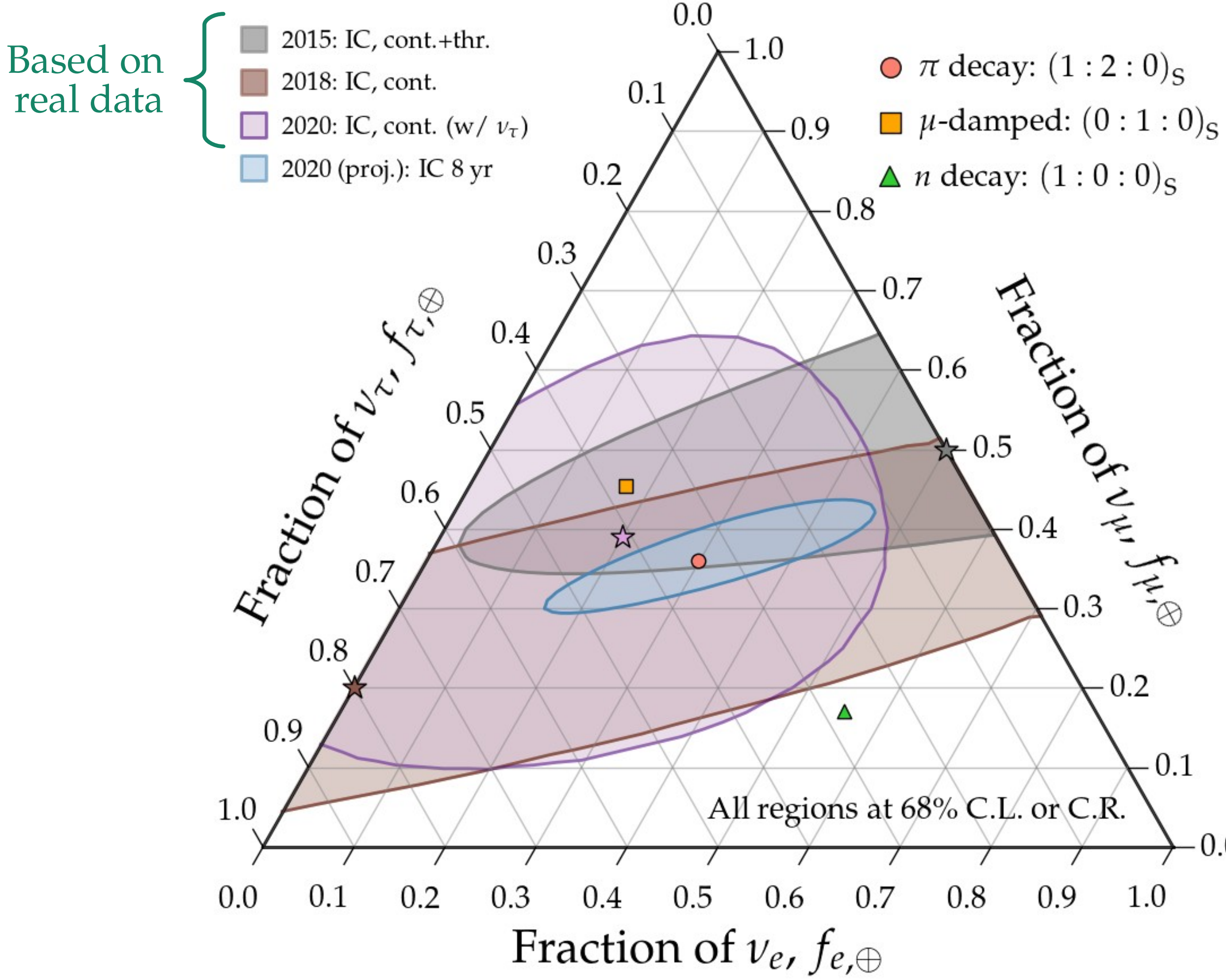
Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

Measuring flavor composition: 2015–2040



Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

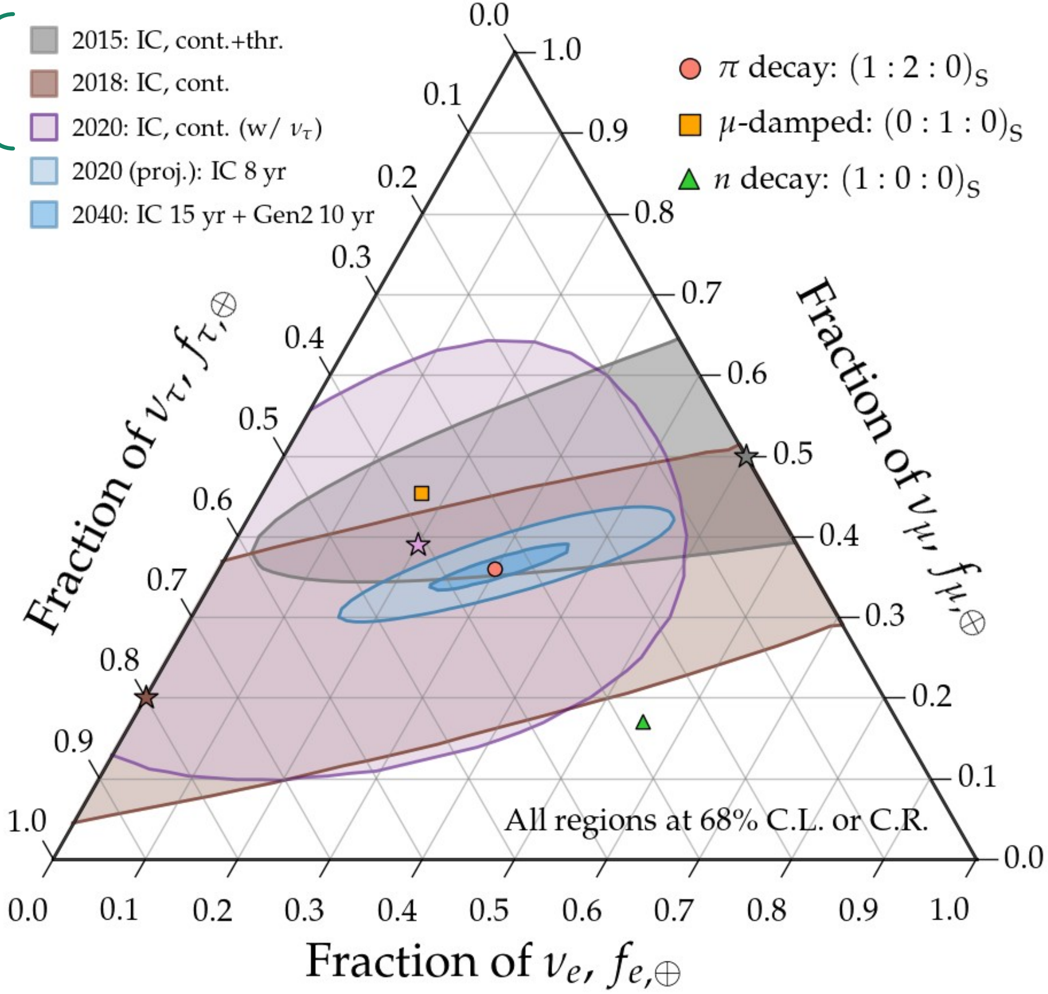
Measuring flavor composition: 2015–2040



Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

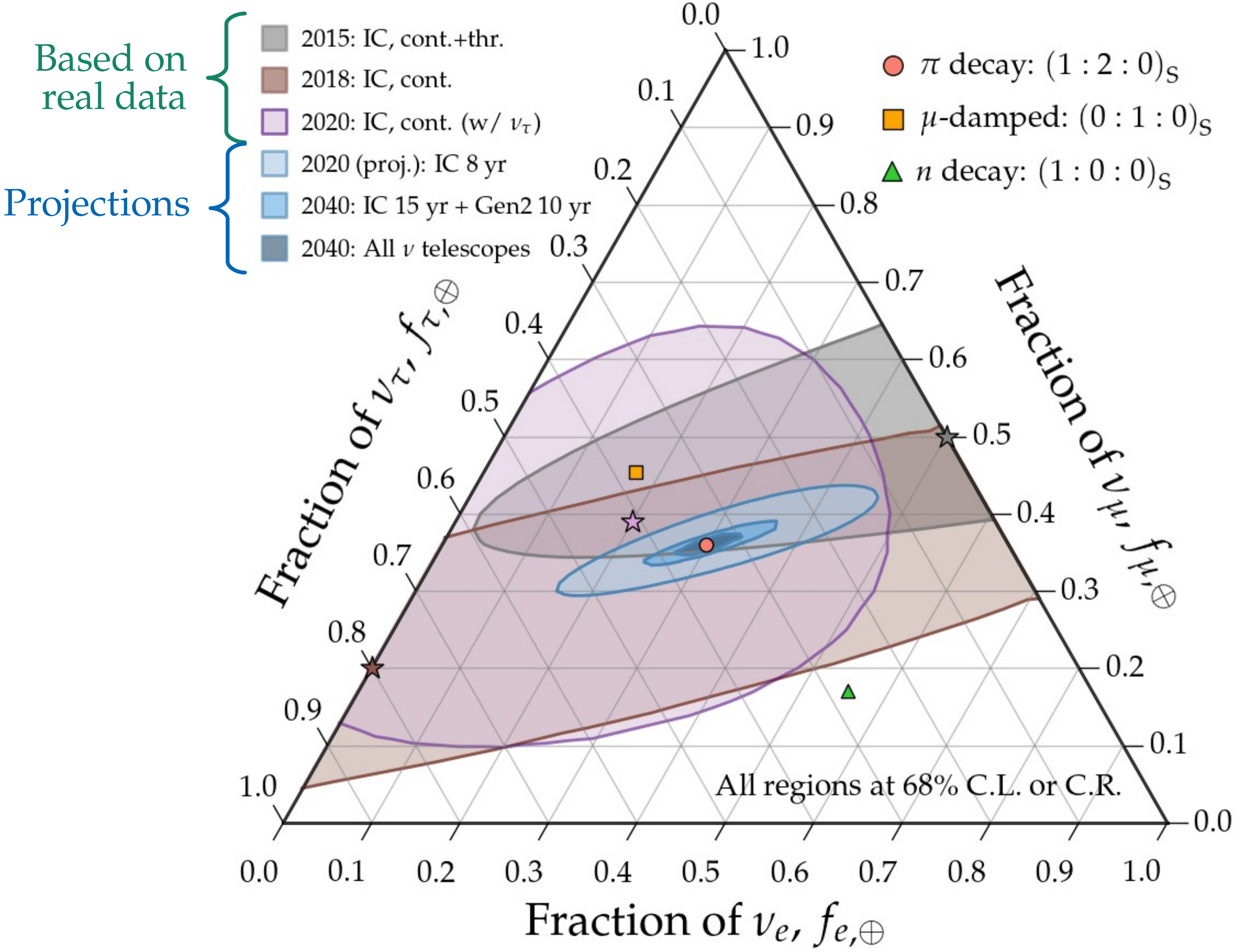
Measuring flavor composition: 2015–2040

Based on real data



Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

Measuring flavor composition: 2015–2040



Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

Today

TeV–PeV ν

Today

TeV–PeV ν

Turn predictions
into data-driven tests

Today

TeV–PeV ν

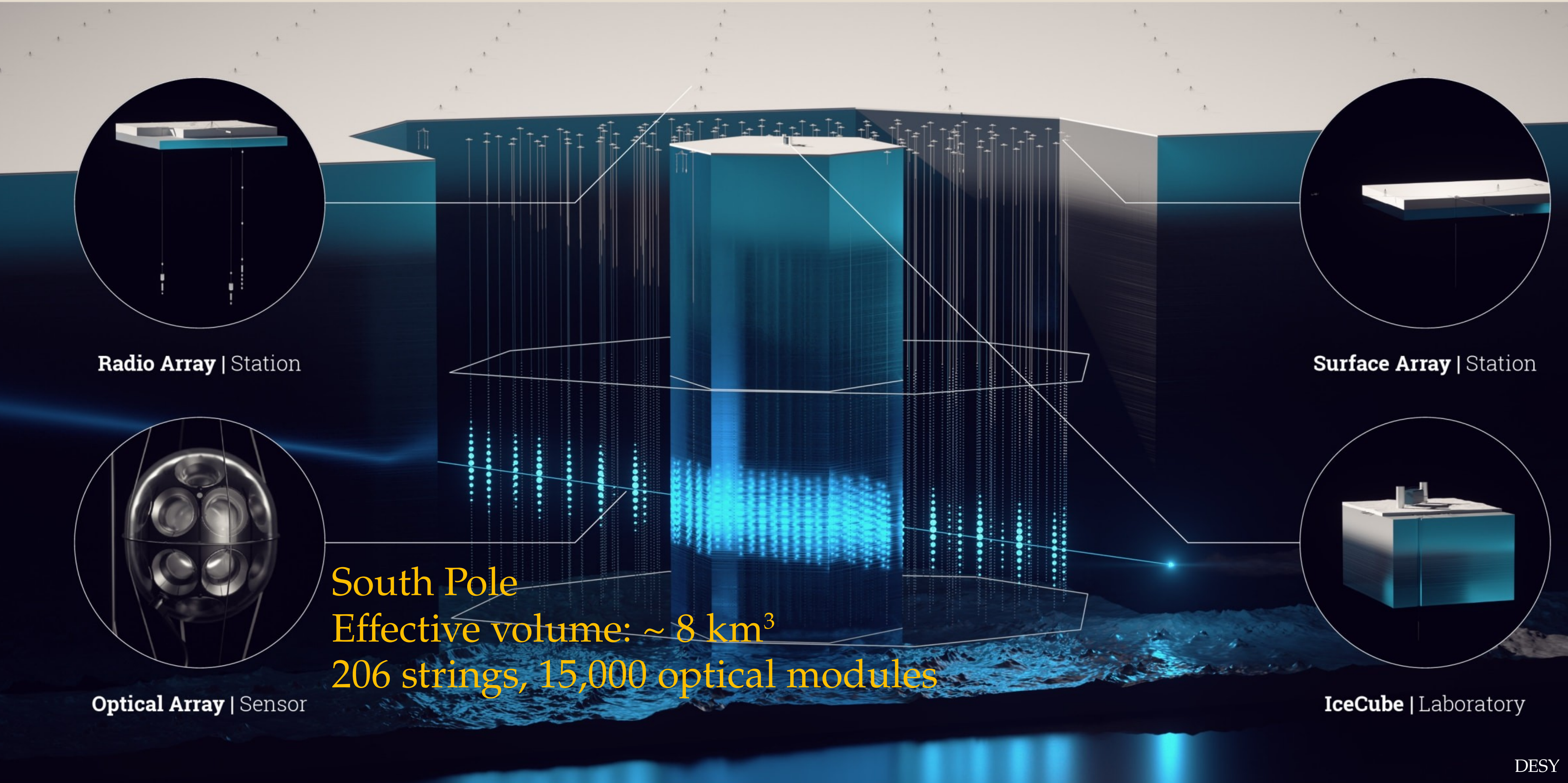
Turn predictions
into data-driven tests

Key developments:
Bigger detectors → larger statistics
Better reconstruction
Smaller astrophysical uncertainties

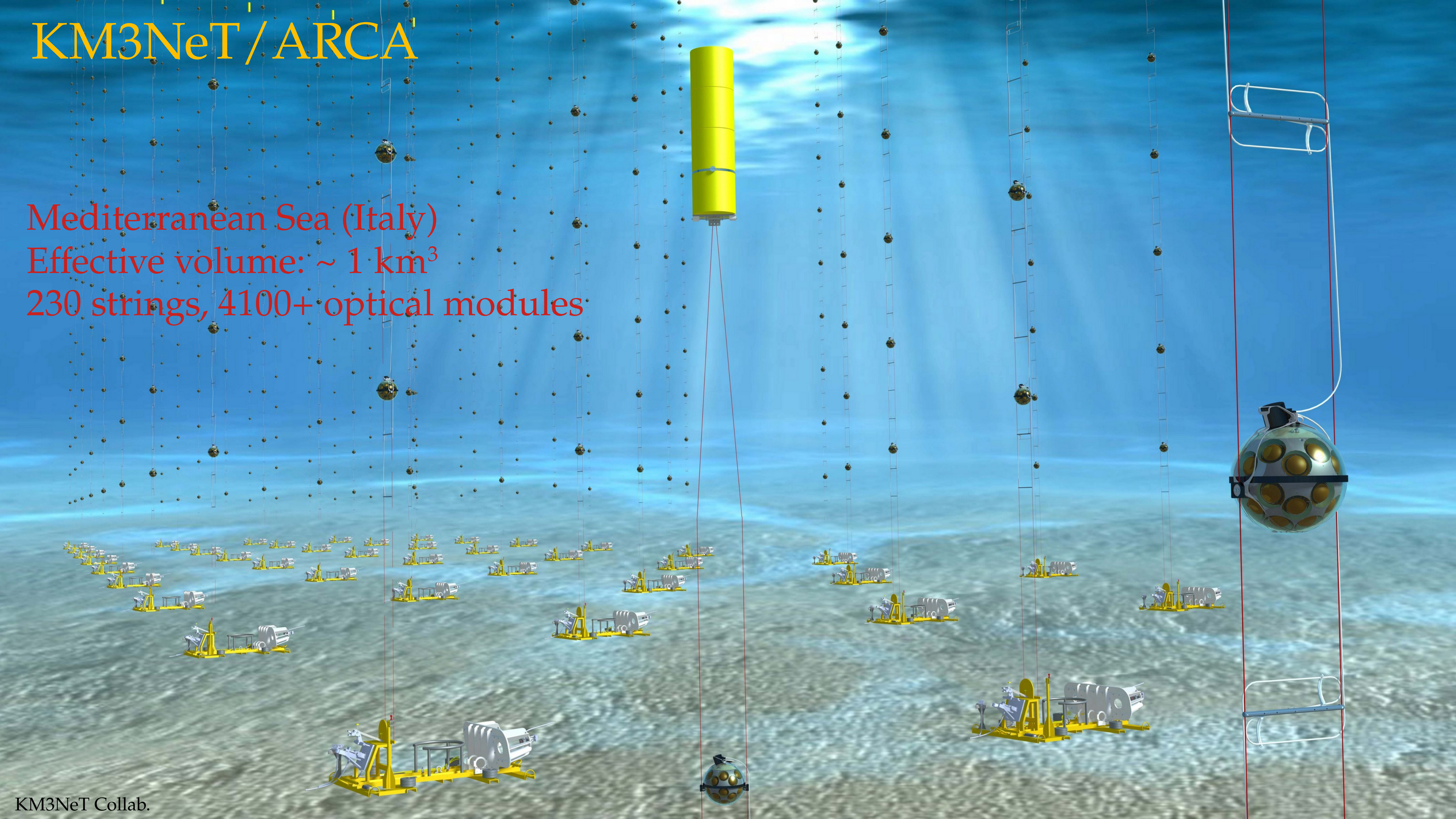
TeV–PeV
ν telescopes
2030s

Strebe/Wikipedia

IceCube-Gen2



DESY


KM3NeT/ARCA
Mediterranean Sea (Italy)
Effective volume: ~ 1 km³
230 strings, 4100+ optical modules
KM3NeT Collab.

Baikal-GVD

Lake Baikal, Russia
Effective volume: ~ 1.5 km3
90 strings, 1000+ optical modules

JINR

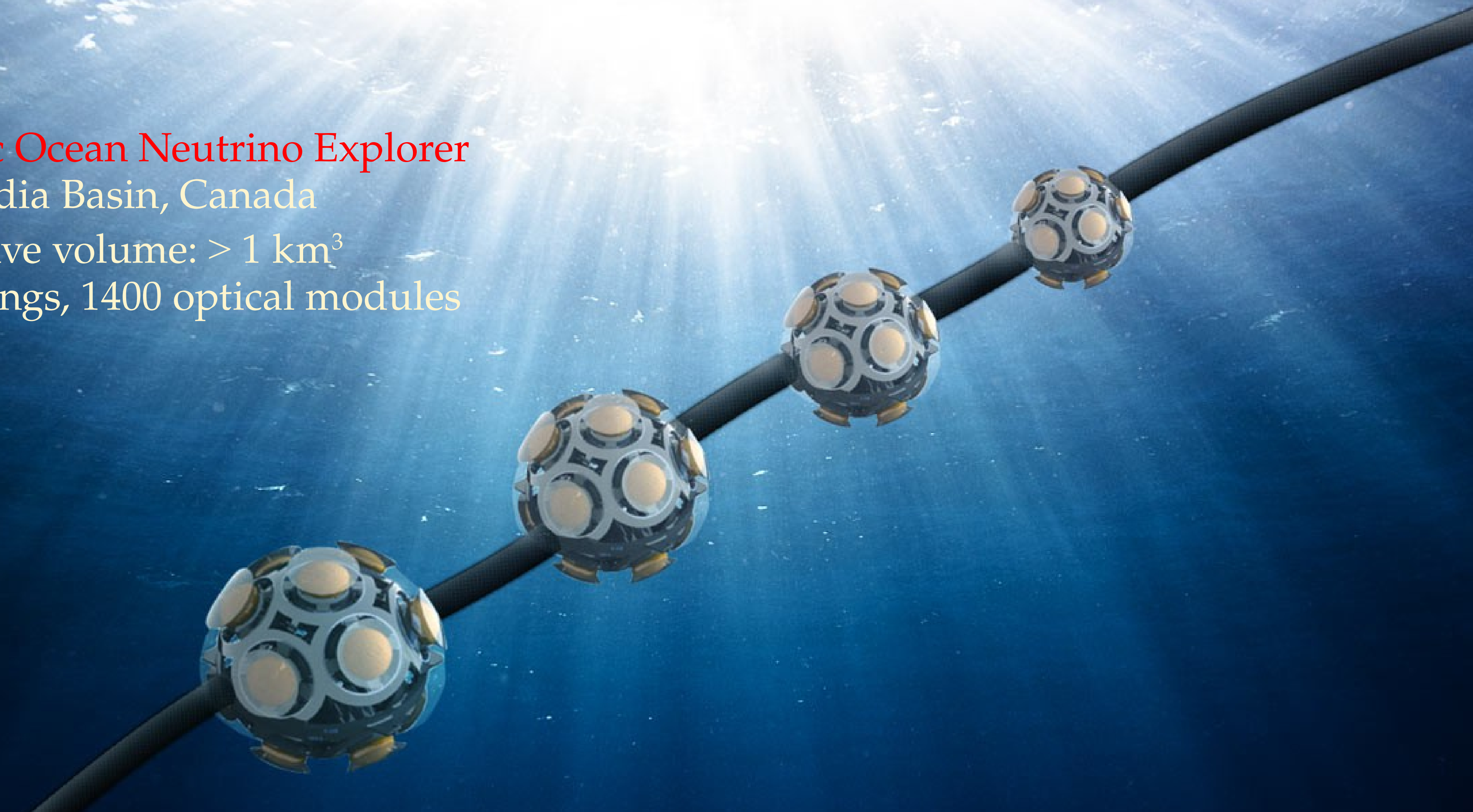
P-ONE

Pacific Ocean Neutrino Explorer

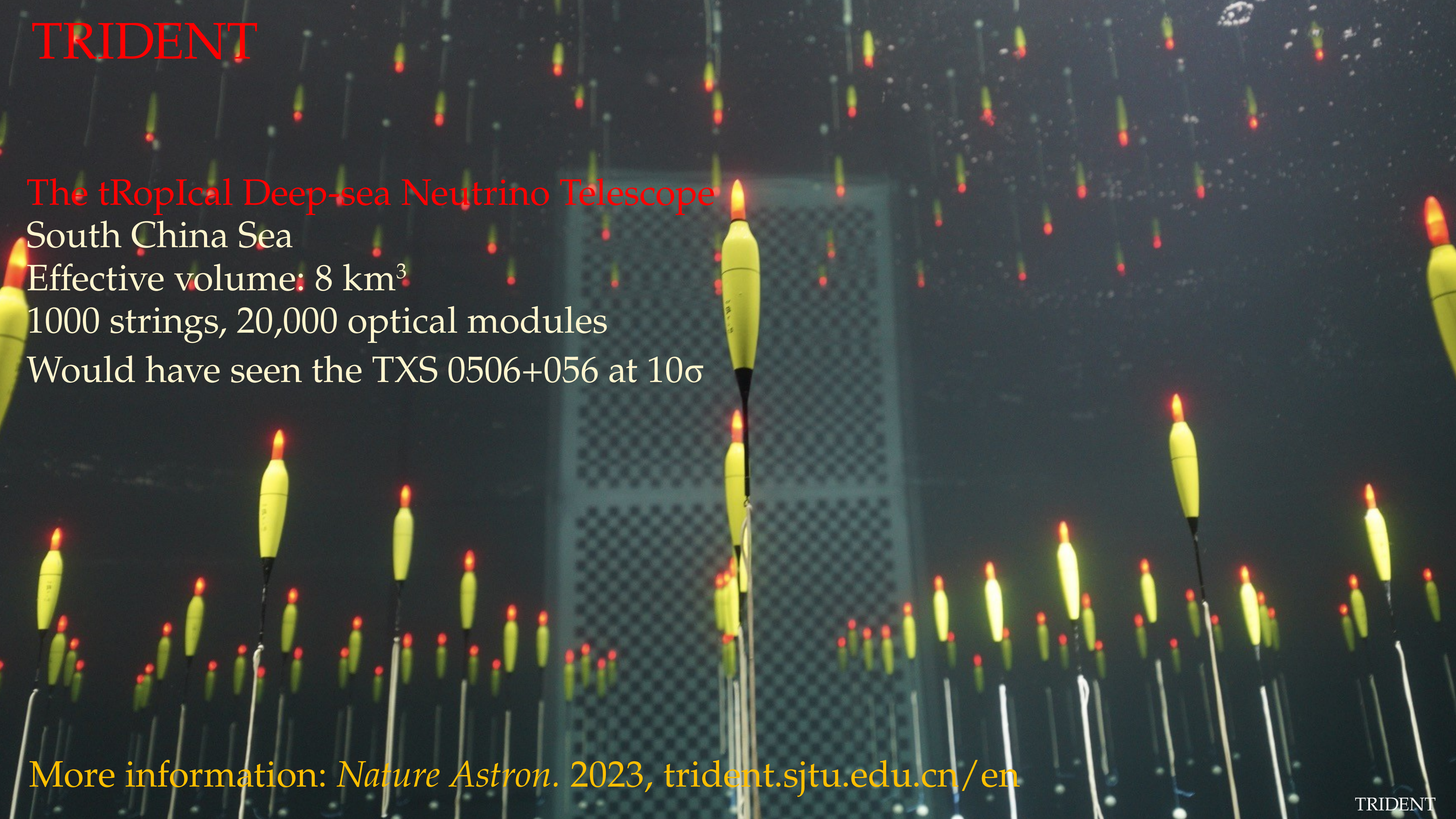
Cascadia Basin, Canada

Effective volume: $> 1\ \mathrm{km}^3$

70 strings, 1400 optical modules



TUM / Marc Blazewicz



TRIDENT

The tRopIcal Deep-sea Neutrino Telescope

South China Sea

Effective volume: 8 km^3

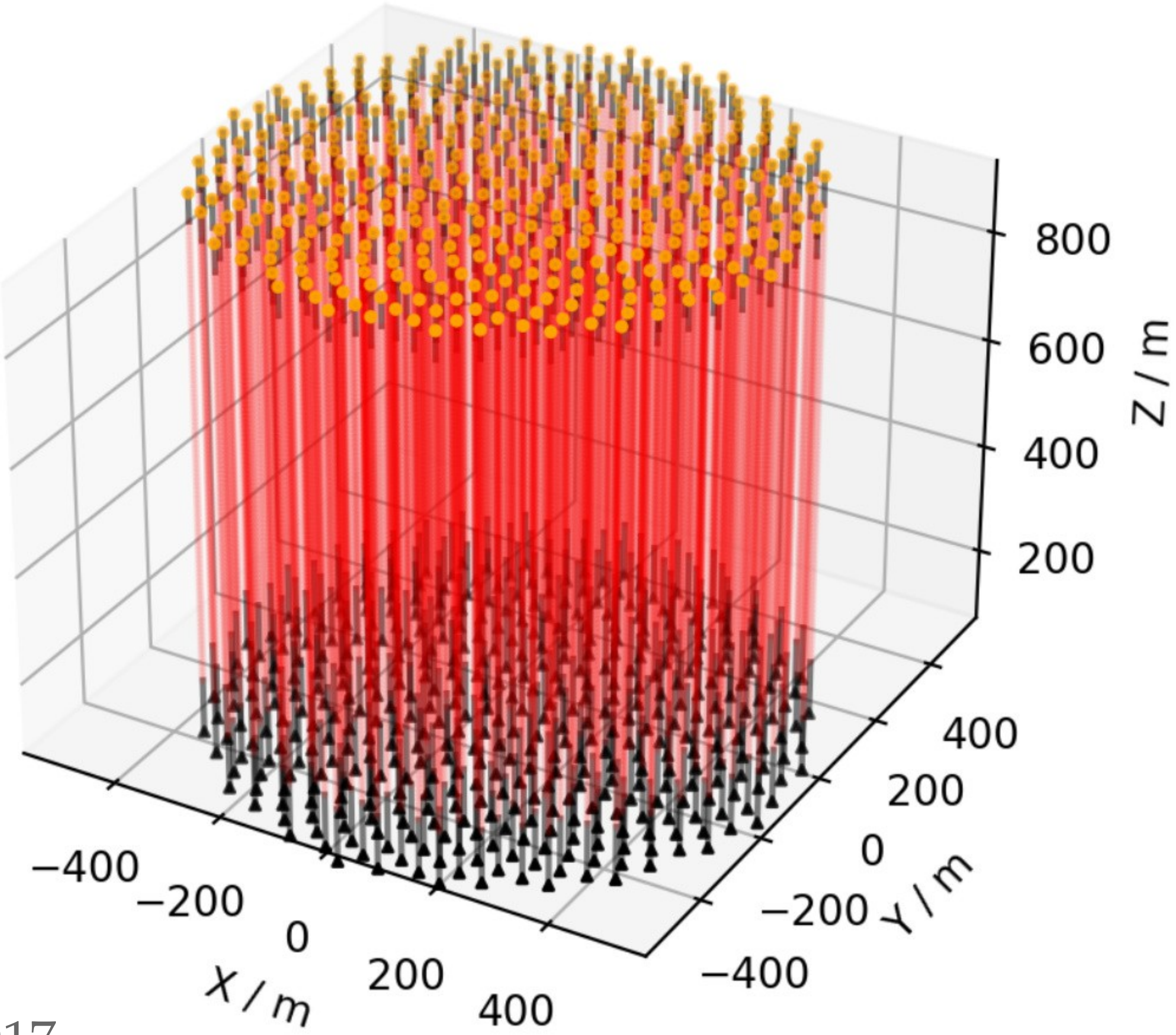
1000 strings, 20,000 optical modules

Would have seen the TXS 0506+056 at 10σ

More information: *Nature Astron.* 2023, trident.sjtu.edu.cn/en

TRIDENT

NEON



Neutrino Observatory in the Nanhai
South China Sea
Effective volume: 0.8 km^3 (dense!)
400 strings, 40,000 optical modules

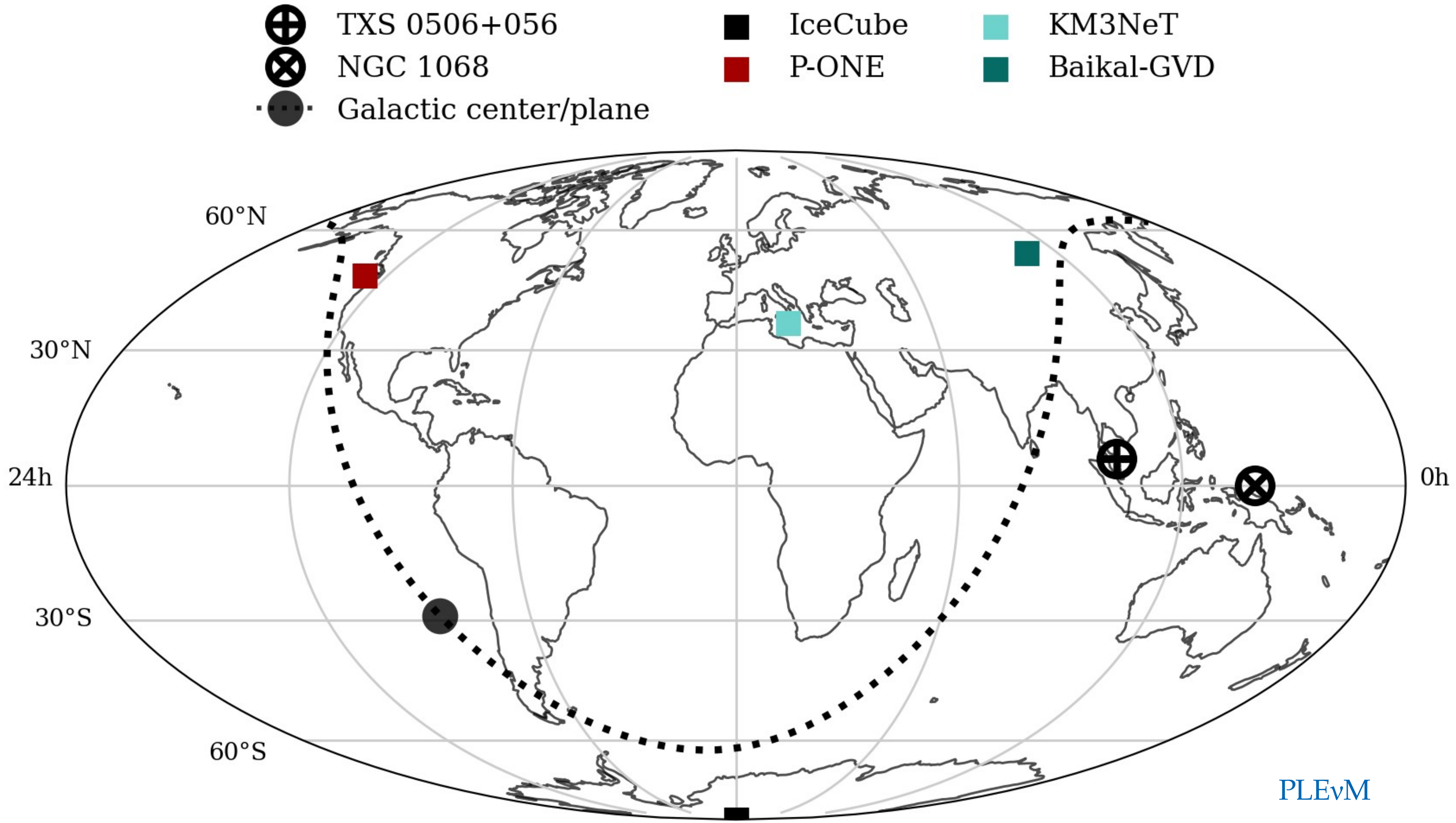
More information: PoS (ICRC2023) 1017

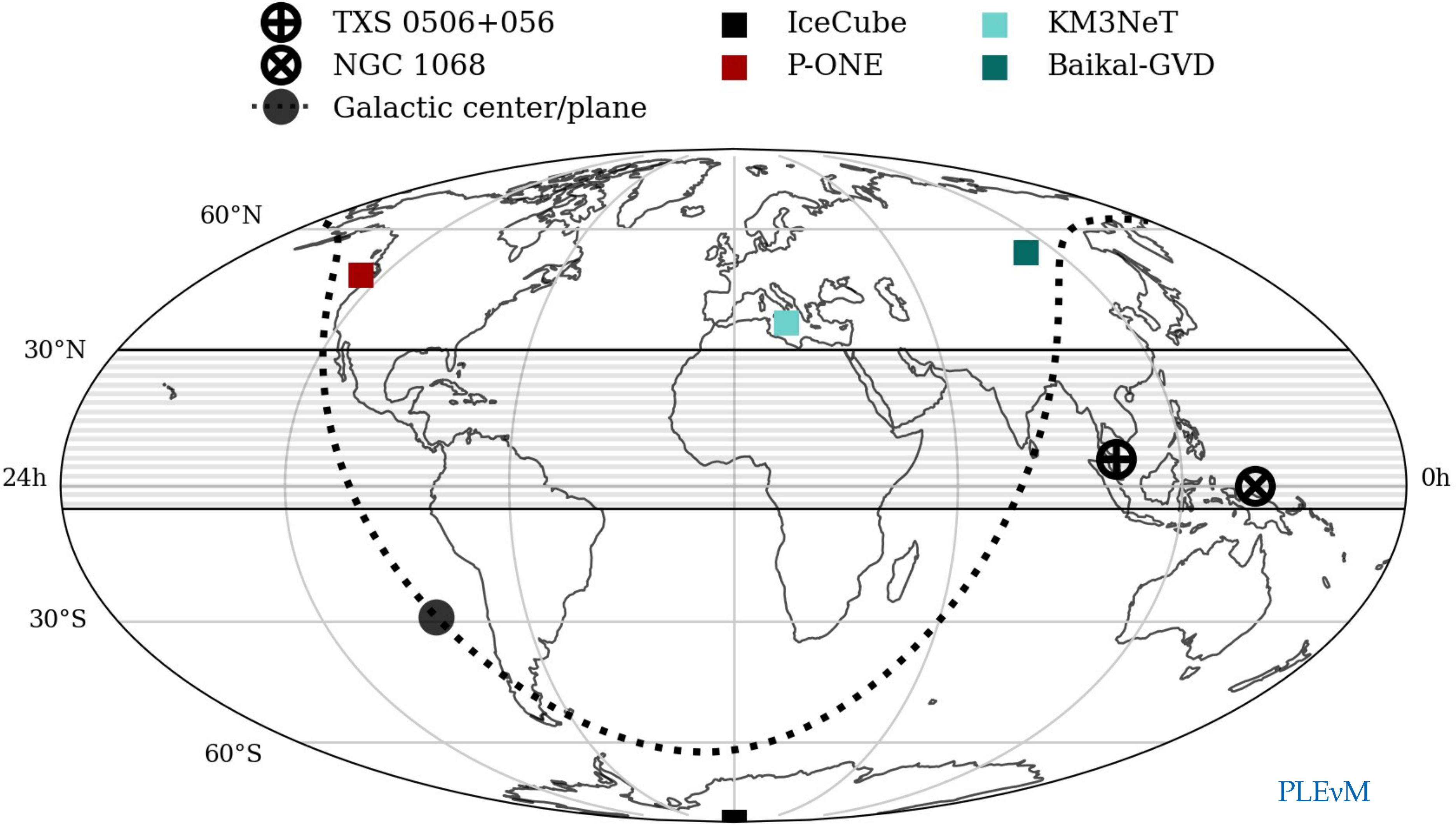
HUNT

High-energy Underwater Neutrino Telescope
South China Sea or Lake Baikal
Effective volume: 30 km^3
2304 strings, 55,296 optical modules

Muon track angular resolution
as good as 0.05°
(for tracks of 6 km in length)

More information: hunt.ihep.ac.cn


TXS 0506+056
NGC 1068
Galactic center/plane
IceCube
P-ONE
KM3NeT
Baikal-GVD
60°N
30°N
24h
0h
30°S
60°S
PLEνM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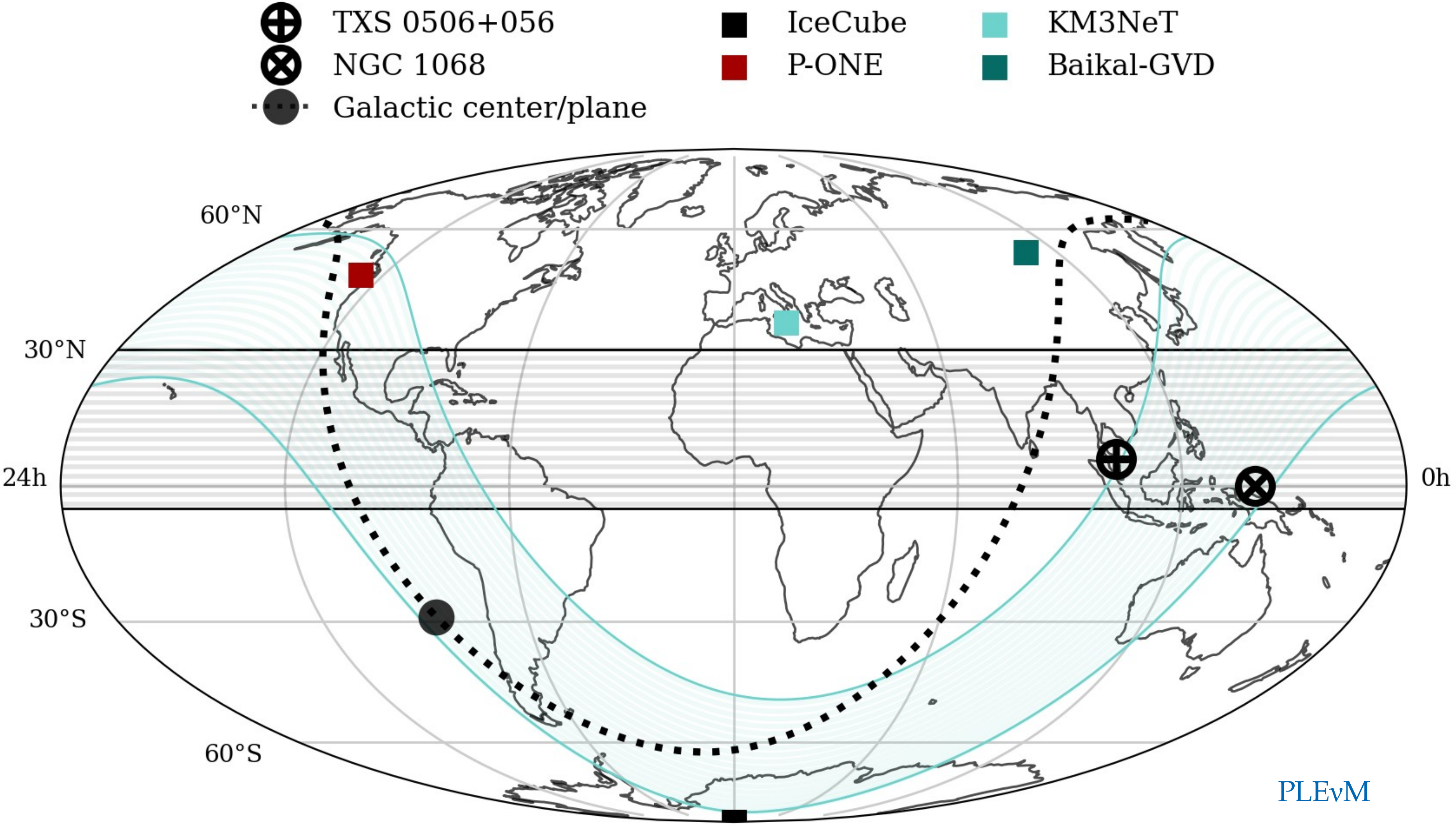
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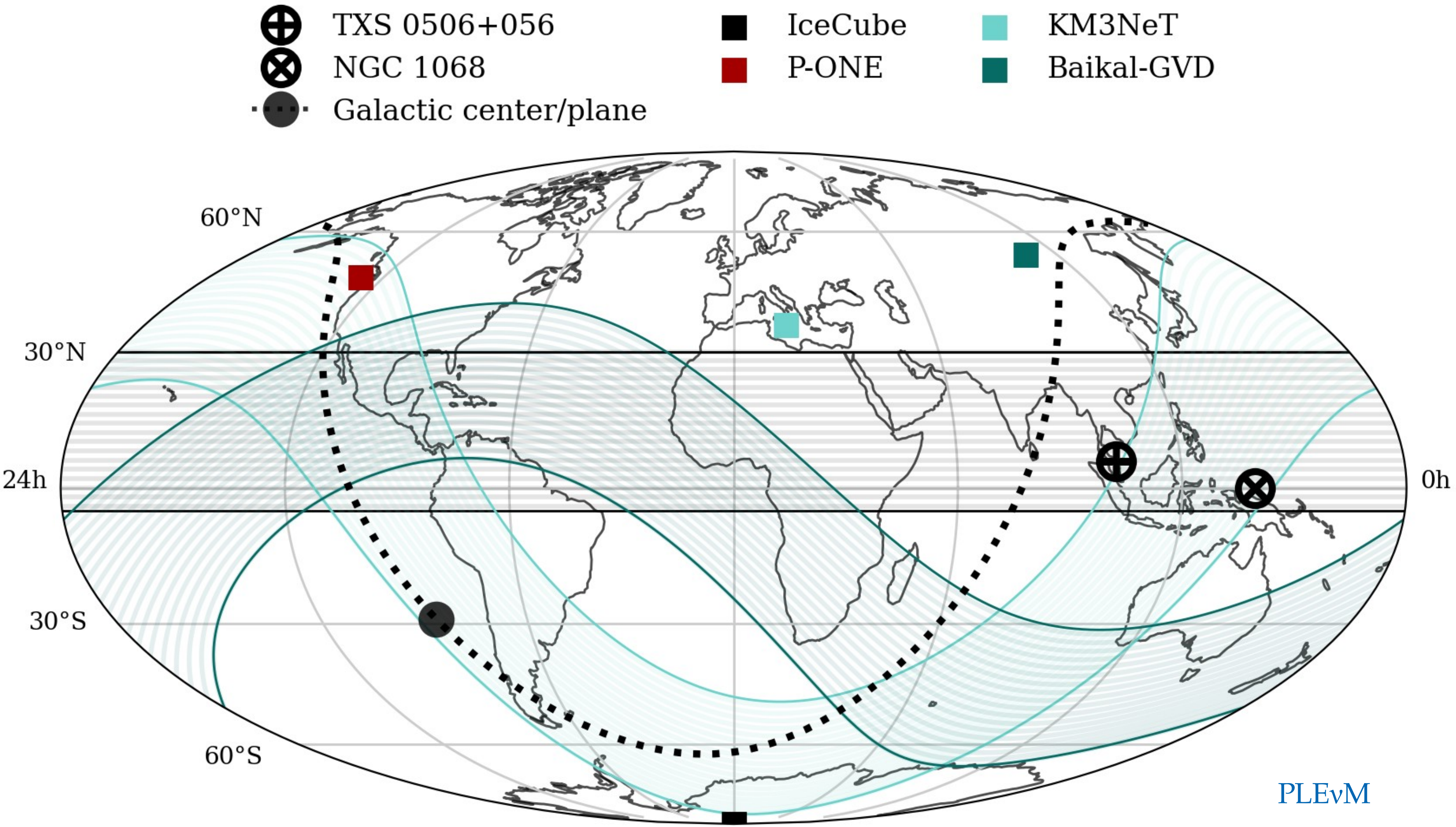
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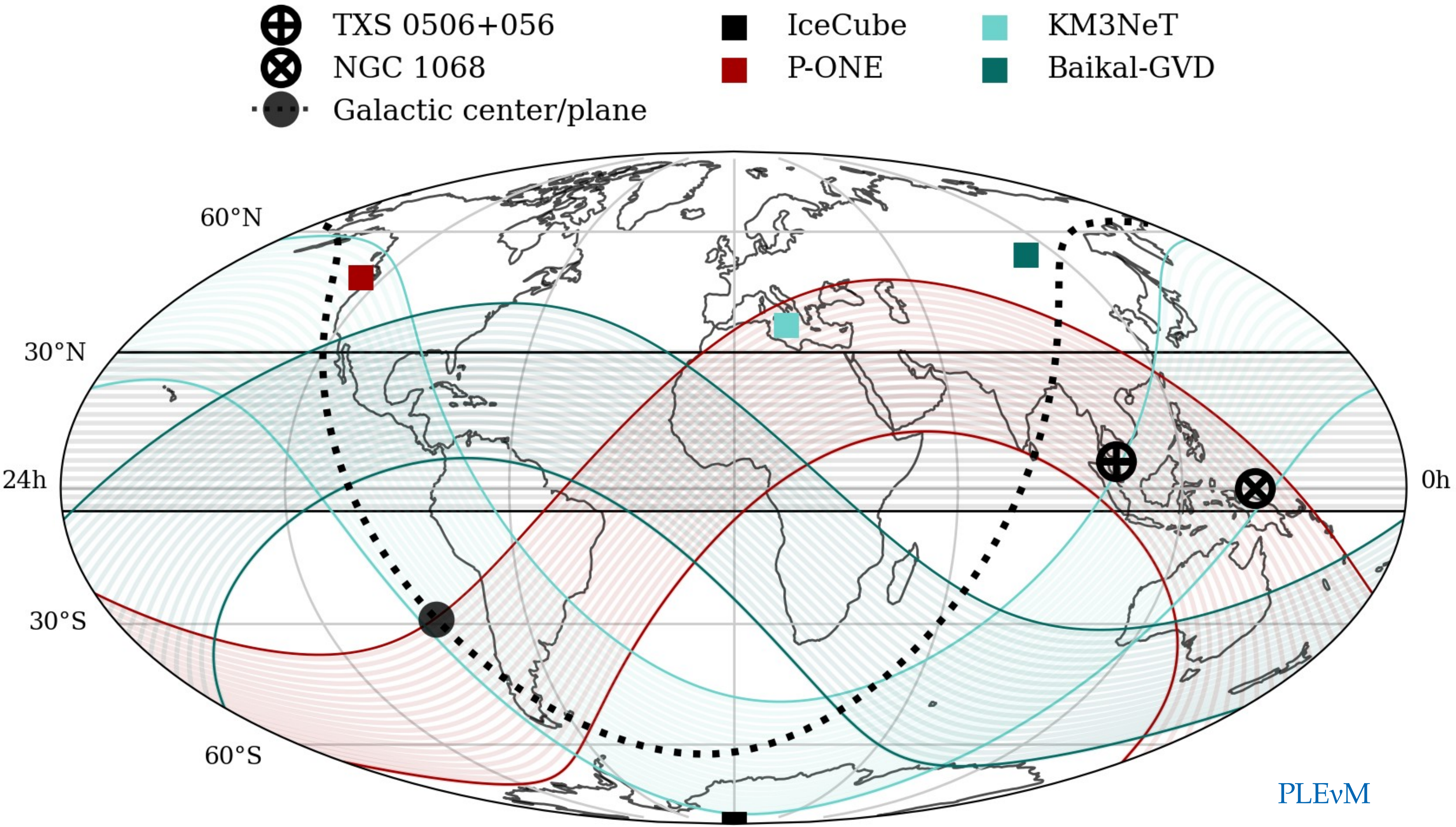
PLEνM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30°S
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Today

TeV–PeV ν

Turn predictions
into data-driven tests

Key developments:

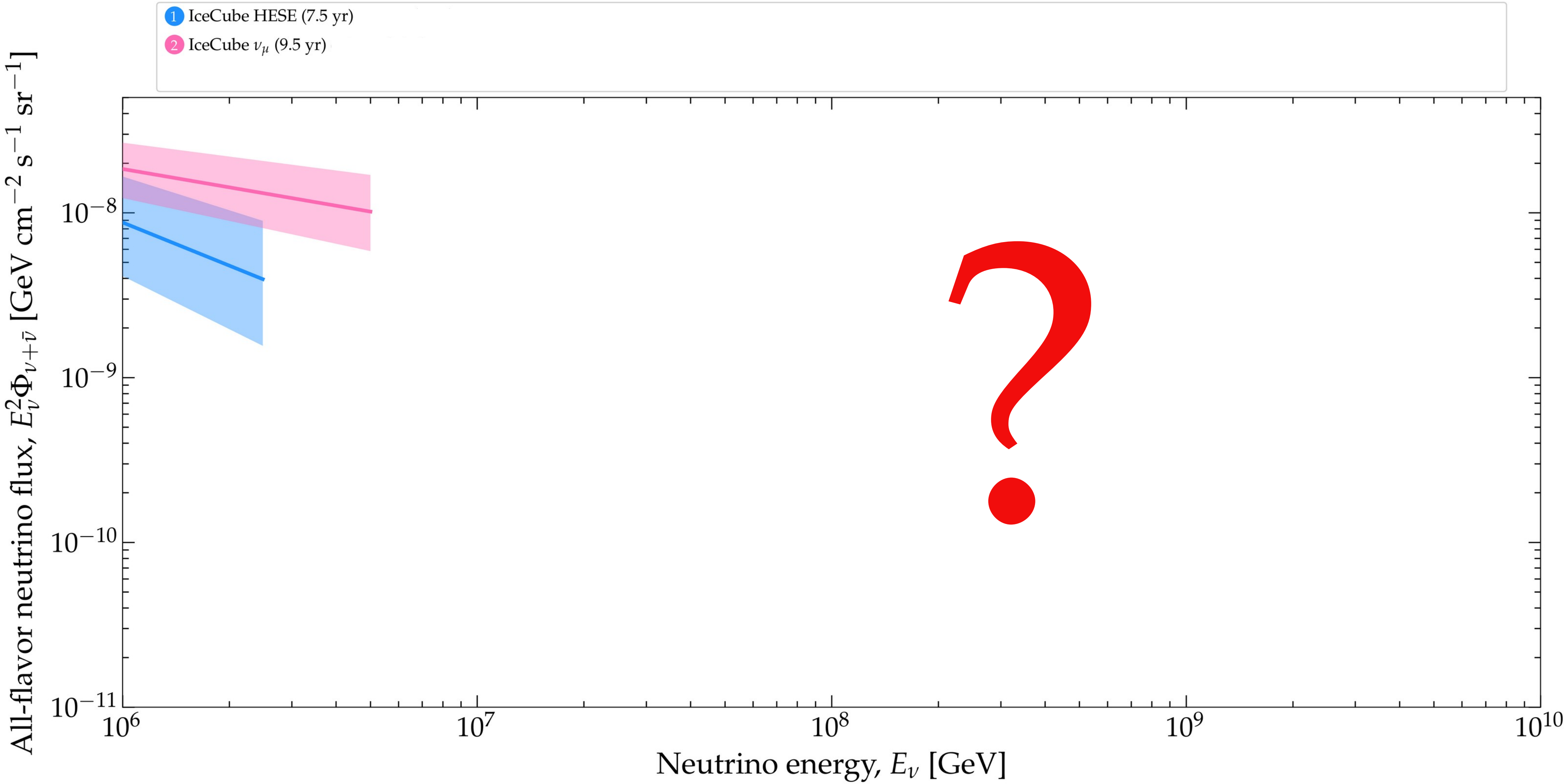
Bigger detectors → larger statistics

Better reconstruction

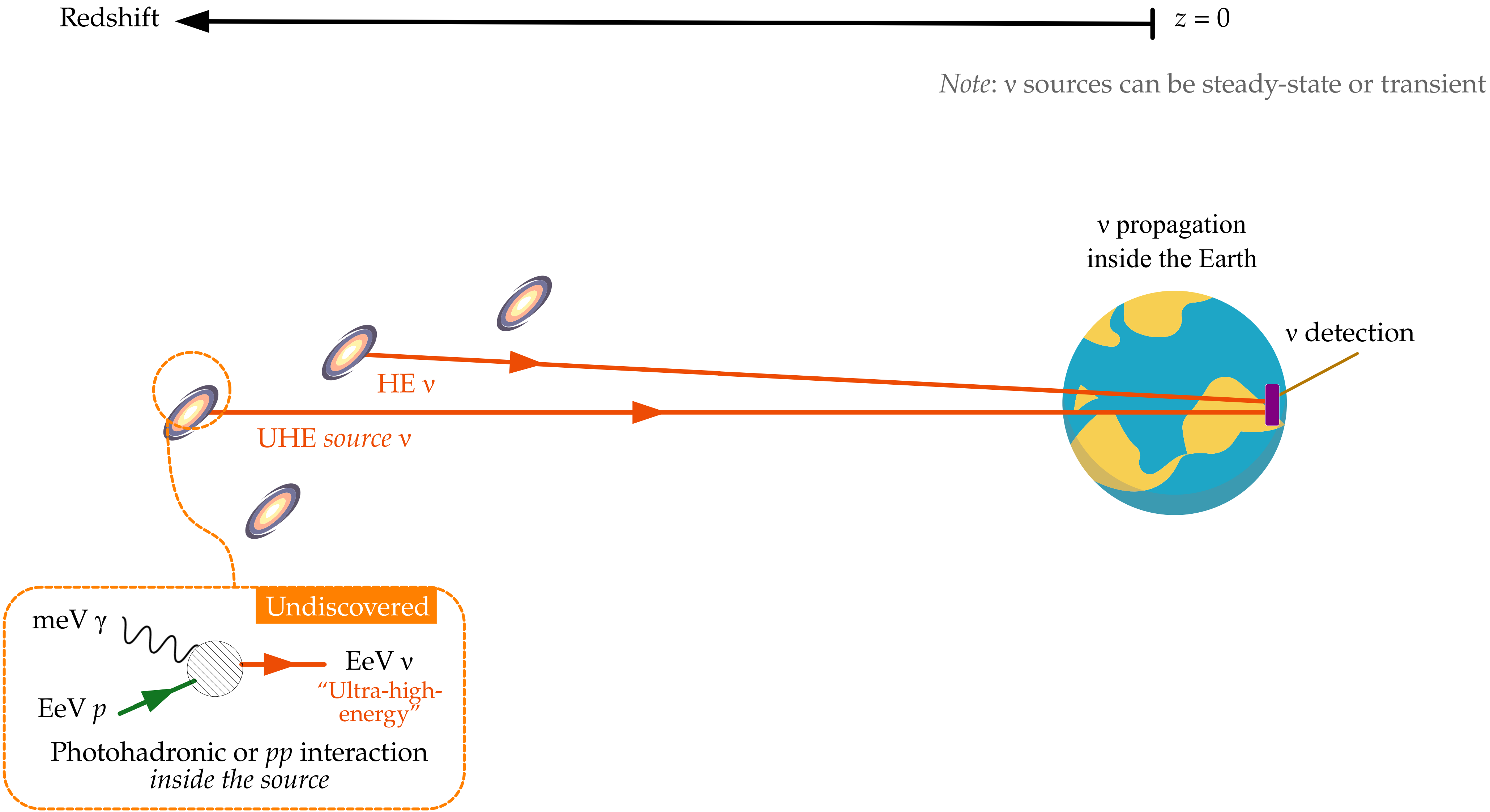
Smaller astrophysical uncertainties

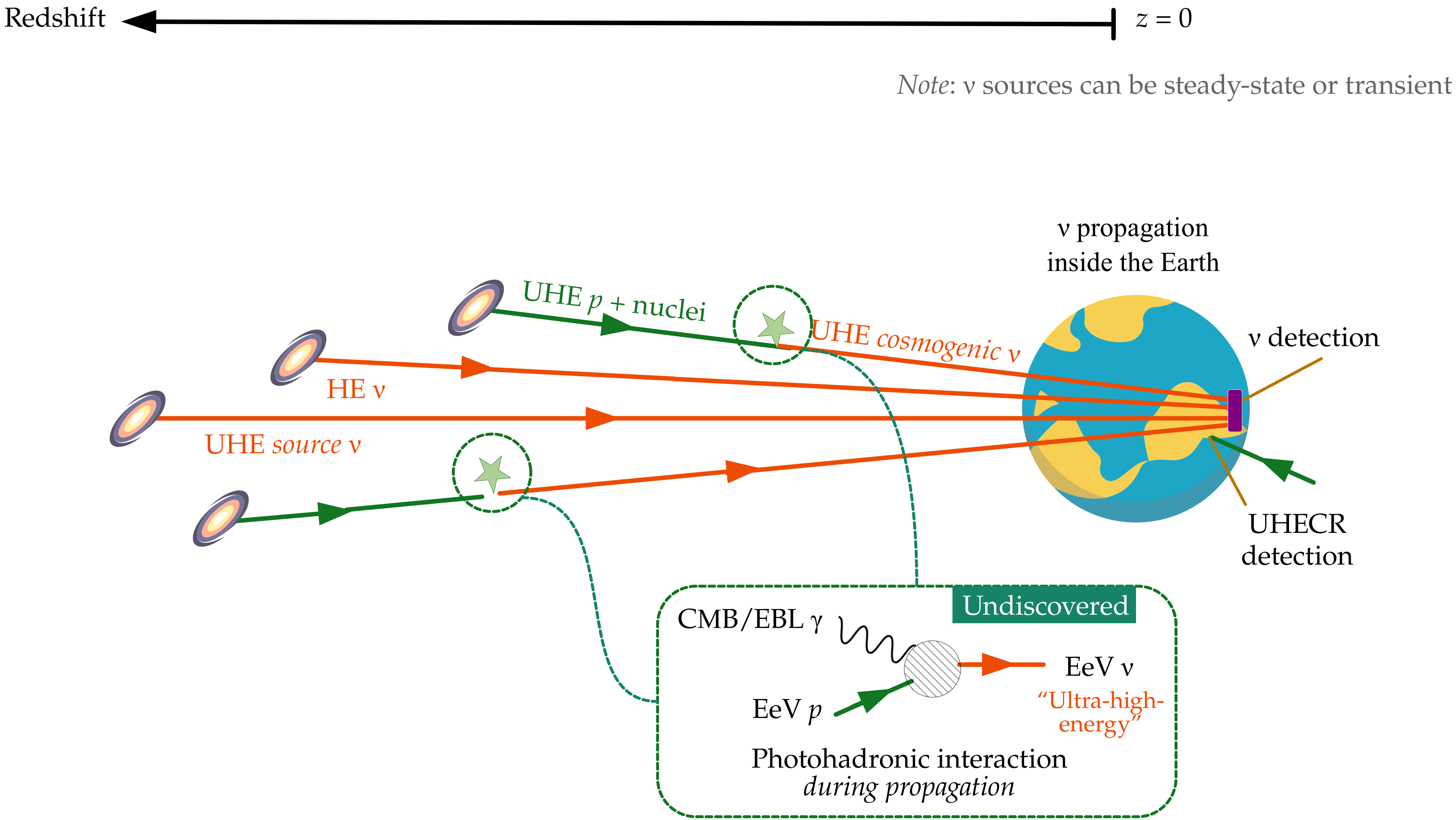
Next decade

> 100-PeV ν

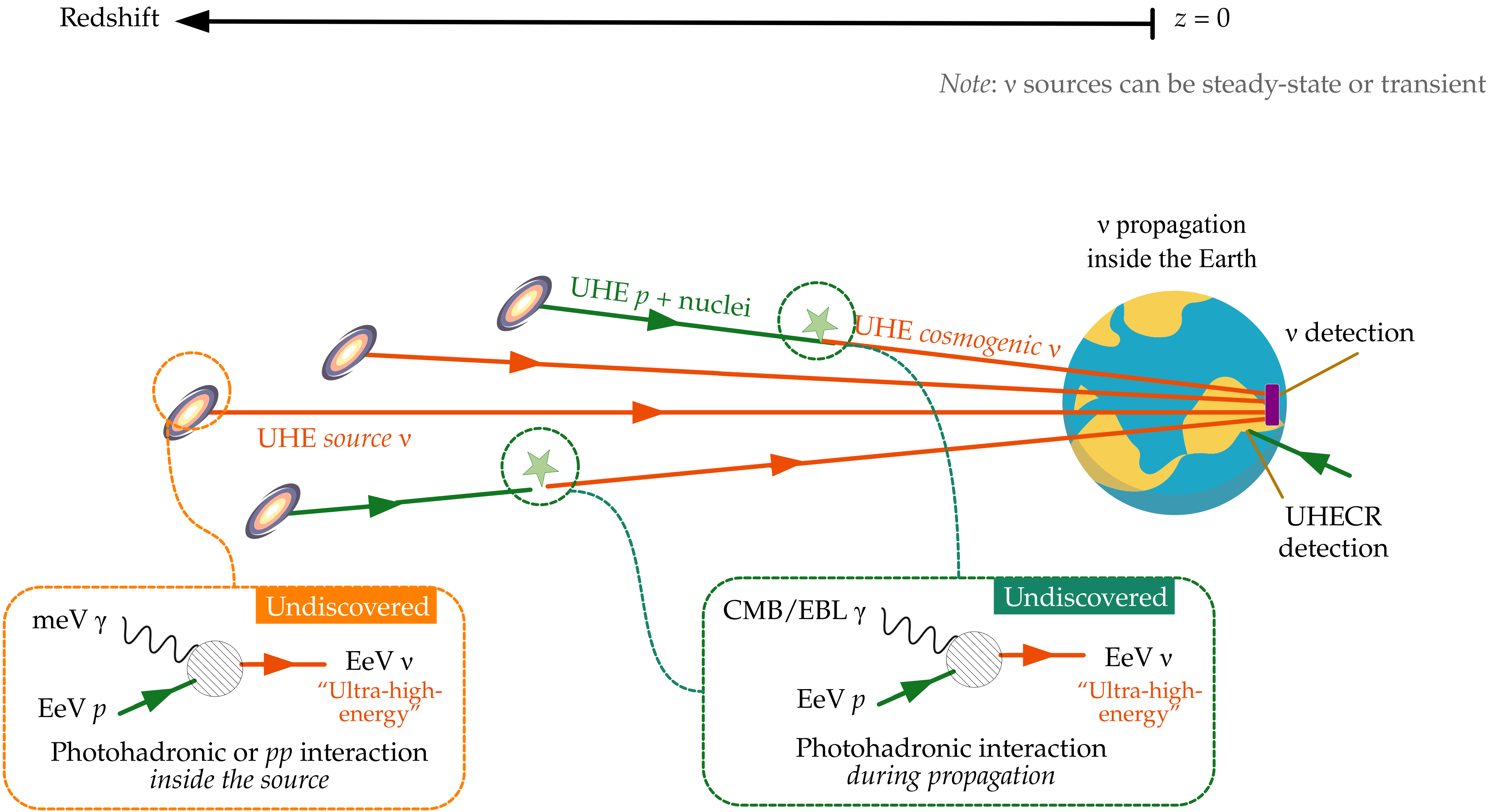


1 IceCube HESE (7.5 yr)
2 IceCube ν_μ (9.5 yr)
All-flavor neutrino flux, $E_\nu^2 \Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^6
10^7
10^8
10^9
10^{10}
Neutrino energy, E_ν [GeV]
?

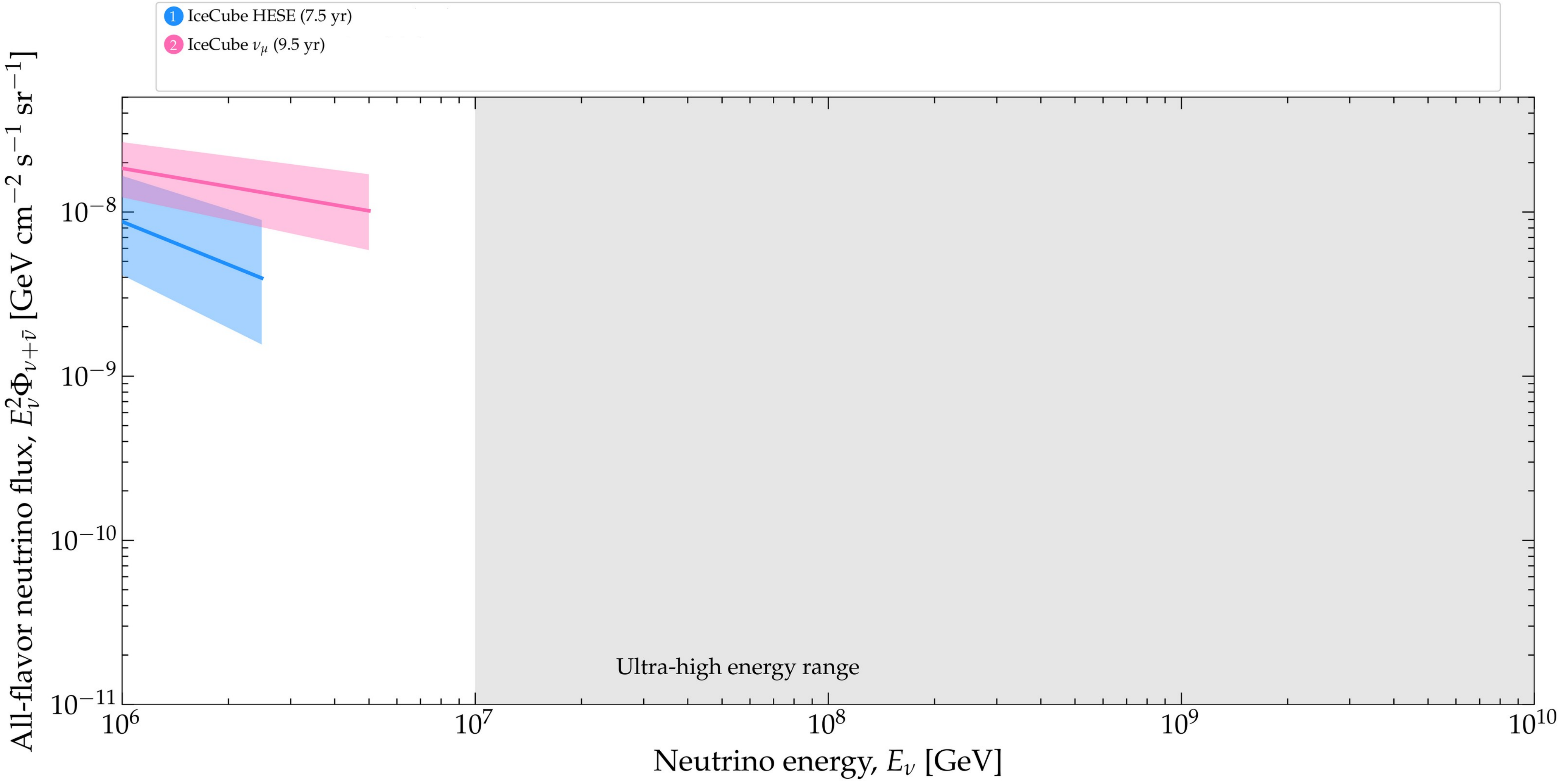

Redshift
$z = 0$
Note: ν sources can be steady-state or transient
ν propagation inside the Earth
ν detection
HE ν
UHE *source* ν
Undiscovered
meV γ
EeV ν
"Ultra-high-energy"
EeV *p*
Photohadronic or *pp* interaction *inside the source*

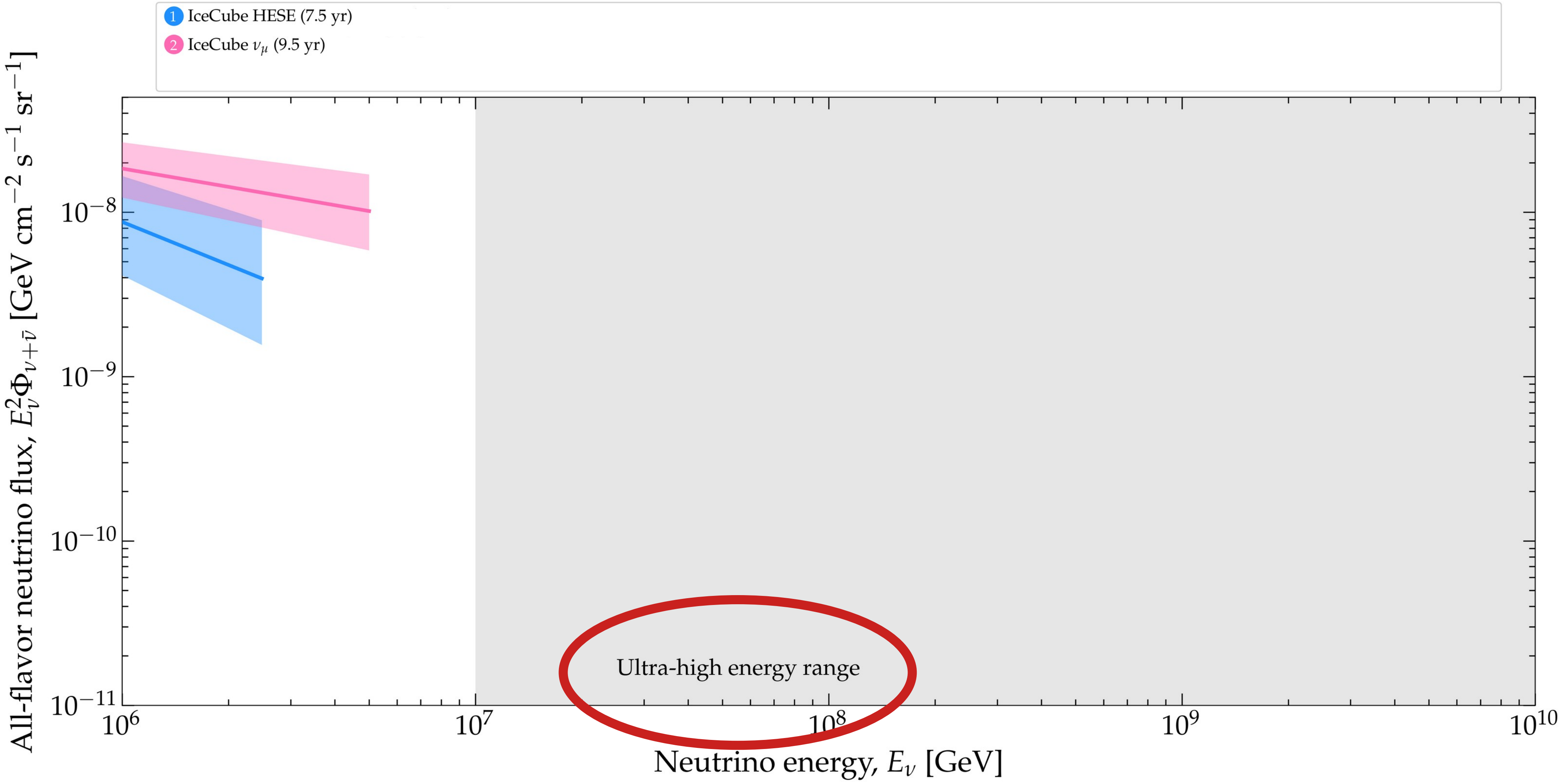


Redshift
$z = 0$
Note: ν sources can be steady-state or transient
ν propagation inside the Earth
UHE p + nuclei
UHE *cosmogenic* ν
HE ν
UHE *source* ν
ν detection
UHECR detection
Undiscovered
CMB/EBL γ
EeV ν
EeV p
"Ultra-high-energy"
Photohadronic interaction *during propagation*



Redshift
z = 0
Note: ν sources can be steady-state or transient
ν propagation
inside the Earth
UHE p + nuclei
UHE cosmogenic ν
ν detection
UHE source ν
UHECR
detection
Undiscovered
meV γ
EeV ν
"Ultra-high-
energy"
EeV p
Photohadronic or pp interaction
inside the source
Undiscovered
CMB/EBL γ
EeV ν
"Ultra-high-
energy"
EeV p
Photohadronic interaction
during propagation


1 IceCube HESE (7.5 yr)
2 IceCube ν_μ (9.5 yr)
All-flavor neutrino flux, $E_\nu^2 \Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
10^{-8}
10^{-9}
10^{-10}
10^{-11}
Ultra-high energy range
10^6
10^7
10^8
10^9
10^{10}
Neutrino energy, E_ν [GeV]



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Ultra-high energy range
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^{6}
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Neutrino energy, E_ν [GeV]
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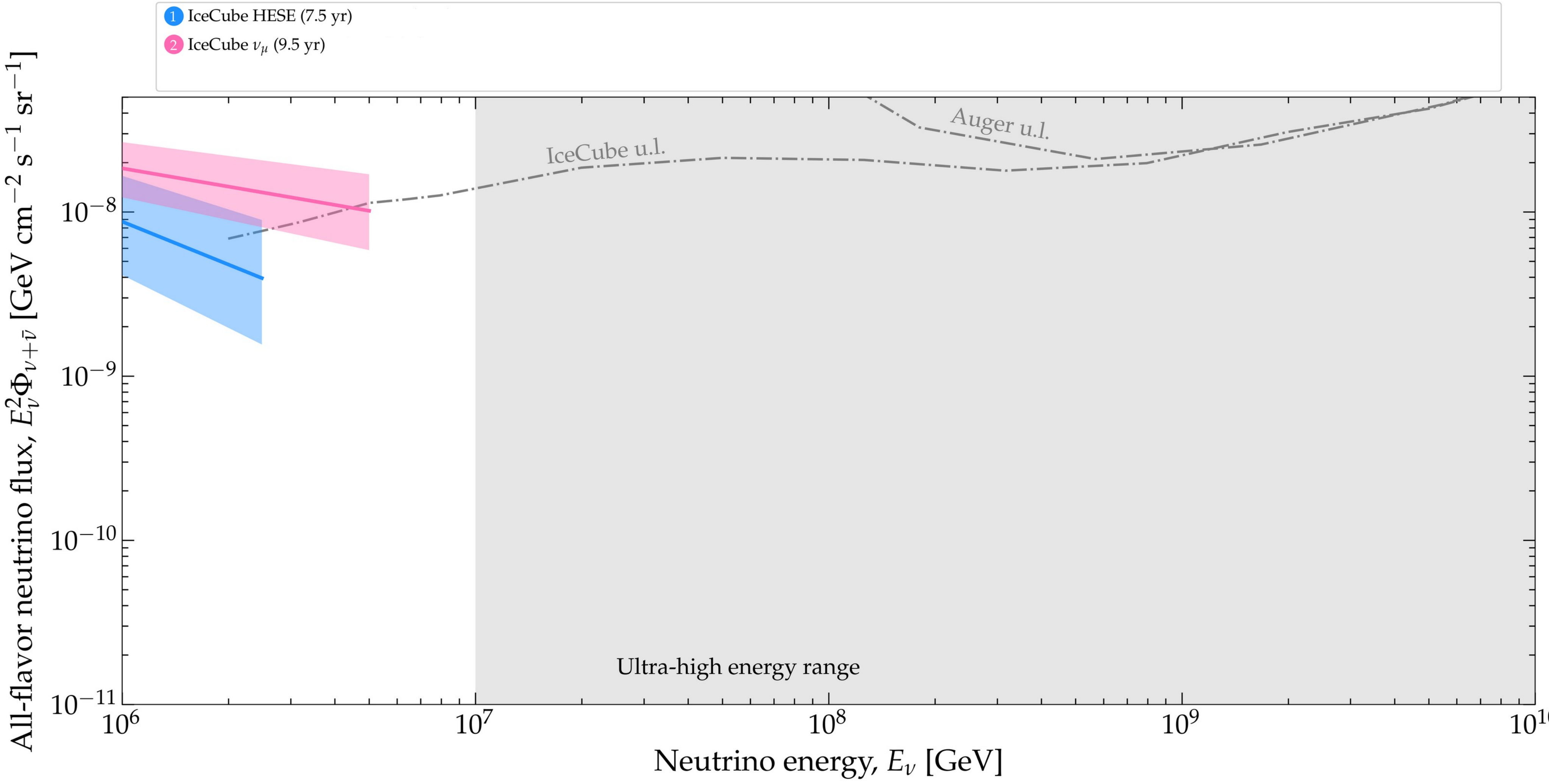

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All-flavor neutrino flux, $E_\nu^2 \Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
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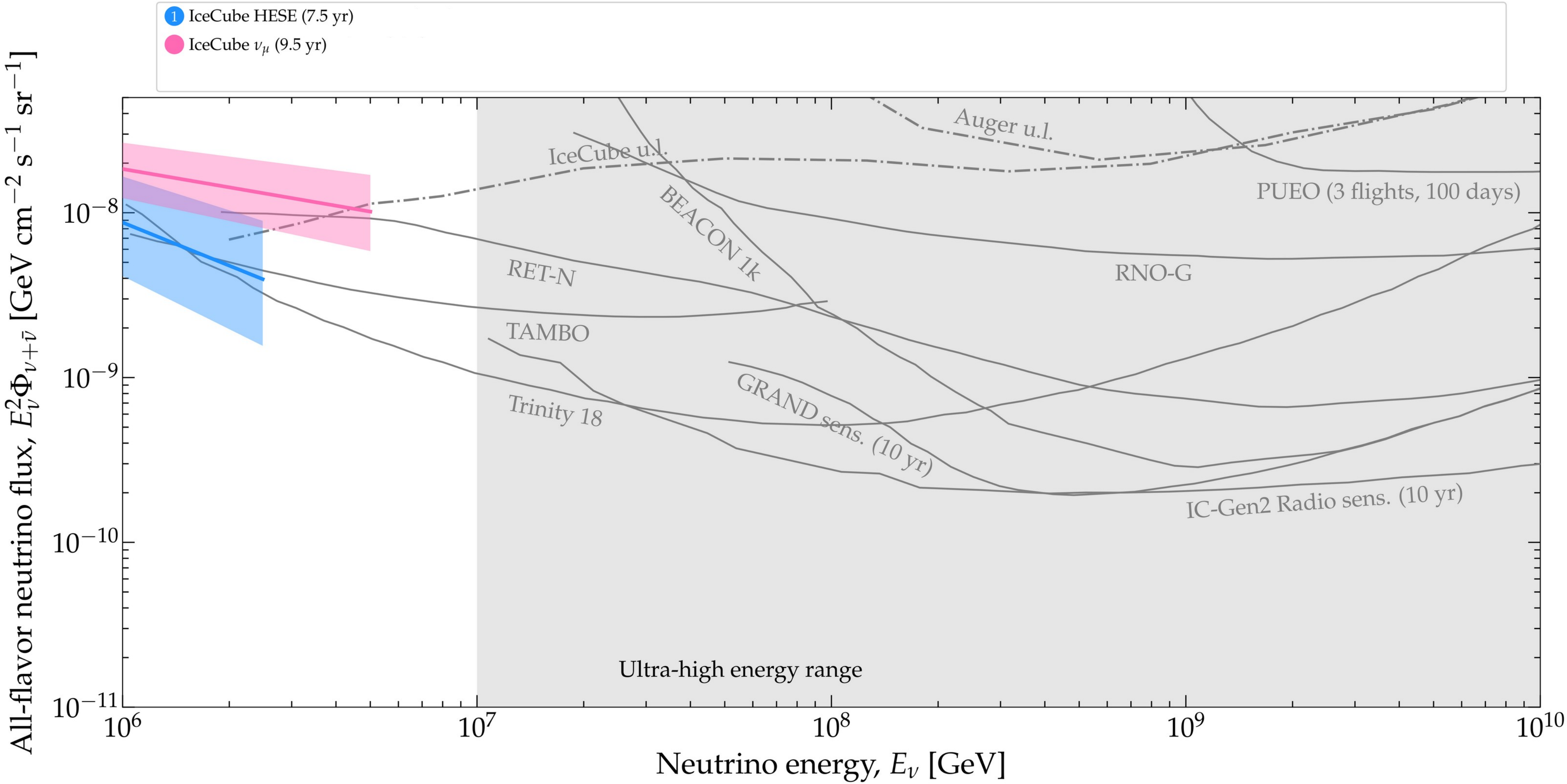
10^{-9}
10^{-10}
10^{-11}
IceCube u.l.
Auger u.l.
Ultra-high energy range
10^6
10^7
10^8
10^9
10^{10}
Neutrino energy, E_ν [GeV]



1 IceCube HESE (7.5 yr)
IceCube ν_μ (9.5 yr)
Auger u.l.
IceCube u.l.
BEACON 1k
PUEO (3 flights, 100 days)
RET-N
RNO-G
TAMBO
GRAND sens. (10 yr)
Trinity 18
IC-Gen2 Radio sens. (10 yr)
Ultra-high energy range
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^{6}
10^{7}
10^{8}
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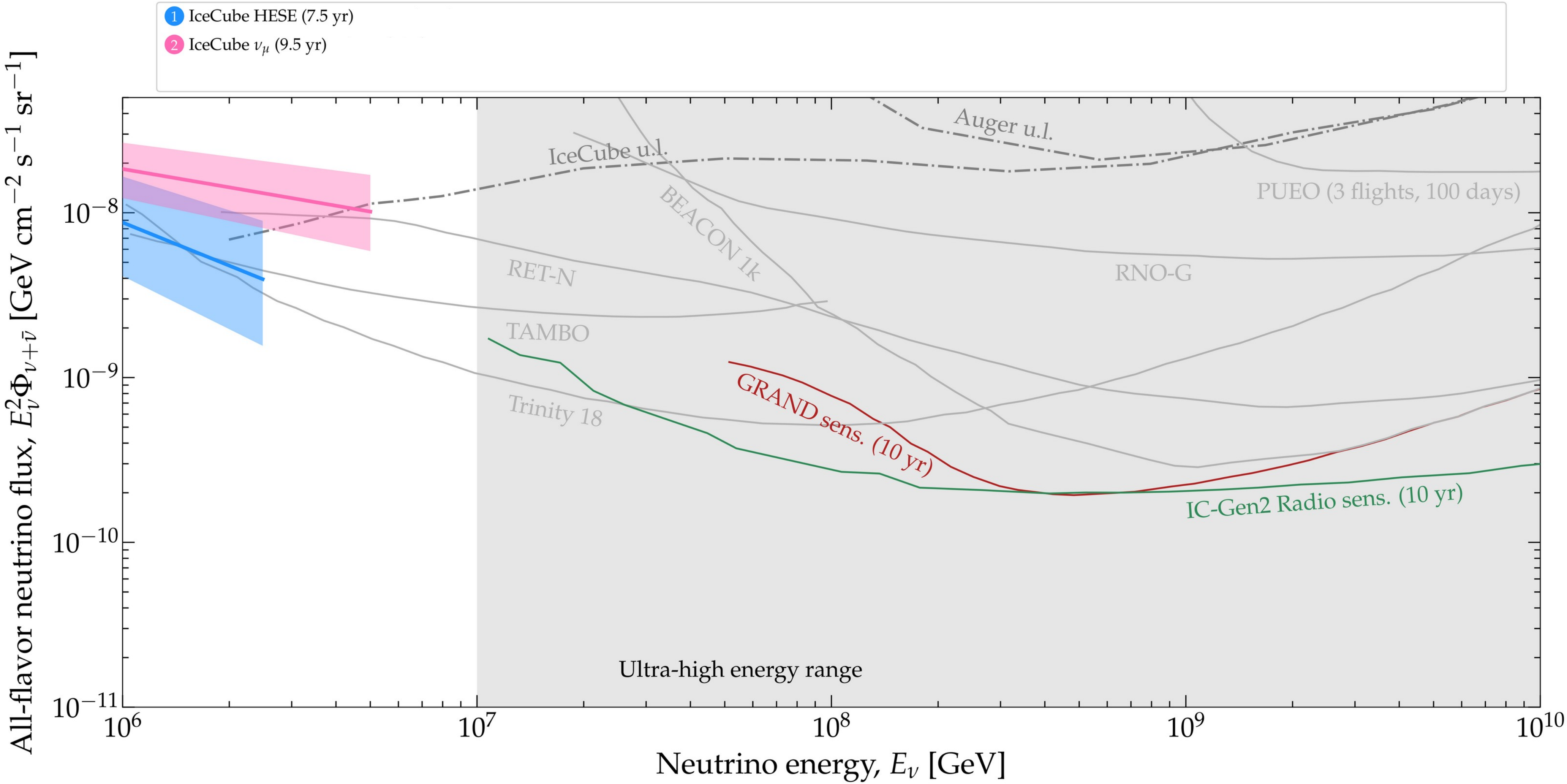
10^{10}
Neutrino energy, E_ν [GeV]
All-flavor neutrino flux, $E_\nu^2\Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]



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10^{-8}
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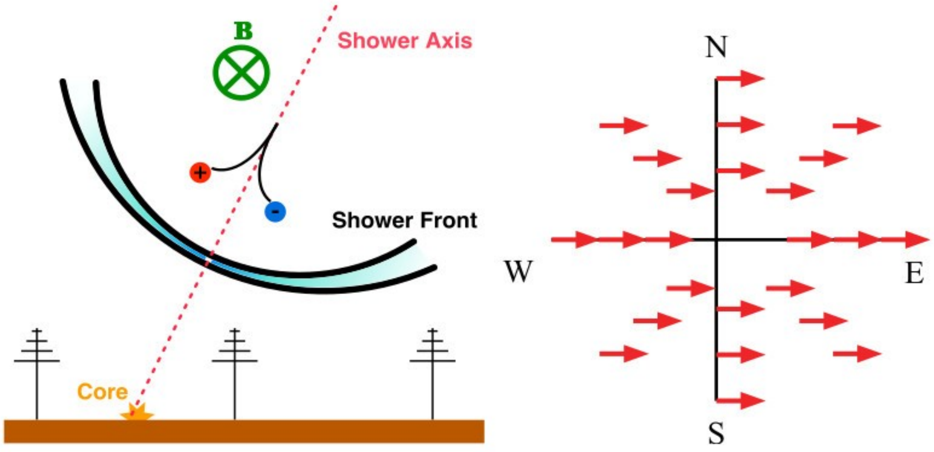
TAMBO
Trinity 18
GRAND sens. (10 yr)
IC-Gen2 Radio sens. (10 yr)
Ultra-high energy range

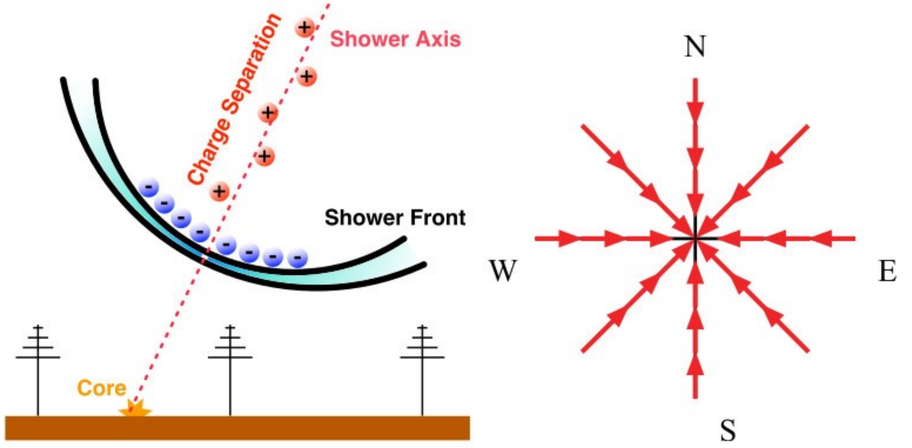
Radio emission: geomagnetic and Askaryan

Geomagnetic



- Time-varying transverse current
- Linearly polarized parallel to Lorentz force
- Dominant in air showers

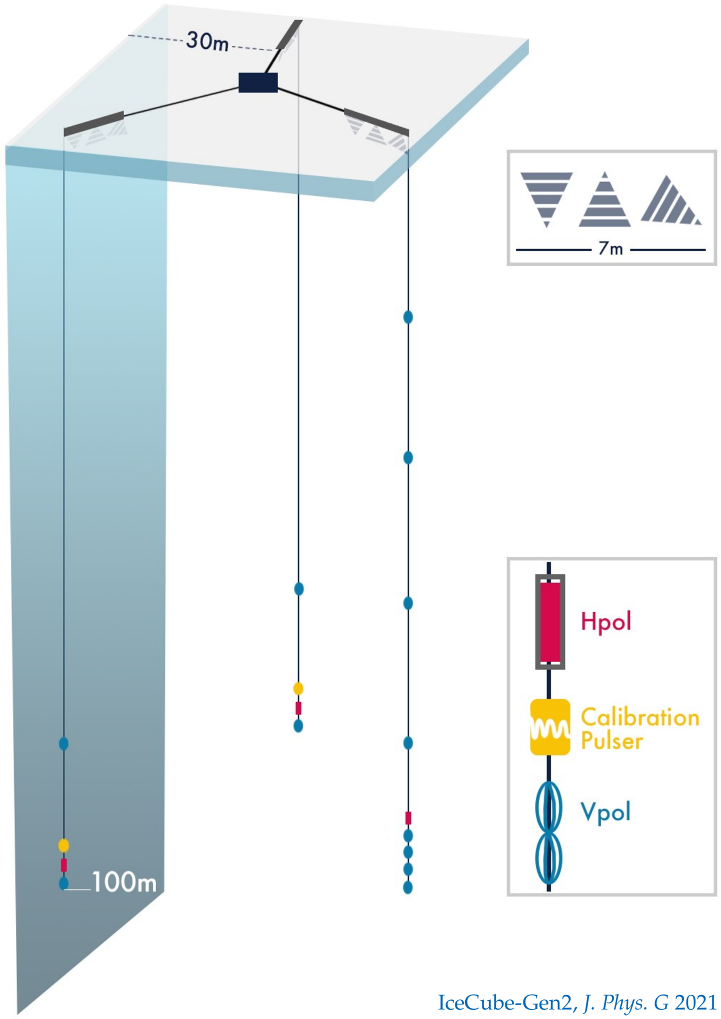
Askaryan



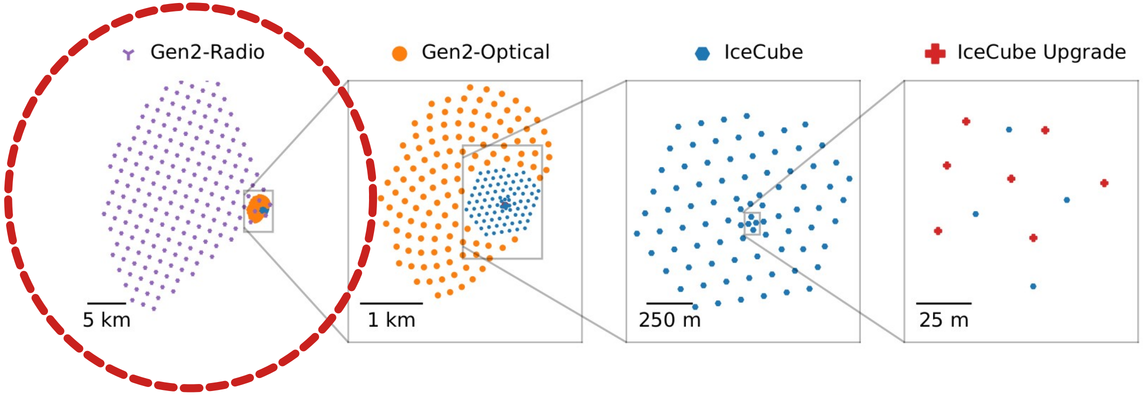
- Time-varying negative-charge ~20% excess
- Linearly polarized towards axis
- Sub-dominant in air showers

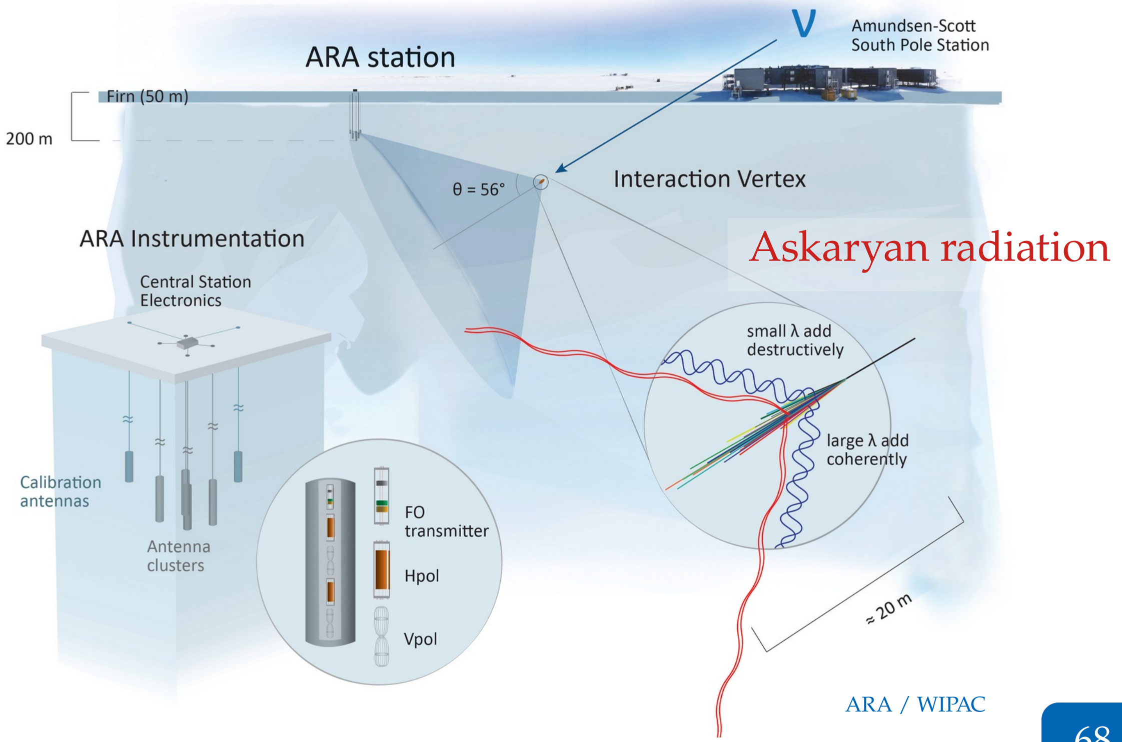
Figures by H. Schoorlemmer and K. D. de Vries

IceCube-Gen2 Radio

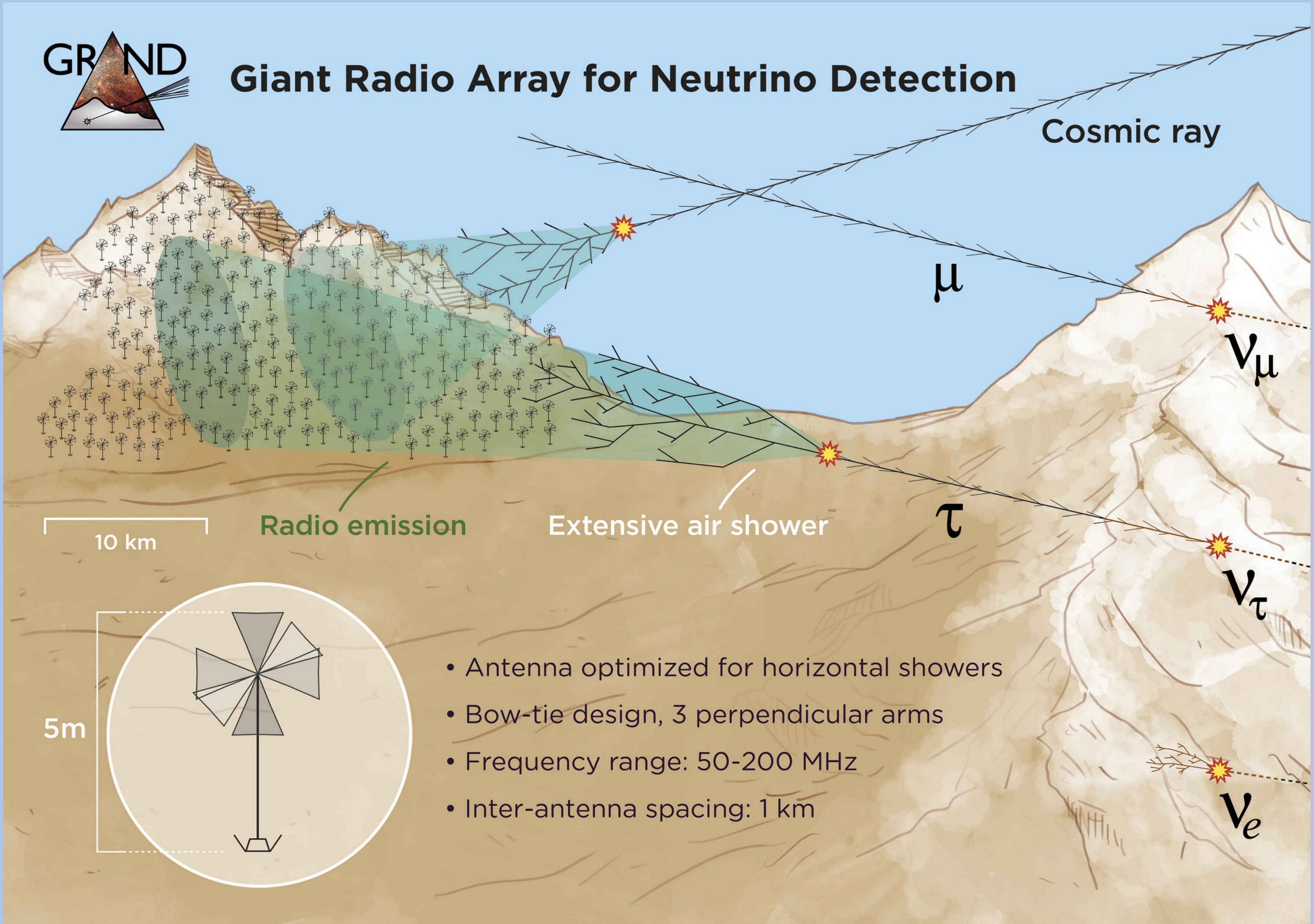


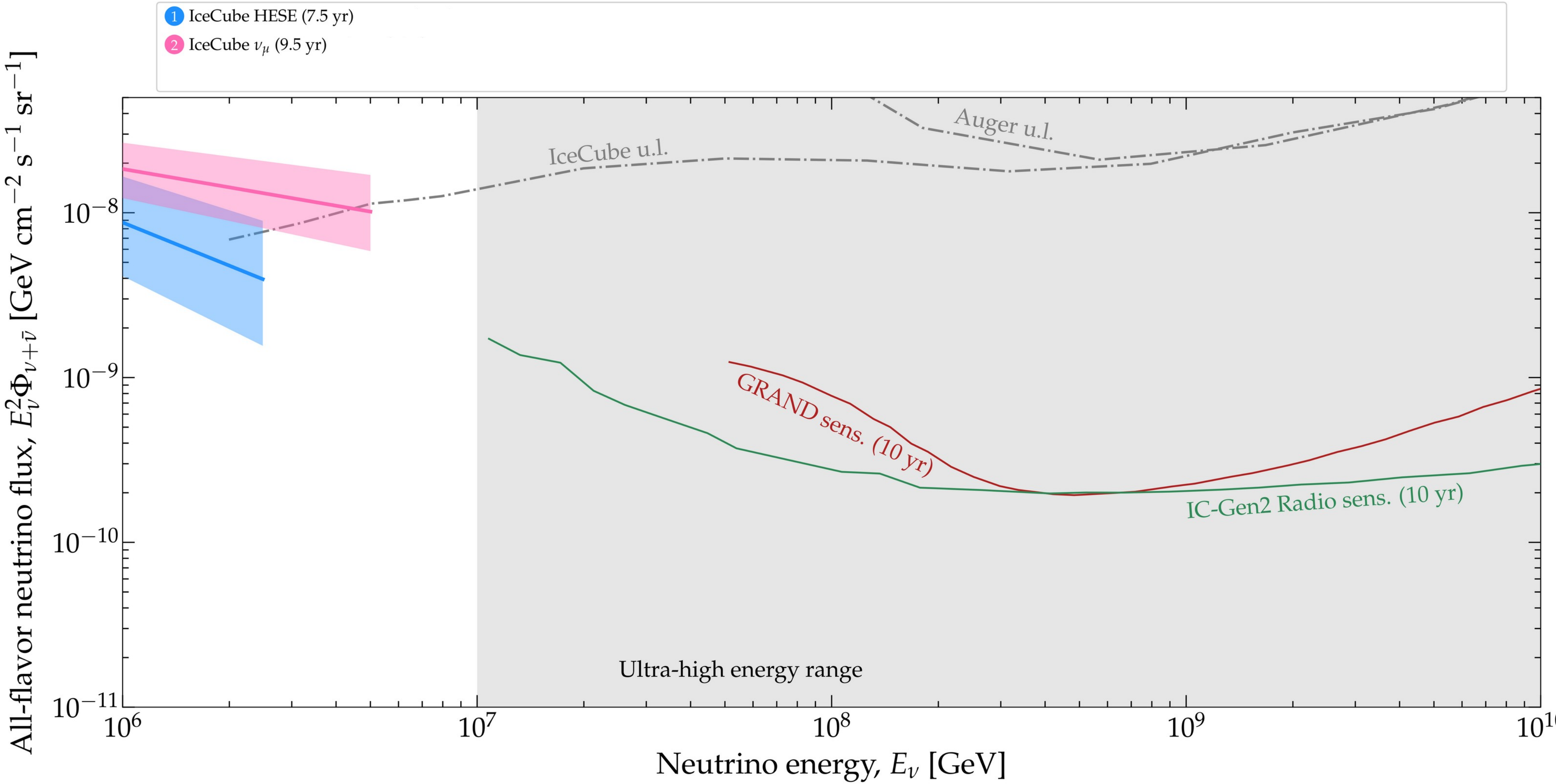
IceCube-Gen2, *J. Phys. G* 2021





ARA / WIPAC


GRAND
Giant Radio Array for Neutrino Detection
Cosmic ray
μ
ν_μ
τ
ν_τ
ν_e
10 km
Radio emission
Extensive air shower
5m
• Antenna optimized for horizontal showers
• Bow-tie design, 3 perpendicular arms
• Frequency range: 50-200 MHz
• Inter-antenna spacing: 1 km


IceCube HESE (7.5 yr)
IceCube ν_μ (9.5 yr)
IceCube u.l.
Auger u.l.
GRAND sens. (10 yr)
IC-Gen2 Radio sens. (10 yr)
Ultra-high energy range
All-flavor neutrino flux, $E_\nu^2 \Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
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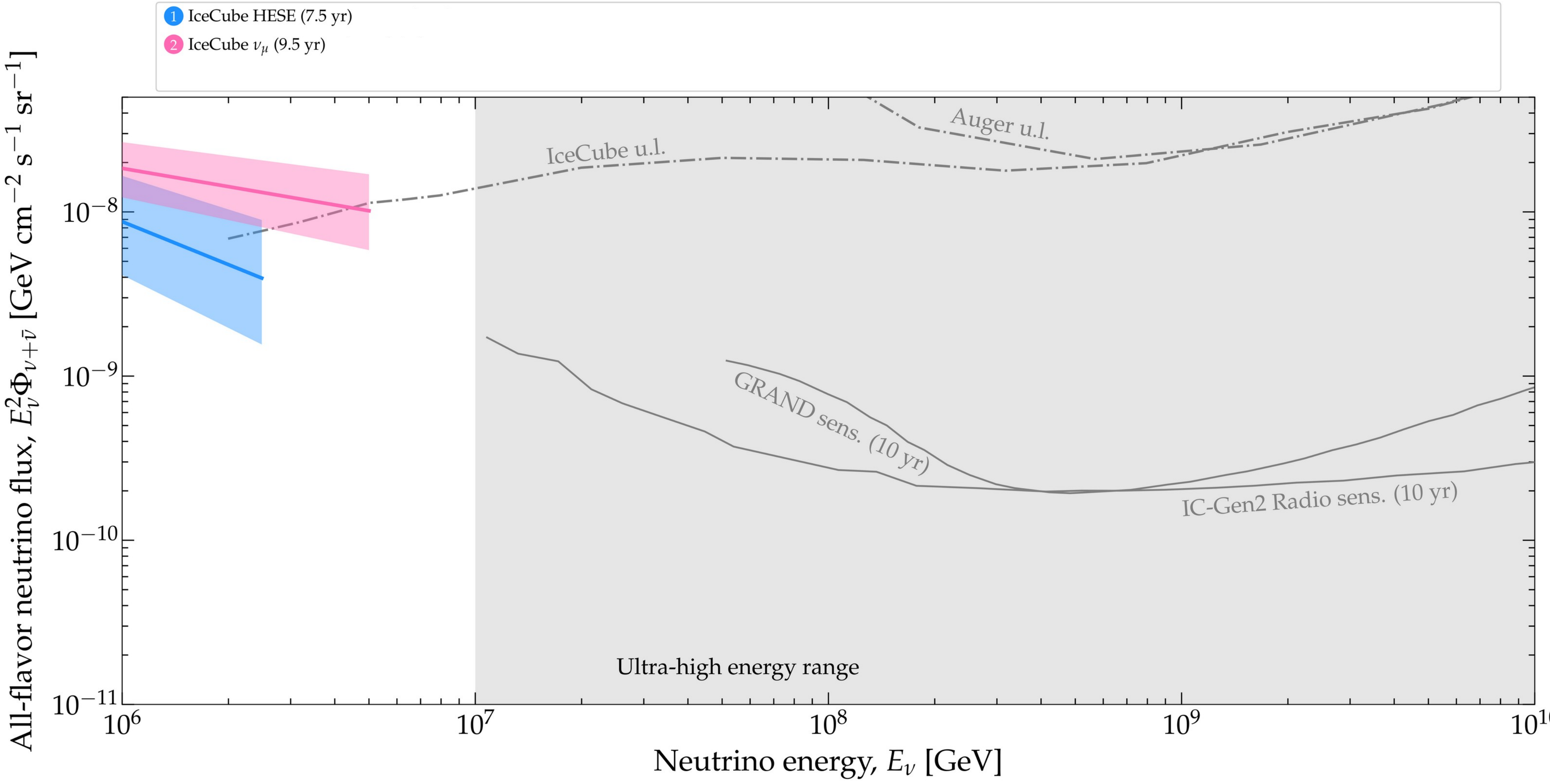
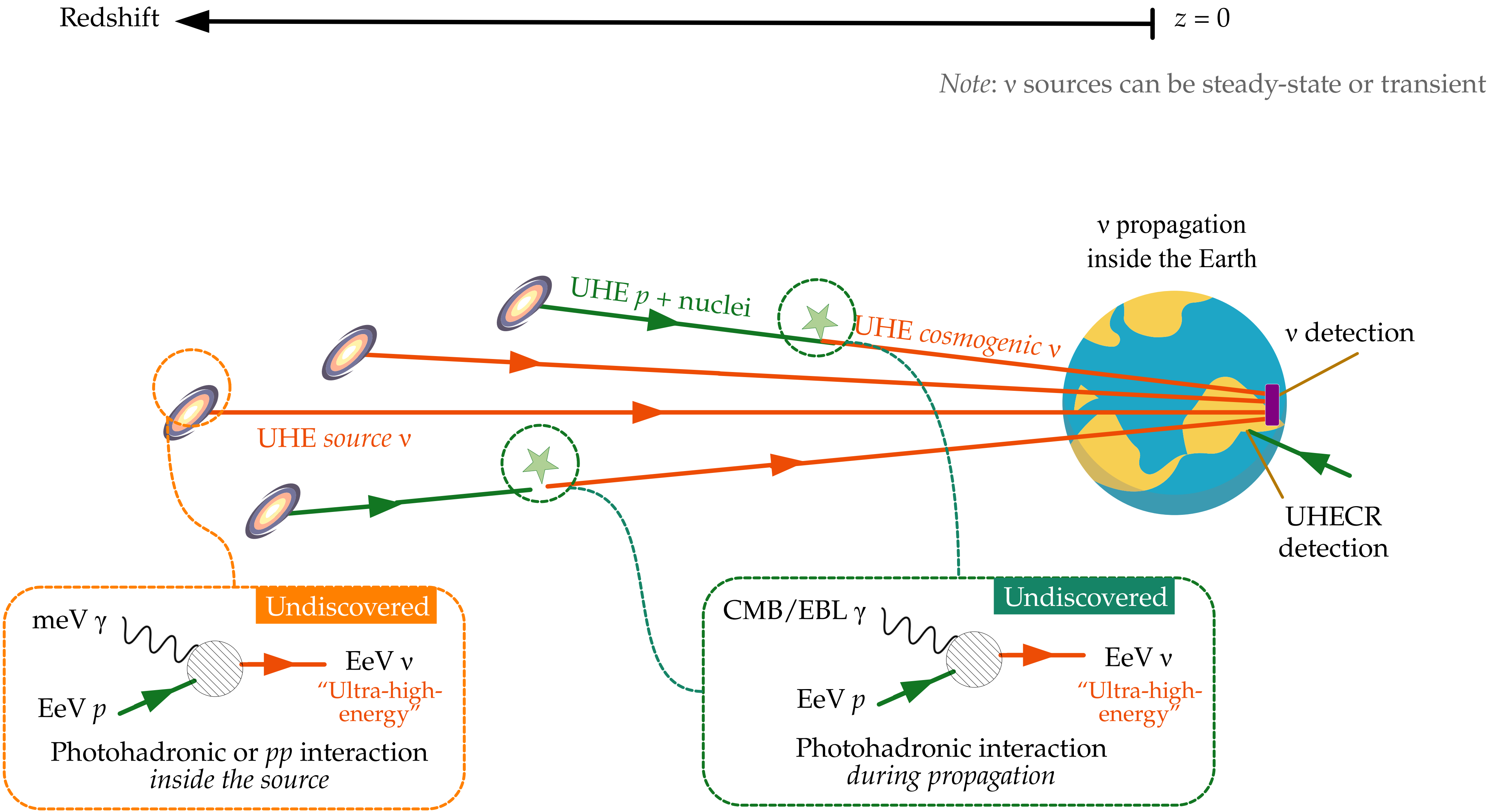
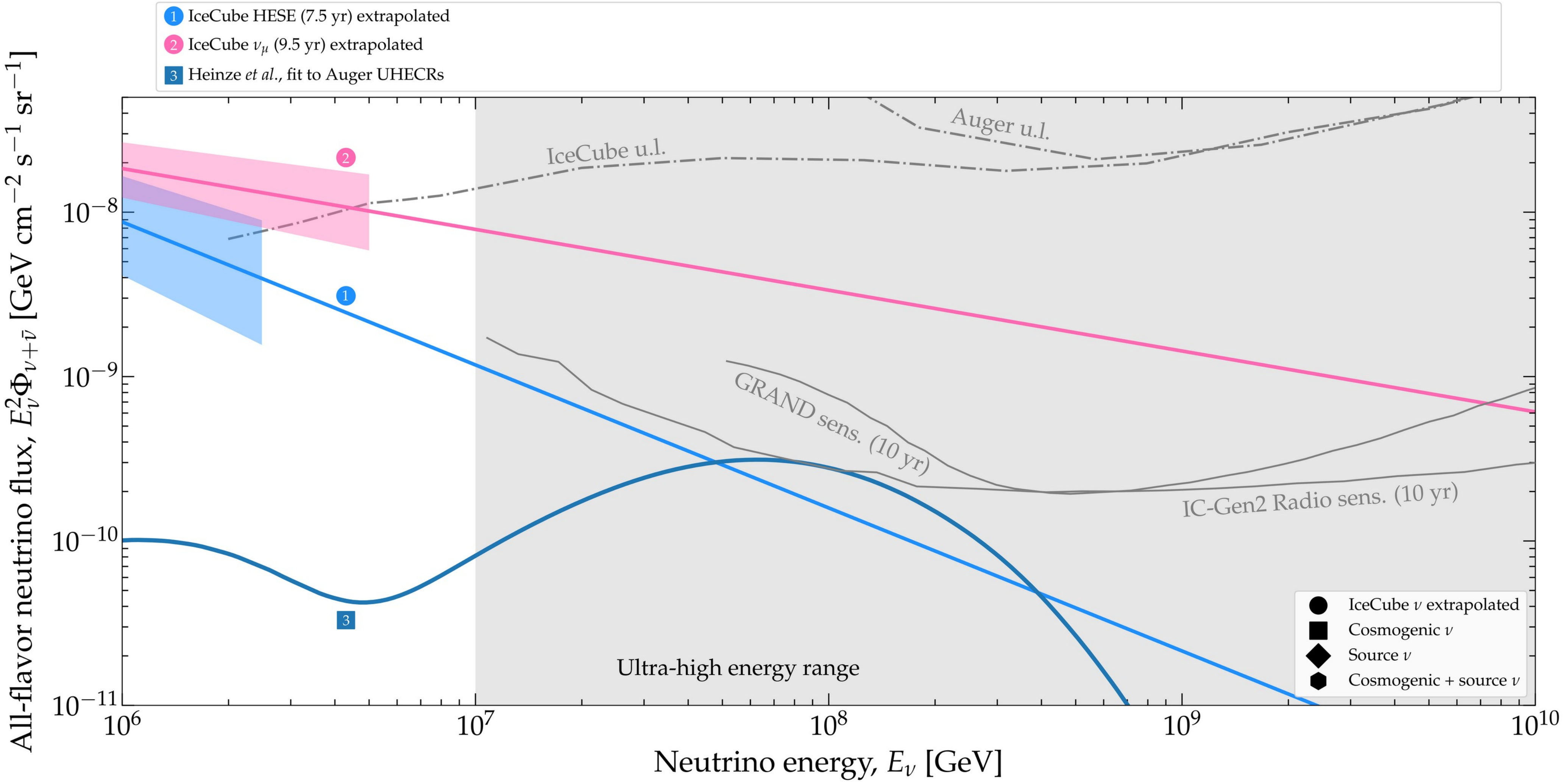


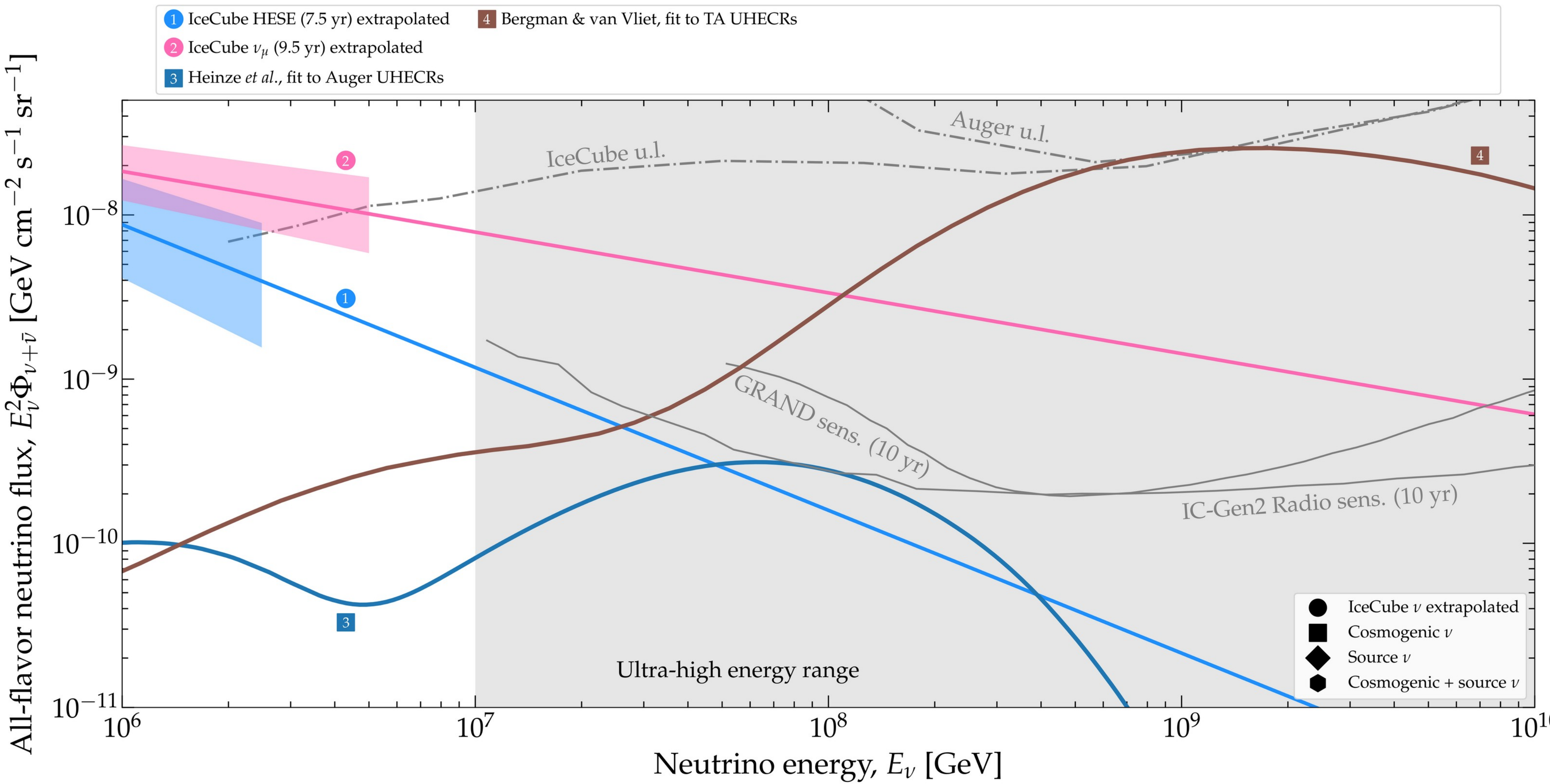
Valera, **MB**, Glaser, *PRD* 2023



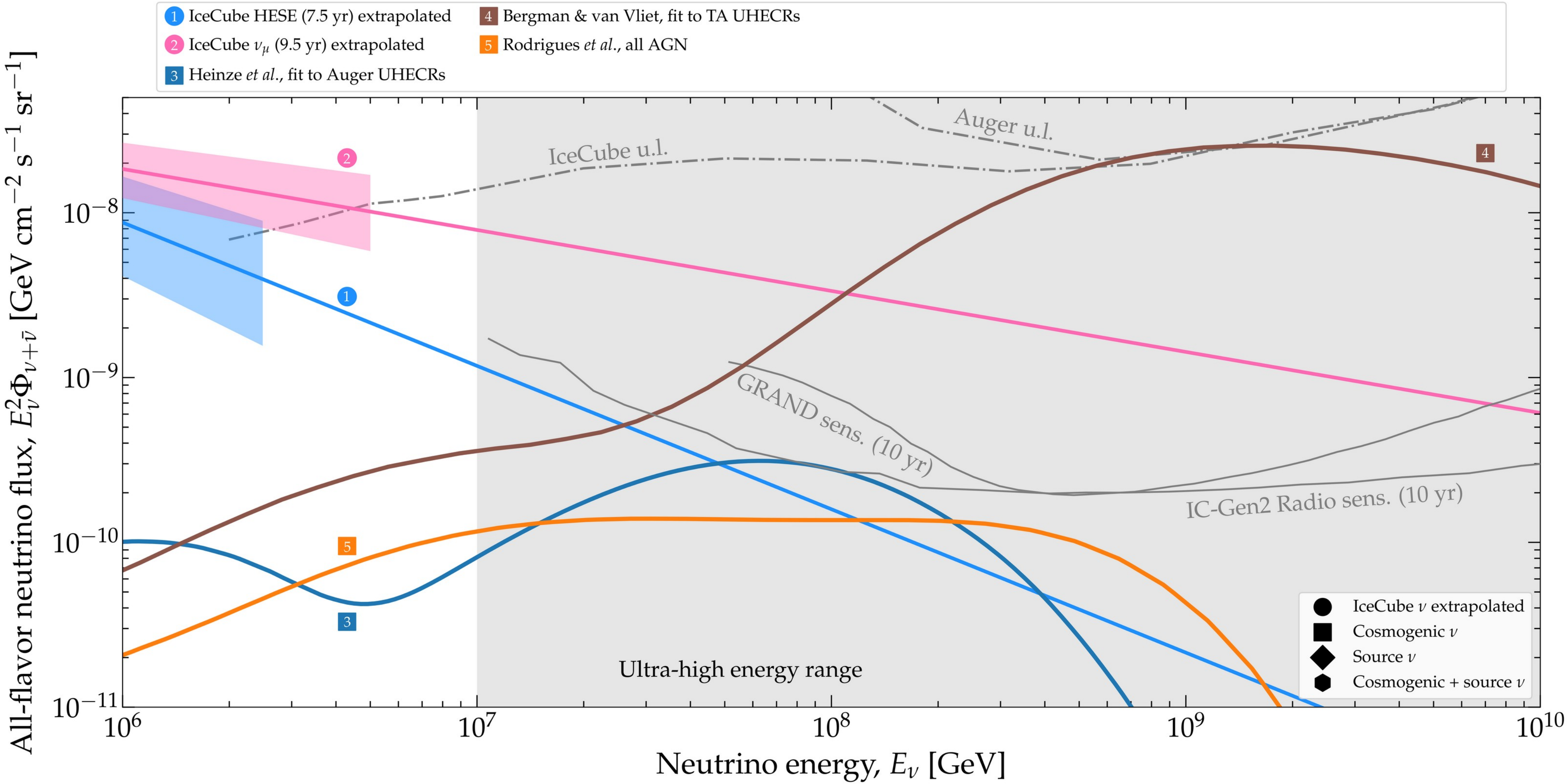
Redshift
$z = 0$
Note: ν sources can be steady-state or transient
ν propagation inside the Earth
UHE *p* + nuclei
UHE *cosmogenic* ν
ν detection
UHE *source* ν
UHECR detection
Undiscovered
meV γ
EeV ν
EeV *p*
"Ultra-high-energy"
Photohadronic or *pp* interaction
inside the source
Undiscovered
CMB/EBL γ
EeV ν
EeV *p*
"Ultra-high-energy"
Photohadronic interaction
during propagation



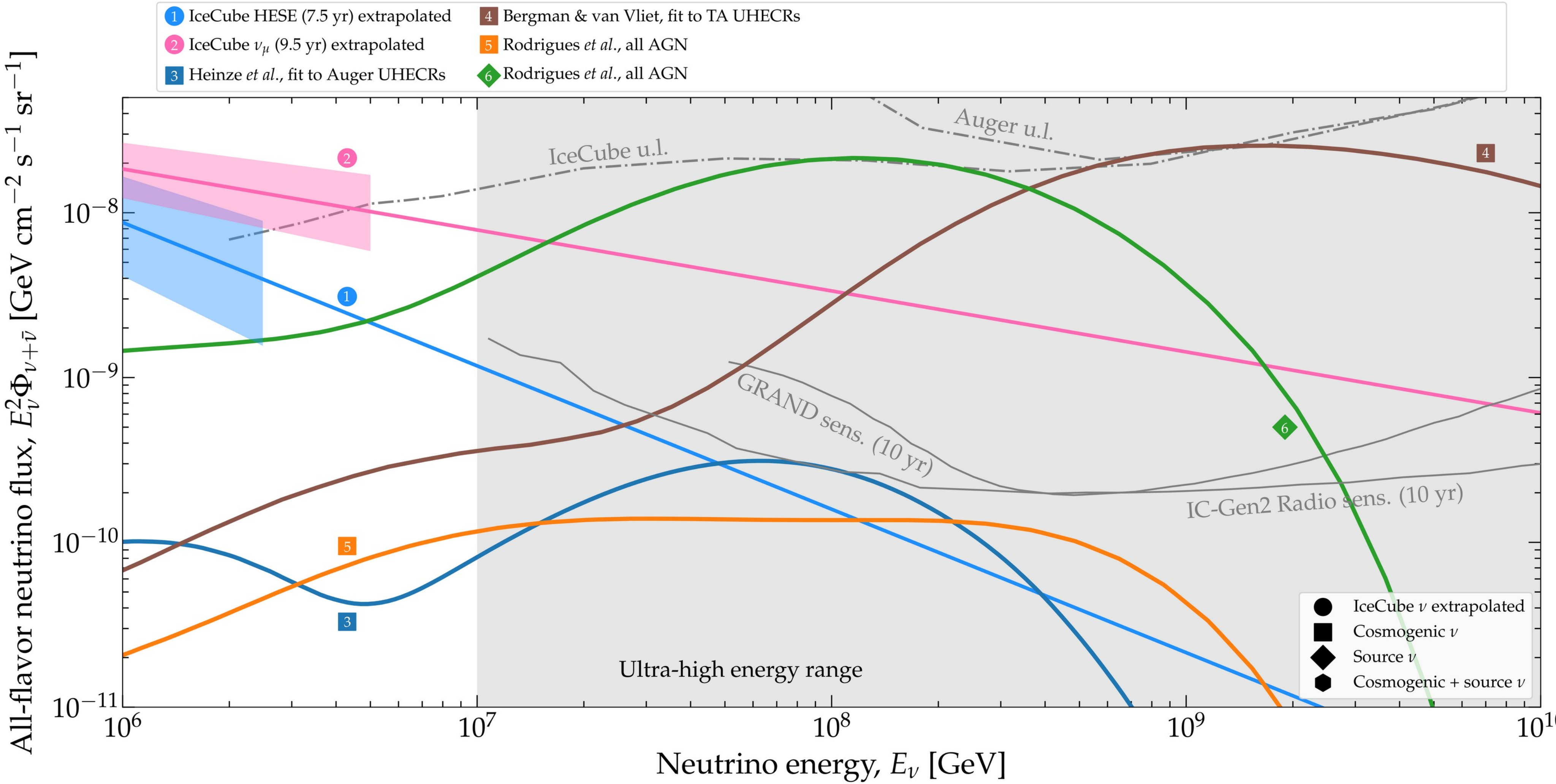
Valera, **MB**, Glaser, *PRD* 2023



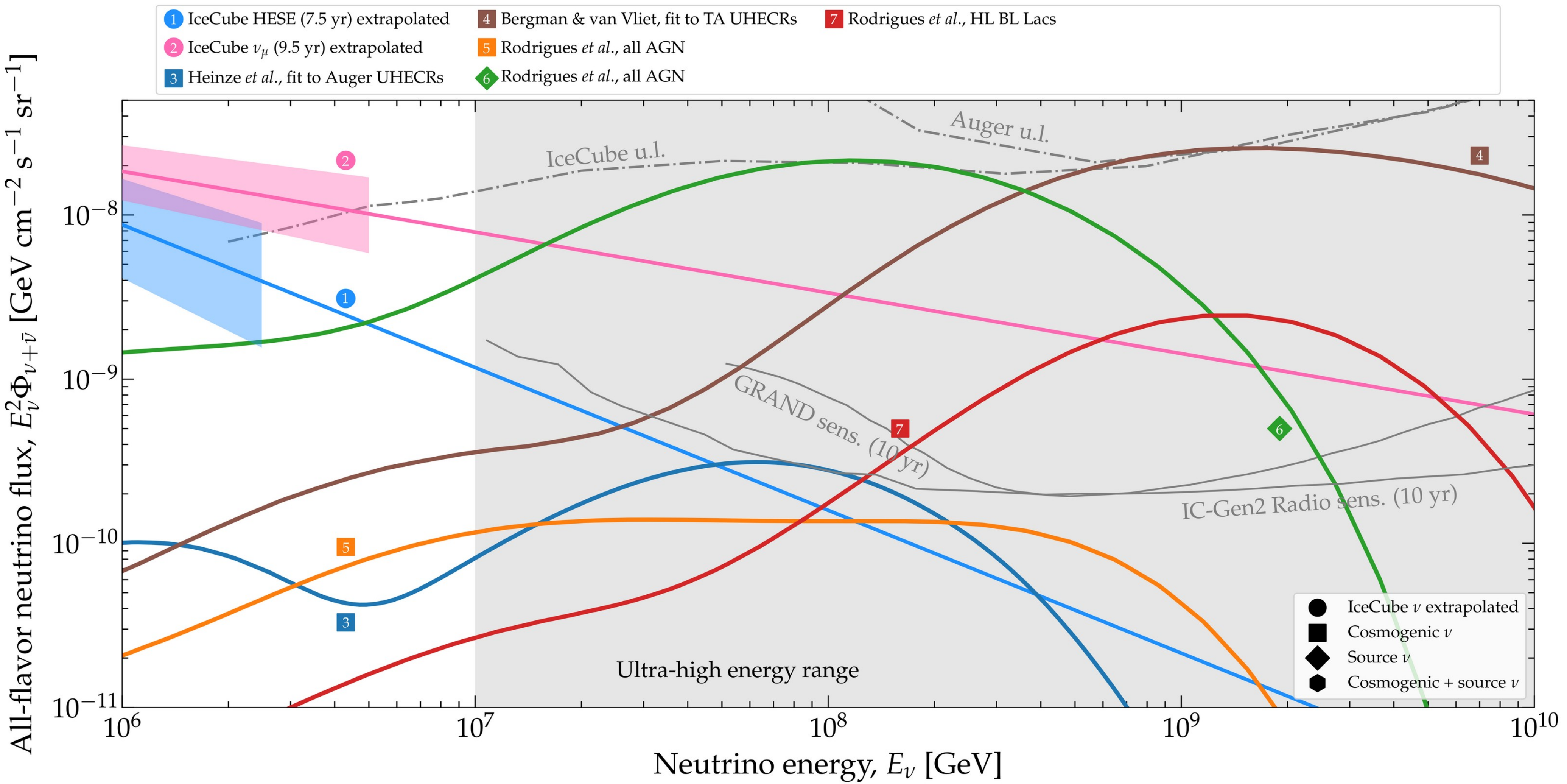
Valera, **MB**, Glaser, *PRD* 2023



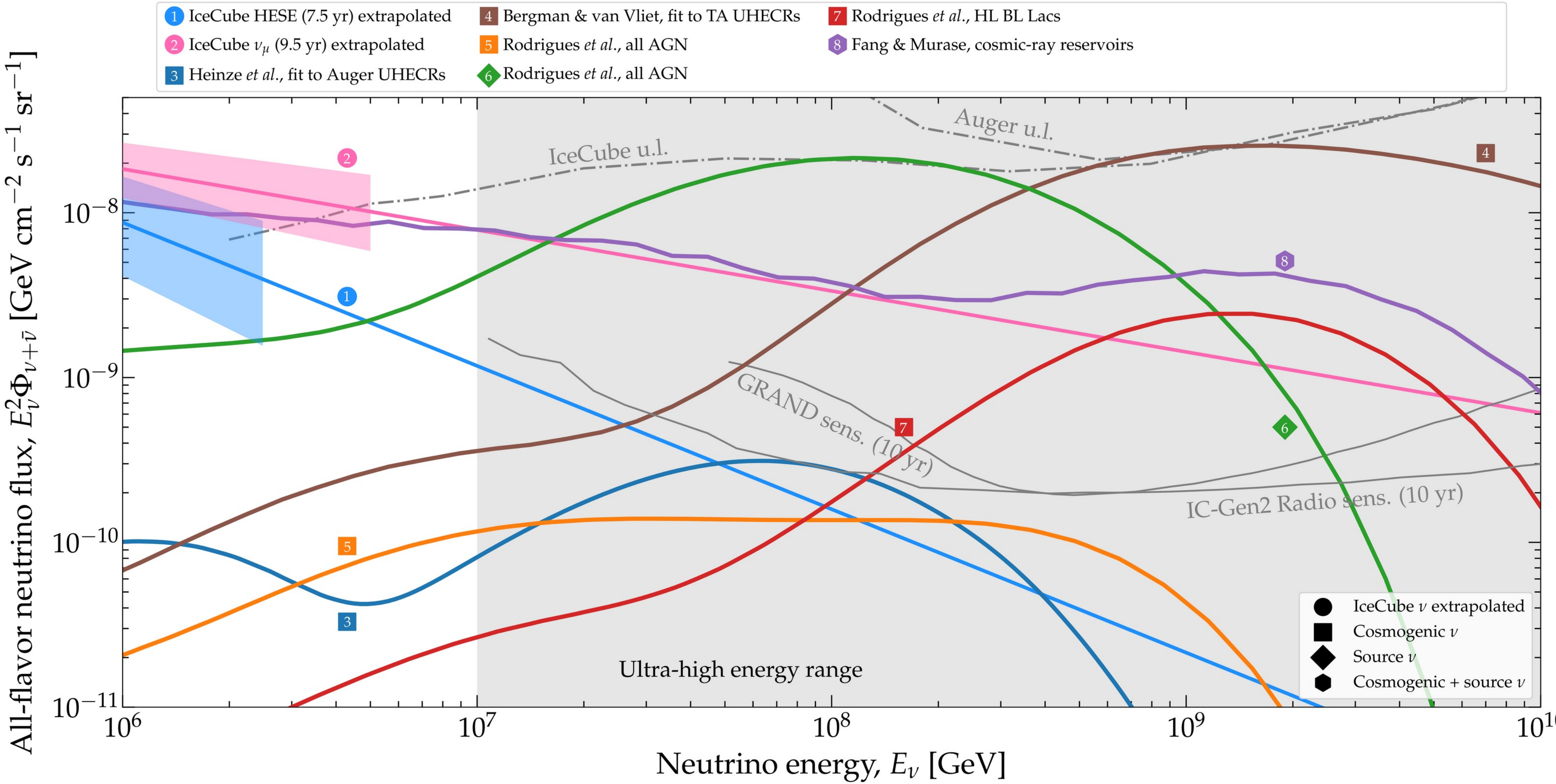
1 IceCube HESE (7.5 yr) extrapolated
2 IceCube ν_μ (9.5 yr) extrapolated
3 Heinze *et al.*, fit to Auger UHECRs
4 Bergman & van Vliet, fit to TA UHECRs
5 Rodrigues *et al.*, all AGN
Auger u.l.
IceCube u.l.
GRAND sens. (10 yr)
IC-Gen2 Radio sens. (10 yr)
Ultra-high energy range
IceCube ν extrapolated
Cosmogenic ν
Source ν
Cosmogenic + source ν
All-flavor neutrino flux, $E_\nu^2\Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
Neutrino energy, E_ν [GeV]



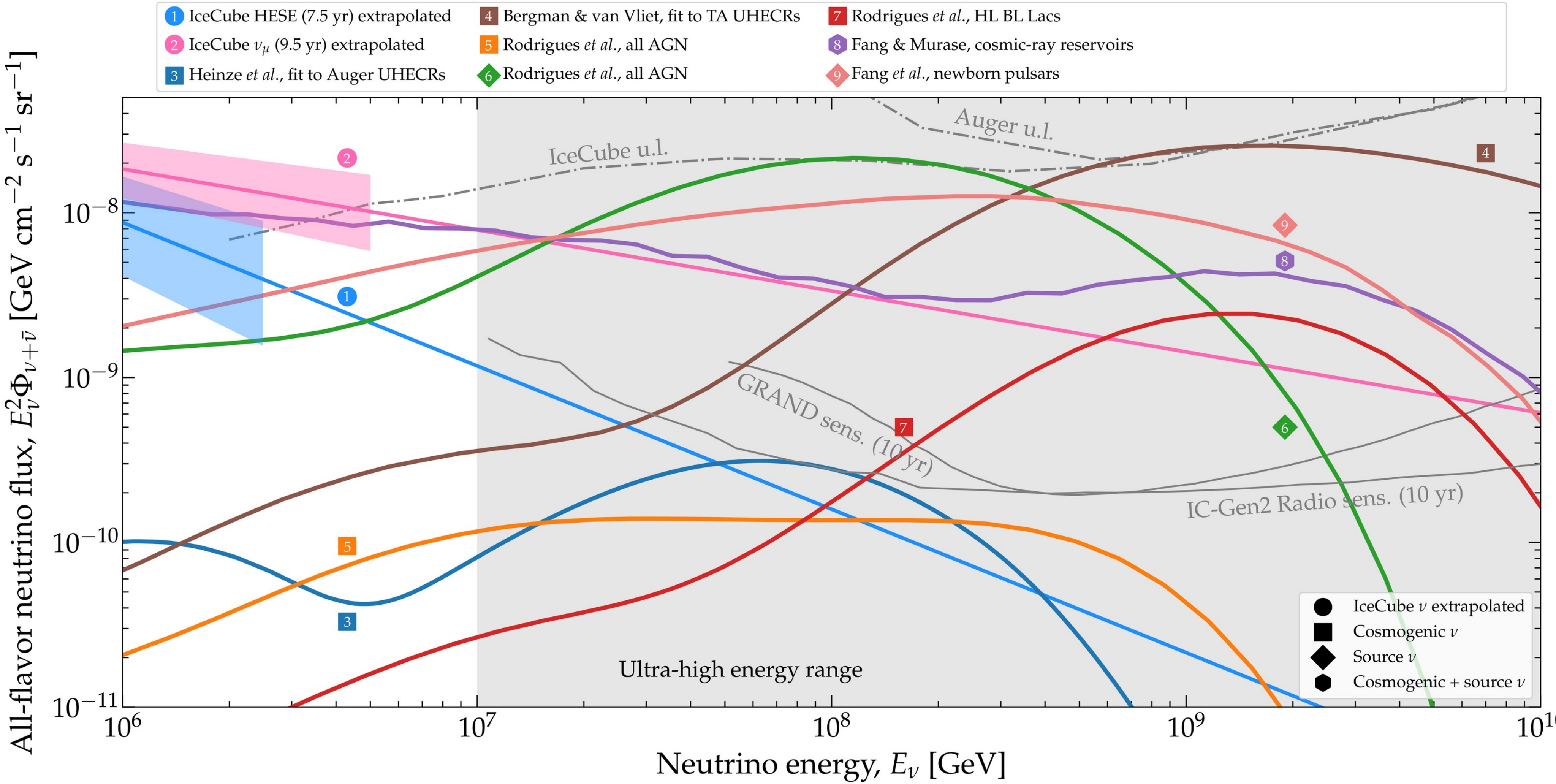
Valera, **MB**, Glaser, *PRD* 2023



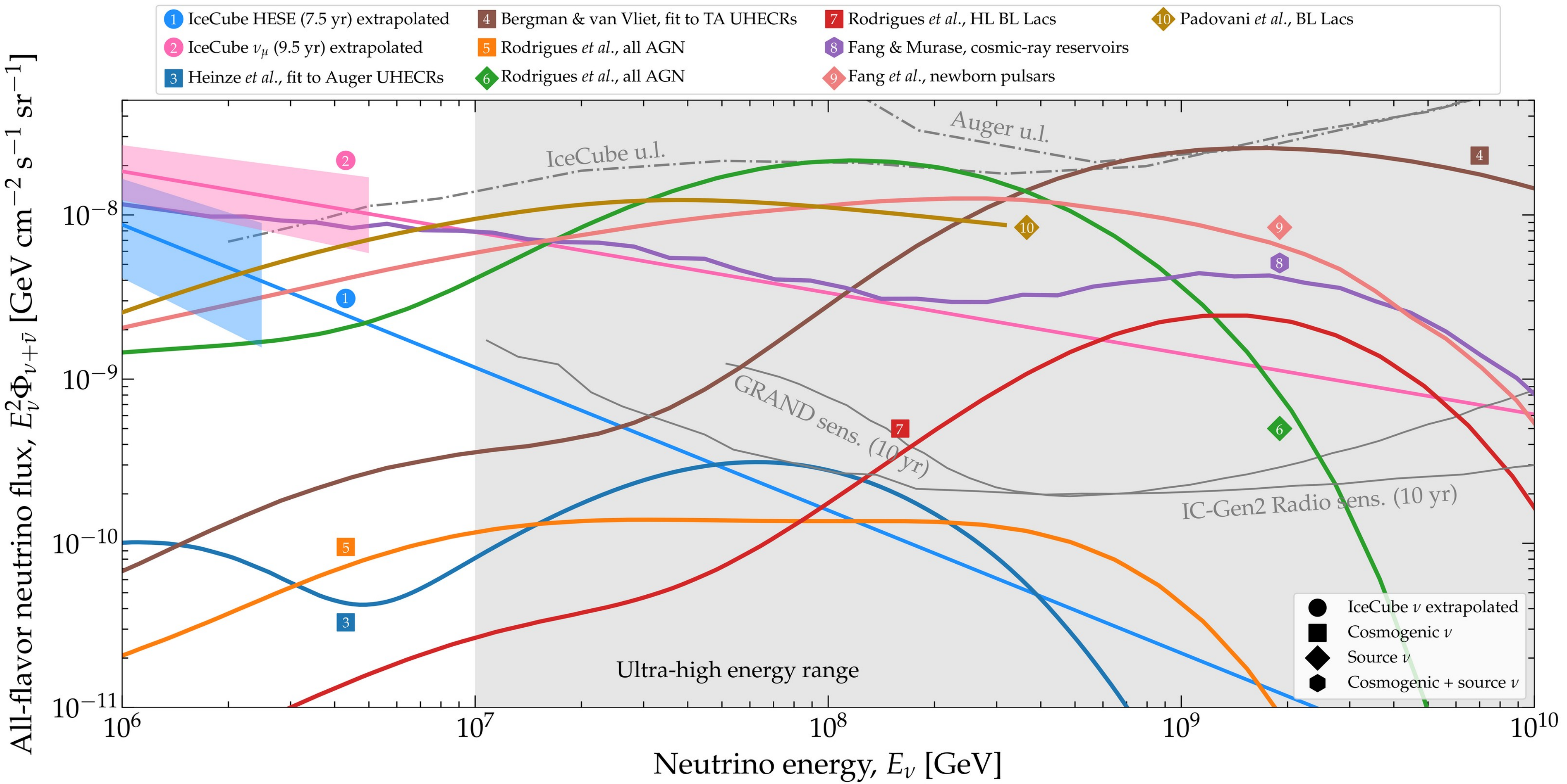
Valera, **MB**, Glaser, *PRD* 2023



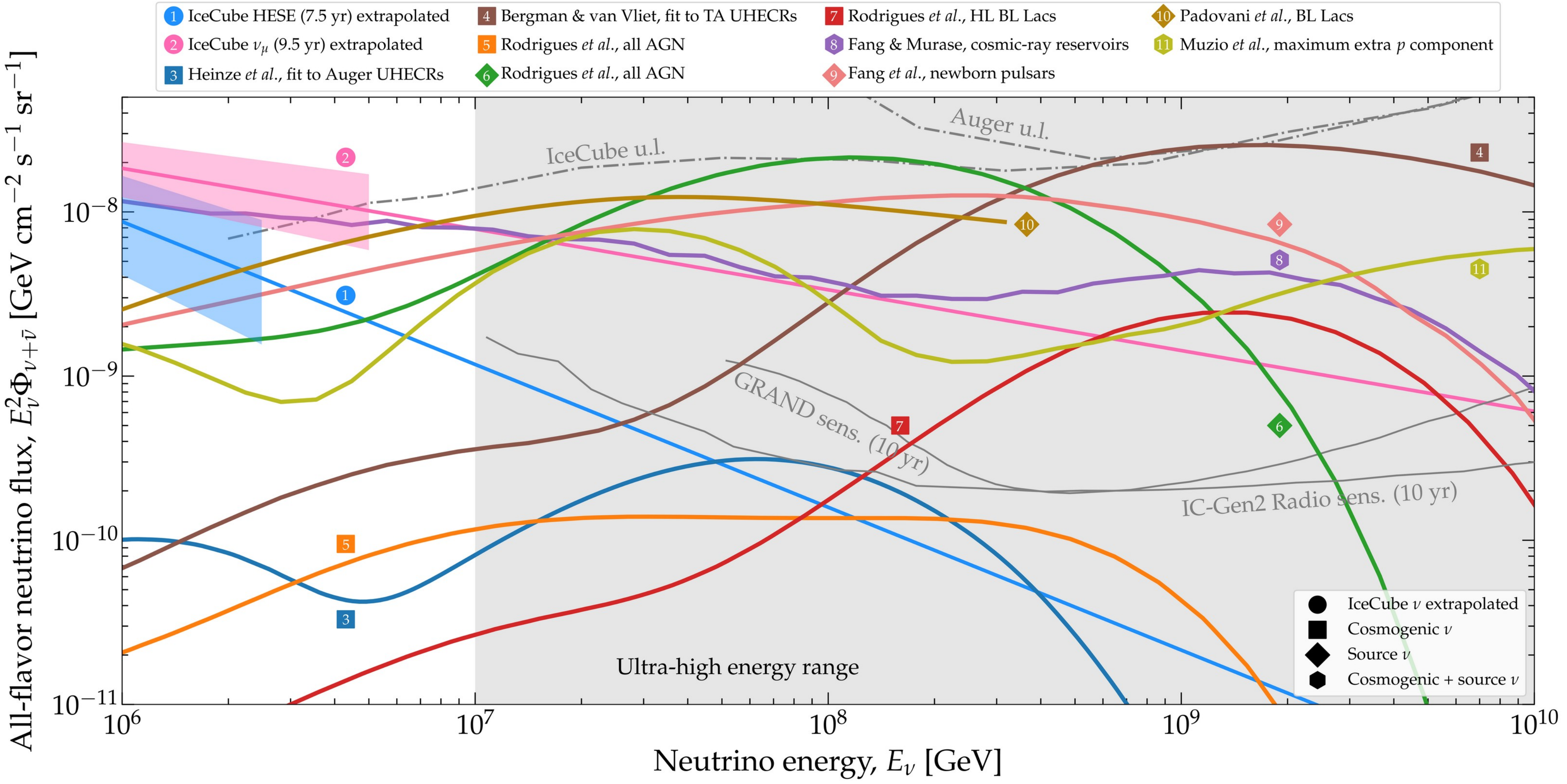
Valera, **MB**, Glaser, *PRD* 2023



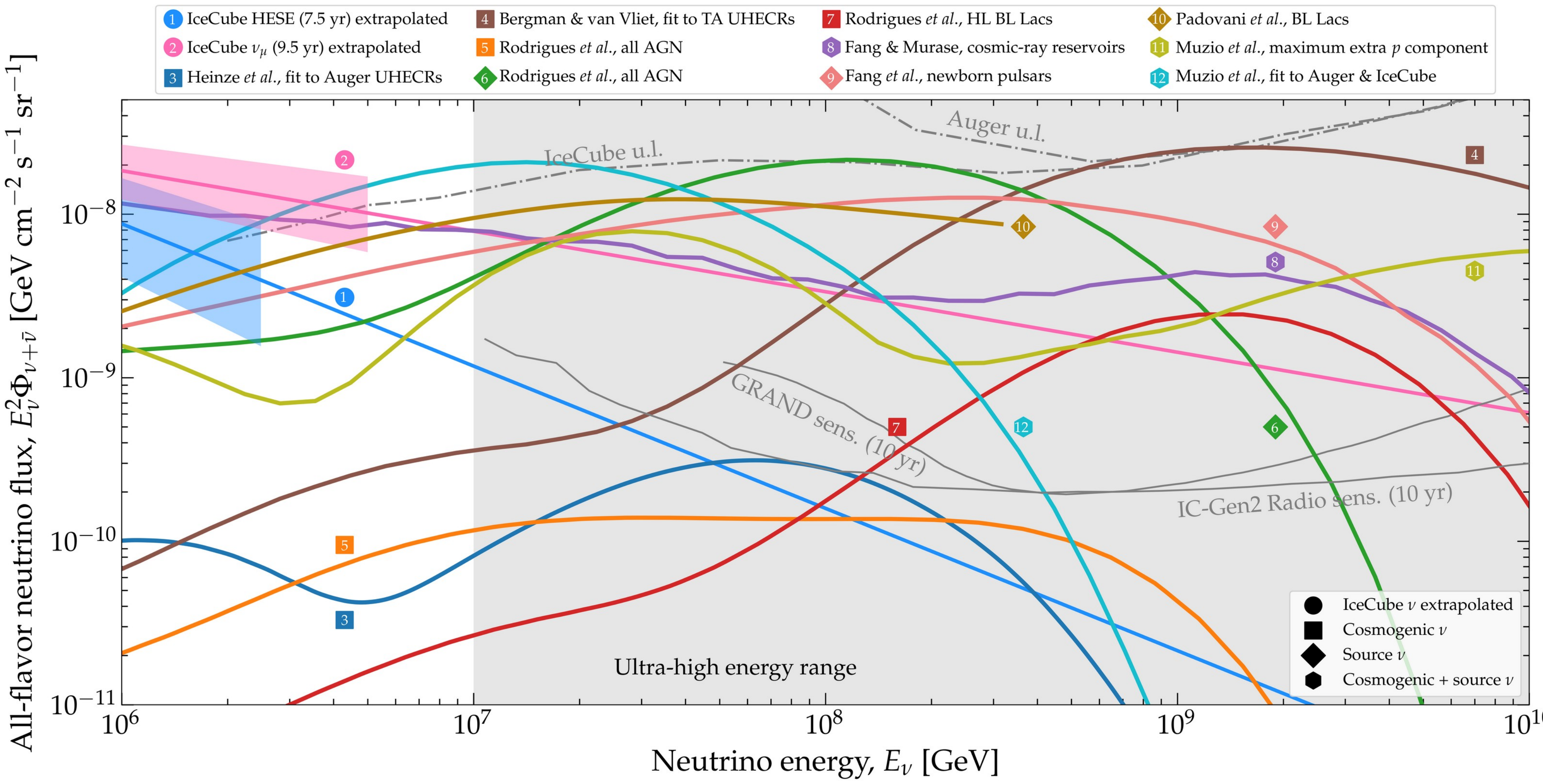
Valera, **MB**, Glaser, *PRD* 2023



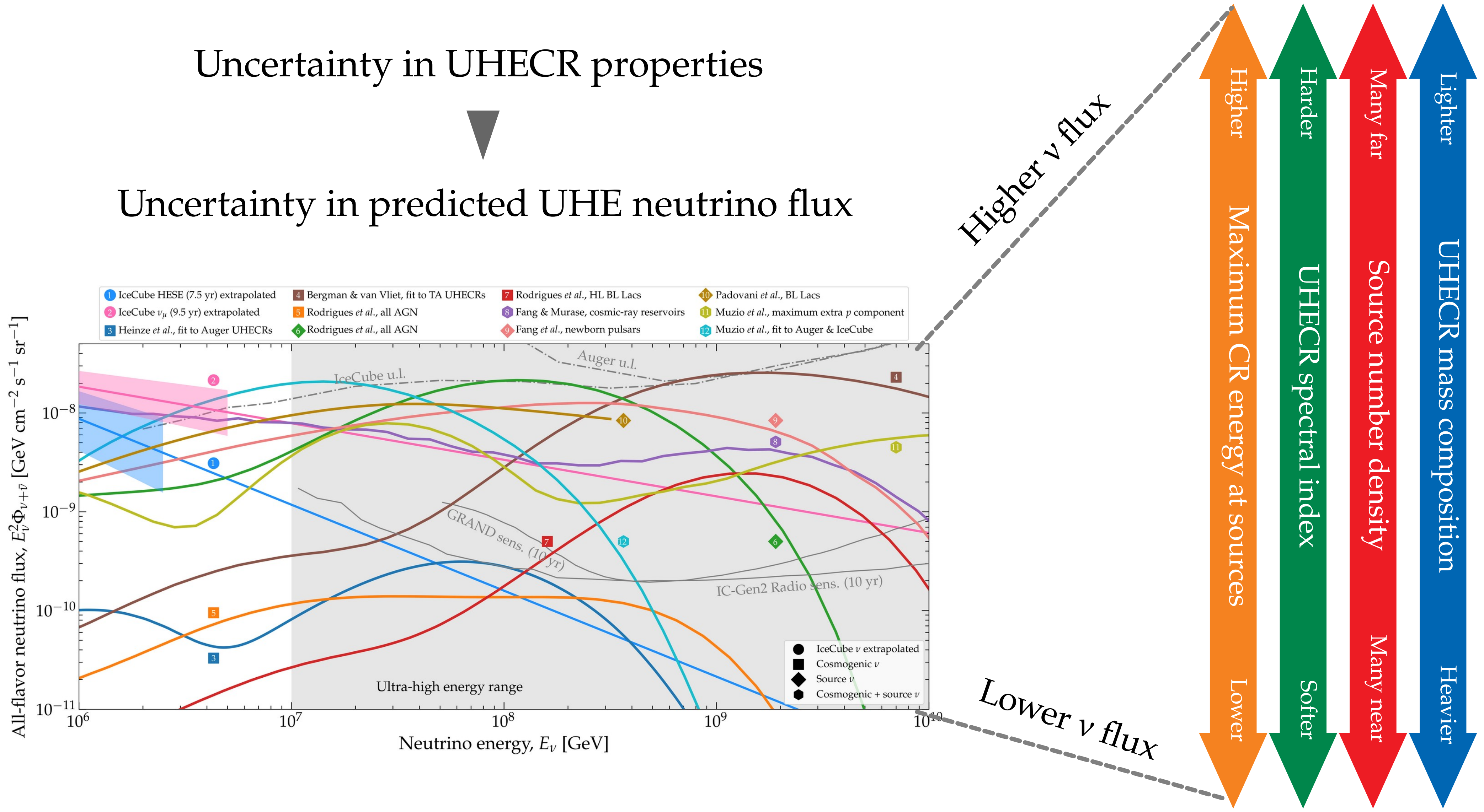
Valera, **MB**, Glaser, *PRD* 2023



Valera, **MB**, Glaser, *PRD* 2023



Valera, **MB**, Glaser, *PRD* 2023



Uncertainty in UHECR properties
Uncertainty in predicted UHE neutrino flux
1 IceCube HESE (7.5 yr) extrapolated
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3 Heinze et al., fit to Auger UHECRs
4 Bergman & van Vliet, fit to TA UHECRs
5 Rodrigues et al., all AGN
6 Rodrigues et al., all AGN
7 Rodrigues et al., HL BL Lacs
8 Fang & Murase, cosmic-ray reservoirs
9 Fang et al., newborn pulsars
10 Padovani et al., BL Lacs
11 Muzio et al., maximum extra p component
12 Muzio et al., fit to Auger & IceCube
IceCube u.l.
Auger u.l.
GRAND sens. (10 yr)
IC-Gen2 Radio sens. (10 yr)
Ultra-high energy range
IceCube ν extrapolated
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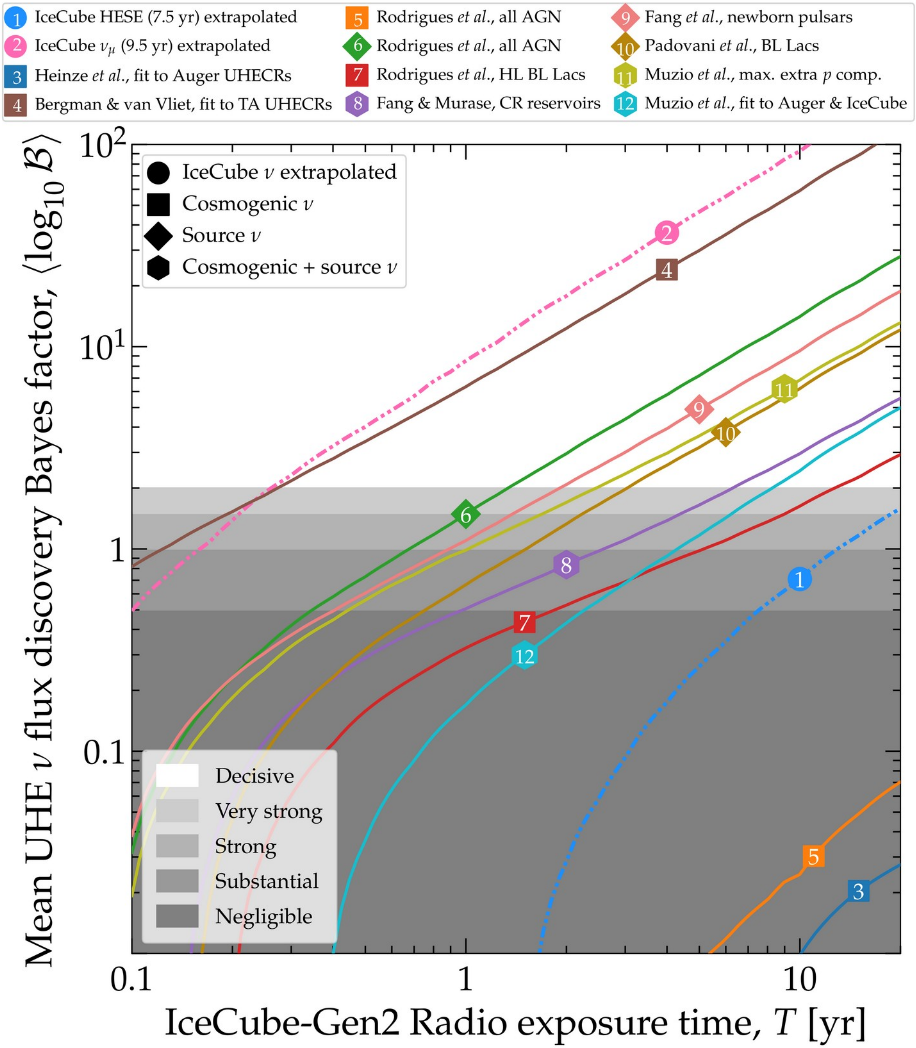
10^{-11}
10^{-10}
10^{-9}
10^{-8}
10^{6}
10^{7}
10^{8}
10^{9}
10^{10}
Higher ν flux
Lower ν flux
Higher
Maximum CR energy at sources
Lower
Harder
UHECR spectral index
Softer
Many far
Source number density
Many near
Lighter
UHECR mass composition
Heavier

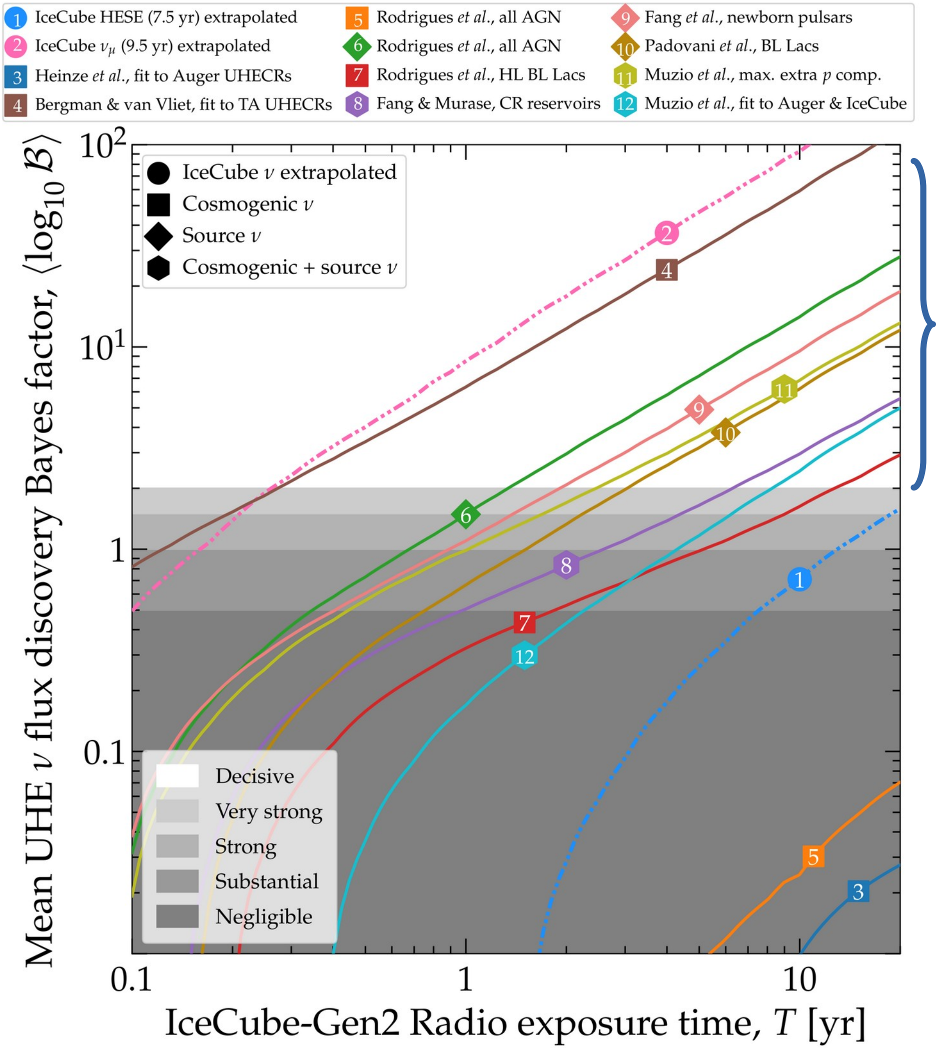
Discovering the diffuse flux of UHE neutrinos



Valera, **MB**, Glaser, *PRD* 2023

Discovering the diffuse flux of UHE neutrinos

Bayes factor compares signal+bkg. *vs.* bkg.-only



Large Bayes factor = decisive flux discover

Valera, **MB**, Glaser, *PRD* 2023

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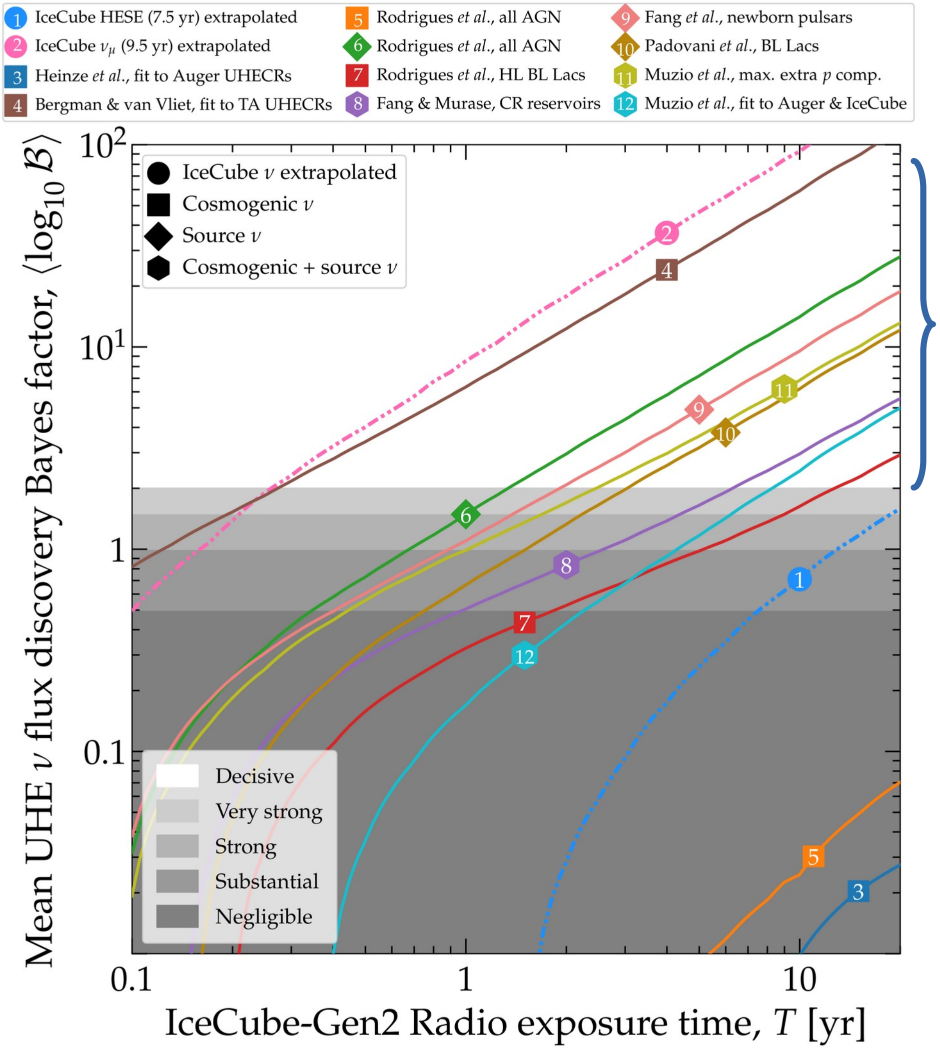


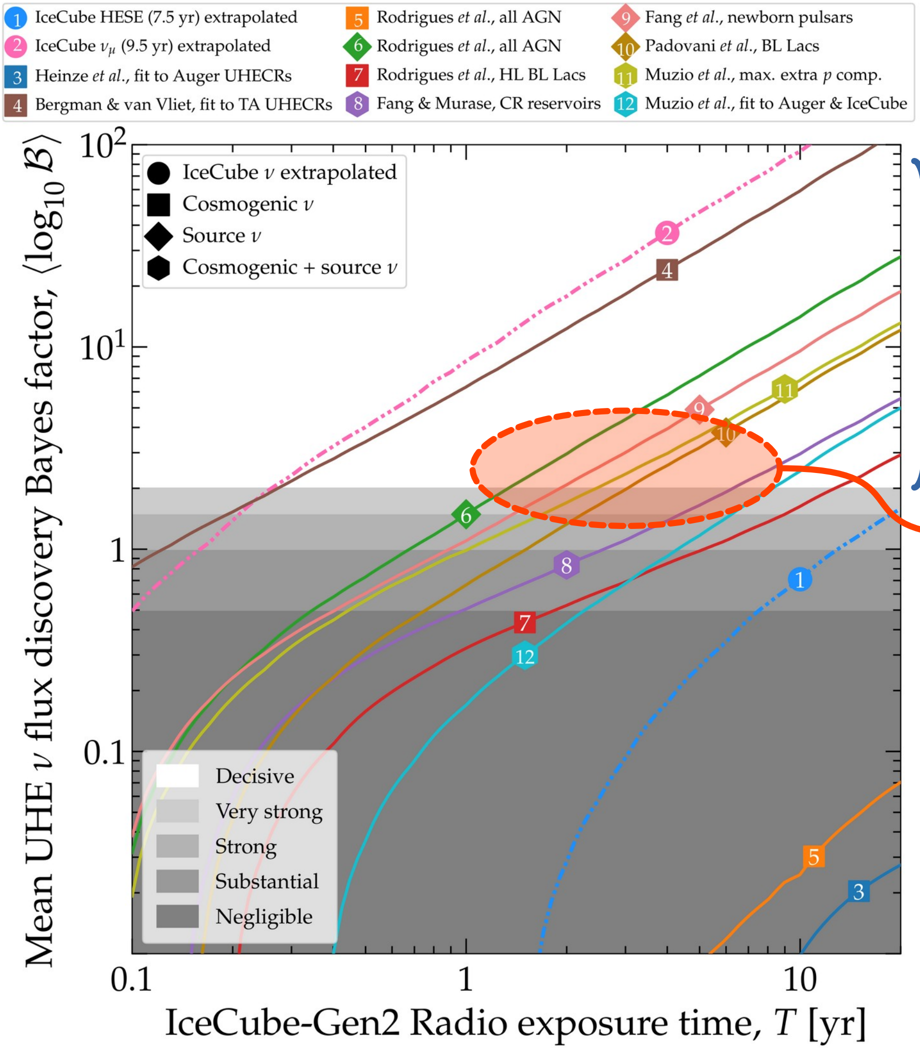
Large Bayes factor
=
decisive flux discover

Forecasts are state-of-the-art:
- Neutrino propagation inside Earth
- Detailed simulation of radio in ice
- Detailed antenna response
- Detector energy & angular resolution
- Statistical fluctuations

Valera, **MB**, Glaser, *PRD* 2023

Discovering the diffuse flux of UHE neutrinos

Bayes factor compares signal+bkg. *vs.* bkg.-only



Large Bayes factor = decisive flux discover

Most flux models are discoverable with a few years

Forecasts are state-of-the-art:
- Neutrino propagation inside Earth
- Detailed simulation of radio in ice
- Detailed antenna response
- Detector energy & angular resolution
- Statistical fluctuations

Valera, **MB**, Glaser, *PRD* 2023

Today

TeV–PeV ν

Turn predictions
into data-driven tests

Key developments:
Bigger detectors → larger statistics
Better reconstruction
Smaller astrophysical uncertainties

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Next decade

> 100-PeV ν

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Make predictions for
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Made robust and meaningful by accounting
for all relevant particle and astrophysics uncertainties

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<u>Key developments</u>:
Bigger detectors → larger statistics
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Smaller astrophysical uncertainties

Next decade

> 100-PeV ν

Make predictions for
a new energy regime

<u>Key developments</u>:
Discovery
New detection techniques
Better UHE ν flux predictions

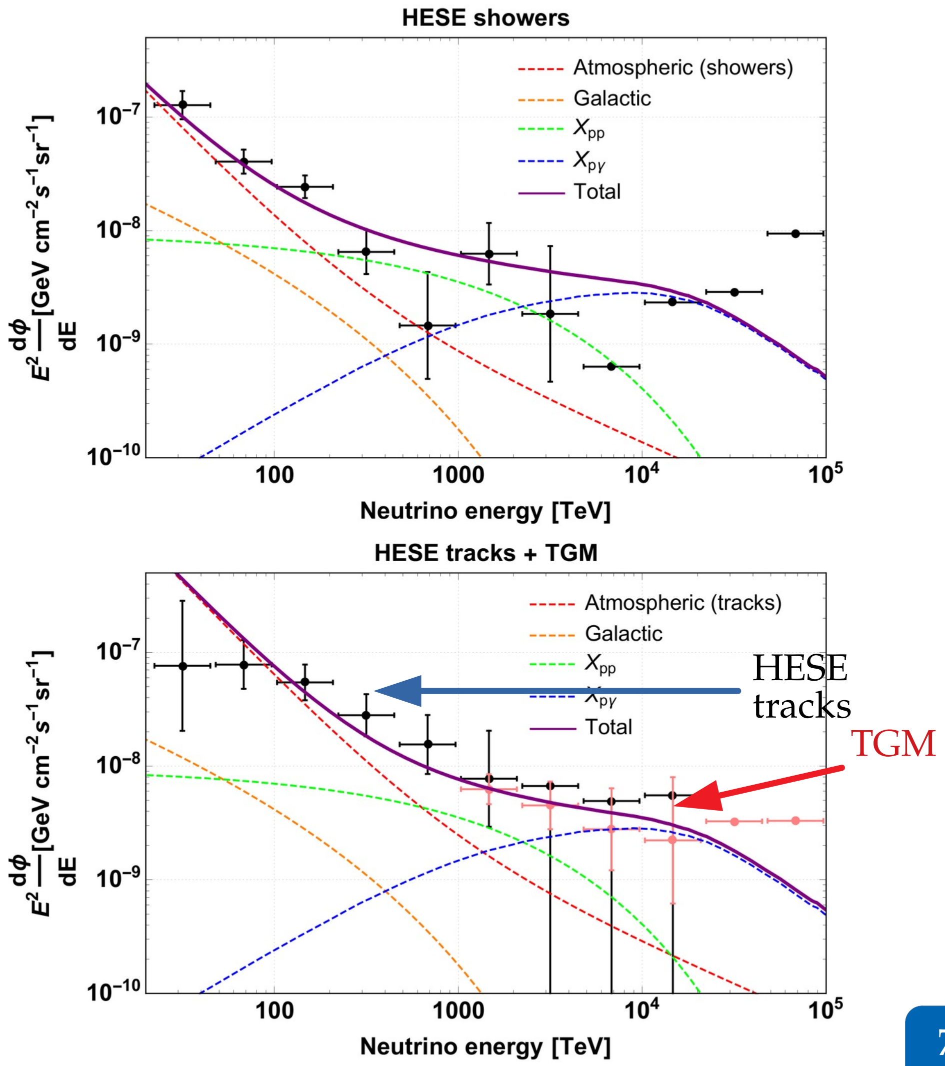
Similar to the evolution of cosmology to a high-precision field in the 1990s

Made robust and meaningful by accounting
for all relevant particle and astrophysics uncertainties

Backup slides

Multi-component model of astrophysical neutrinos

- Four diffuse components:
  - Residual atmospheric (0.2–0.5 PeV): Conv. ($E^{-3.7}$) & prompt ($E^{-2.7}$) ν + muons
  - Galactic ν (≲ PeV): *pp* with disc gas ($E^{-2.6}$), confined to |b| < 5°, |l| < 45°
  - Extragalactic ν from *pp*, *Ap*: *á la* starbursts (E^{-2})
  - Extragalactic ν from *pγ*, *Aγ*: *á la* TDE (peaked around a few PeV)
- Simultaneous fit to HESE showers, tracks, through-going muons (TGM)



Palladino & Winter, *A&A* 2018

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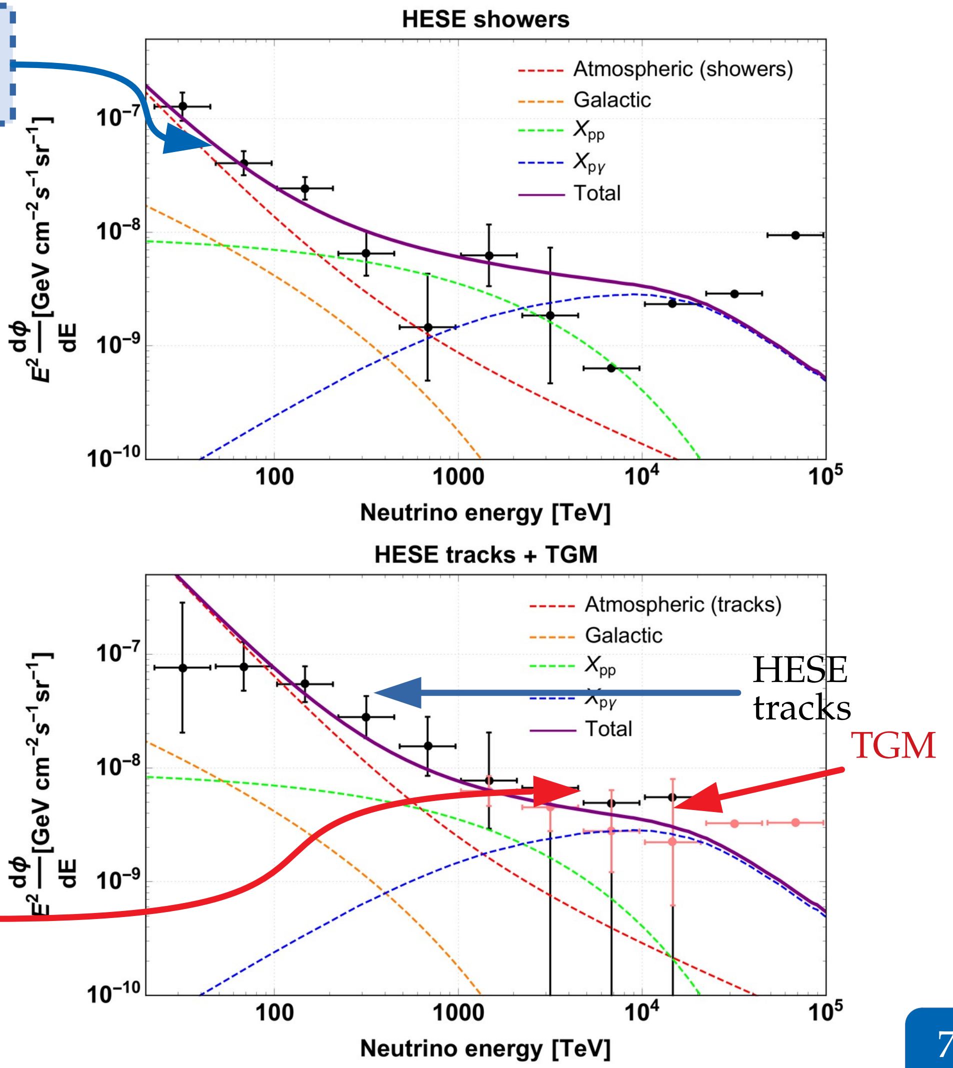
Describes astrophysical ν better at low energies



Describes astrophysical ν better at high energies

Palladino & Winter, *A&A* 2018

Multi-component model of astrophysical neutrinos

- Four diffuse components:
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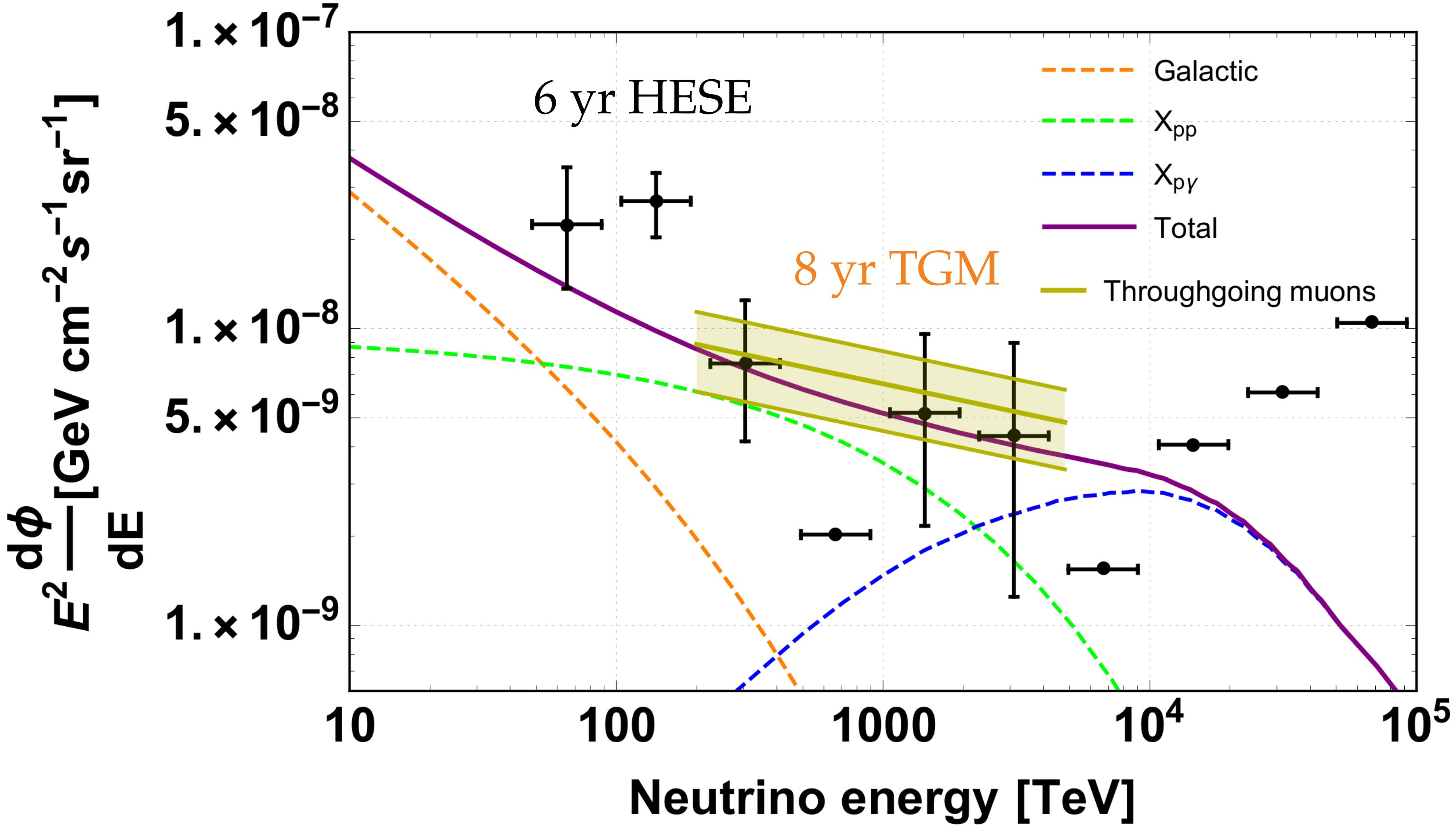
Palladino & Winter, *A&A* 2018

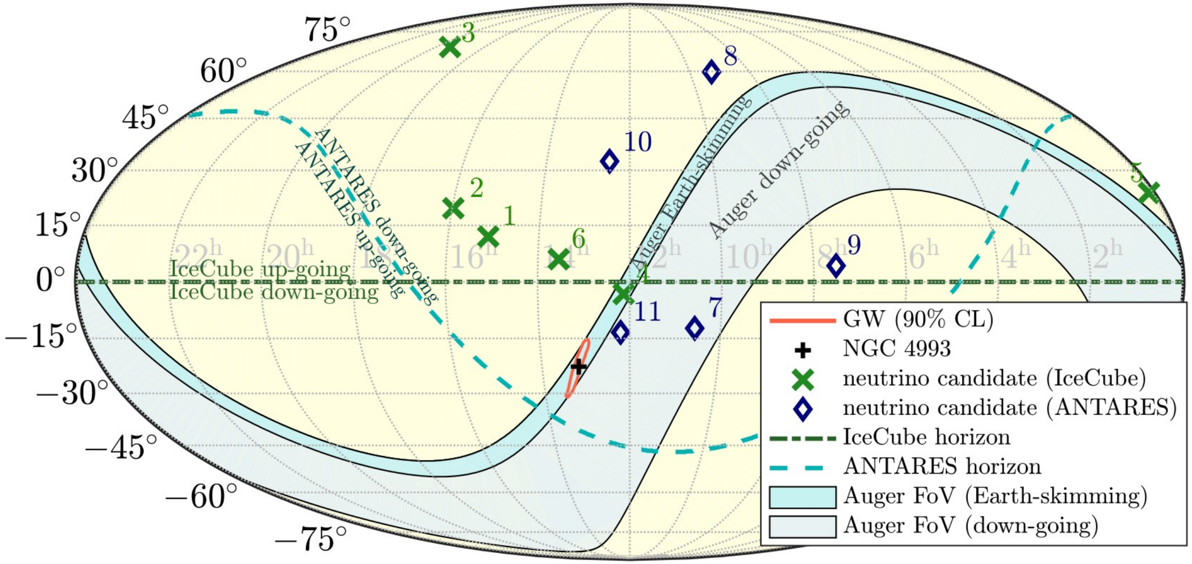
Multi-component model of astrophysical neutrinos

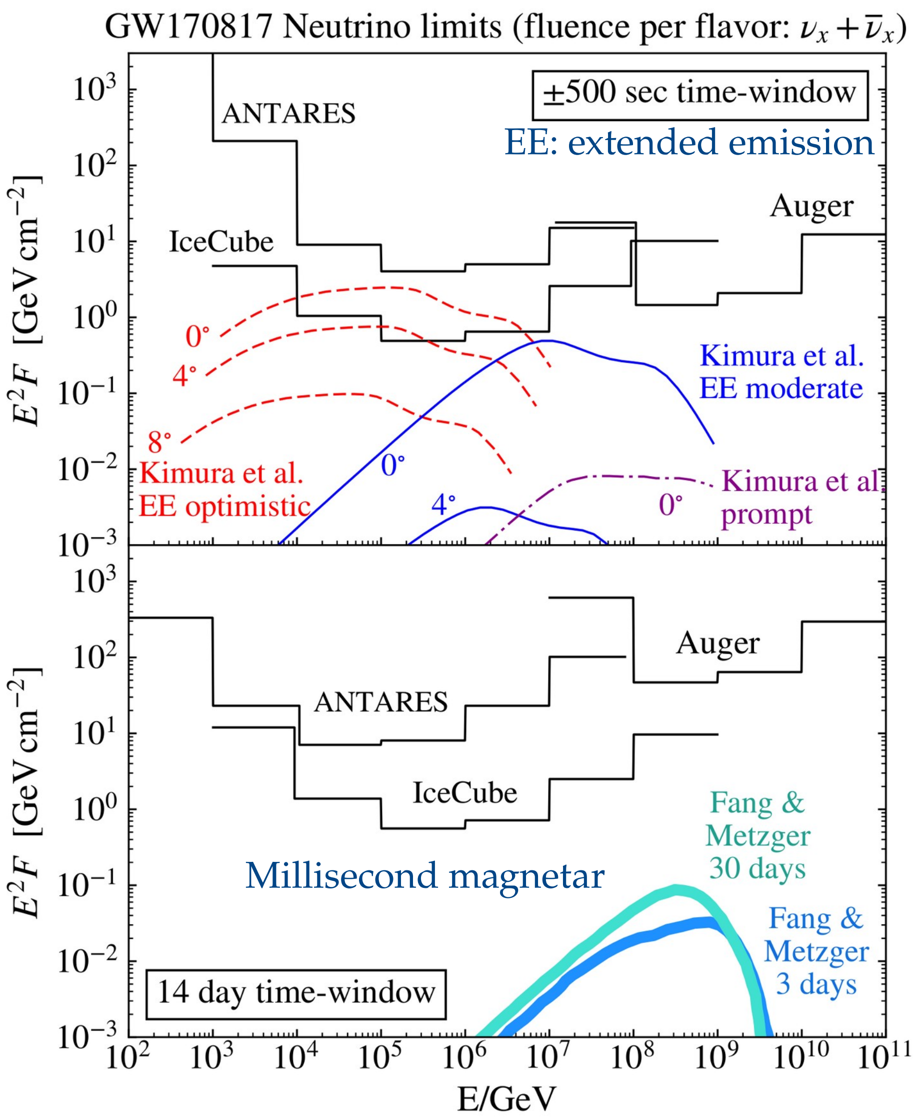


Palladino & Winter, *A&A* 2018

GW170817 (NS-NS merger)

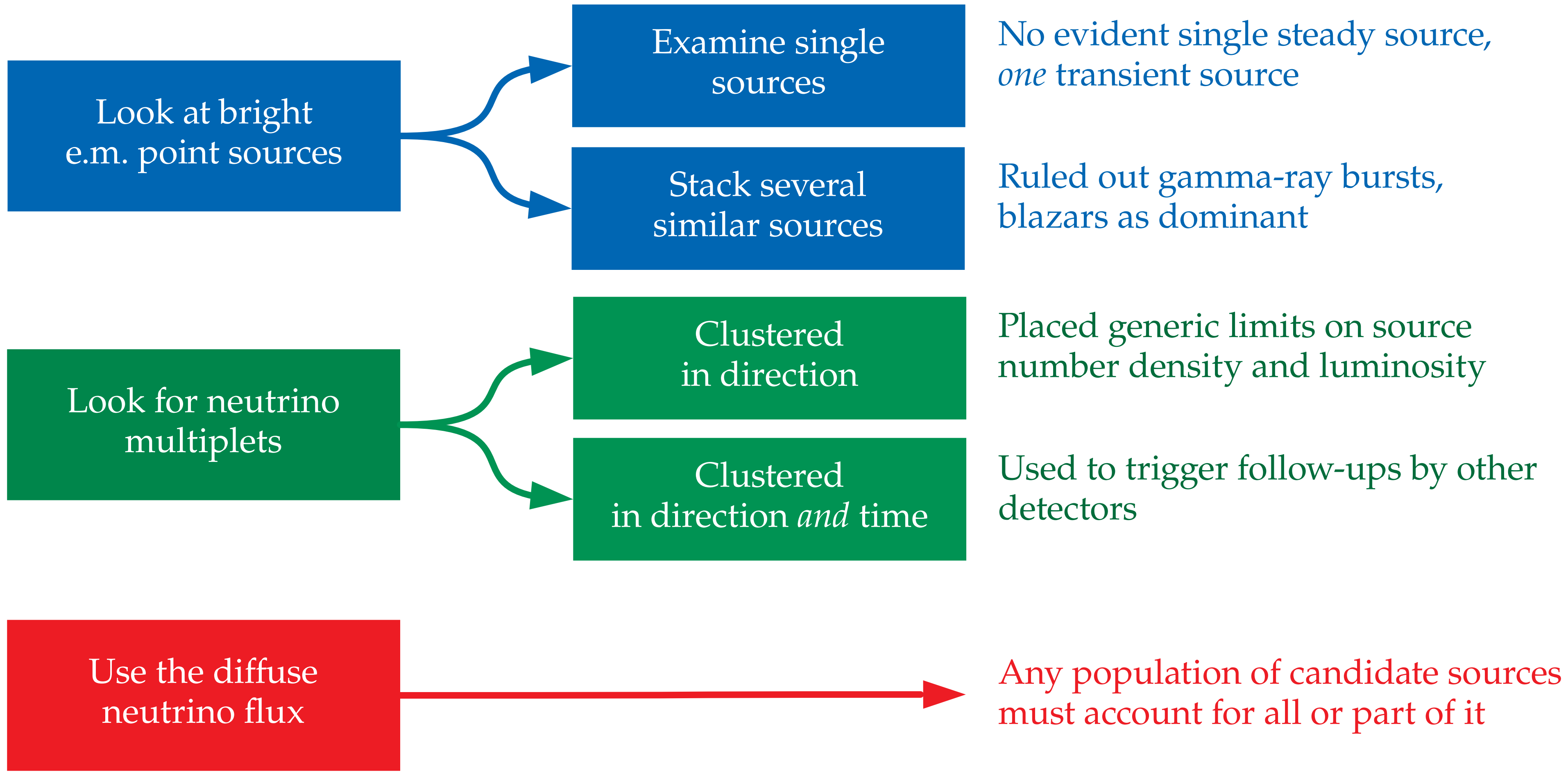
- Short GRB seen in *Fermi*-GBM, INTEGRAL
- Neutrino search by IceCube, ANTARES, and Auger
- MeV–EeV neutrinos, 14-day window
- Non-detection consistent with off-axis





ANTARES, IceCube, Pierre Auger Collab., *ApJL* 2017

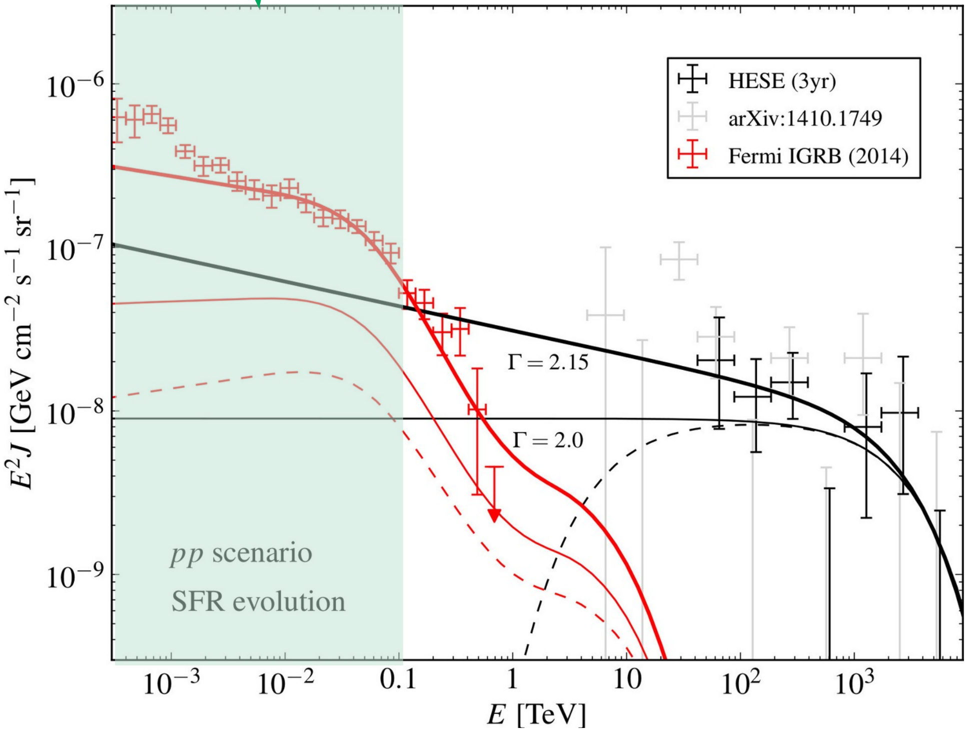
Three Strategies to Reveal Sources Using TeV–PeV ν



Constraints from the gamma-ray background

- Production via *pp*: ν and gamma-ray spectra follow the CR spectrum $E^{-\Gamma}$
- Gamma-ray interactions on the CMB make them pile up at GeV
- *Fermi* gamma-ray background is not exceeded only if $\Gamma < 2.2$
- But IceCube found Γ = 2.5–2.7
- Therefore, production via *pp* is disfavored between 10–100 TeV

Murase, Ahlers, Lacki, *PRD* 2013



Using high-energy neutrinos as magnetometers

If sources have strong magnetic fields, charged particles cool via synchrotron:

$$p + \gamma_{\text{target}} \longrightarrow \Delta^+ \longrightarrow \begin{cases} p + \pi^0, & \text{Br} = 2/3 \\ n + \pi^+, & \text{Br} = 1/3 \end{cases}$$

$$\pi^0 \longrightarrow \gamma + \gamma$$

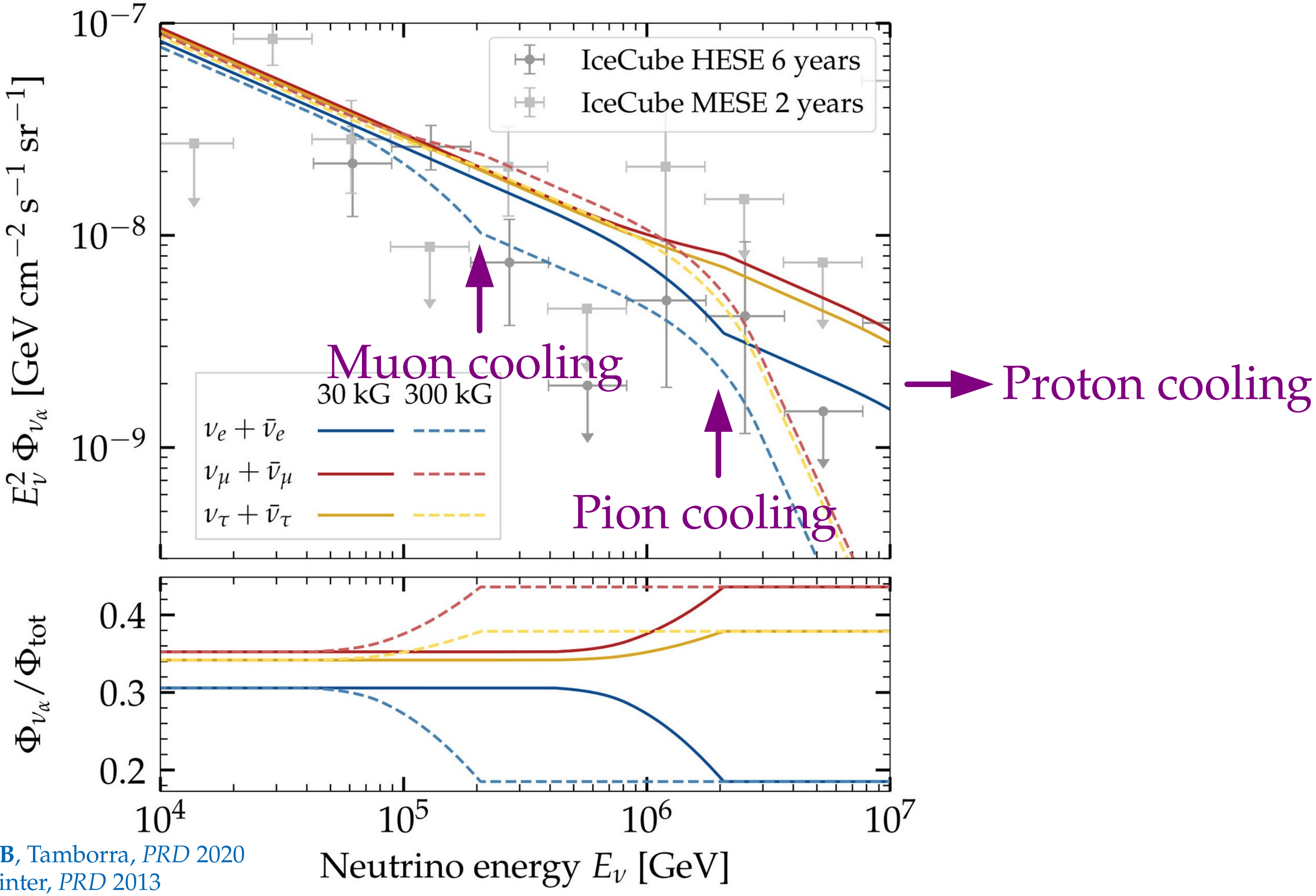
$$\pi^+ \longrightarrow \mu^+ + \nu_\mu \longrightarrow \bar{\nu}_\mu + e^+ + \nu_e + \nu_\mu$$

$$n \text{ (escapes)} \longrightarrow p + e^- + \bar{\nu}_e$$

MB, Tamborra, *PRD* 2020
Winter, *PRD* 2013

Using high-energy neutrinos as magnetometers

If sources have strong magnetic fields, charged particles cool via synchrotron:



MB, Tamborra, *PRD* 2020
Winter, *PRD* 2013

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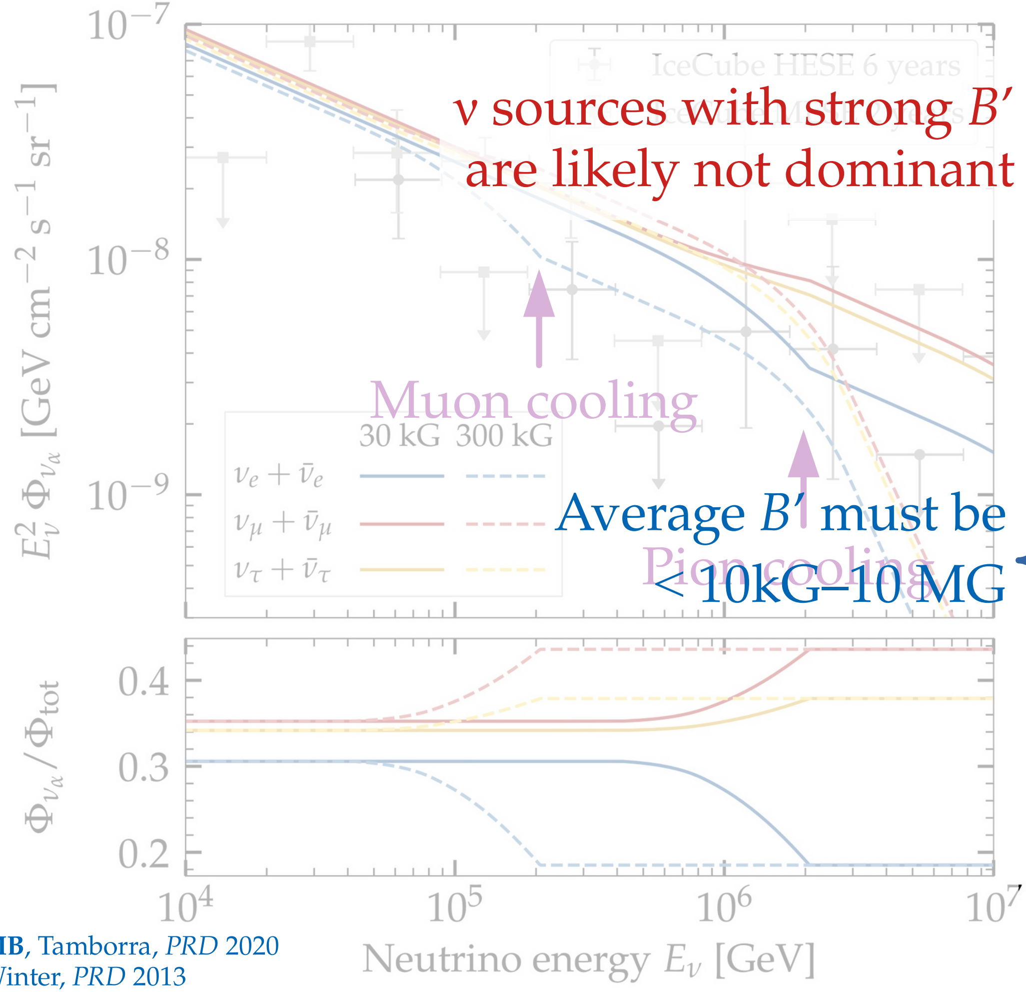
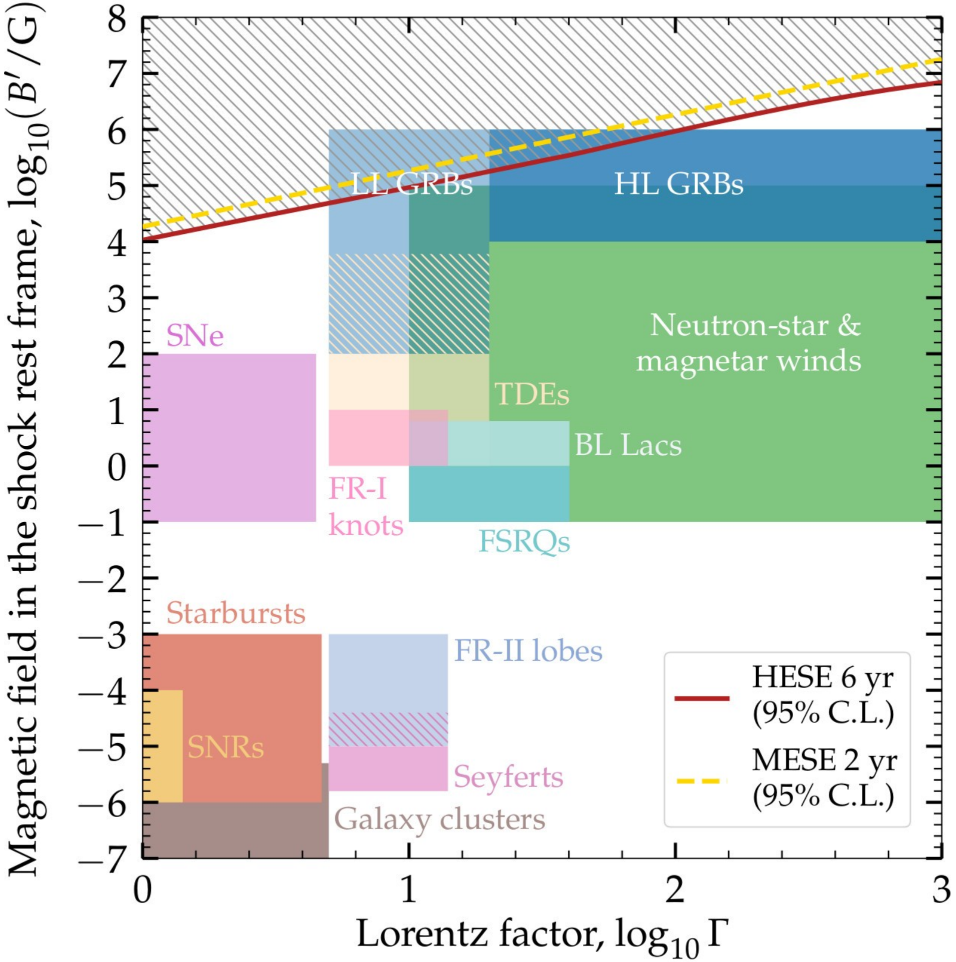


ν sources with strong B' are likely not dominant

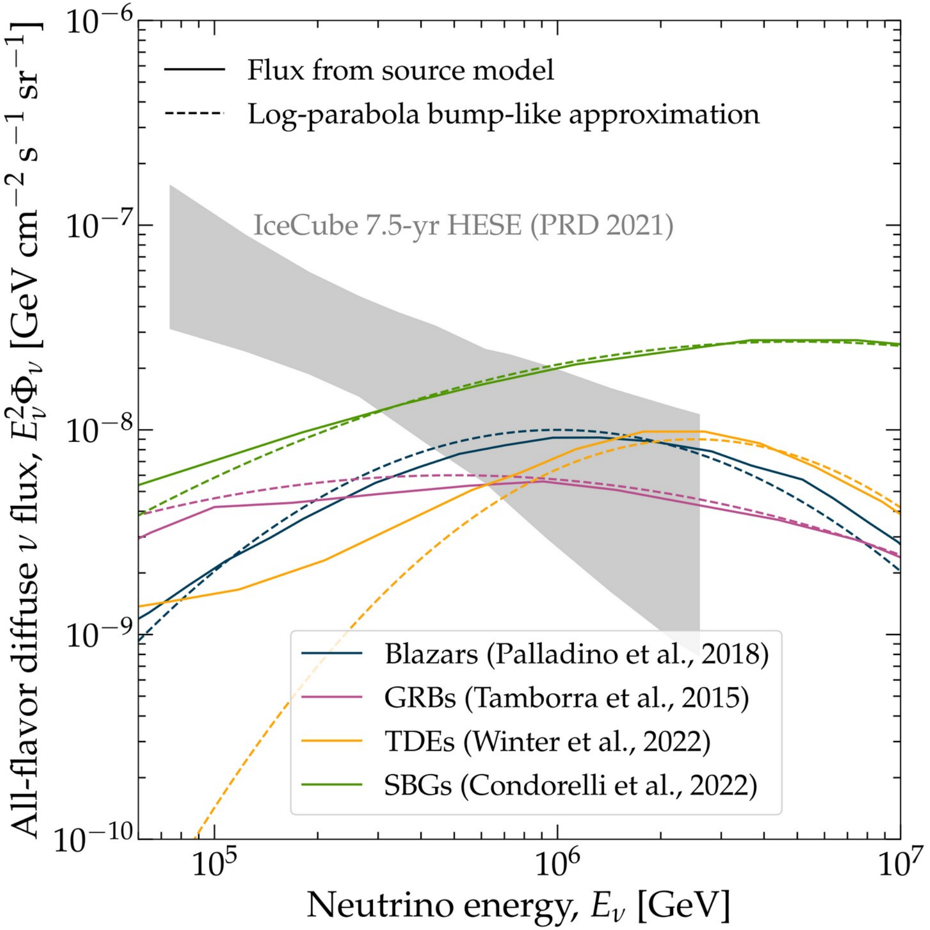
Average B' must be < 10kG–10 MG



MB, Tamborra, *PRD* 2020
Winter, *PRD* 2013

Bump-hunting in the diffuse flux of high-energy neutrinos

Bump-like spectra can reveal the presence of ν production via $p\gamma$:



Fiorillo, **MB**, *PRD* 2023

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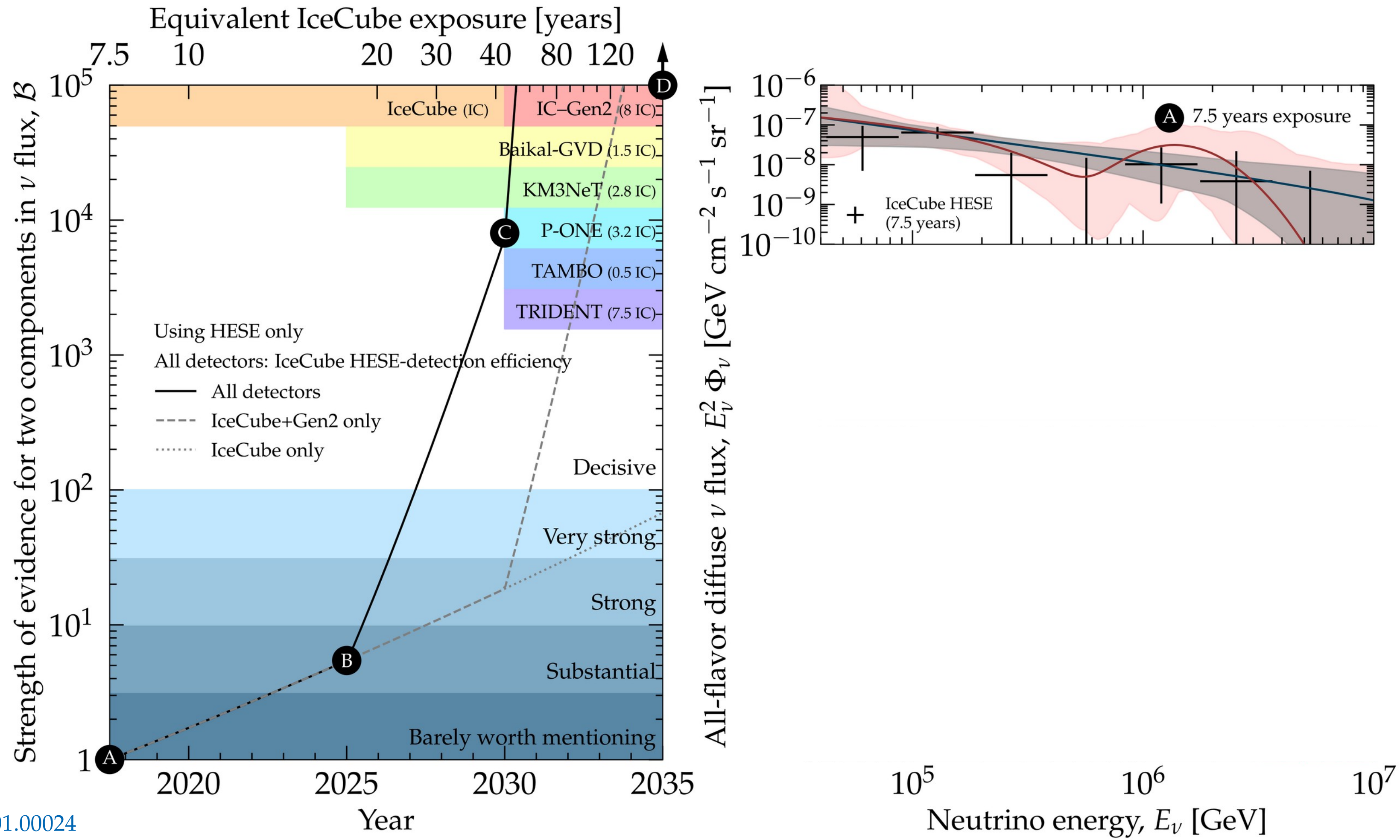

Fiorillo, **MB**, 2301.00024

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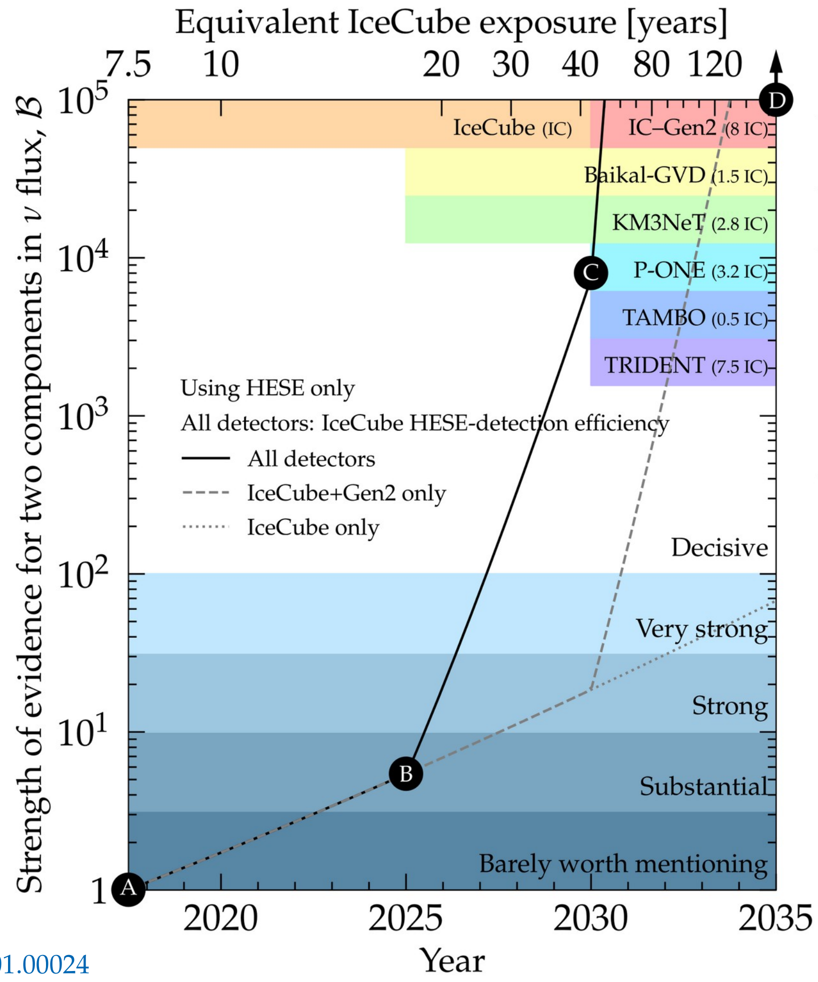

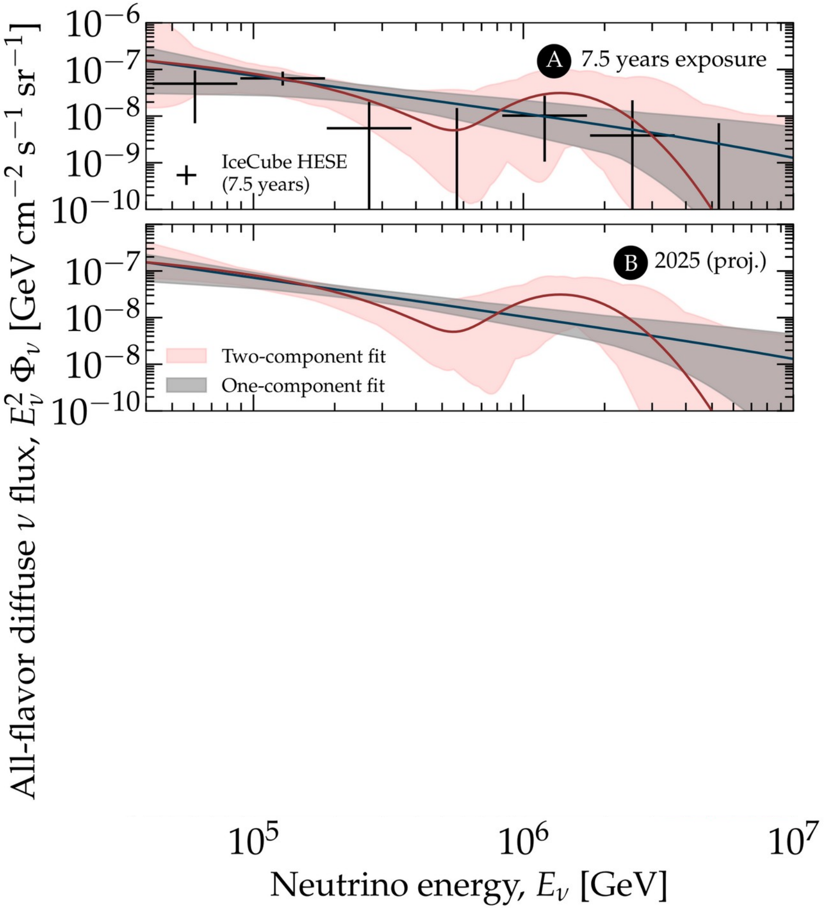

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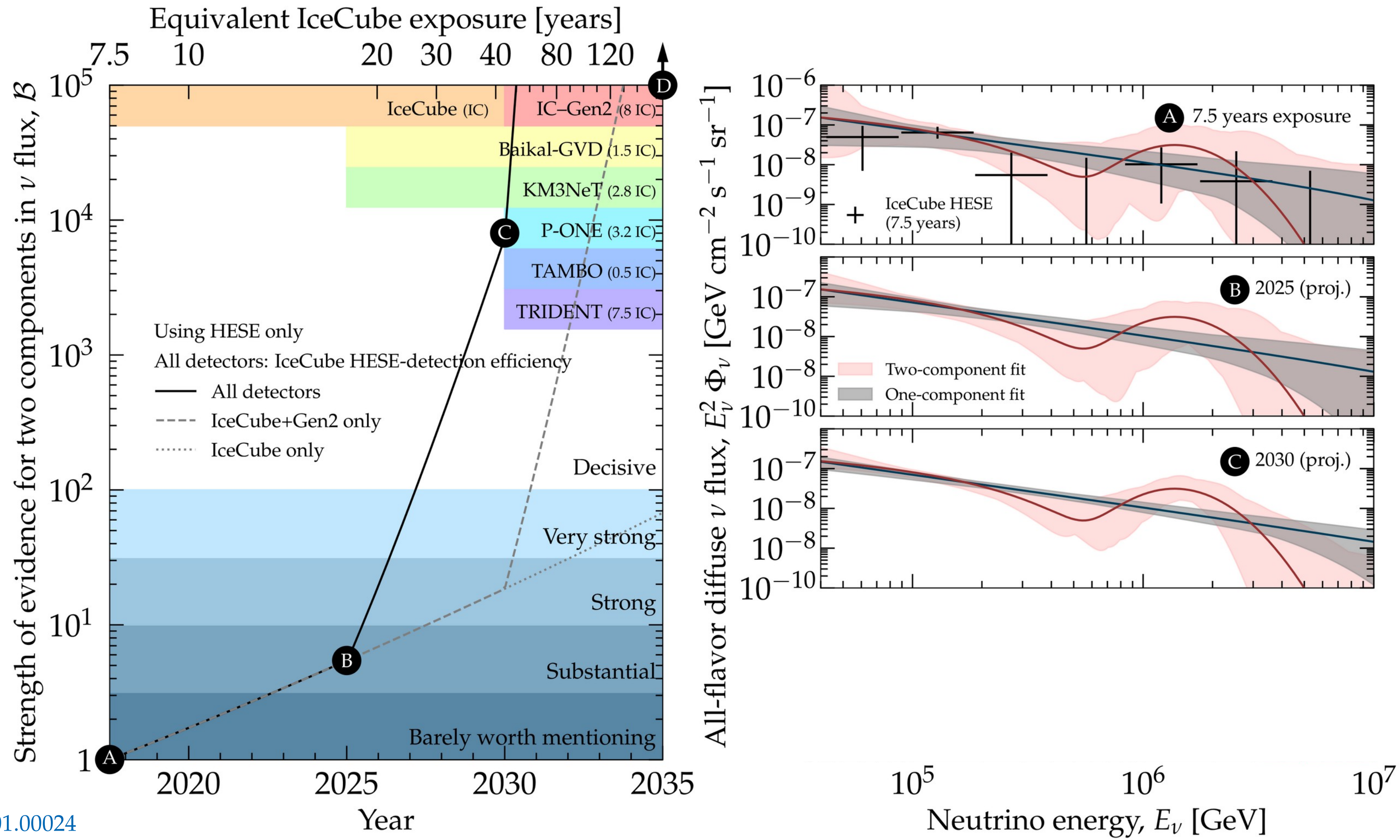

Fiorillo, **MB**, 2301.00024

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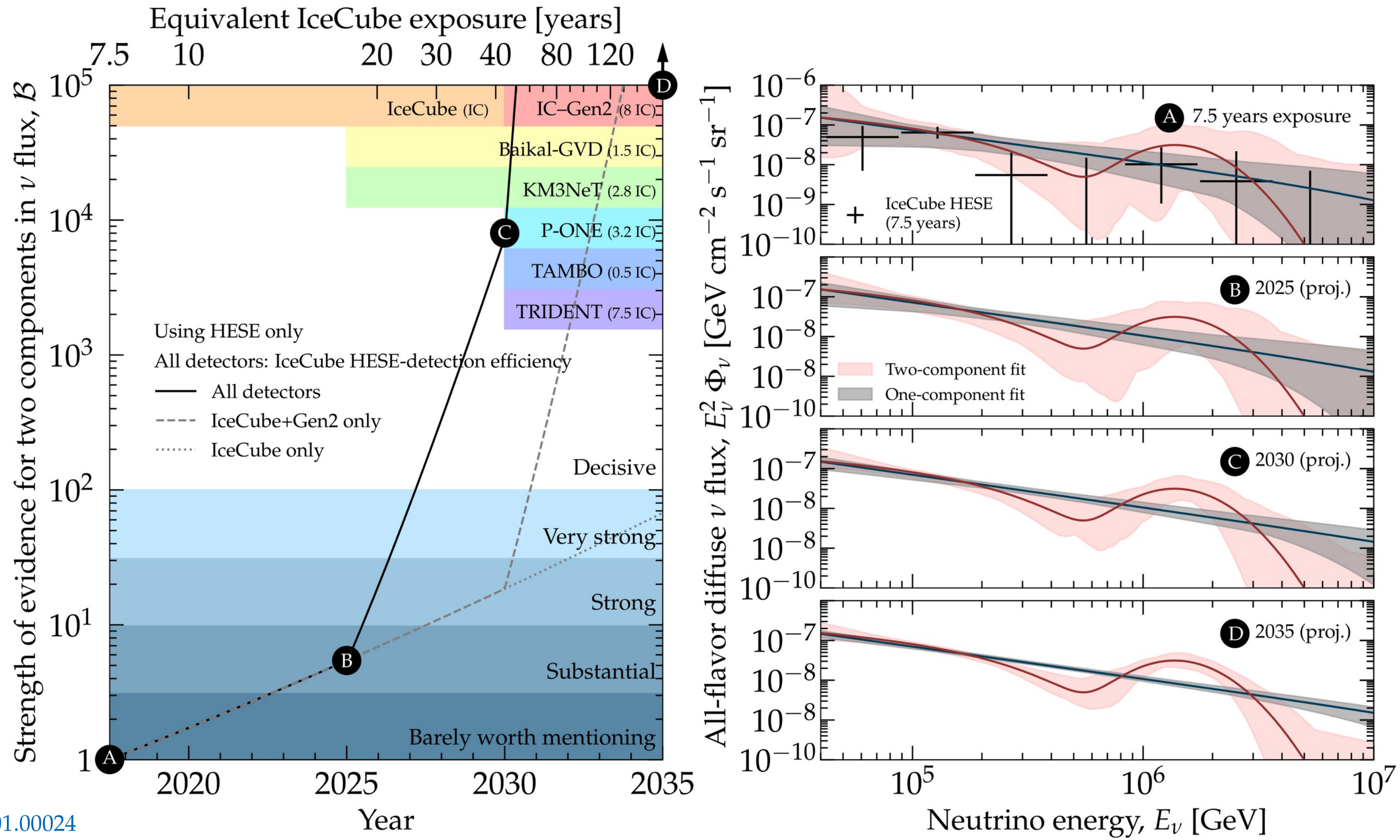


Fiorillo, **MB**, 2301.00024

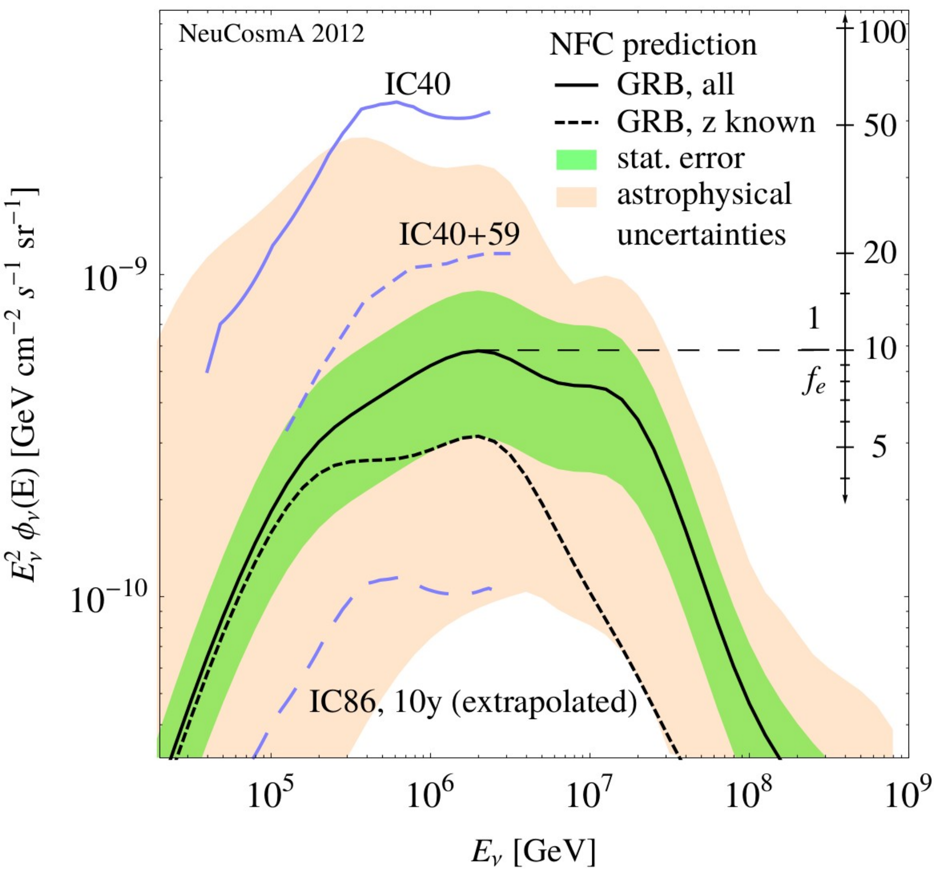
Diffuse flux of neutrinos from GRBs

- How do we estimate it?
- Compute the expected ν fluence from a sample of N_{obs} observed GRBs
- Stack the fluences to obtain the total F_ν
- Quasi diffuse flux:

$$\phi_\nu(E_\nu) = F_\nu(E_\nu)\,\frac{1}{4\pi}\,\frac{1}{N_{\mathrm{obs}}}\,\frac{667\ \mathrm{bursts}}{\mathrm{yr}}$$

($N_{\mathrm{obs}} = 117$ in the plot)

S. Hümmer, P. Baerwald, & W. Winter, *PRL* 2012



Are GRBs still good UHECR source candidates?

- High-luminosity bursts: Not so much
- Low-luminosity bursts: Yes!

	HL GRBs	LL GRBs
Luminosity (erg s^{-1})	$> 10^{49}$	$< 10^{49}$
Rate (Gpc^{-3} yr^{-1})	1	300 (predicted)
Survival of heavy nuclei in jet?	Unlikely	Likely
Can explain IceCube ν?	No	Yes

D. Boncioli, D. Biehl, & W. Winter, *ApJ* 2019; B.T. Zhang *et al.*, *PRD* 2018

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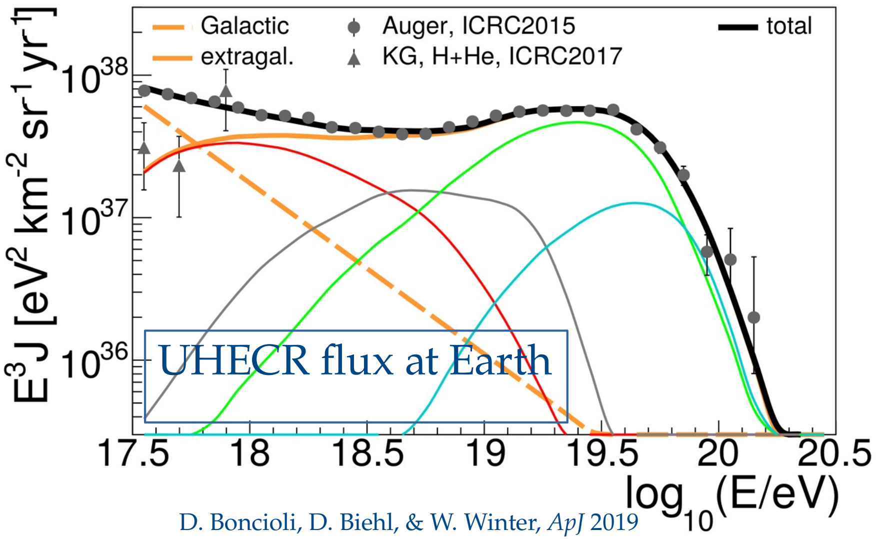

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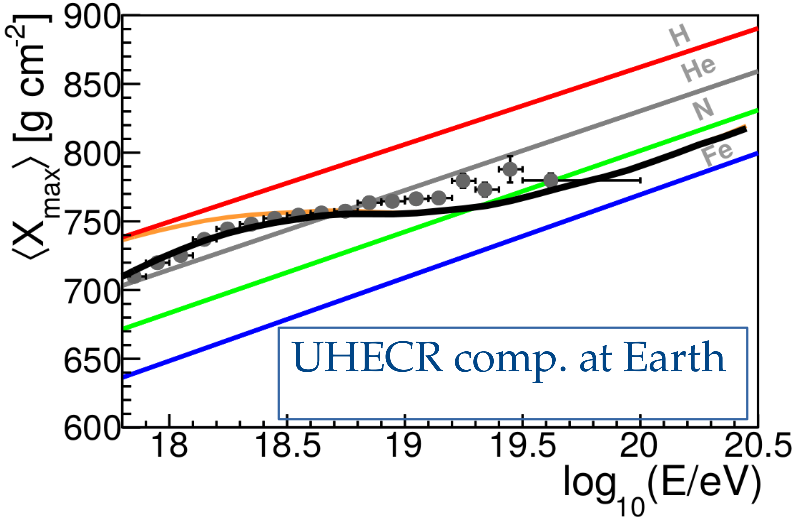

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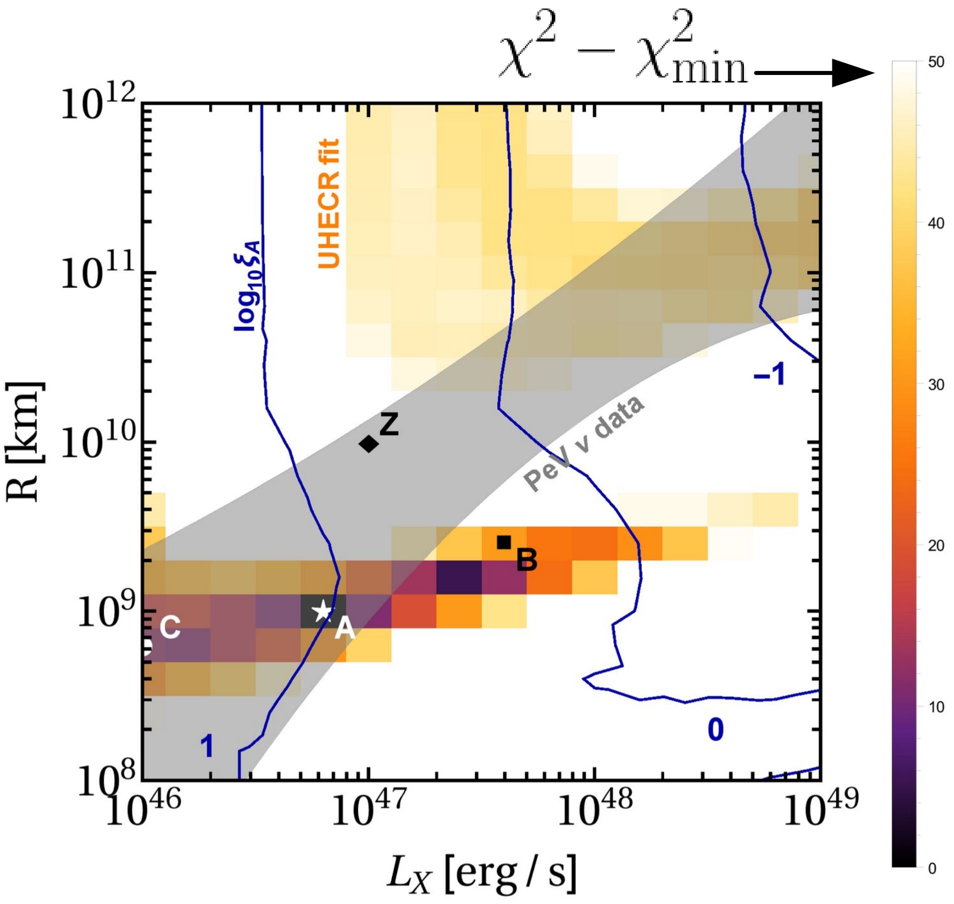


D. Boncioli, D. Biehl, & W. Winter, *ApJ* 2019; B.T. Zhang *et al.*, *PRD* 2018

Neutrino zenith angle distribution

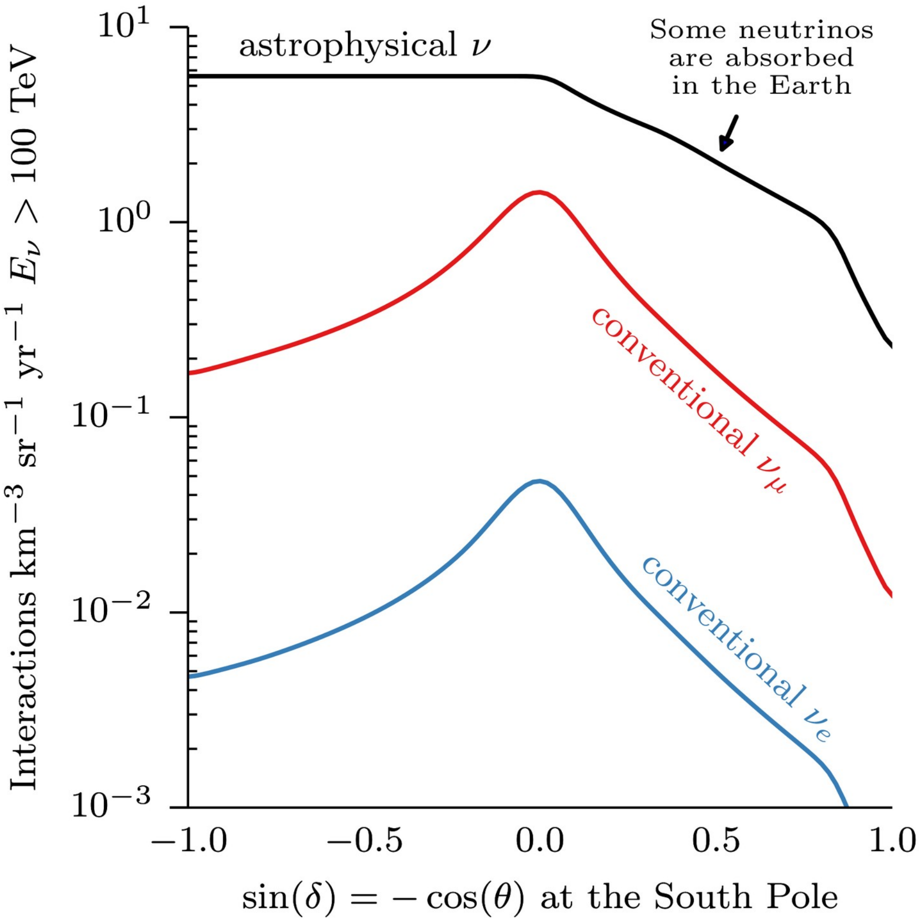
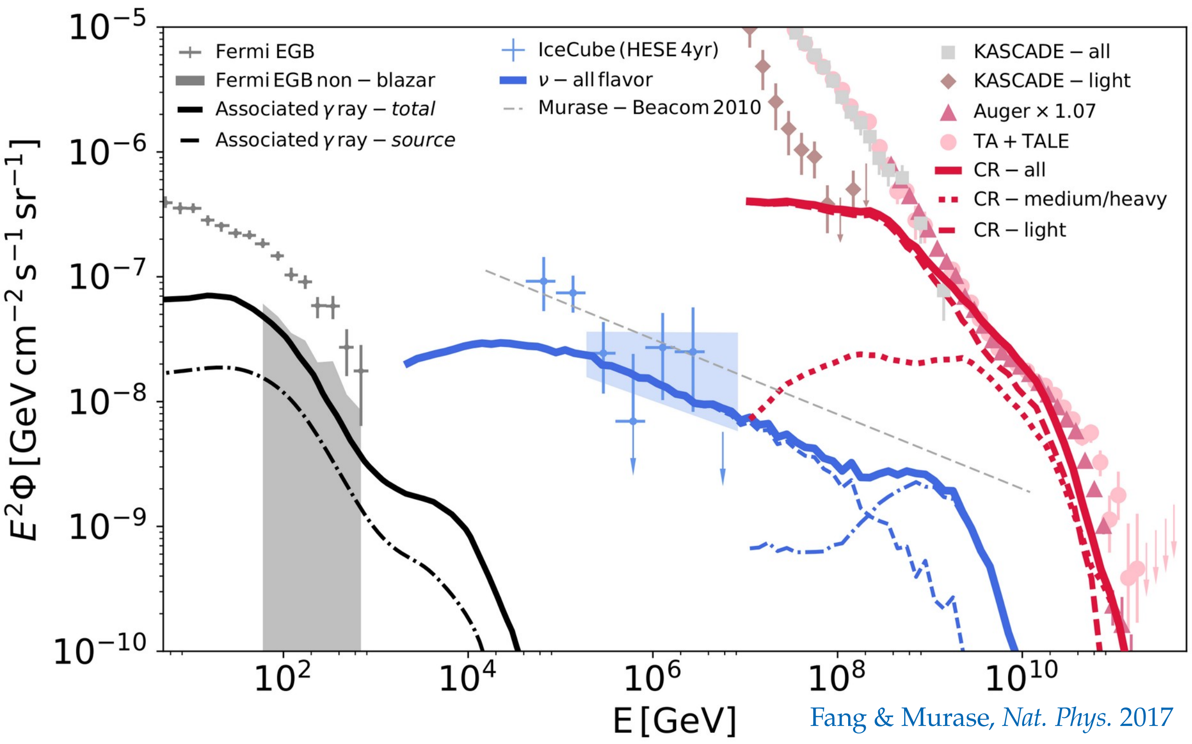


Figure by
Jakob Van Santen
ICRC 2017

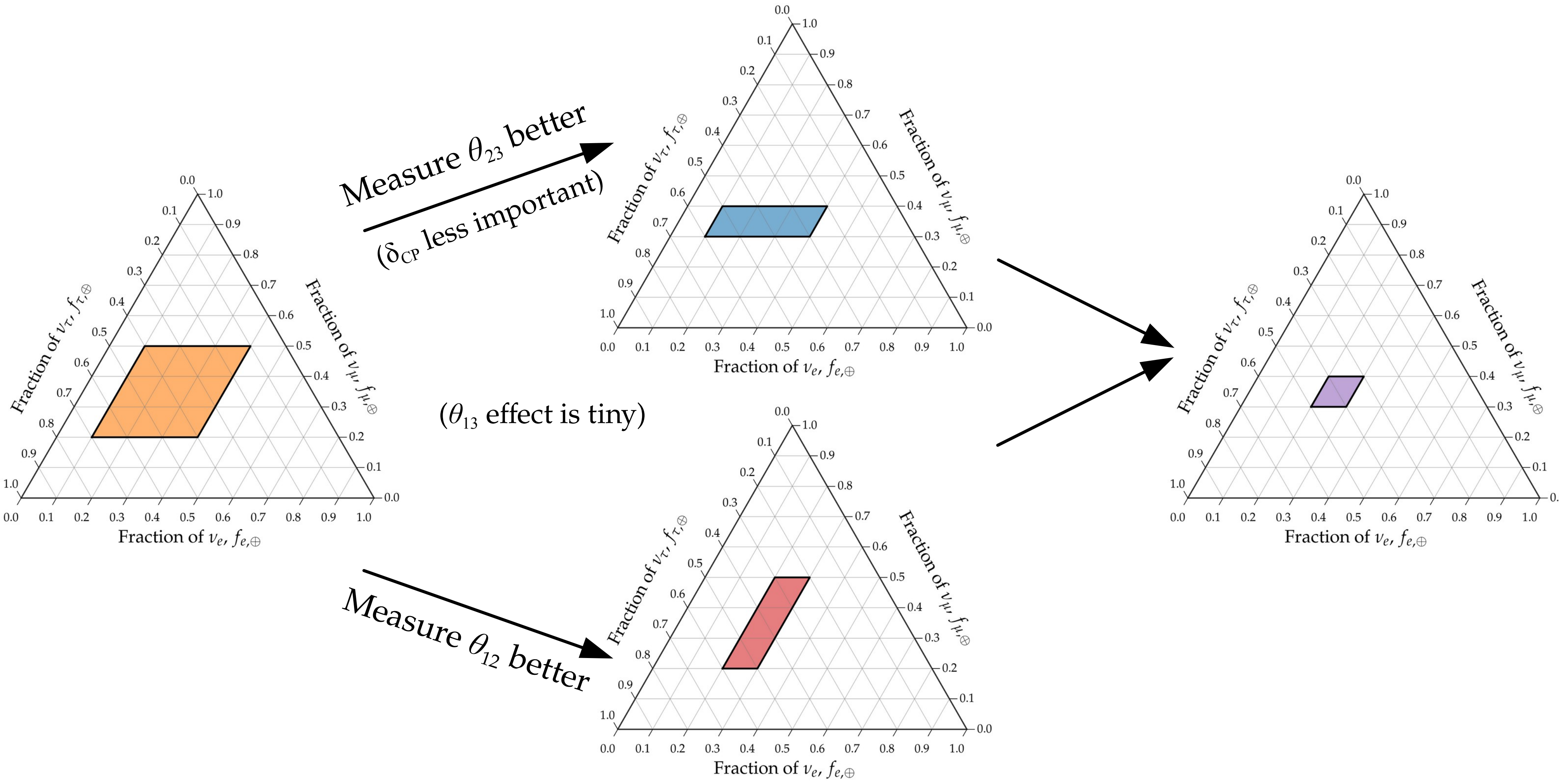
Grand-unified ν–UHECR–gamma-ray model

- Black-hole jets in galaxy clusters accelerate cosmic rays
- UHECRs make ν and γ in the magnetized cluster medium
- UHECRs above 0.1 EeV escape
- Consistent w/ observed UHECR spectrum, composition, isotropy
- Explains IceCube neutrinos
- Explains non-blazar *Fermi* EGB



Fang & Murase, *Nat. Phys.* 2017

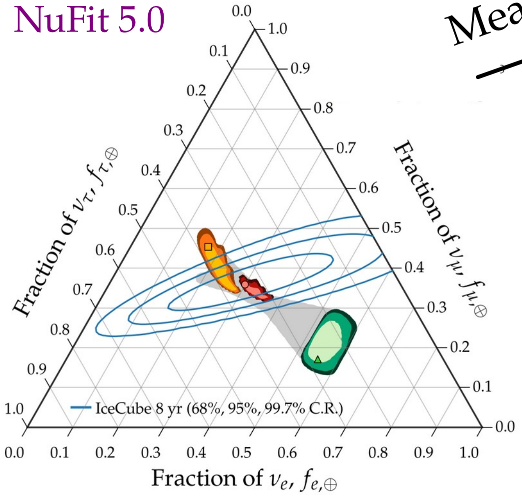
How knowing the mixing parameters better helps



How knowing the mixing parameters better helps

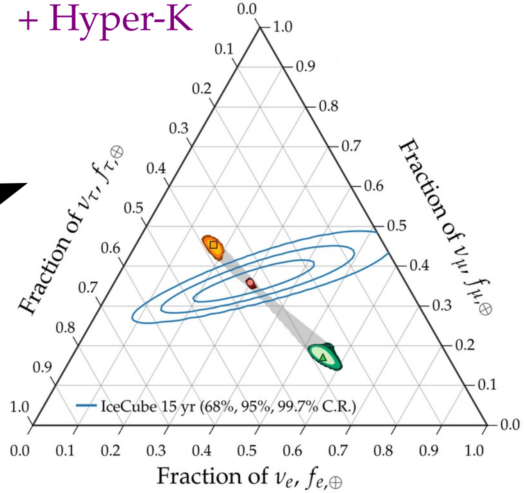
2020

NuFit 5.0



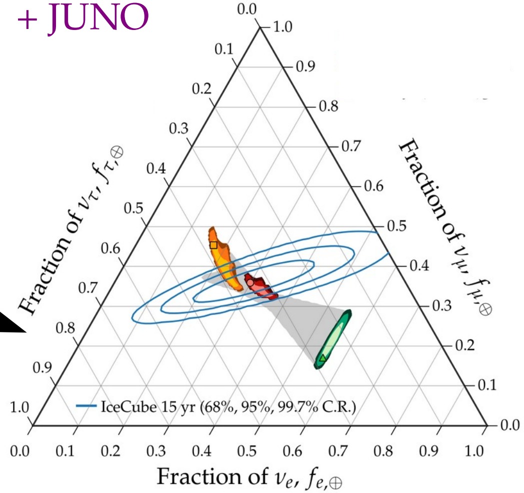
Measure θ_{23} better

+ Hyper-K



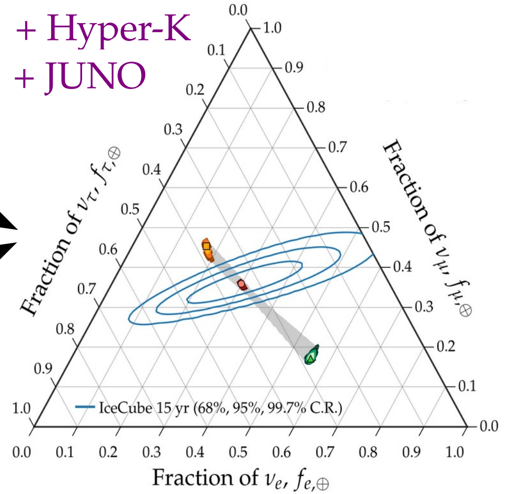
Measure θ_{12} better

+ JUNO



~2030

+ Hyper-K
+ JUNO

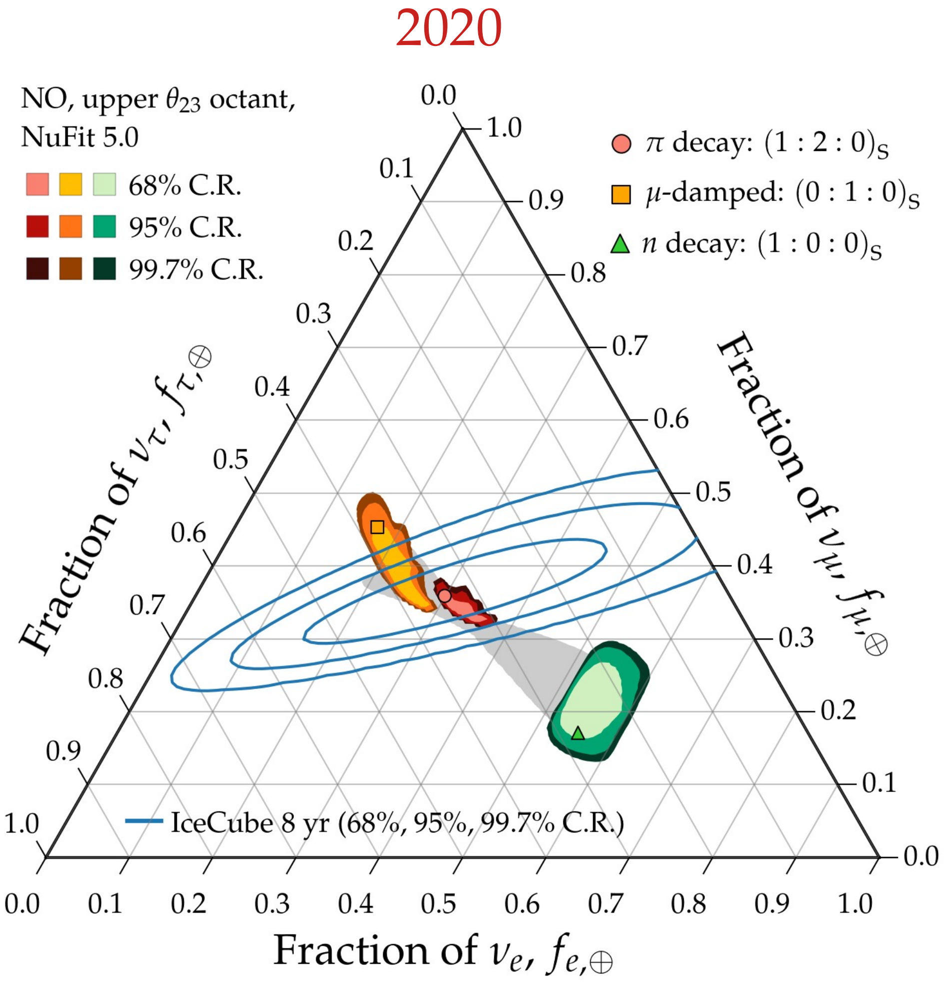


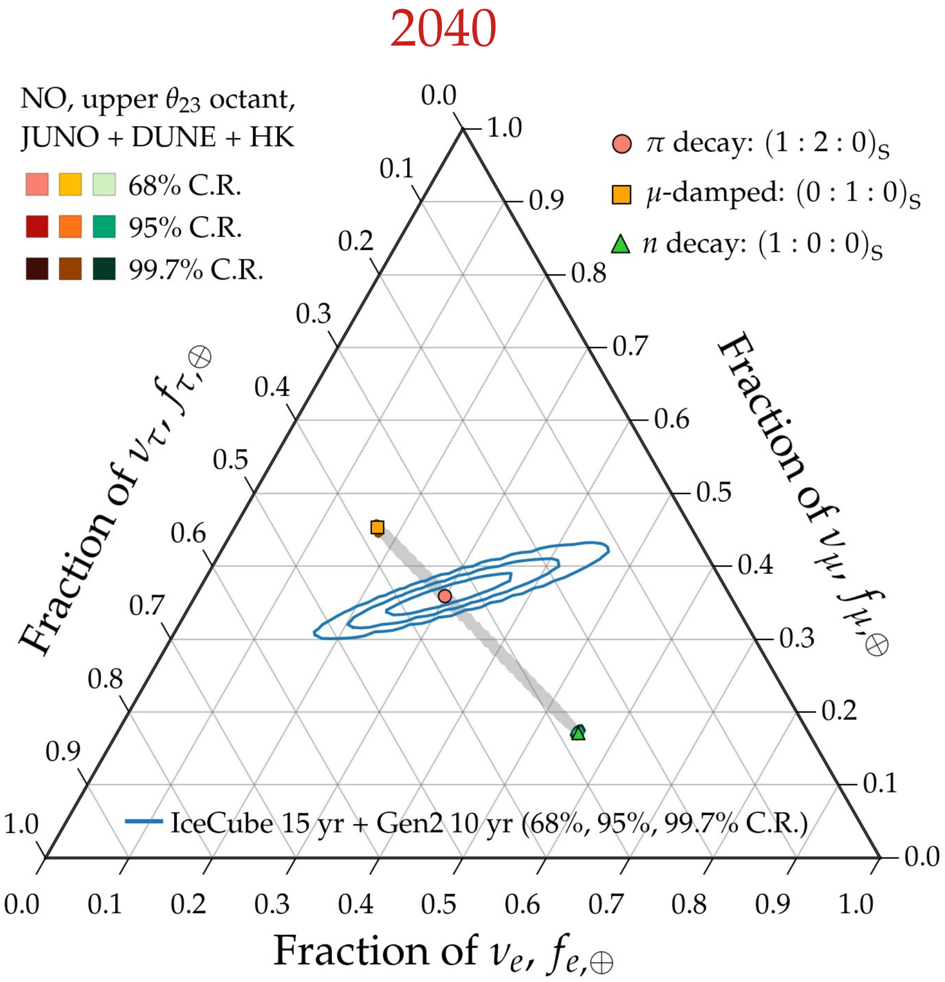
In our results:
JUNO + Hyper-K + DUNE

Marginal improvement til 2040

Song, Li, Argüelles, **MB**, Vincent, *JCAP* 2021

Theoretically palatable regions: 2020 → 2040





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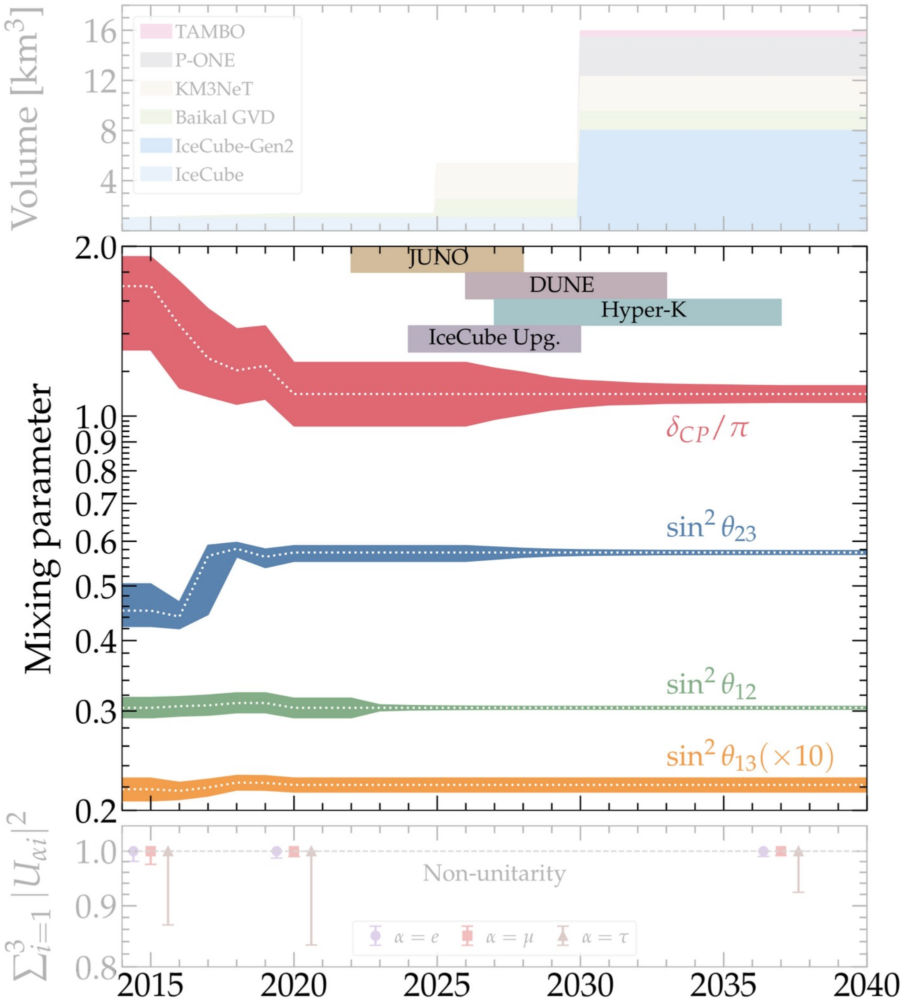

We can compute the oscillation probability more precisely:

$$f_{\alpha,\oplus} = \sum_{\beta=e,\mu,\tau} P_{\beta\alpha} f_{\beta,\mathrm{S}}$$

So we can convert back and forth between source and Earth more precisely

Flavor at the Earth: *theoretically palatable regions*

Theoretically palatable flavor regions

$\equiv$

MB, Beacom, Winter, *PRL* 2015

Allowed regions of flavor ratios at Earth derived from oscillations

Note:
The original palatable regions were frequentist [**MB**, Beacom, Winter, *PRL* 2015]; the new ones are Bayesian

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Ingredient #1:

Flavor ratios at the source,

$(f_{e,S}, f_{\mu,S}, f_{\tau,S})$

Fix at one of the benchmarks
(pion decay, muon-damped, neutron decay)

or

Explore all possible combinations

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Probability density of mixing parameters $(\theta_{12}, \theta_{23}, \theta_{13}, \delta_{\mathrm{CP}})$

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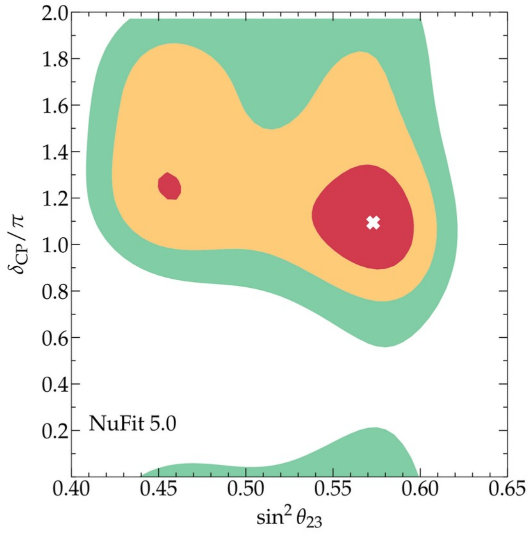
Esteban *et al.*, *JHEP* 2020
www.nu-fit.org



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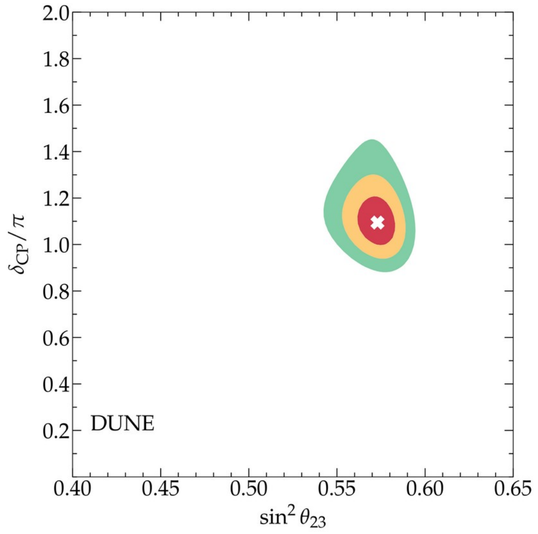
Esteban *et al.*, *JHEP* 2020
www.nu-fit.org

Post-2020: Build our own profiles using simulations of JUNO, DUNE, Hyper-K

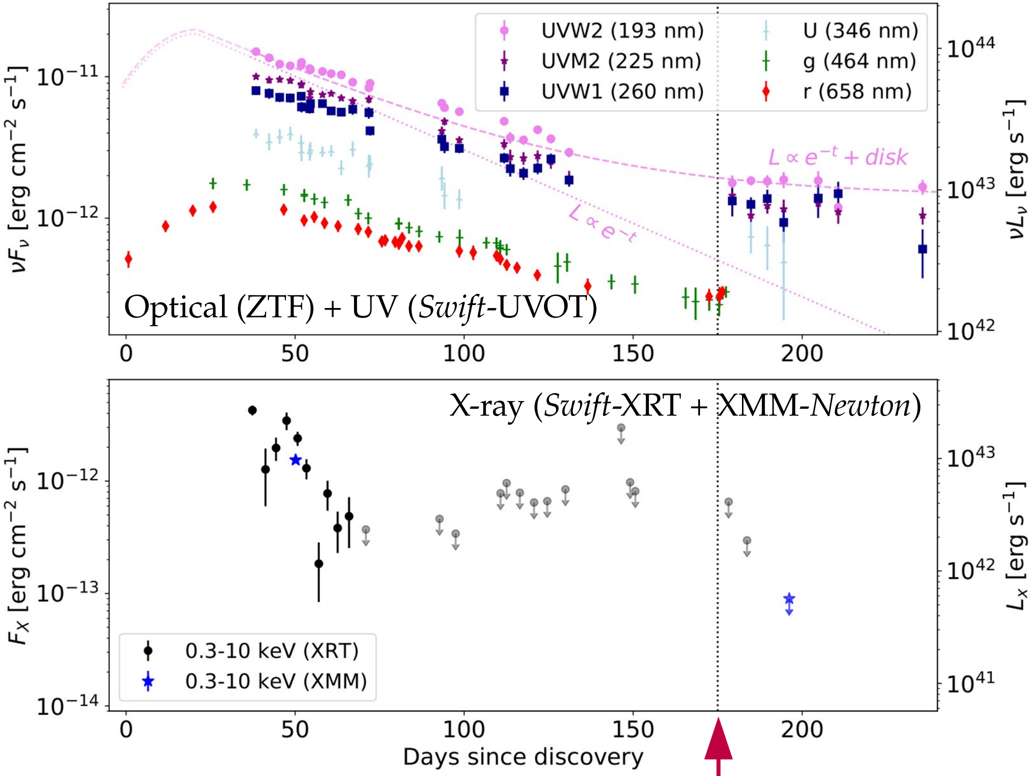
An *et al.*, *J. Phys. G* 2016
DUNE, 2002.03005
Huber, Lindner, Winter, *Nucl. Phys. B* 2002



An apparent TDE neutrino source

Radio-emitting TDE AT2019dsg coincident with neutrino event IC191001A:

AT2019dsg: Apr 9, 2019 / $z = 0.051$ (230 Mpc) / $M_{BH} = 3 \times 10^7\, M_\odot$

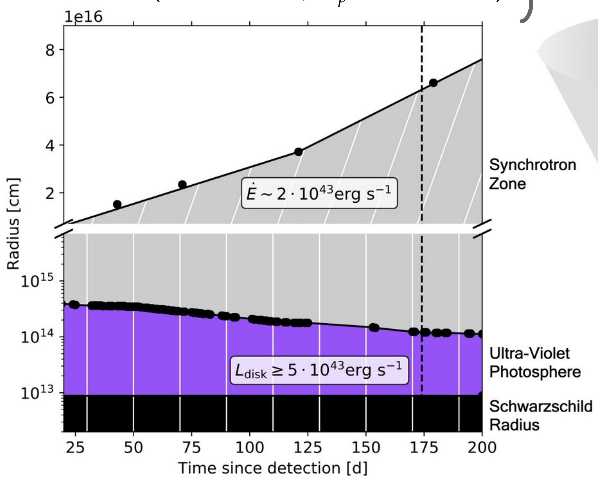


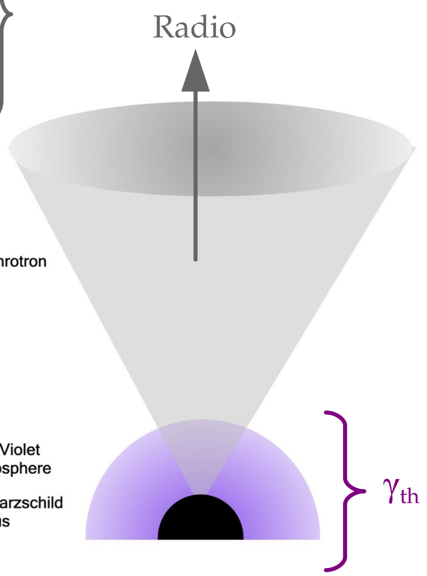
IC191001A, ~200 TeV

Multi-zone model:

From radio:
mildly relativistic expansion
($v/c \sim 0.2$) + acceleration

p and e accelerated here
($B = 0.07$ G, $E_p < 160$ PeV)





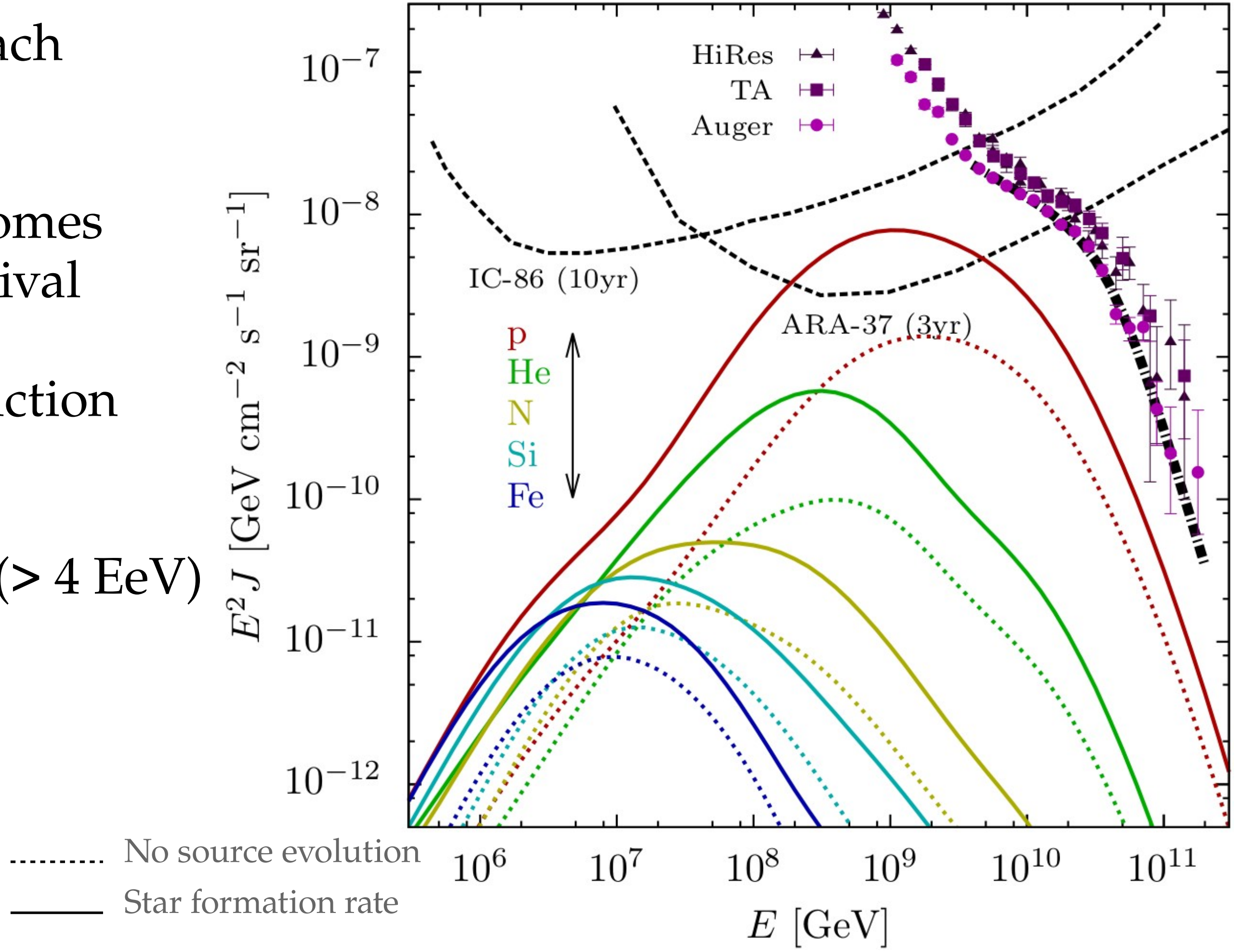
$p + \gamma_{th}$ (or p) $\rightarrow \nu$

Stein *et al.*, *Nature Astron.* 2021 – See also: Winter & Lunardini, *Nature Astron.* 2021; Murase, Kimura, Zhang, Oikonomou, Petropoulou, *ApJ* 2020

The Cosmogenic Neutrino Floor

Ahlers & Halzen, *PRD* 2012

- In a nucleus A of energy E, each nucleon has energy E/A
- Minimal cosmogenic ν flux comes from maximizing nuclei survival
- *I.e.*, from minimizing p production from photo-disintegration
- ν fluxes from UHECR nuclei (> 4 EeV) are presently beyond reach

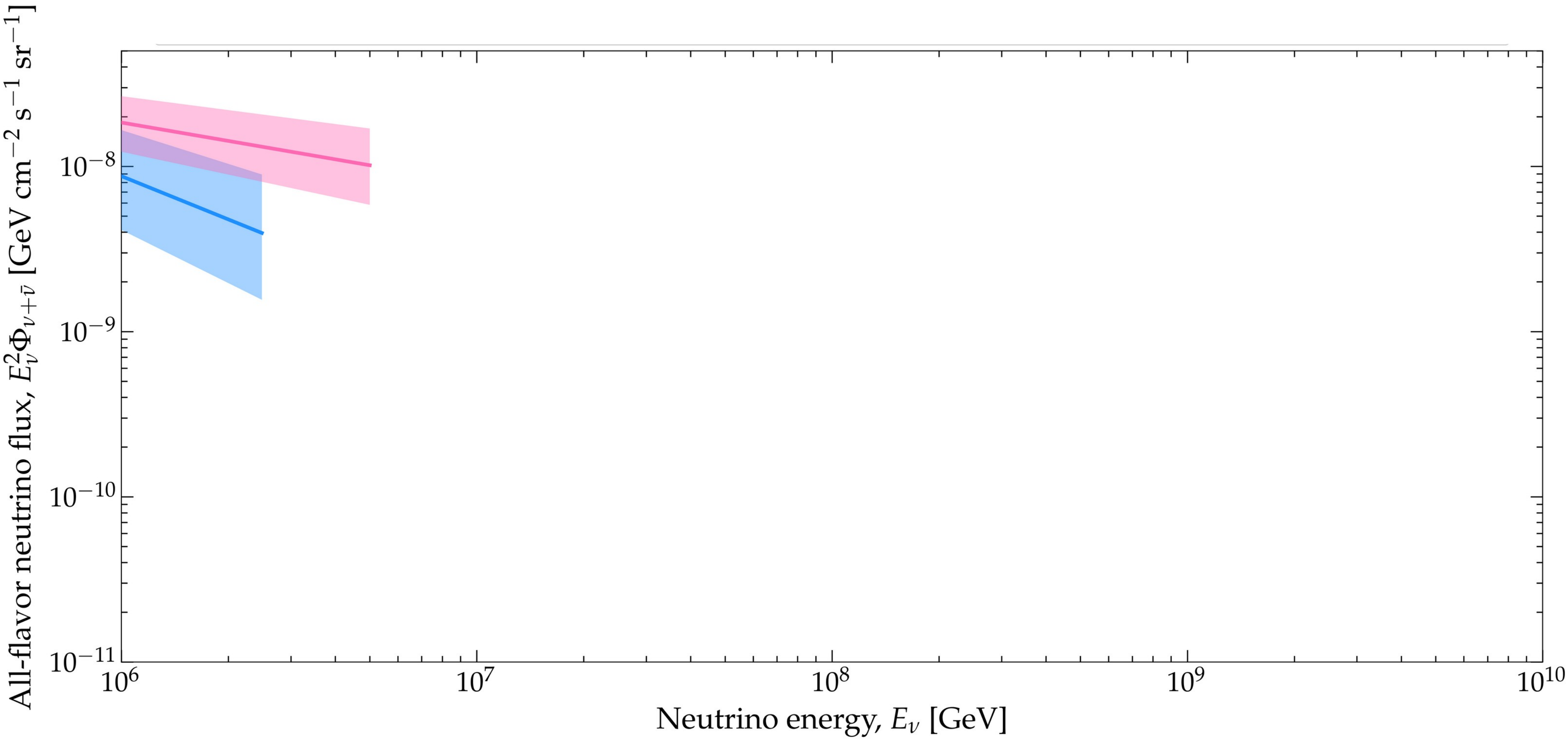


Many TeV–EeV ν telescopes in planning for 2020–2040

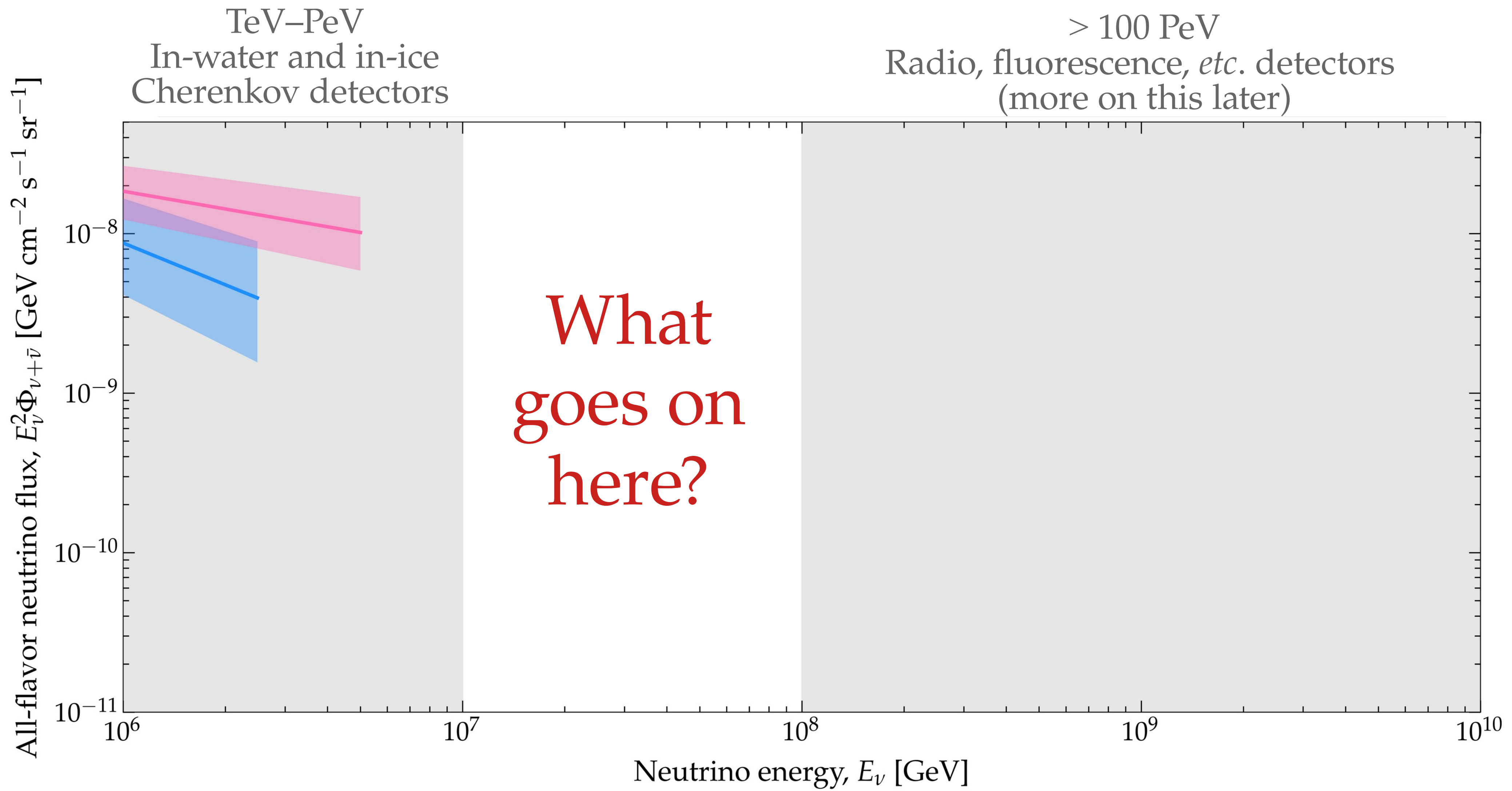
				Flavor		Technique			Neutrino Target					Geometry					
Experiments	Phase & Online Date	Energy Range	Site	Tau	All Flavor	Optical / UV	Radio	Showers	H_2O	Atmosphere	Earth's limb	Topography	Lunar Regolith	Embedded	Planar Arrays	Valley	Mountains	Balloon	Satellite
IceCube	2010	TeV-EeV	South Pole		✓	✓			✓					✓					
KM3NeT	2021	TeV-PeV	Mediteranean		✓	✓			✓					✓					
Baikal-GVD	2021	TeV-PeV	Lake Baikal		✓	✓			✓					✓					
P-ONE	2020	TeV-PeV	Pacific Ocean		✓	✓			✓					✓					
IceCube-Gen2	2030+	TeV-EeV	South Pole		✓	✓	✓		✓					✓					
ARIANNA	2014	>30 PeV	Moore's Bay		✓		✓		✓					✓					
ARA	2011	>30 PeV	South Pole		✓		✓		✓					✓					
RNO-G	2021	>30 PeV	Greenland		✓		✓		✓					✓					
RET-N	2024	PeV-EeV	Antarctica		✓		✓		✓					✓					
ANITA	2008,2014,2016	EeV	Antarctica	✓	✓		✓		✓		✓							✓	
PUEO	2024	EeV	Antarctica	✓	✓		✓		✓		✓							✓	
GRAND	2020	EeV	China / Worldwide	✓			✓			✓	✓	✓			✓		✓		
BEACON	2018	EeV	CA, USA/ Worldwide	✓			✓				✓	✓					✓		
TAROGE-M	2018	EeV	Antarctica	✓			✓				✓	✓					✓		
SKA	2029	>100 EeV	Australia		✓		✓						✓		✓				
Trinity	2022	PeV-EeV	Utah, USA	✓		✓					✓						✓		
POEMMA		>20 PeV	Satellite	✓	✓	✓				✓	✓								✓
EUSO-SPB	2022	EeV	New Zealand	✓		✓					✓							✓	
Pierre Auger	2008	EeV	Argentina	✓	✓			✓		✓	✓	✓			✓				
AugerPrime	2022	EeV	Argentina	✓	✓		✓	✓		✓	✓	✓			✓				
Telescope Array	2008	EeV	Utah, USA	✓	✓			✓		✓					✓				
TAx4		EeV	Utah, USA	✓	✓			✓											
TAMBO	2025-2026	PeV-EeV	Peru	✓				✓				✓				✓			

Operational		Date full operations began
Prototype		Date protoype operations began or begin
Planning		Projected full operations

Abraham *et al.* (inc. **MB**), *J. Phys. G: Nucl. Part. Phys.* 59, 11 (2022) [2203.05591]



All-flavor neutrino flux, $E_\nu^2\Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^{6}
10^{7}
10^{8}
10^{9}
10^{10}
Neutrino energy, E_ν [GeV]


TeV–PeV
In-water and in-ice
Cherenkov detectors
> 100 PeV
Radio, fluorescence, etc. detectors
(more on this later)
What goes on here?
All-flavor neutrino flux, $E_\nu^2 \Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
10^{-8}
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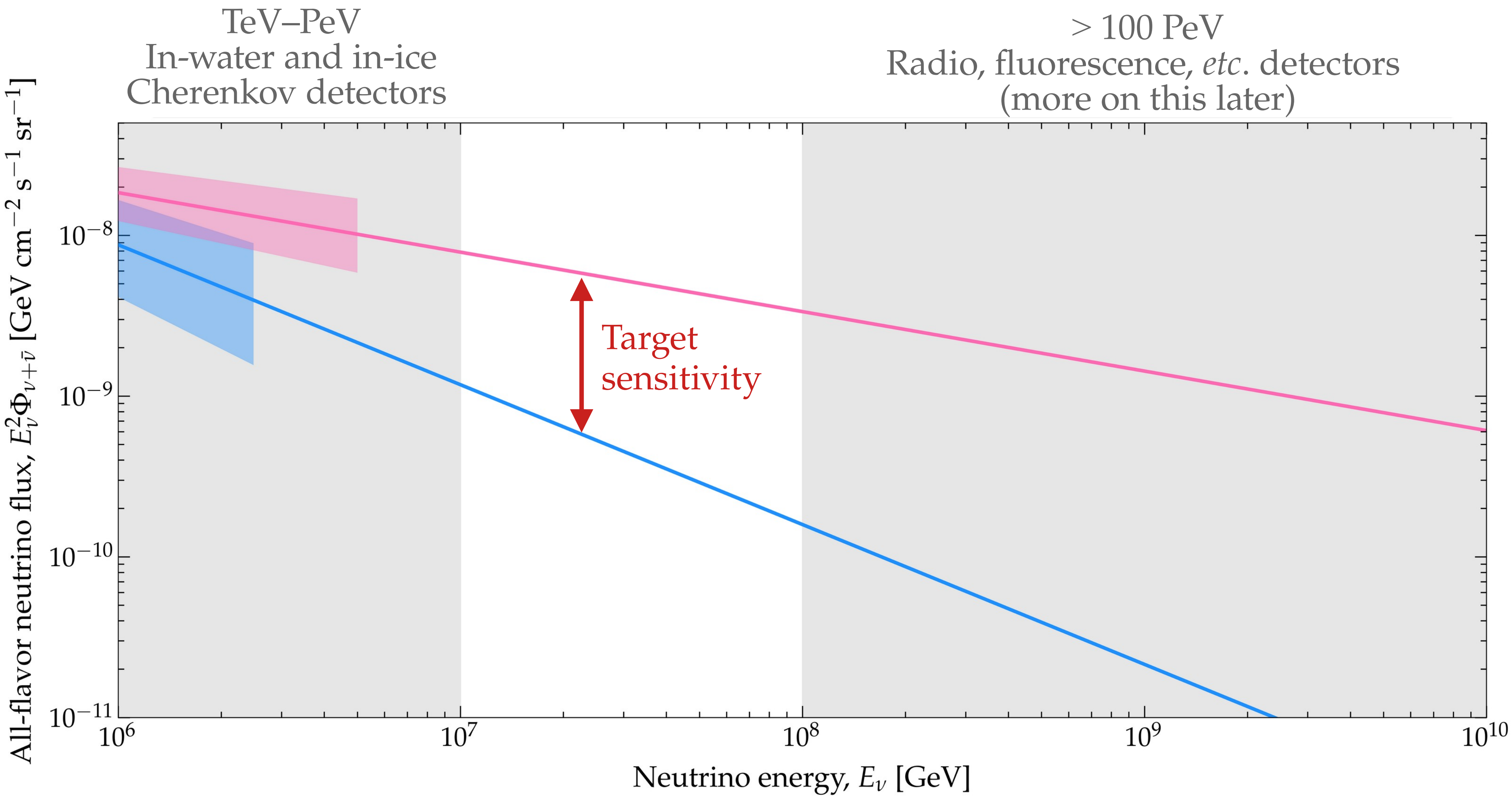

TeV–PeV
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Target
sensitivity
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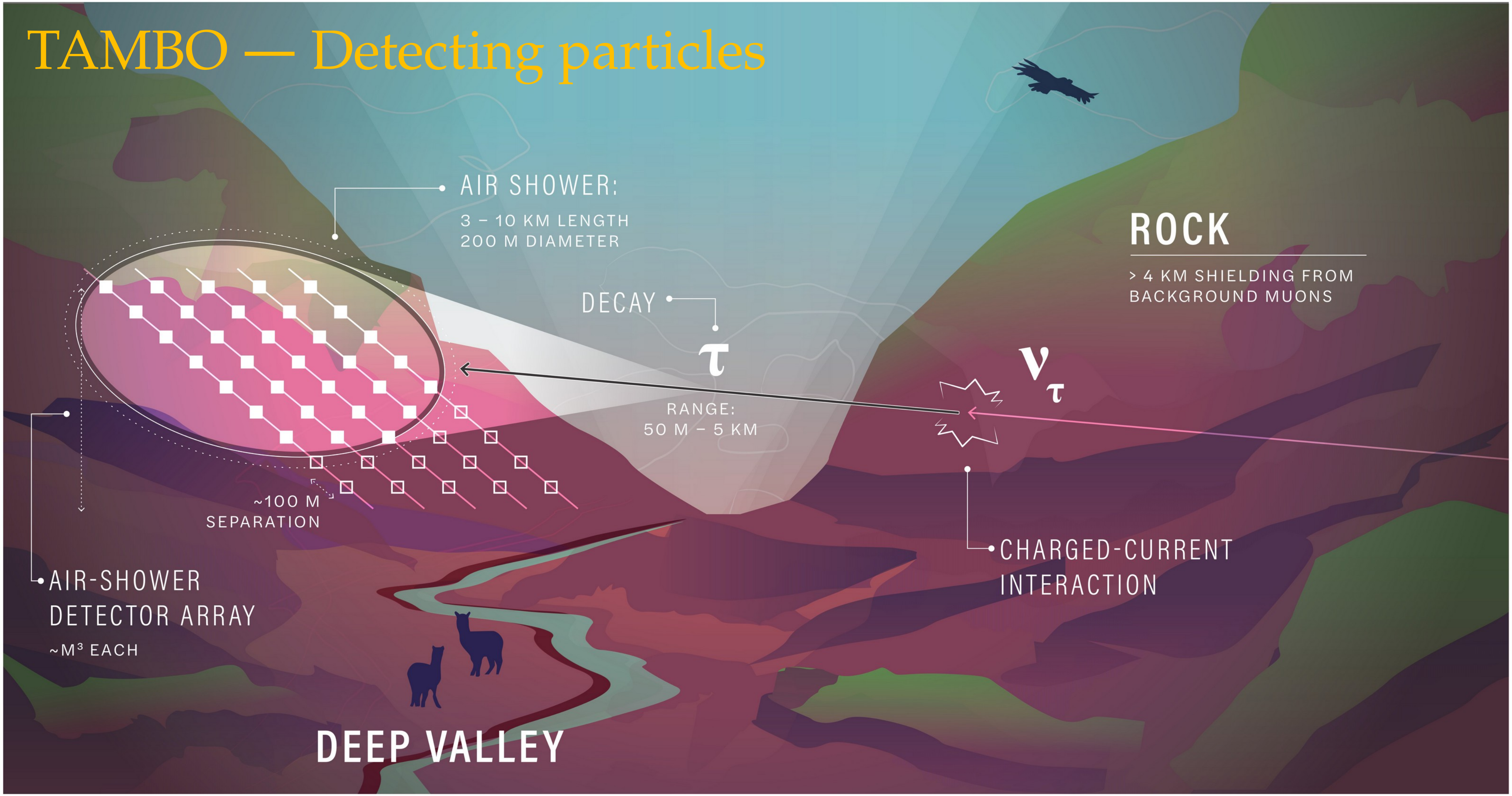
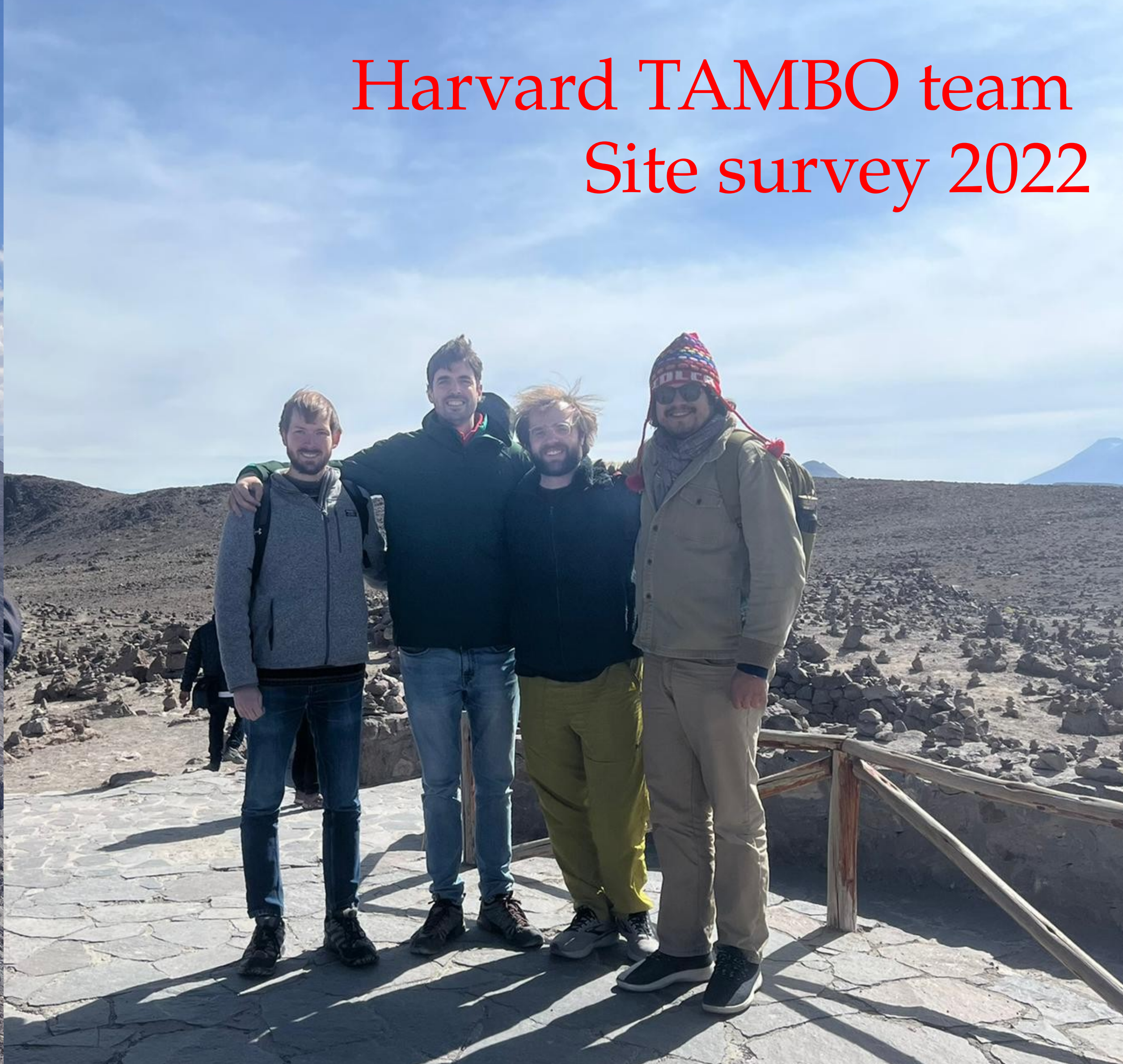
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^{6}
10^{7}
10^{8}
10^{9}
10^{10}
Neutrino energy, E_ν [GeV]

TAMBO — Detecting particles



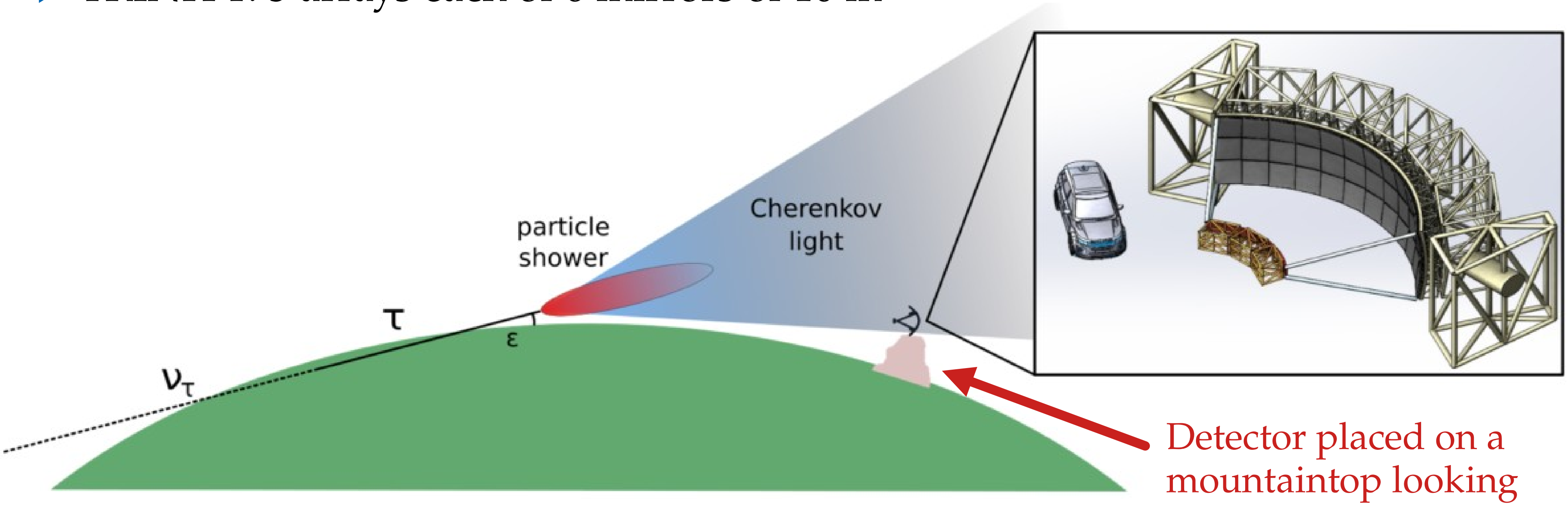
TAU AIR-SHOWER MOUNTAIN-BASED OBSERVATORY (TAMBO) · COLCA VALLEY, PERU


Colca Valley, Peru
2000 (bottom) to 4000 (top) m.a.s.l.

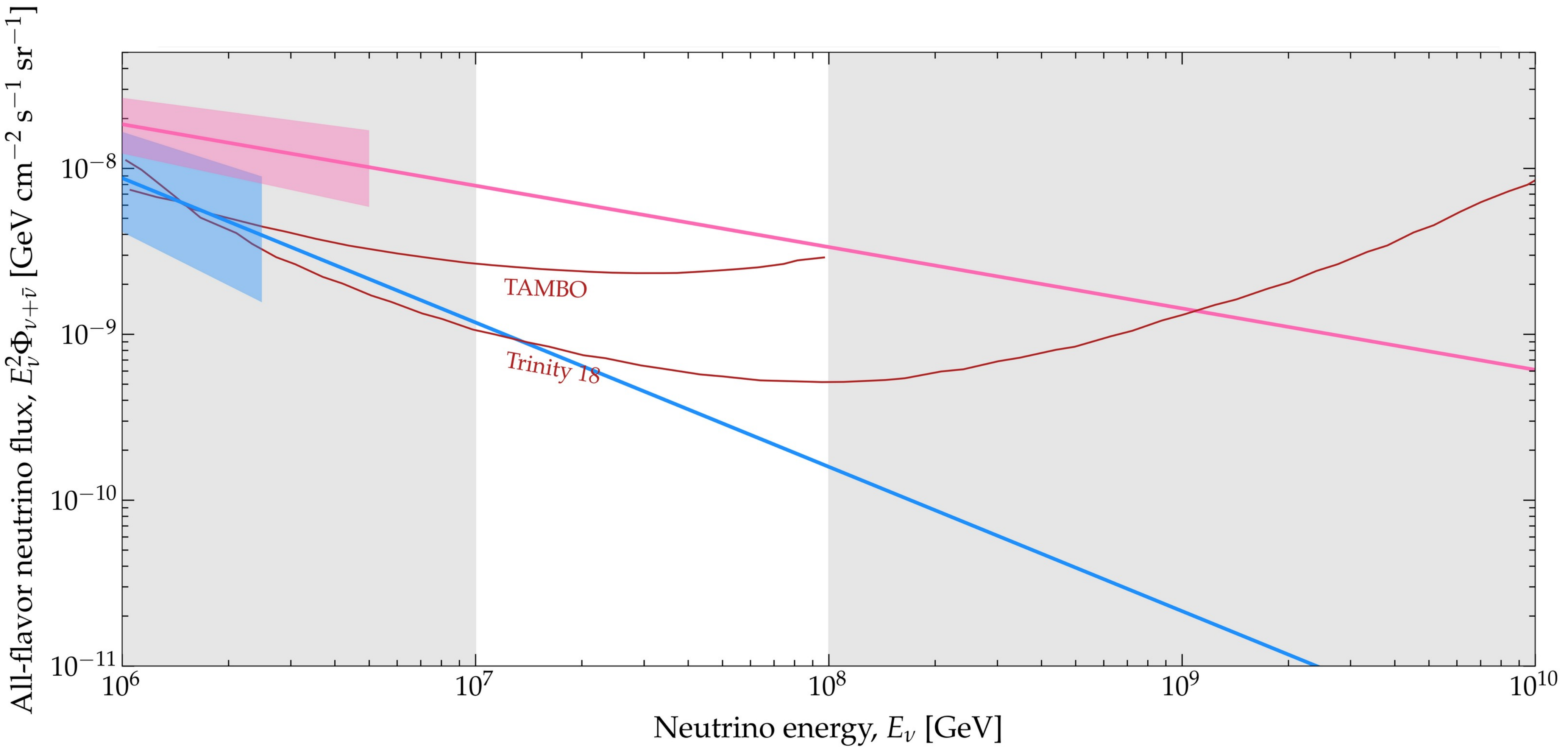


Harvard TAMBO team
Site survey 2022

TRINITY — Detecting Cherenkov light

- Atmospheric Cherenkov imaging applied to PeV neutrinos
- Pioneered by MAGIC (pointing at Atlantic), ASHRA, and NTA (Mauna Kea)
- TRINITY: 3 arrays each of 6 mirrors of 10 m^2



TRINITY Demonstrator
Frisco Peak, UT



All-flavor neutrino flux, $E_\nu^2\Phi_{\nu+\bar{\nu}}$ [GeV cm^{-2} s^{-1} sr^{-1}]
10^{-8}
10^{-9}
10^{-10}
10^{-11}
10^{6}
10^{7}
10^{8}
10^{9}
10^{10}
Neutrino energy, E_ν [GeV]
TAMBO
Trinity 18